Introduction to Radio Interferometry

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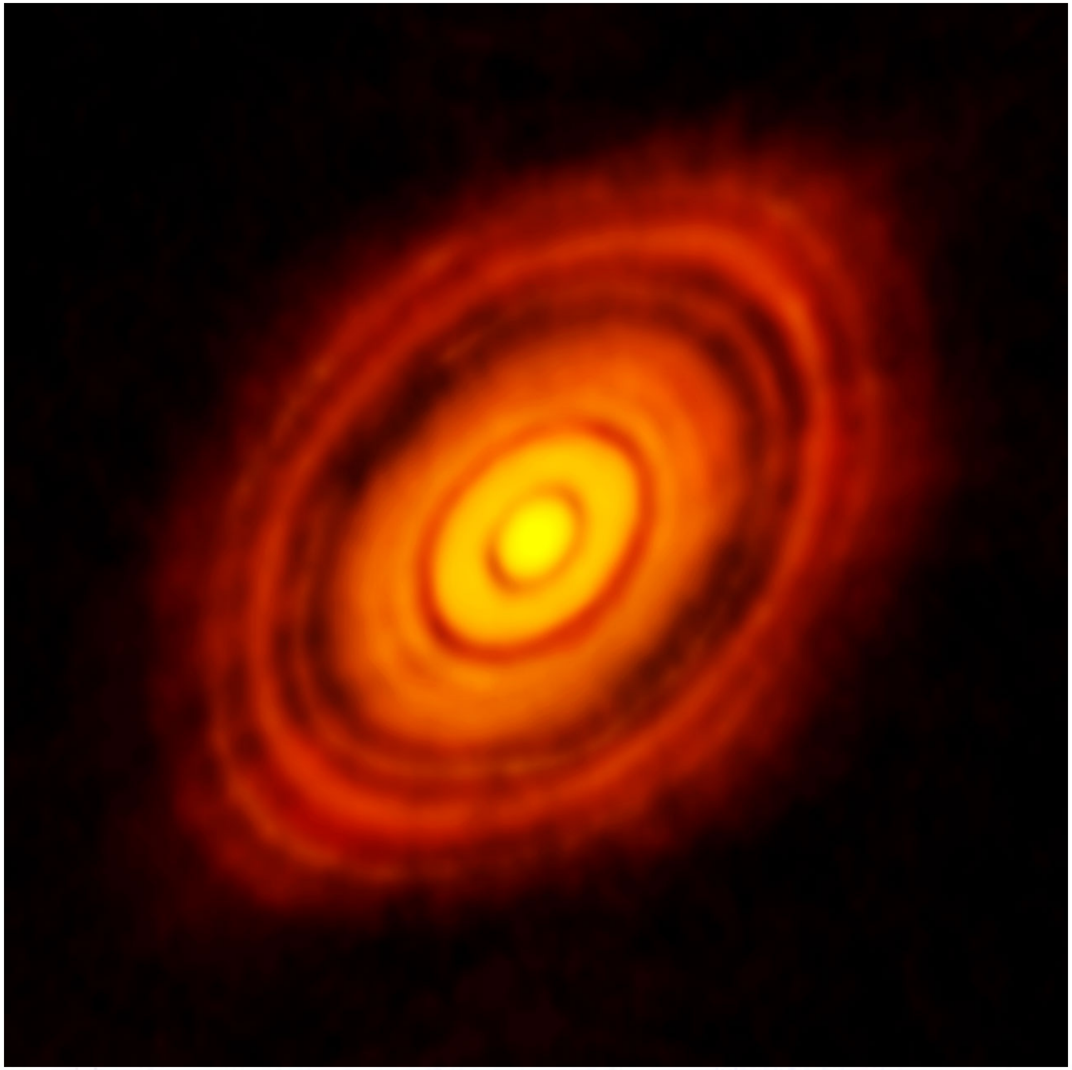
Atacama Large Millimeter/submillimeter Array
Karl G. Jansky Very Large Array
Very Long Baseline Array

What can you observe?

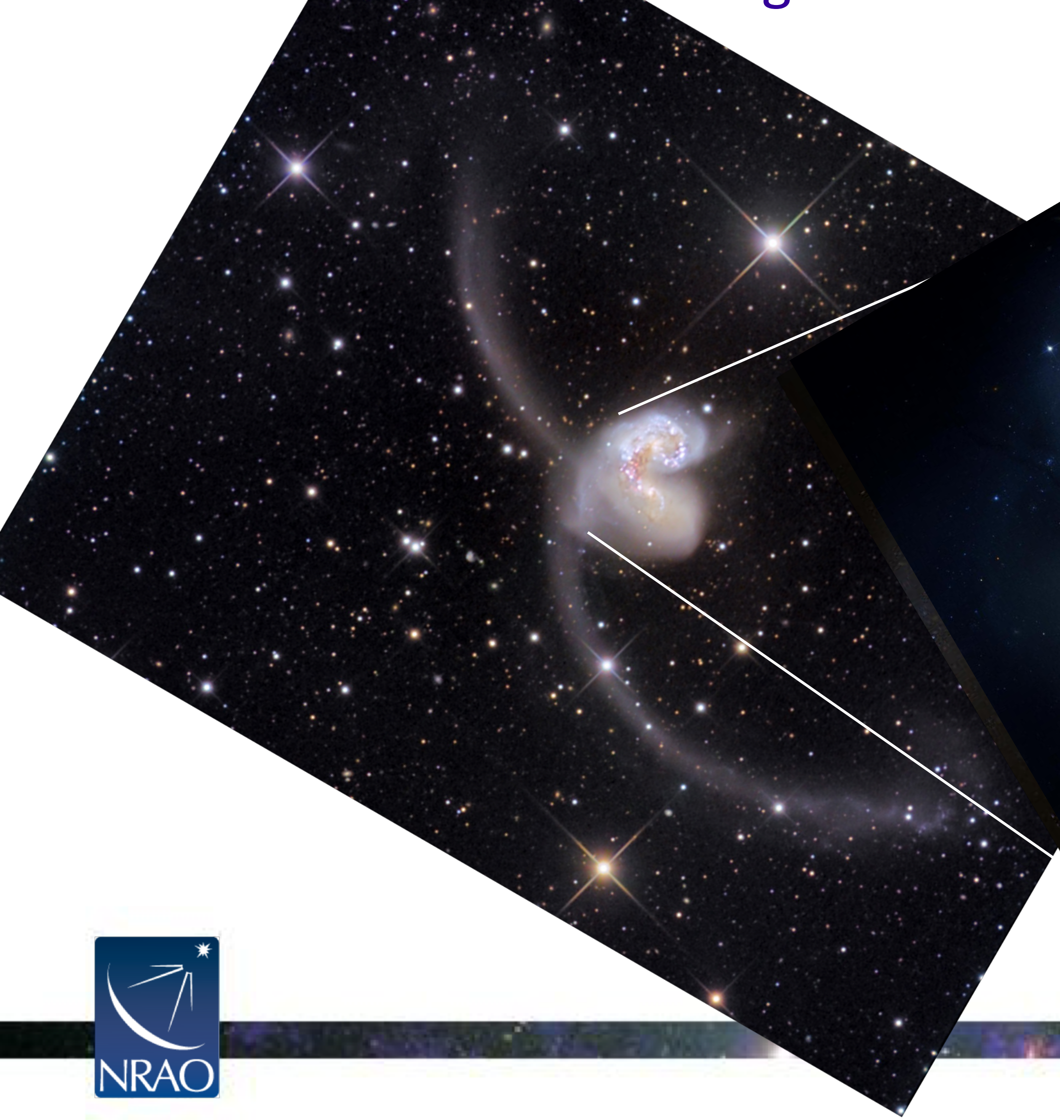
Protoplanetary discs like this one around HL Tauri



Credit: *ALMA (NRAO/ESO/NAOJ)*

What can we observe?

At higher frequencies we can observe a broad range of molecular lines



Resolution of Observations

Angular resolution for most telescopes is ~ λ/D

D is the diameter of the telescope and λ is the wavelength of observation

For the Hubble Space Telescope:

λ ~ 1um / D of 2.4m = resolution ~ 0.13"

To reach that resolution at λ ~1mm, we would need a 2 km-diameter dish!

Instead, we use arrays of smaller dishes to achieve the same high angular resolution at radio frequencies

This is interferometry!

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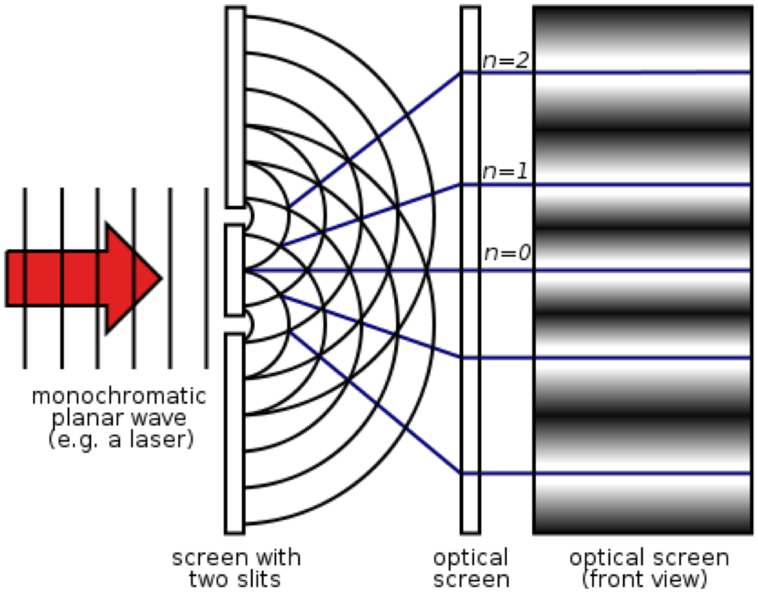
What is an interferometer?

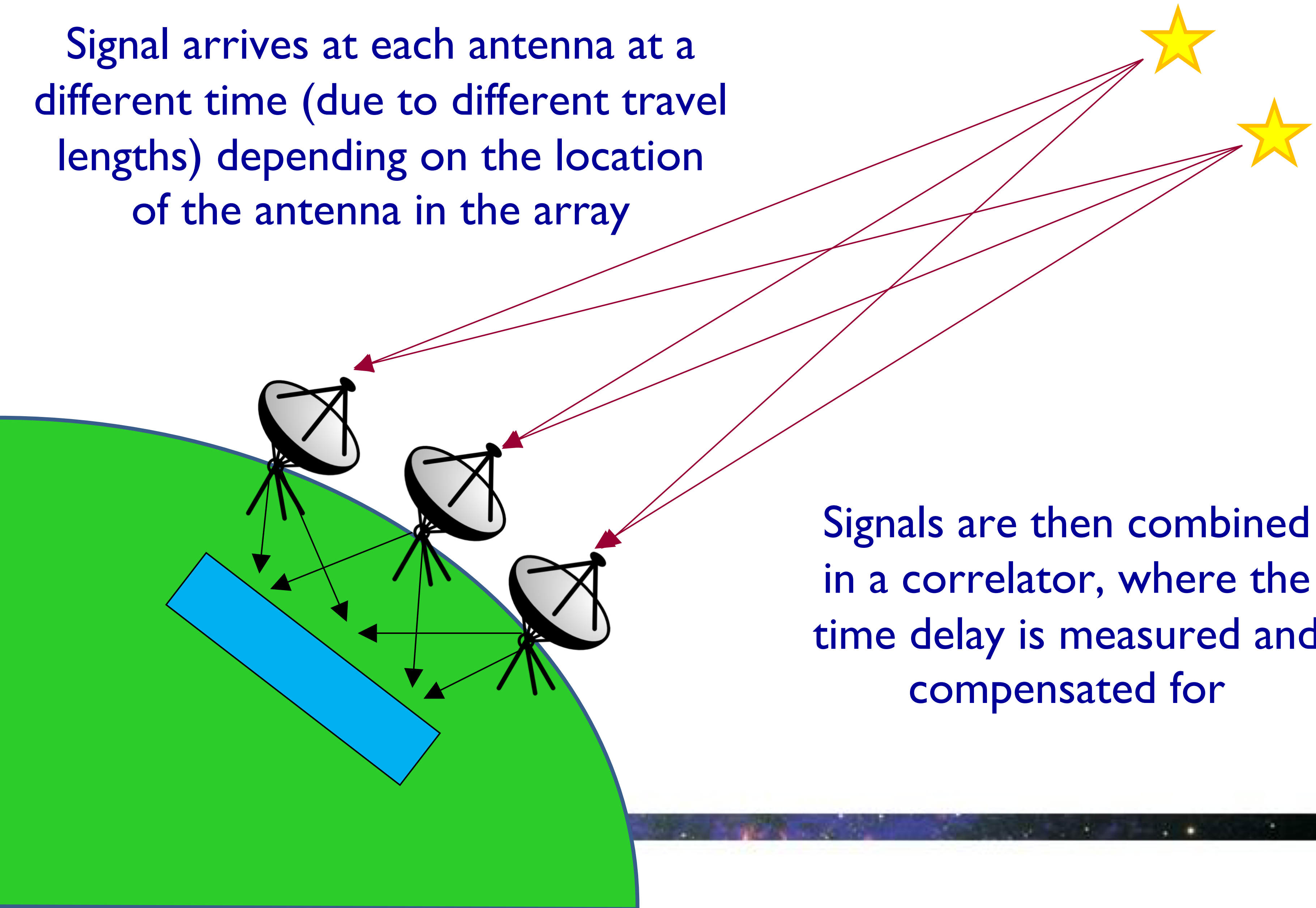
An *interferometer* measures the interference pattern produced by multiple apertures, much like a 2-slit experiment



******However, the interference patterns measured by radio telescopes are produced by* ***multiplying*** *- not adding - the wave signals measured at the different telescopes (i.e. apertures)*

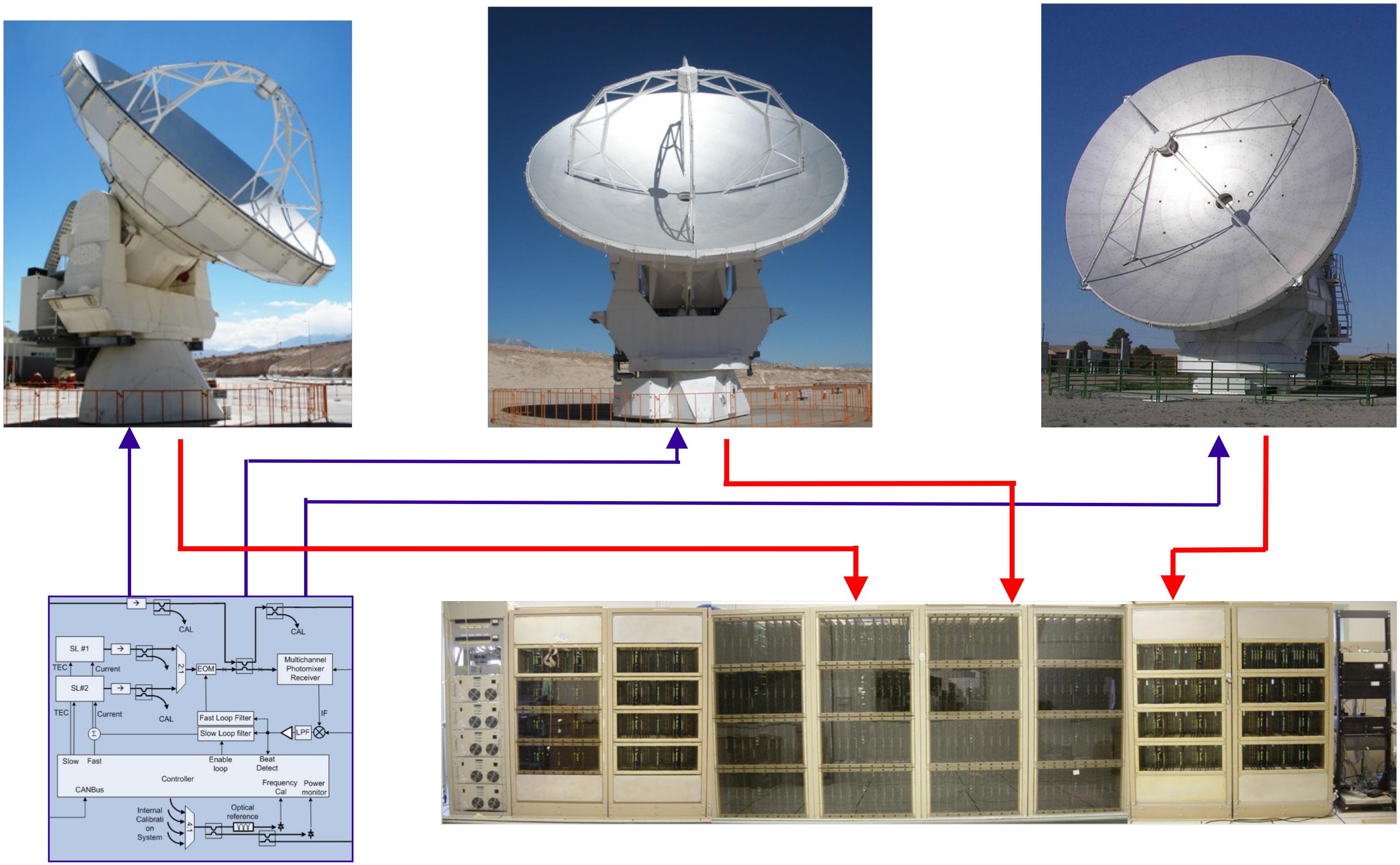
How Do We Use Interferometry?

Signal arrives at each antenna at a different time (due to different travel lengths) depending on the location of the antenna in the array



Signals are then combined in a correlator, where the time delay is measured and compensated for

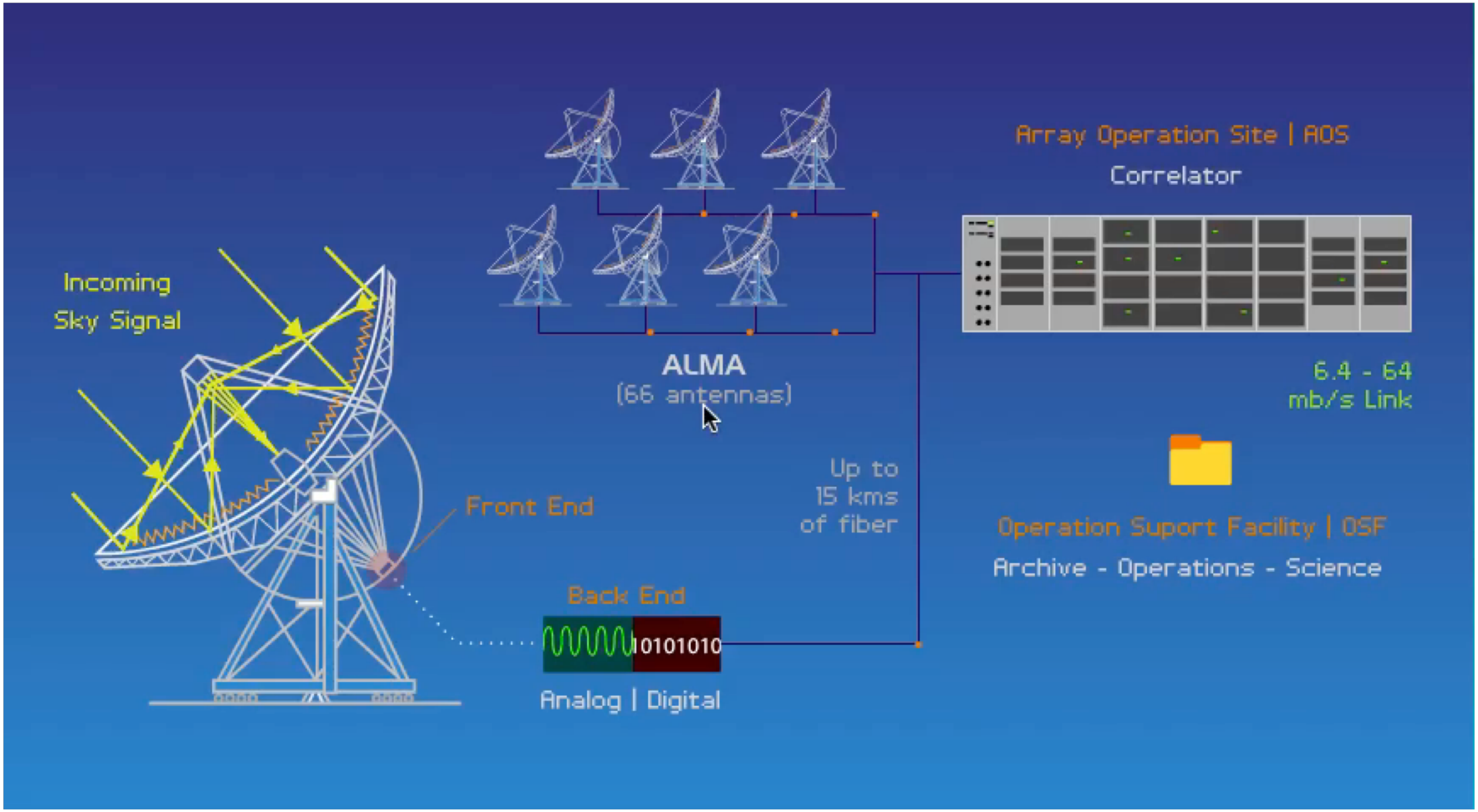
Some Instrument Details



Signals from each antenna are digitized and sent to the correlator for multiplication & averaging

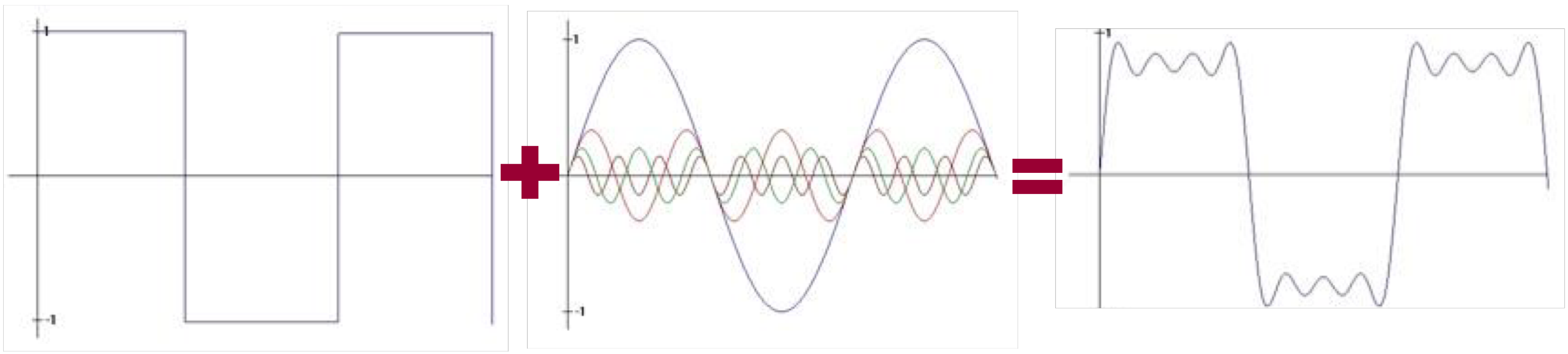
For ~50 antennas, the data rate is 600 GB/sec for the correlator to process

An Interferometer In Action



Introducing the Fourier Transform

Fourier theory states that any well behaved signal (including images) can be expressed as the sum of sinusoids



Reference signal

4 sinusoids

Sum of sinusoids & signal

The Fourier transform is the mathematical tool that decomposes a signal into its sinusoidal components

The Fourier transform contains *all* of the information of the original signal

Visibility and Sky Brightness

The van Cittert-Zernike theorem

Visibility as a function of baseline coordinates (u,v) is the Fourier transform of the sky brightness distribution as a function of the sky coordinates (x,y)

$$V(u,v) \xrightarrow{\textbf{FT}} T(x,y)$$

V(u,v) = the complex visibility function $= \iint \boldsymbol{T(x,y)e^{2\pi i(ux+vy)}dxdy}$

T(x,y) = the sky brightness distribution $= \iint \boldsymbol{V(u,v)e^{-2\pi i(ux+vy)}dudv}$

What Are Visibilities?

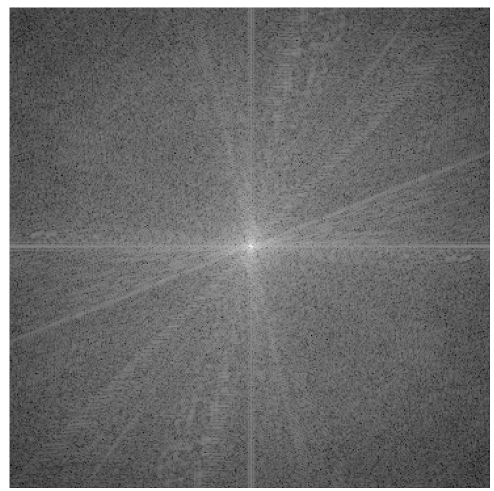
Each V(u,v) contains information on T(x,y) everywhere

Each V(u,v) is a complex quantity

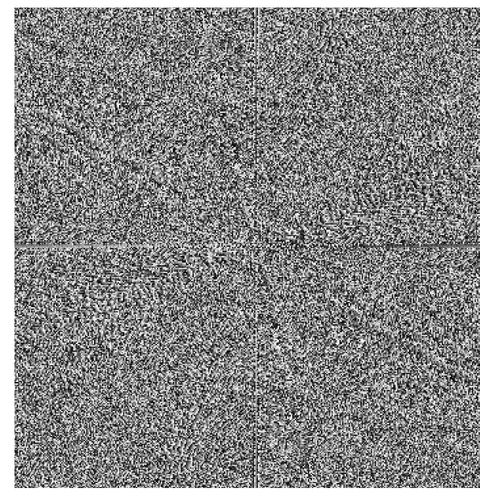
Expressed as (real, imaginary) or (amplitude, phase)

T(x,y)

FT →



V(u,v) amplitude



V(u,v) phase

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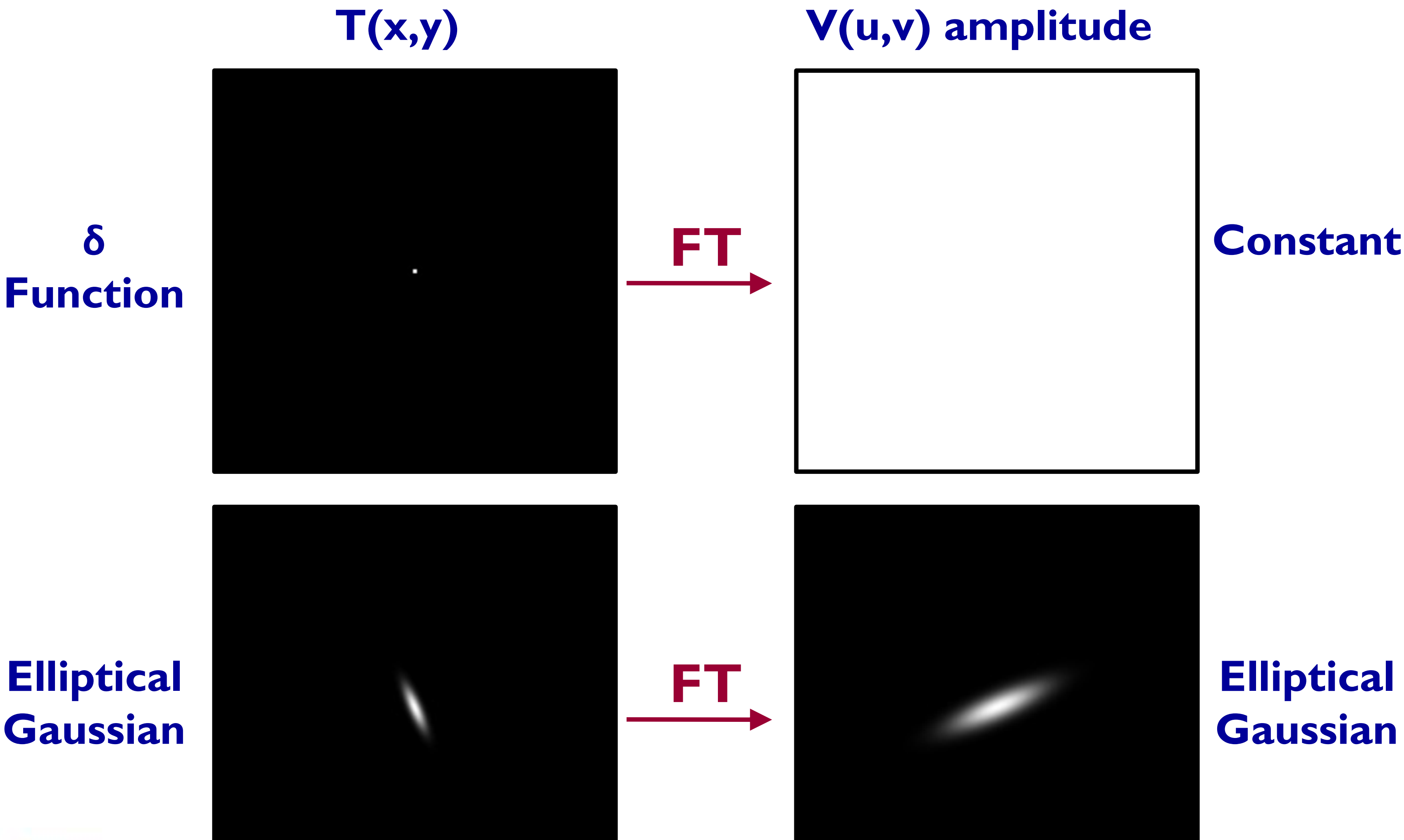
Examples of 2D Fourier Transforms



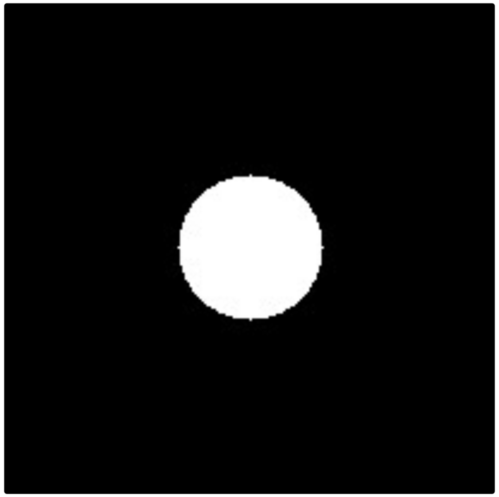
Rules of the Fourier Transform:

Narrow features transform to wide features (and vice versa)

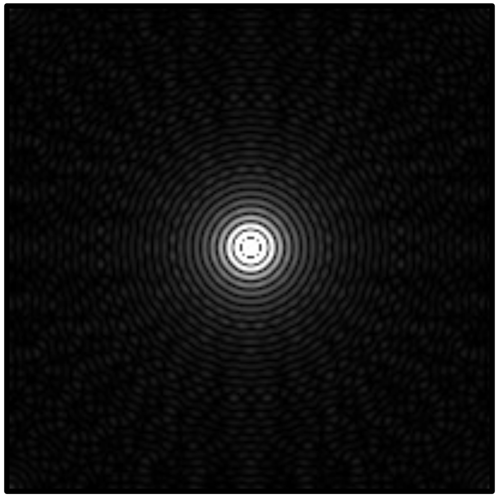
Examples of 2D Fourier Transforms

T(x,y) **V(u,v) amplitude**

Uniform Disk

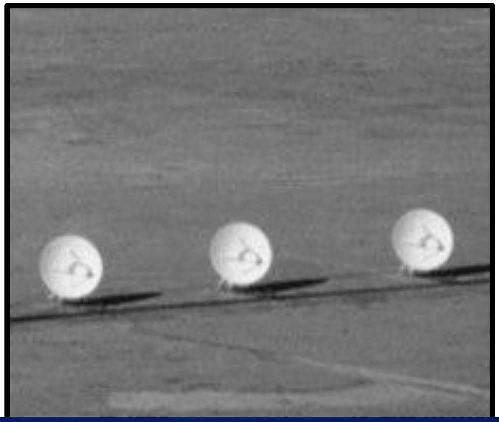


FT →



Bessel Function

VLA Antennas



FT →

Bessel Function!

Rules of the Fourier Transform:

Sharp features (edges) result in many high spatial features

Basics of Aperture Synthesis

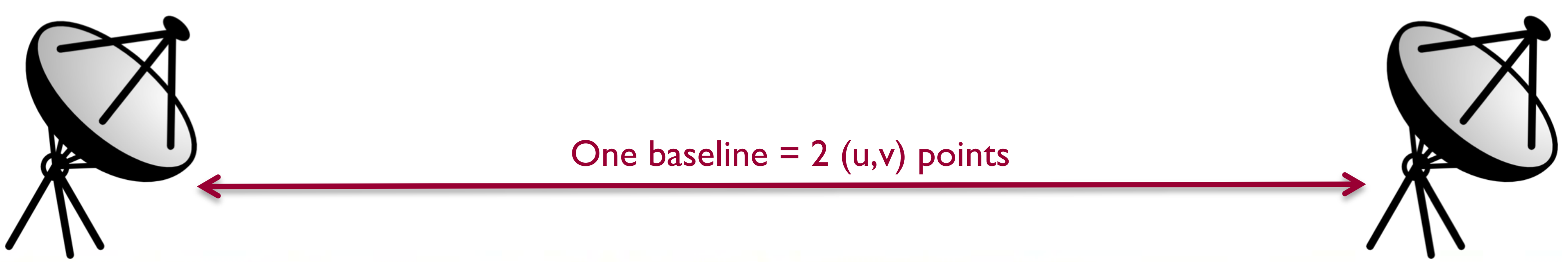
Idea: Sample V(u,v) at an enough (u,v) points using distributed small aperture antennas to synthesize a large aperture antenna of size (u_{max}, v_{max})

One pair of antennas = one baseline

For **N antennas**, we get **N(N-1) samples** at a time

How do we fill in the rest of the (u,v) plane?

1. Earth's rotation
2. Reconfigure physical layout of N antennas



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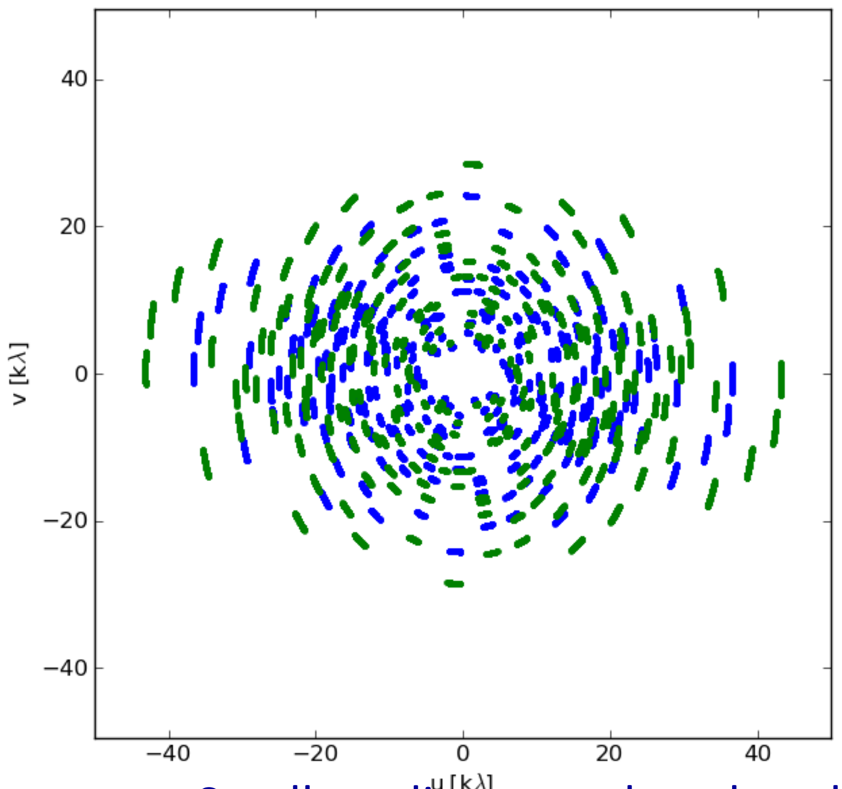
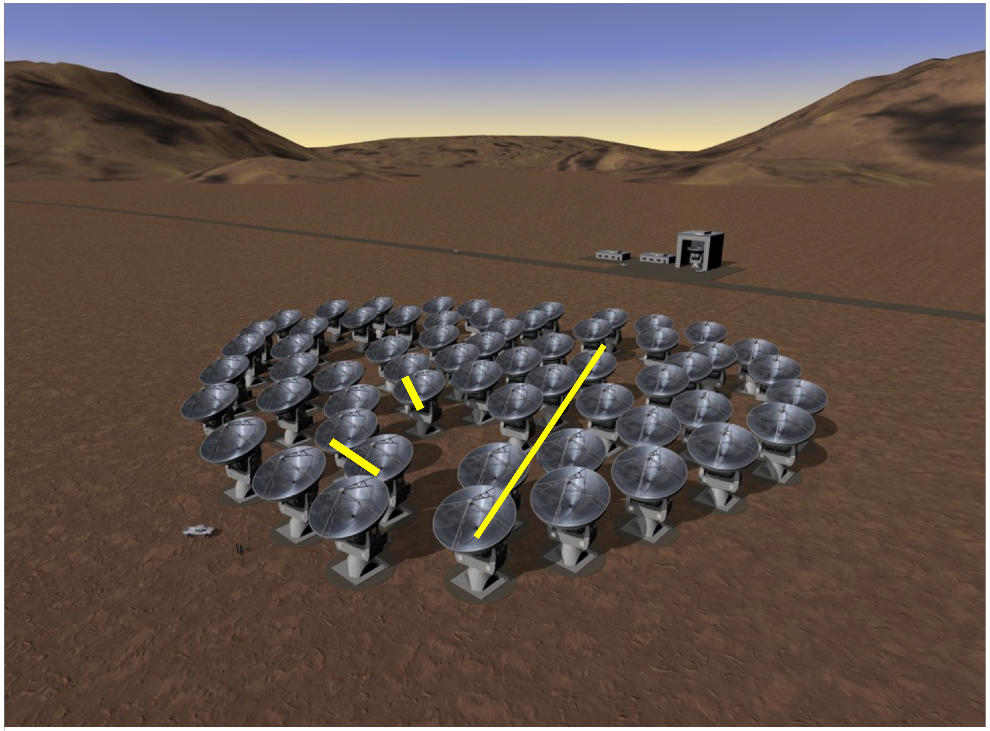
Sampling Function

Each antenna pair samples only one spot; the array cannot sample the entire Fourier/uv domain resulting in an **imperfect image**





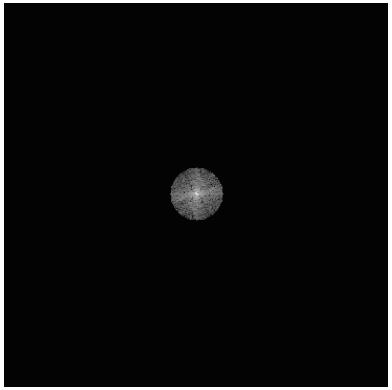
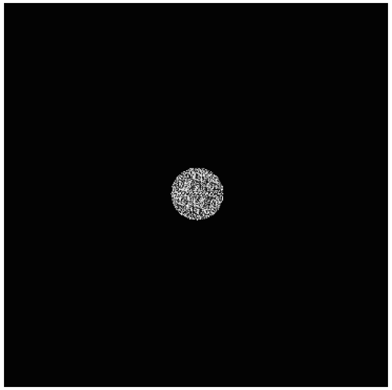
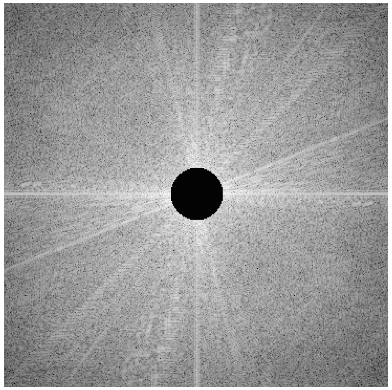
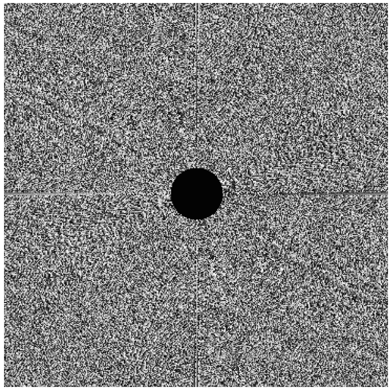
Small uv-distance: short baselines (measure extended emission)

Long uv-distance: long baselines (measure small scale emission)

Orientation of baseline also determines orientation in the uv-plane

Implications of (u,v) Coverage

What does it mean if our (u,v) coverage is not complete?

	V(u,v) amplitude	V(u,v) phase		T(x,y)
Missing High Spatial Frequencies			FT →	
Missing Low Spatial Frequencies			FT →	

Characteristic Angular Scales

Angular resolution of telescope array:

$\sim \lambda/B_{max}$ (B_{max} = longest baseline)

Maximum angular scale:

$\sim \lambda/B_{min}$ (B_{min} = shortest distance between antennas)

Field of view (FOV):

$\sim \lambda/D$ (D = antenna diameter)

*Sources more extended than the FOV can be observed using multiple pointing centers in a mosaic

An interferometer is sensitive to a range of angular sizes: $\lambda/B_{max} < \theta < \lambda/B_{min}$

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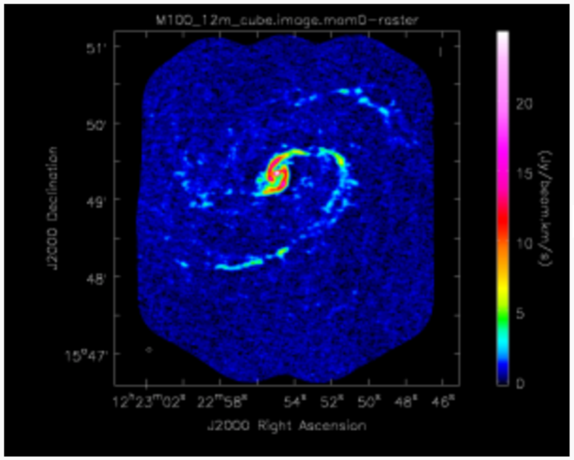
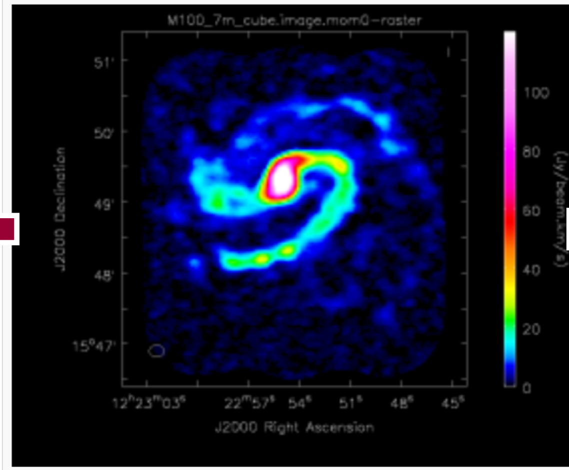
Characteristic Angular Scales: M100

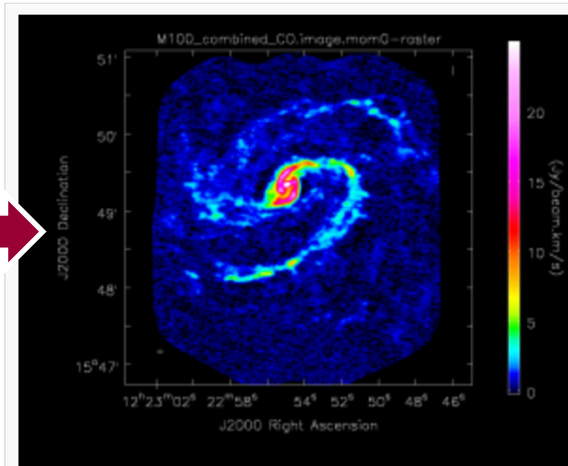
ALMA 12m



ACA 7m



Combined



ALMA 12m shows smaller spatial scales (denser, clumpier emission)
ACA 7m data shows larger spatial scales (diffuse, extended emission)

To get both — you need a combined image!

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Interferometry: Spatial Scales

- The **sensitivity** is given by the number of antennas times their area
- The **field of view** is given by the beam of a single antenna (corresponding to the resolution for a single dish telescope or the primary beam)
- The **resolution** is given by the largest distance between antennas (called the synthesized beam)
- The **largest angular scale** that can be imaged is given by the shortest distance between antennas

Angular Scales — A Proposal Tip!

Interferometers act as spatial filters - shorter baselines are sensitive to larger targets, so remember …

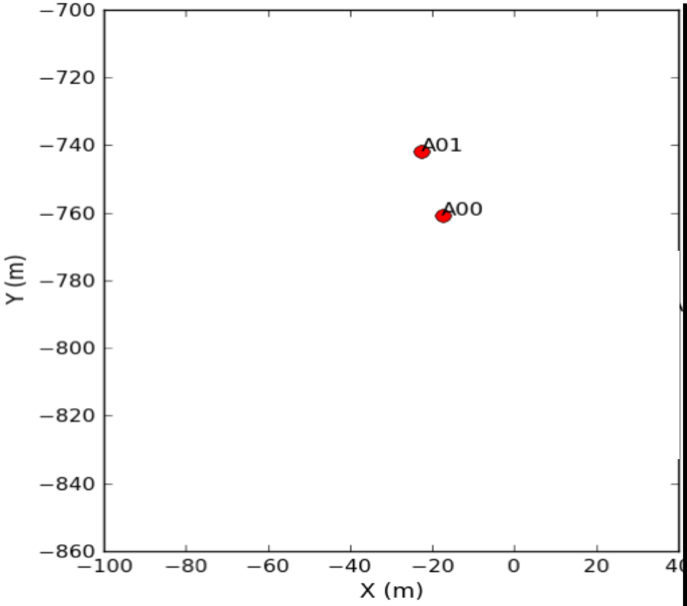
Spatial scales larger than the smallest baseline cannot be imaged

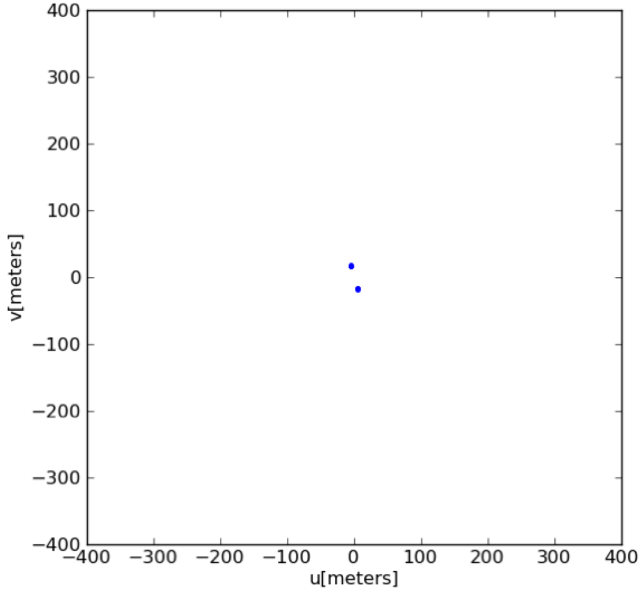
Spatial scales smaller than the largest baseline cannot be resolved

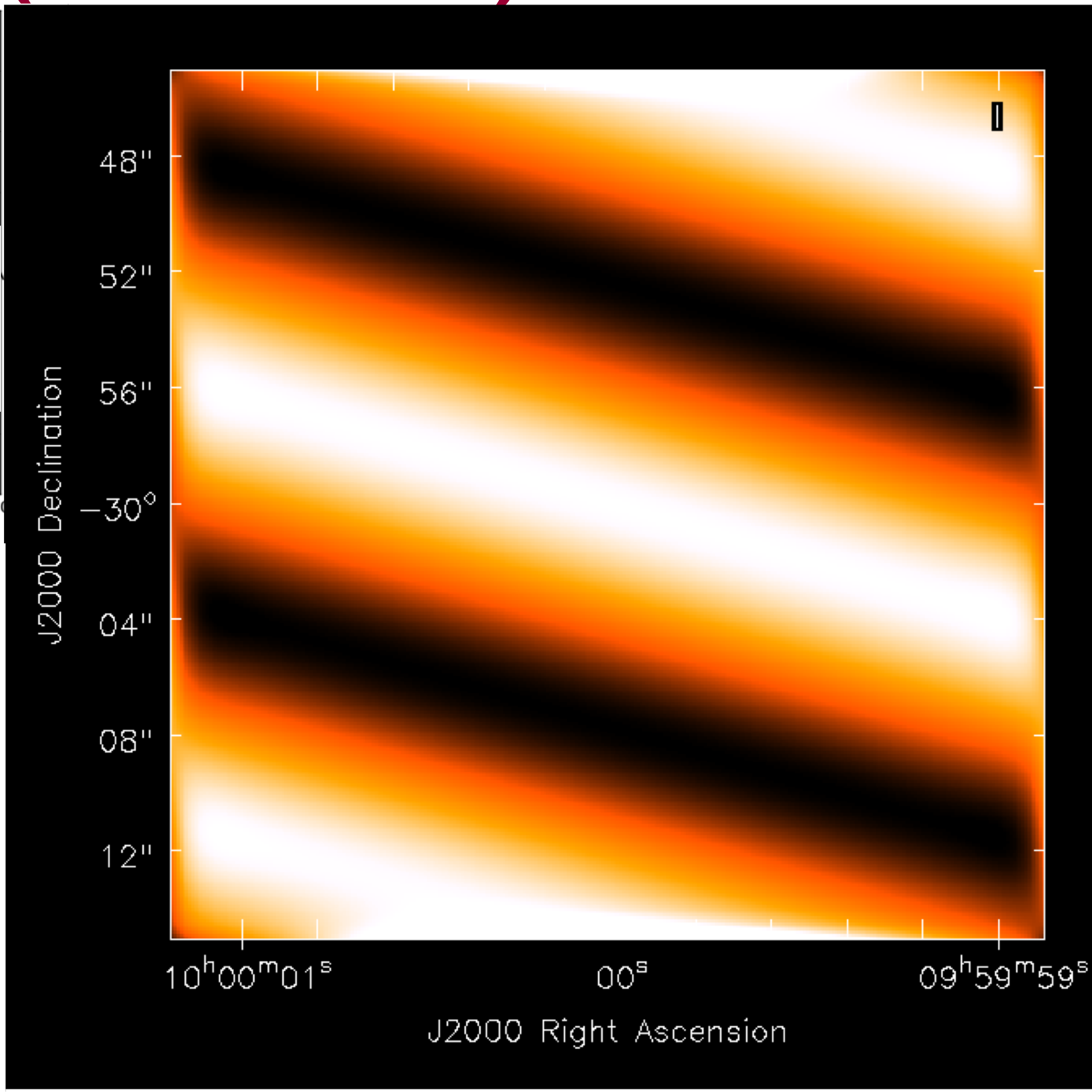
Table A-1: Angular Resolutions (AR) and Maximum Recoverable Scales (MRS) for the Cycle 7 Array configurations

Config	Lmax		Band 3	Band 4	Band 5	Band 6	Band 7	Band 8	Band 9	Band 10
	Lmin		100 GHz	150 GHz	183 GHz	230 GHz	345 GHz	460 GHz	650 GHz	870 GHz
7-m Array	45 m	AR	12.5″	8.4″	6.8″	5.4″	3.6″	2.7″	1.9″	1.4″
	9 m	MRS	66.7″	44.5″	36.1″	29.0″	19.3″	14.5″	10.3″	7.7″
C43-1	161 m	AR	3.4″	2.3″	1.8″	1.5″	1.0″	0.74″	0.52″	0.39″
	15 m	MRS	28.5″	19.0″	15.4″	12.4″	8.3″	6.2″	4.4″	3.3″
C43-2	314 m	AR	2.3″	1.5″	1.2″	1.0″	0.67″	0.50″	0.35″	0.26″
	15 m	MRS	22.6″	15.0″	12.2″	9.8″	6.5″	4.9″	3.5″	2.6″
C43-3	500 m	AR	1.4″	0.94″	0.77″	0.62″	0.41″	0.31″	0.22″	0.16″
	15 m	MRS	16.2″	10.8″	8.7″	7.0″	4.7″	3.5″	2.5″	1.9″
C43-4	784 m	AR	0.92″	0.61″	0.50″	0.40″	0.27″	0.20″	0.14″	0.11″
	15 m	MRS	11.2″	7.5″	6.1″	4.9″	3.3″	2.4″	1.7″	1.3″
C43-5	1.4 km	AR	0.54″	0.36″	0.30″	0.24″	0.16″	0.12″	0.084″	0.063″
	15 m	MRS	6.7″	4.5″	3.6″	2.9″	1.9″	1.5″	1.0″	0.77″
C43-6	2.5 km	AR	0.31″	0.20″	0.16″	0.13″	0.089″	0.067″	0.047″	0.035″
	15 m	MRS	4.1″	2.7″	2.2″	1.8″	1.2″	0.89″	0.63″	0.47″
C43-7	3.6 km	AR	0.21″	0.14″	0.11″	0.092″	0.061″	0.046″	0.033″	0.024″
	64 m	MRS	2.6″	1.7″	1.4″	1.1″	0.75″	0.56″	0.40″	0.30″
C43-8	8.5 km	AR	0.096″	0.064″	0.052″	0.042″	0.028″	N/A	N/A	N/A
	110 m	MRS	1.4″	0.95″	0.77″	0.62″	0.41″			
C43-9	13.9 km	AR	0.057″	0.038″	0.031″	0.025″	0.017″	N/A	N/A	N/A
	368 m	MRS	0.81″	0.54″	0.44″	0.35″	0.24″			
C43-10	16.2 km	AR	0.042″	0.028″	0.023″	0.018″	0.012″	N/A	N/A	N/A
	244 m	MRS	0.50″	0.33″	0.27″	0.22″	0.14″			

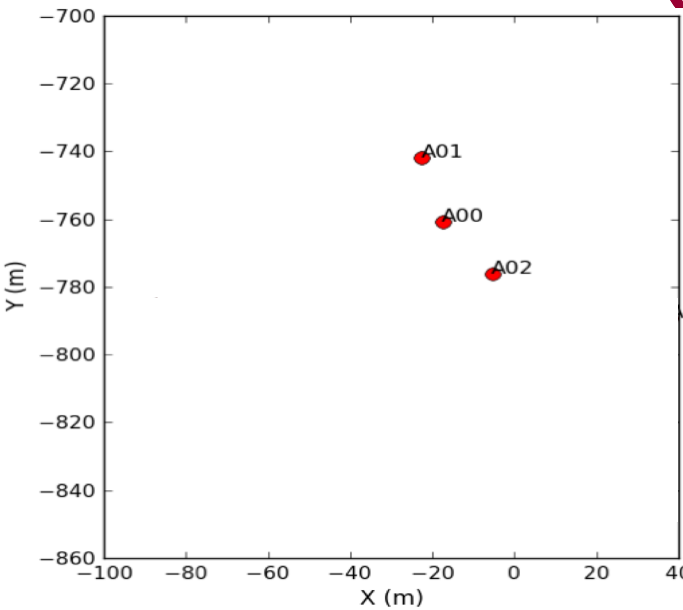
Example: Fringe pattern with 2 Antennas (one baseline)

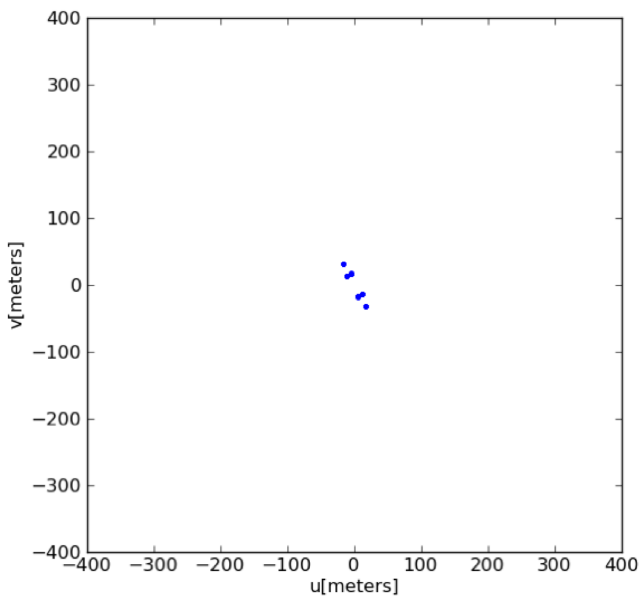


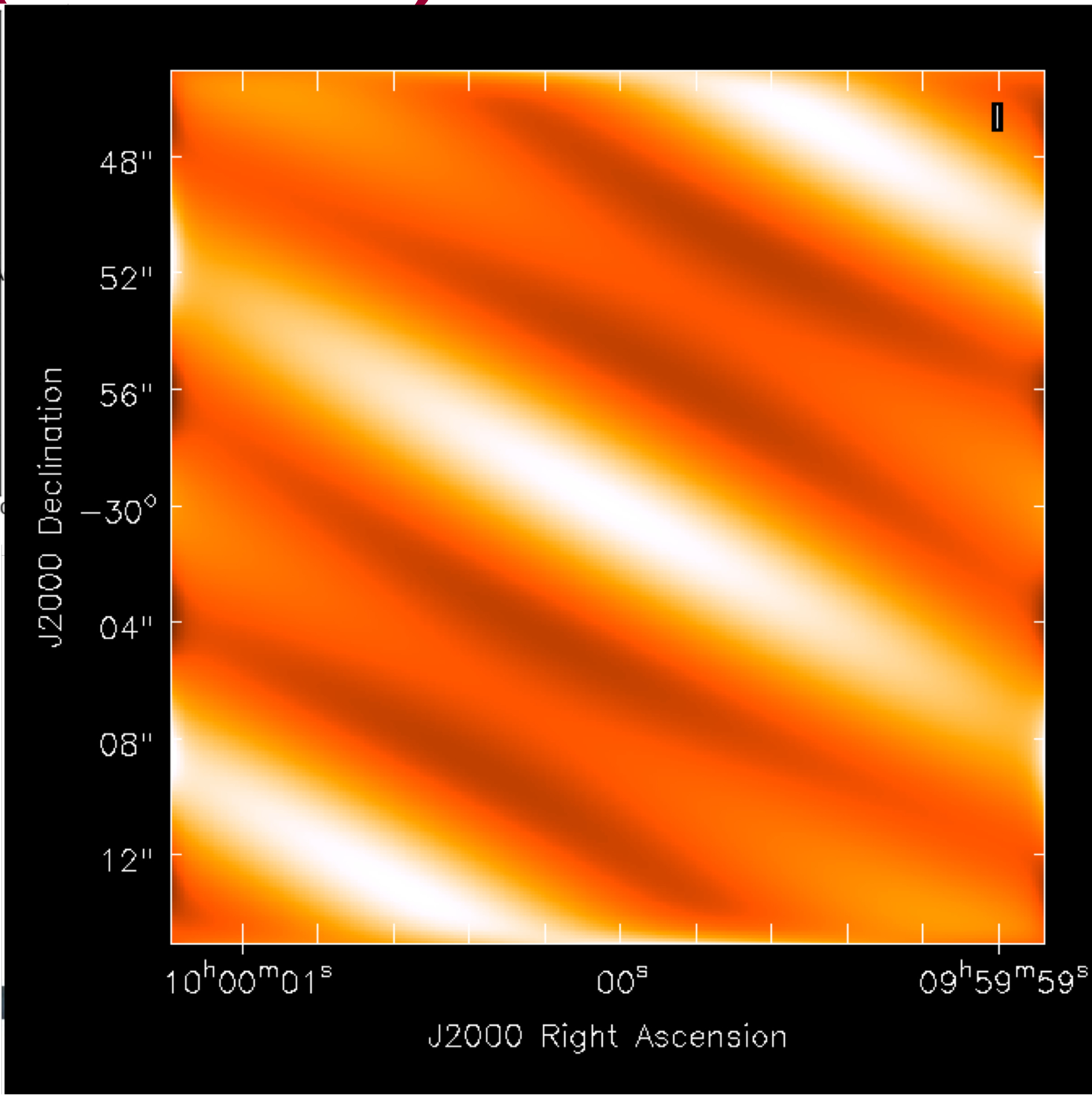




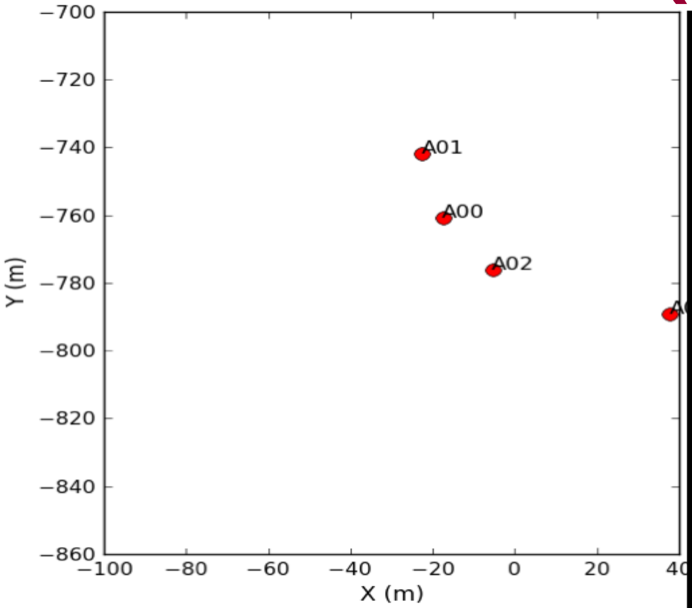
Example: Fringe pattern with 3 Antennas (3 baselines)

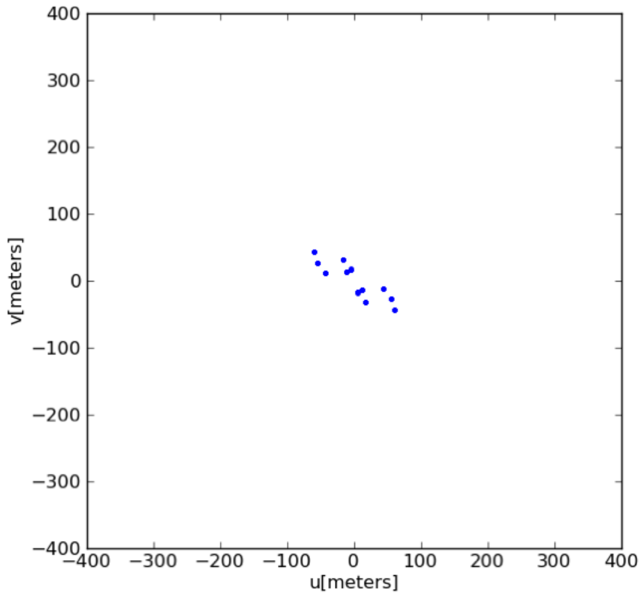


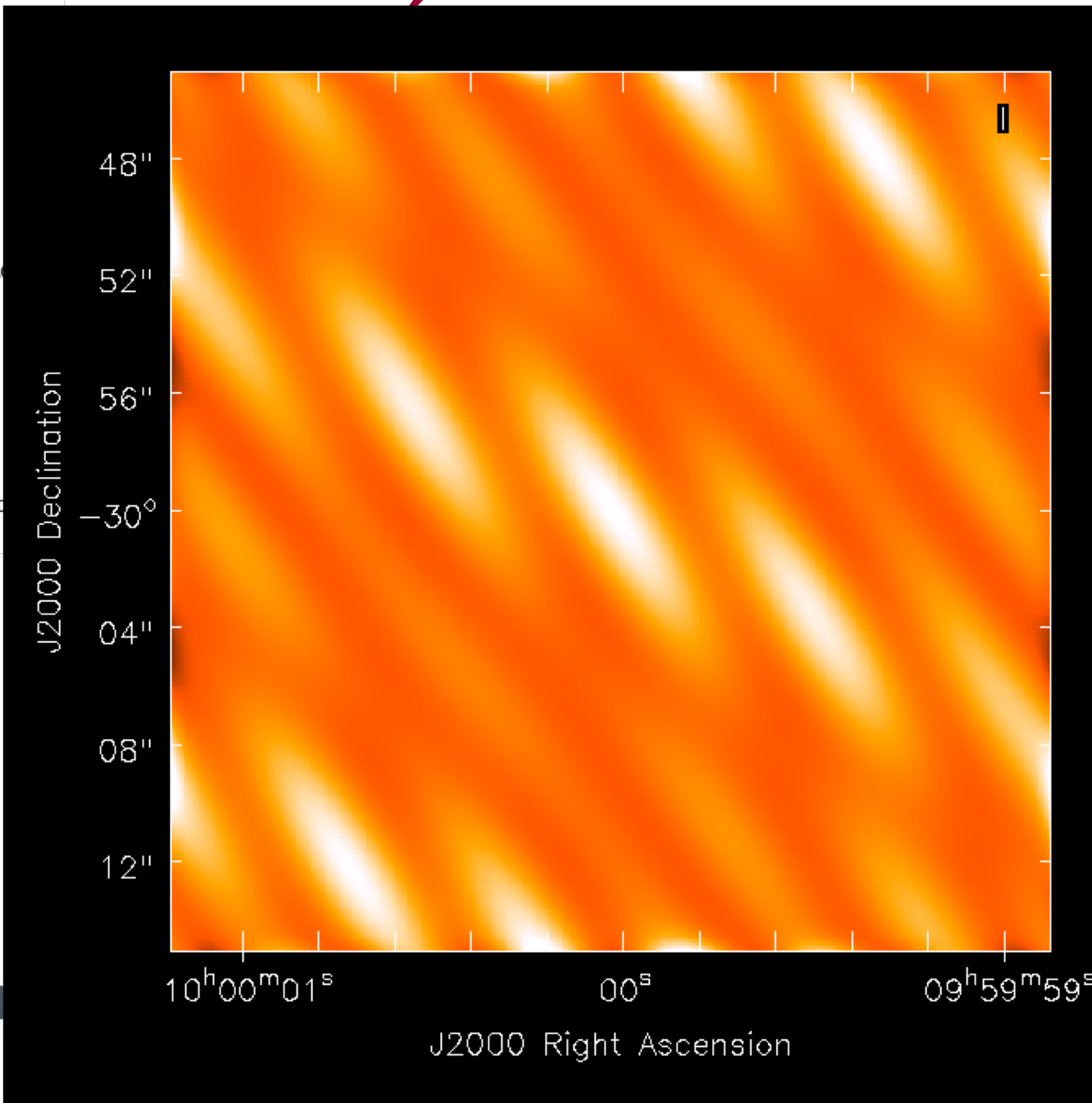




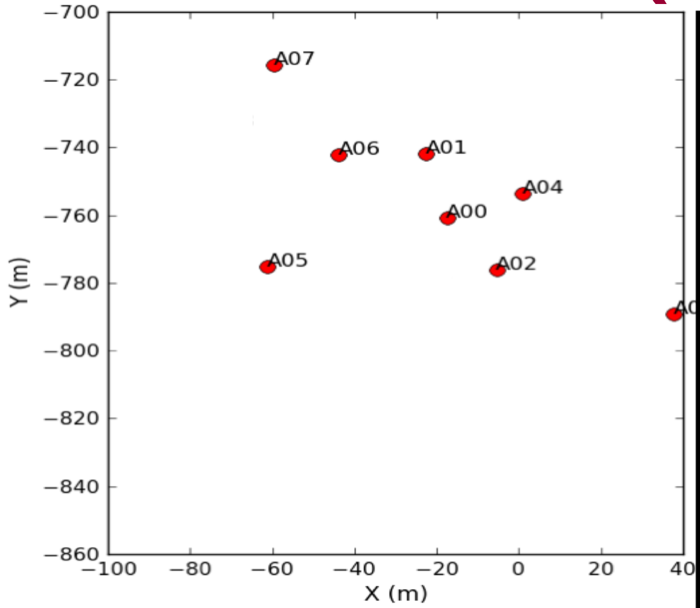
Example: Fringe pattern with 4 Antennas (6 baselines)

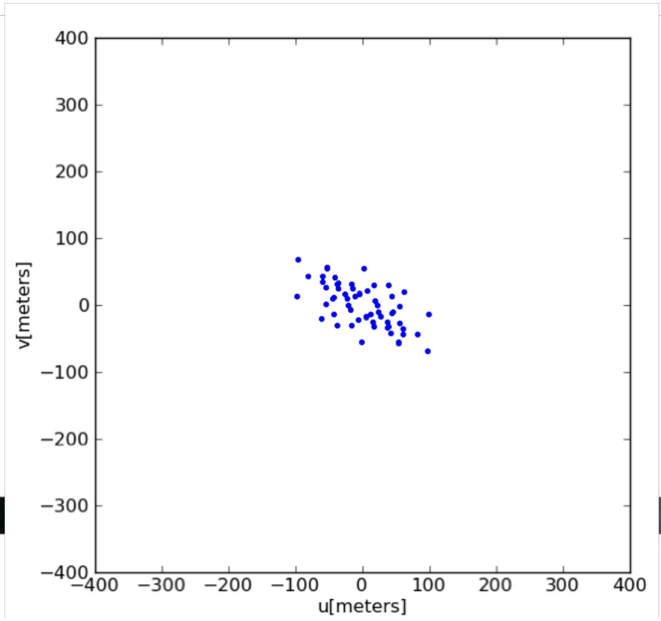


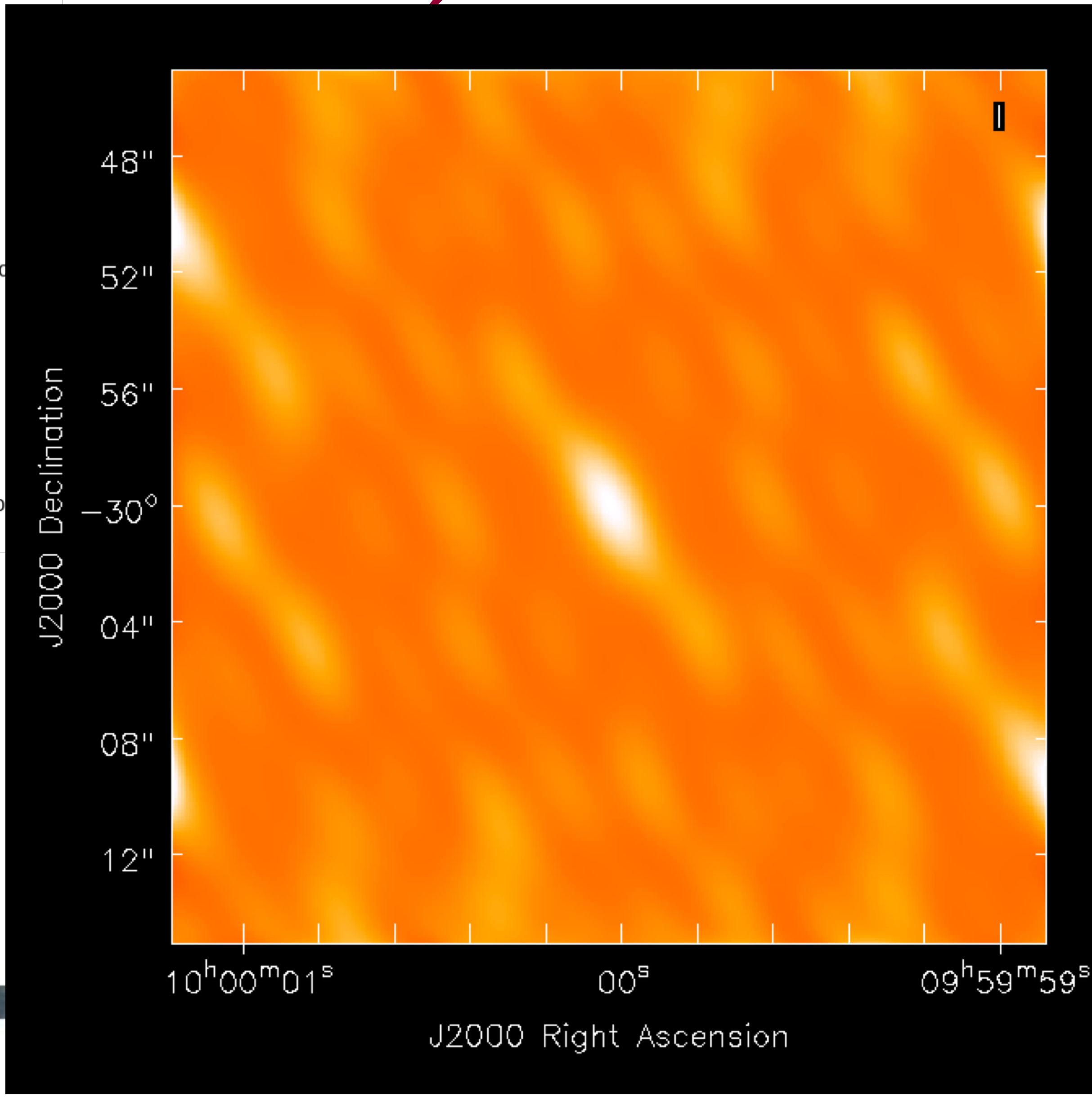




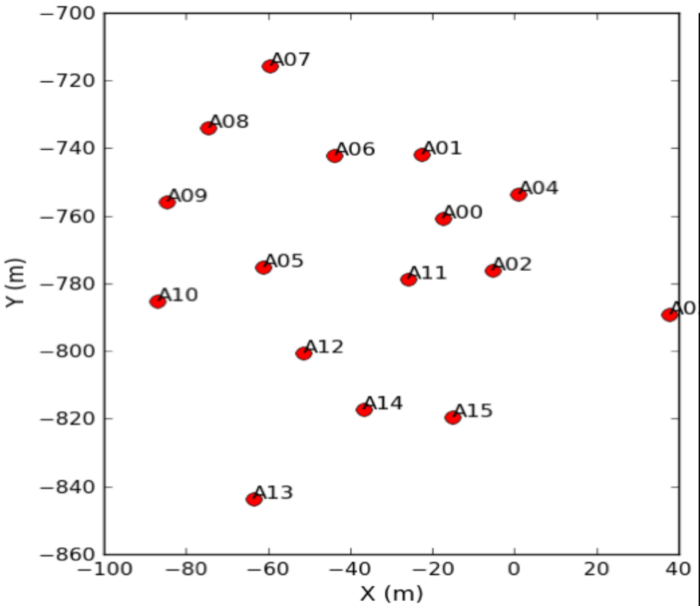
Example: Fringe pattern with 8 Antennas (28 baselines)

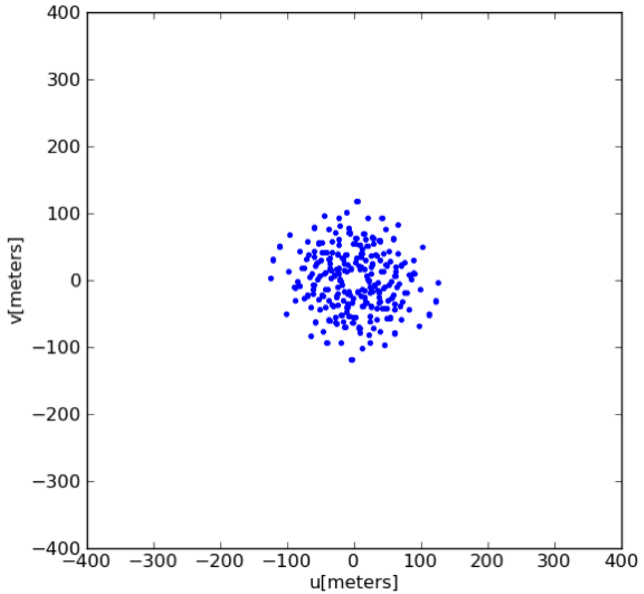


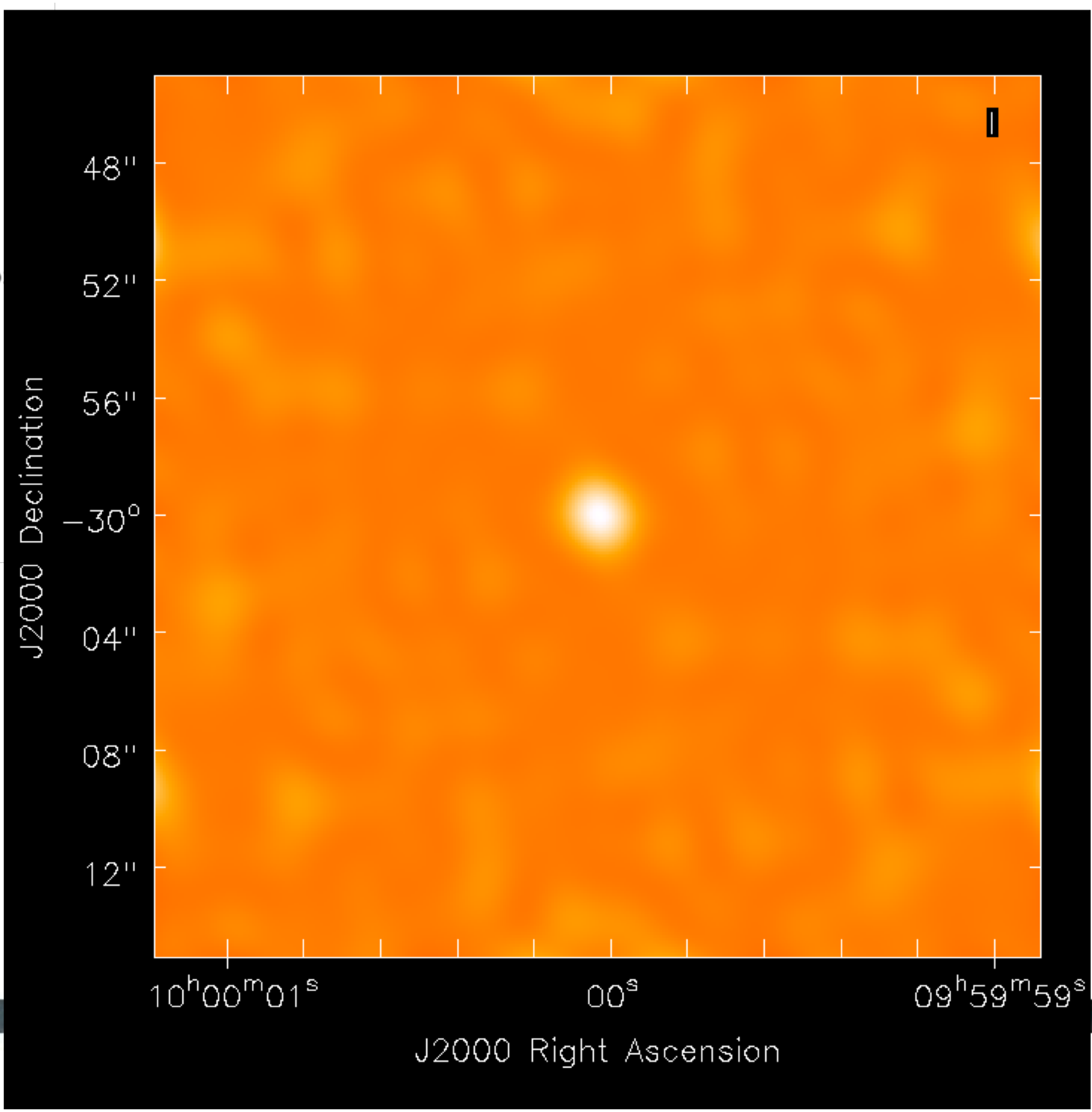




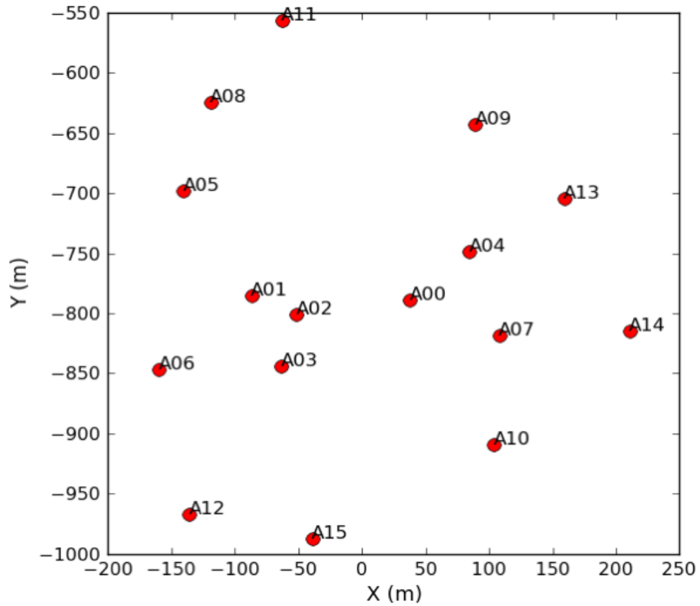
16 Antennas – Compact Configuration

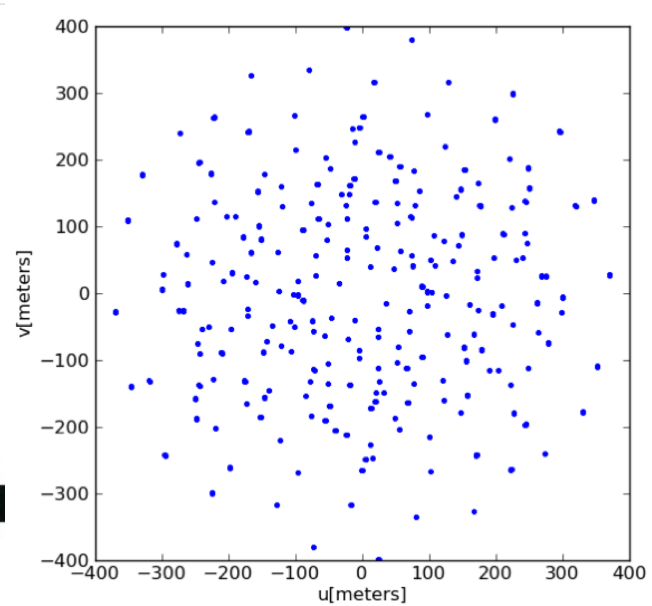


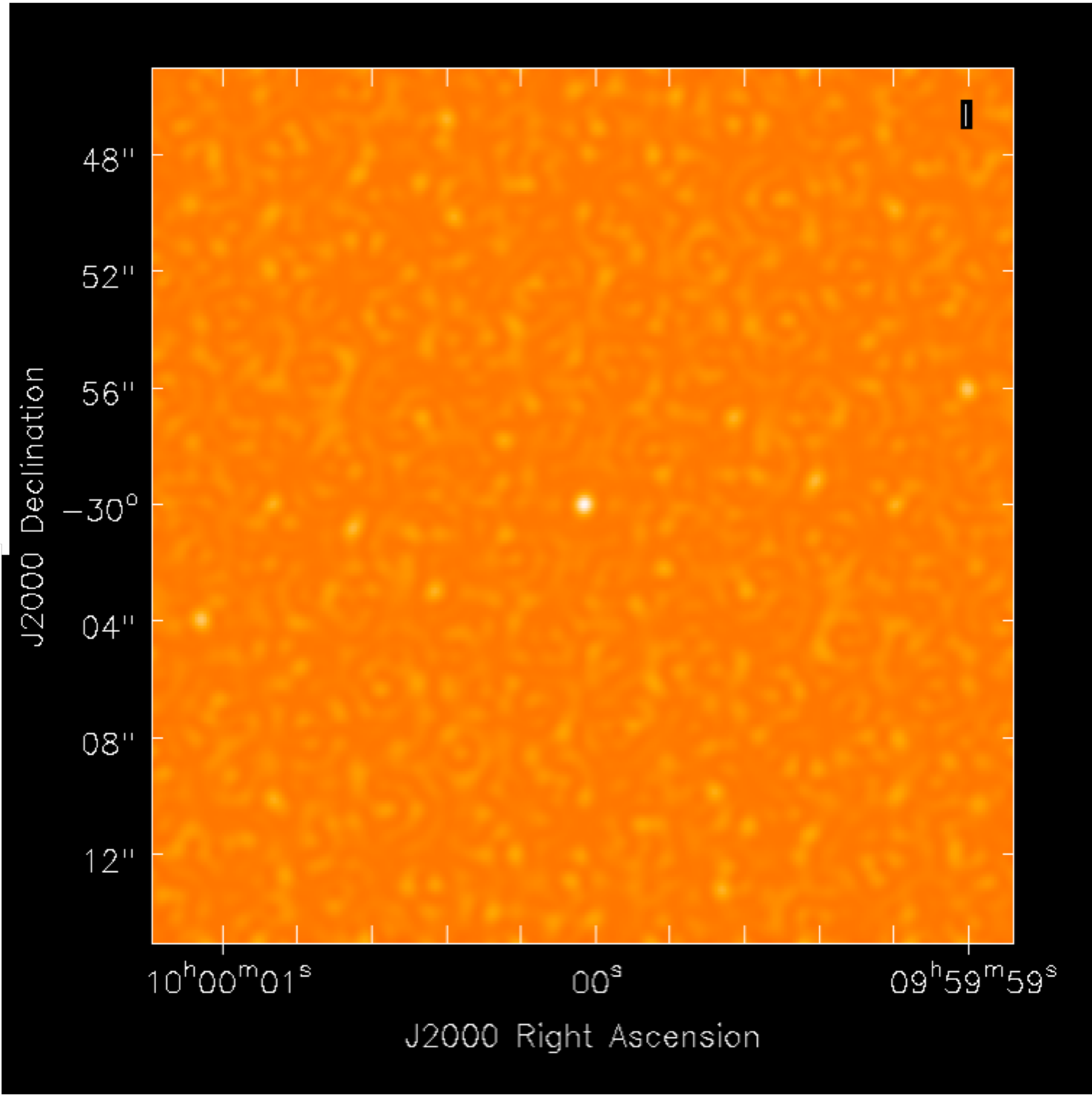




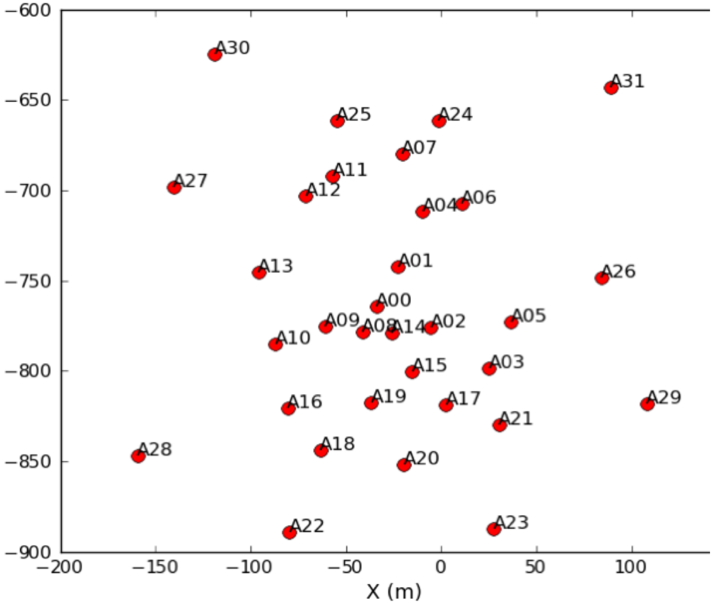
16 Antennas – Extended Configuration

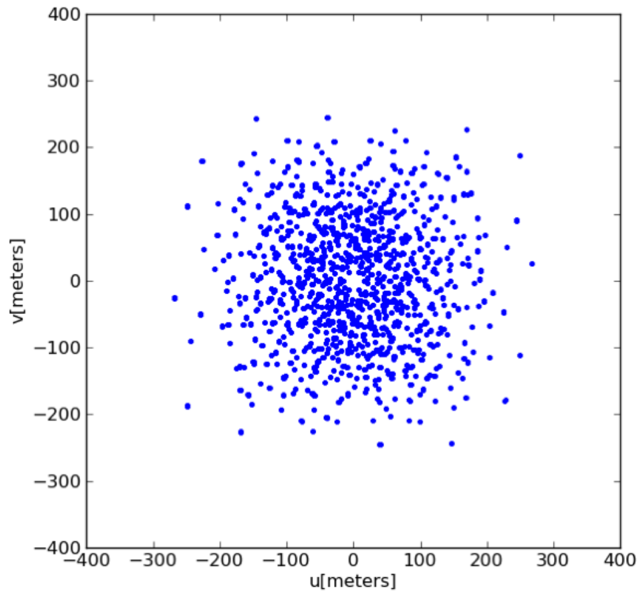


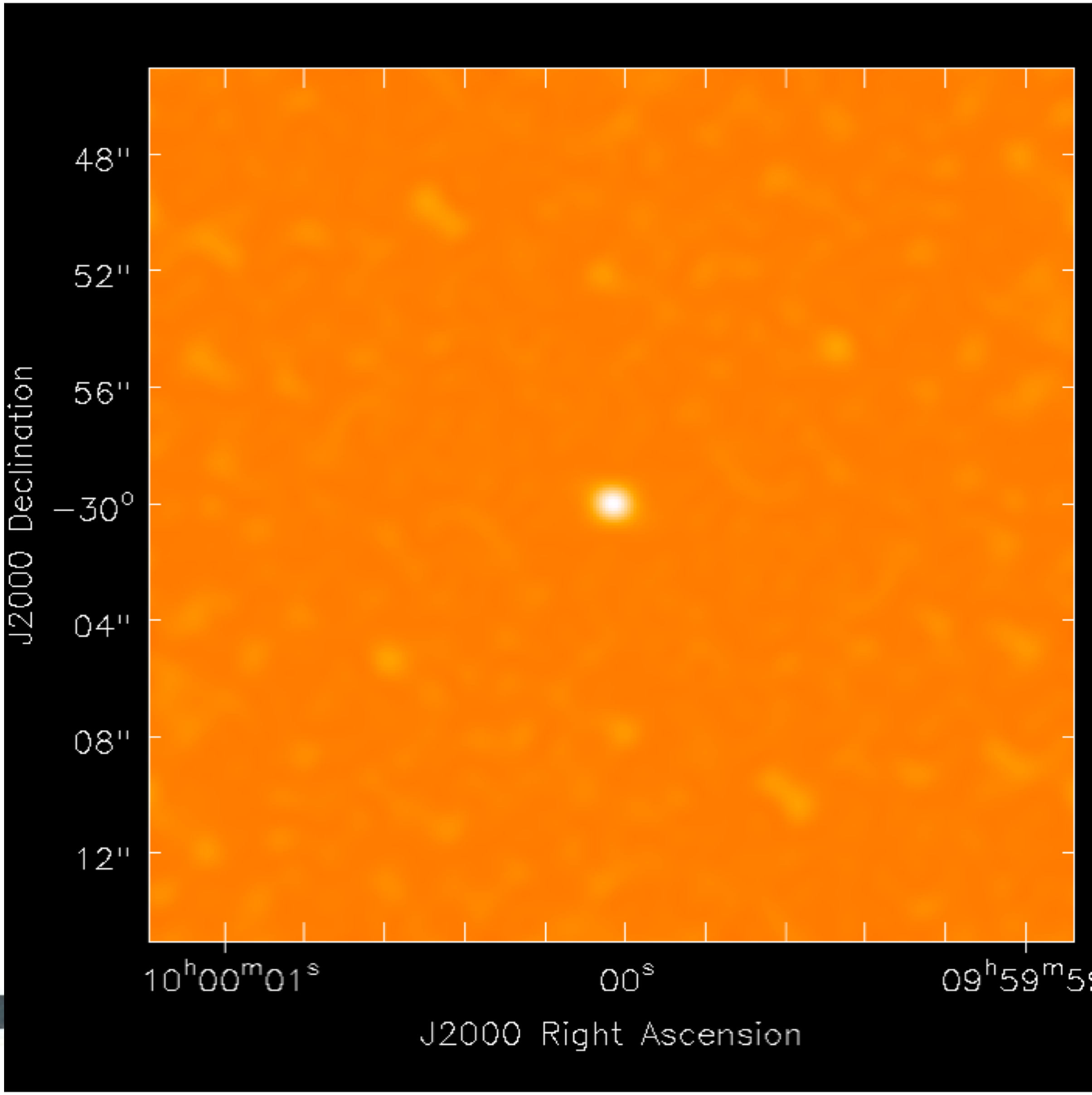




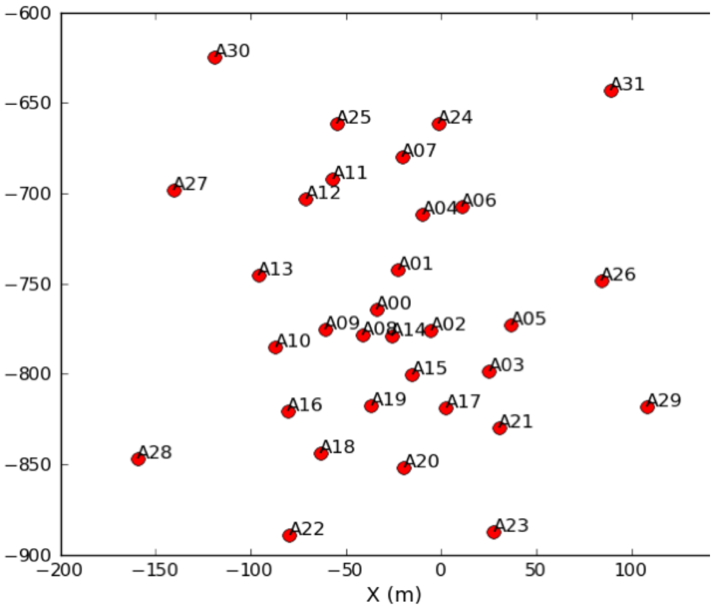
32 Antennas – Instantaneous

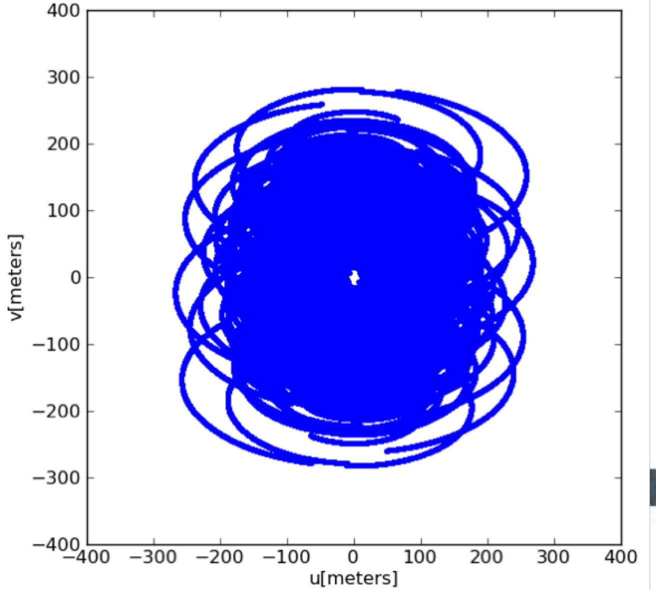


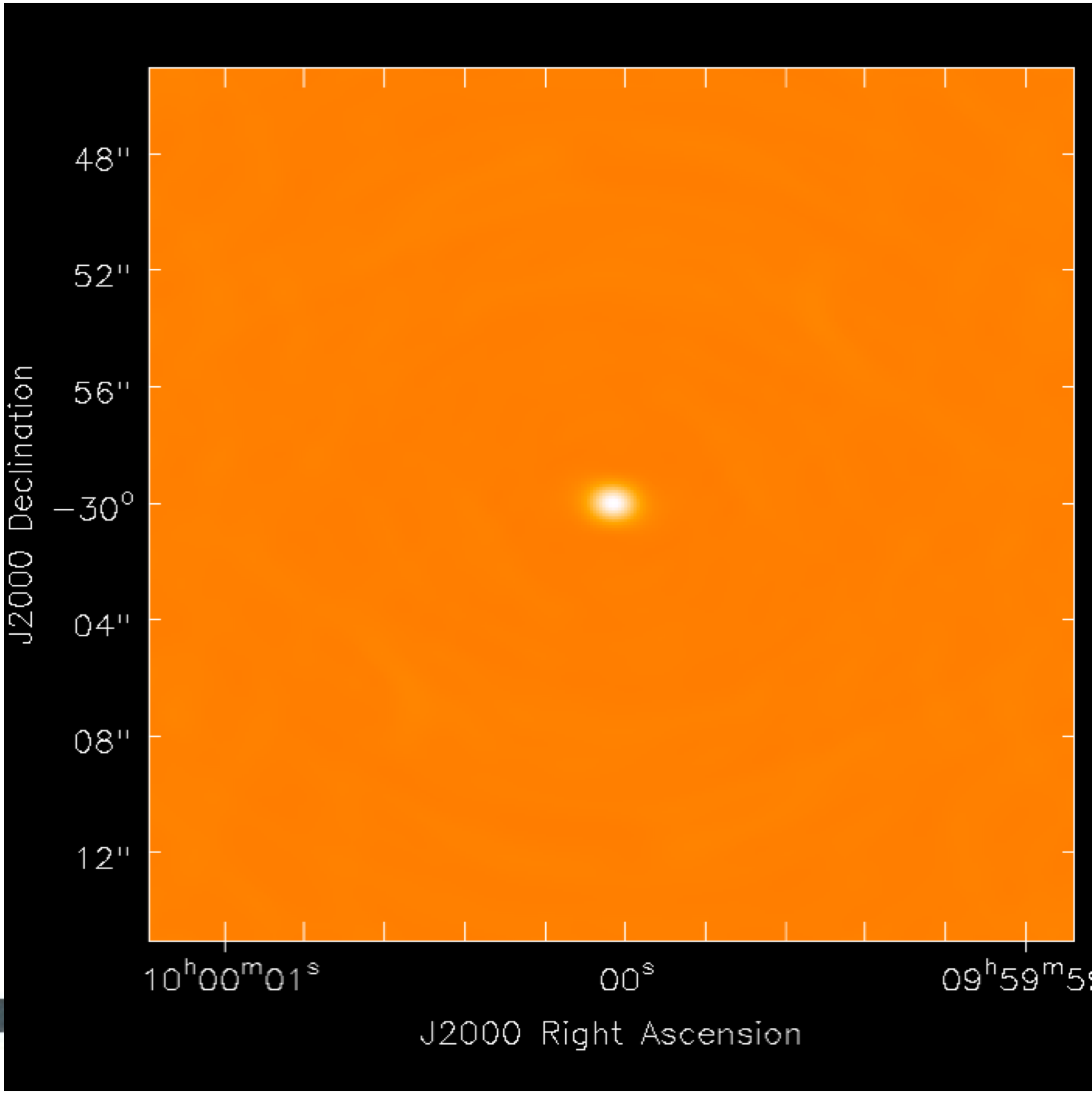




32 Antennas – 8 hours

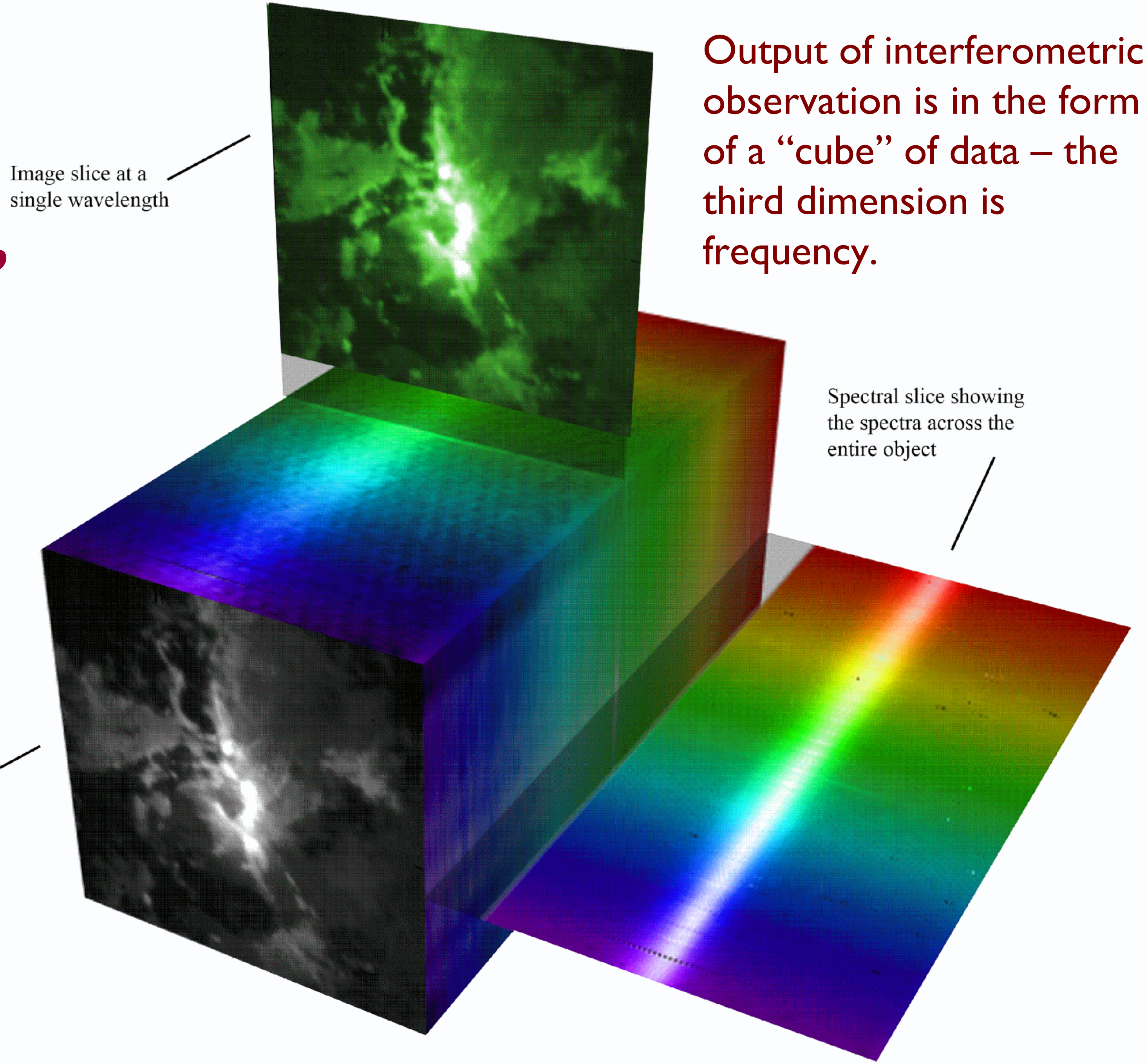






Not only 2D imaging, but 3D

Output of interferometric observation is in the form of a “cube” of data – the third dimension is frequency.



Calibration Process

Calibration is the effort to measure and remove the time-dependent and frequency-dependent atmospheric and instrumental variations.

Steps in calibrating interferometric data:
(Note: You don't have to worry about these in your observational set up!)

- Bandpass calibration (correct frequency-dependent telescope response)
- Phase and amplitude gain calibration (remove effects of atmospheric water vapor and correct time-varying phases/amplitudes)
- Set absolute flux scale

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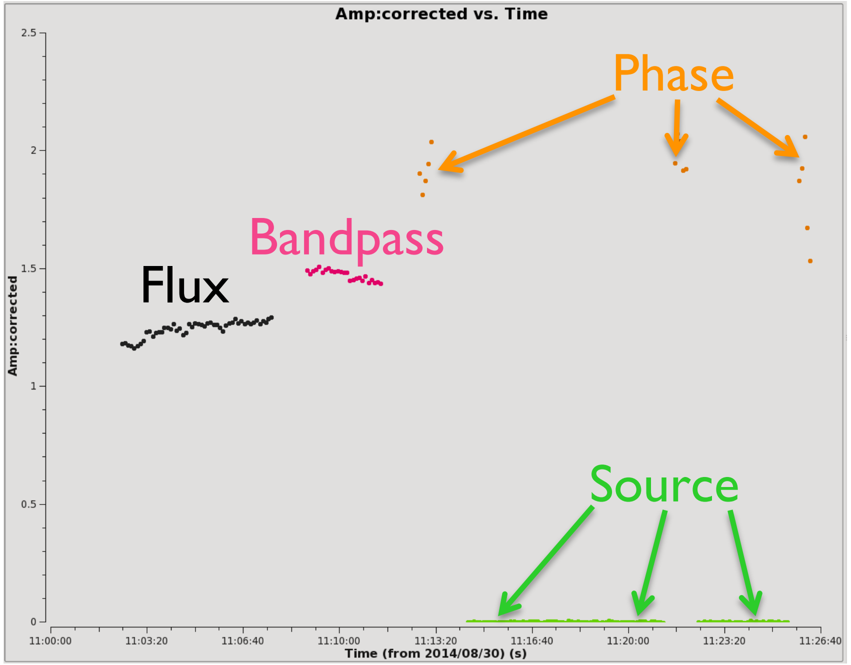
A Brief Word on Calibration

Calibration requirements (Handled by ALMA):

Phase calibrator
Bright quasar near science target
Solves for atmospheric and instrumental variations with time

Bandpass calibrator
Bright quasar
Fixes instrumental effects and variations vs frequency

Absolute flux calibrator
Solar system object or quasar
Used to scale relative amplitudes to absolute value



ALMA Data

- Data delivered after passing Quality Assurance (QA)
- Download data from *Archive Query* and *Request Handler* tools on the ALMA Science Portal
- Delivered data include:
    - Fully calibrated data ("Measurement Set")
    - Calibration tables and diagnostics
    - Preliminary images (better products may be possible with more careful continuum identification & cleaning)

See Sections 11, 12, 14, and Appendix C of ALMA Technical Handbook for details

We'll go into more detail about ALMA datasets this afternoon.

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Good Future References

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Perley, R.A., Schwab, F.R., Bridle, A.H. eds. 1989 ASP Conf. Series 6 "Synthesis Imaging in Radio Astronomy" (San Francisco: ASP)
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www.iram.fr/IRAMFR/IS/IS2008/archive.html

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