# Introduction to Radio Interferometry


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Atacama Large Millimeter/submillimeter Array
Karl G. Jansky Very Large Array
Very Long Baseline Array

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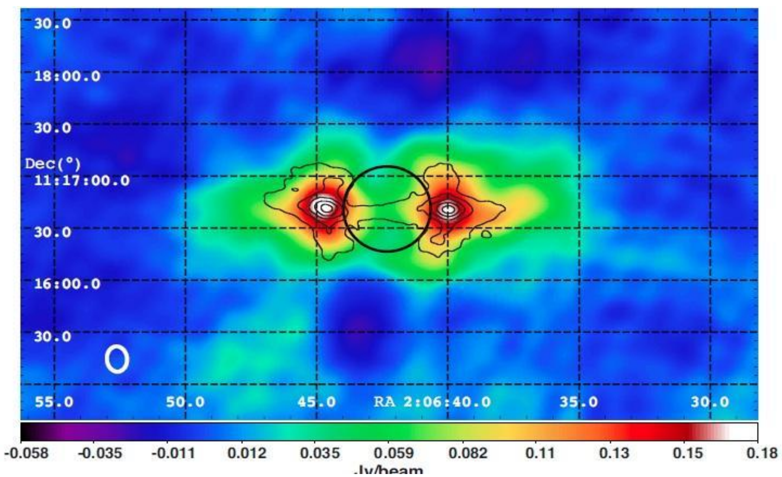
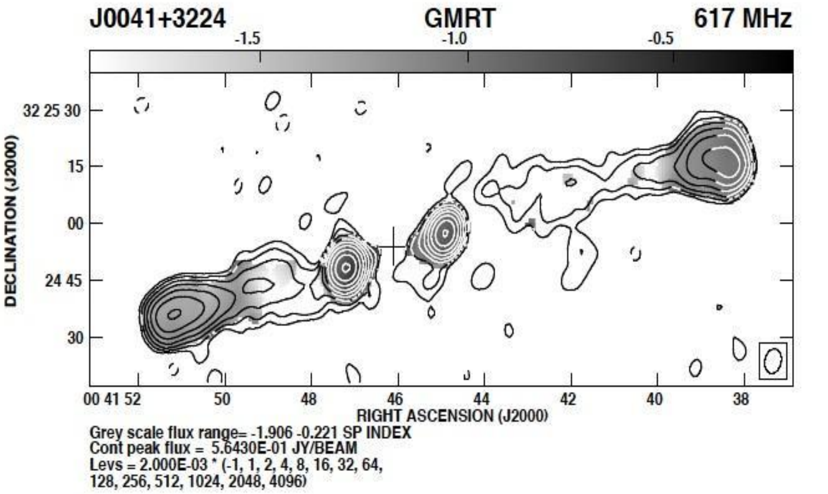
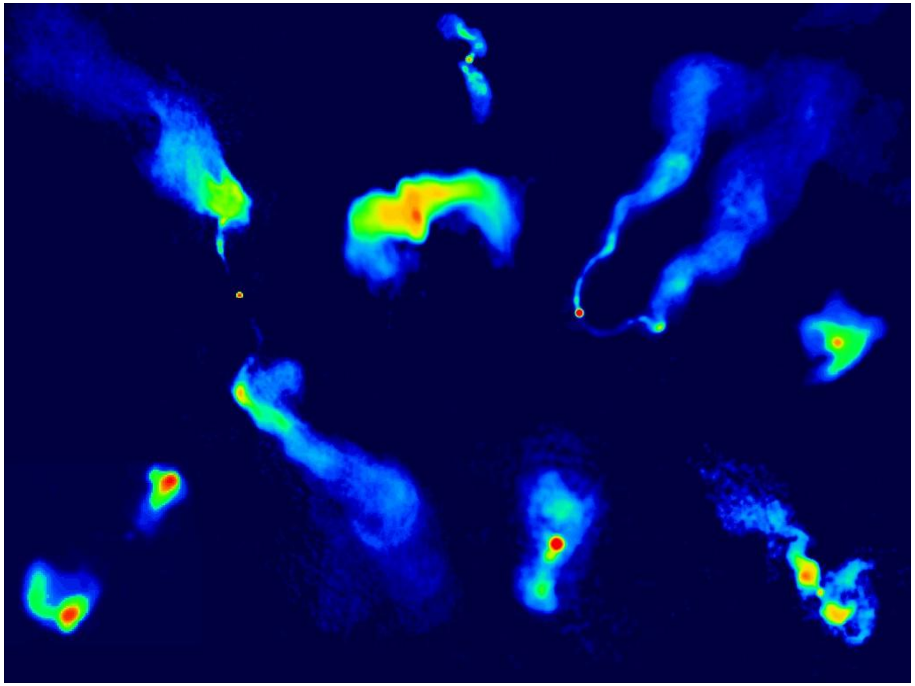
# What can we observe? (MHz-GHz range)

Jupiter's radiation belt at 100MHz




Relic emission from old radio galaxies



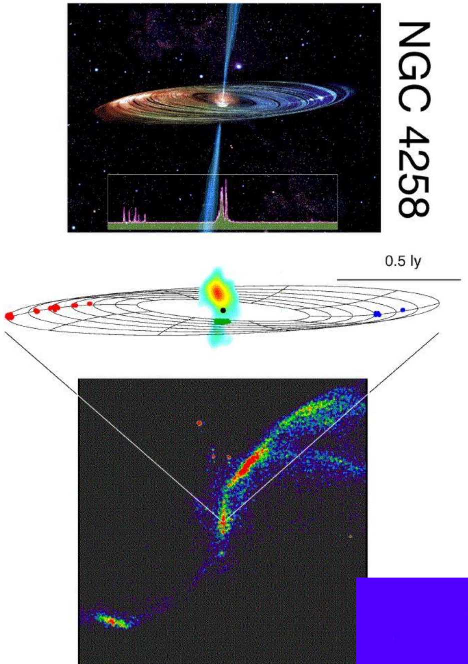
Synchrotron emission from extended radio galaxies (5 GHz)

Images from NRAO Image Gallery: http://images.nrao.edu/

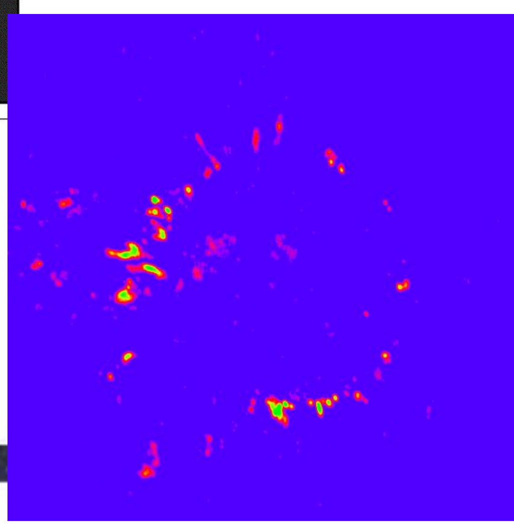
# What can we observe?

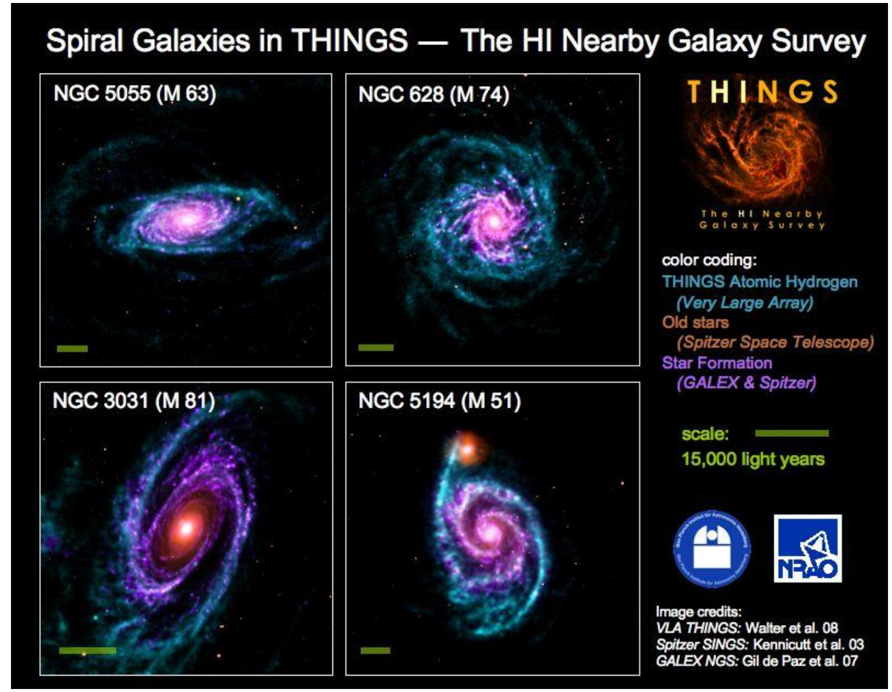
Images from NRAO Image Gallery
http://images.nrao.edu/

## At low frequencies (MHz-GHz):

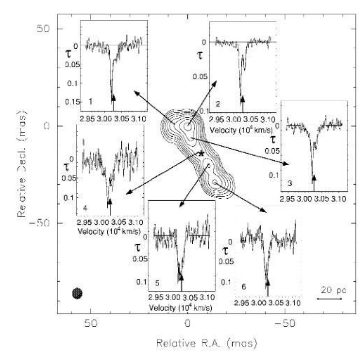


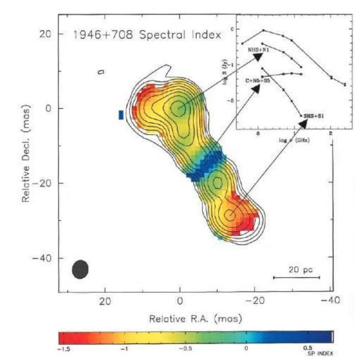

$H_2O$, OH or SiO masers in galaxies and stars





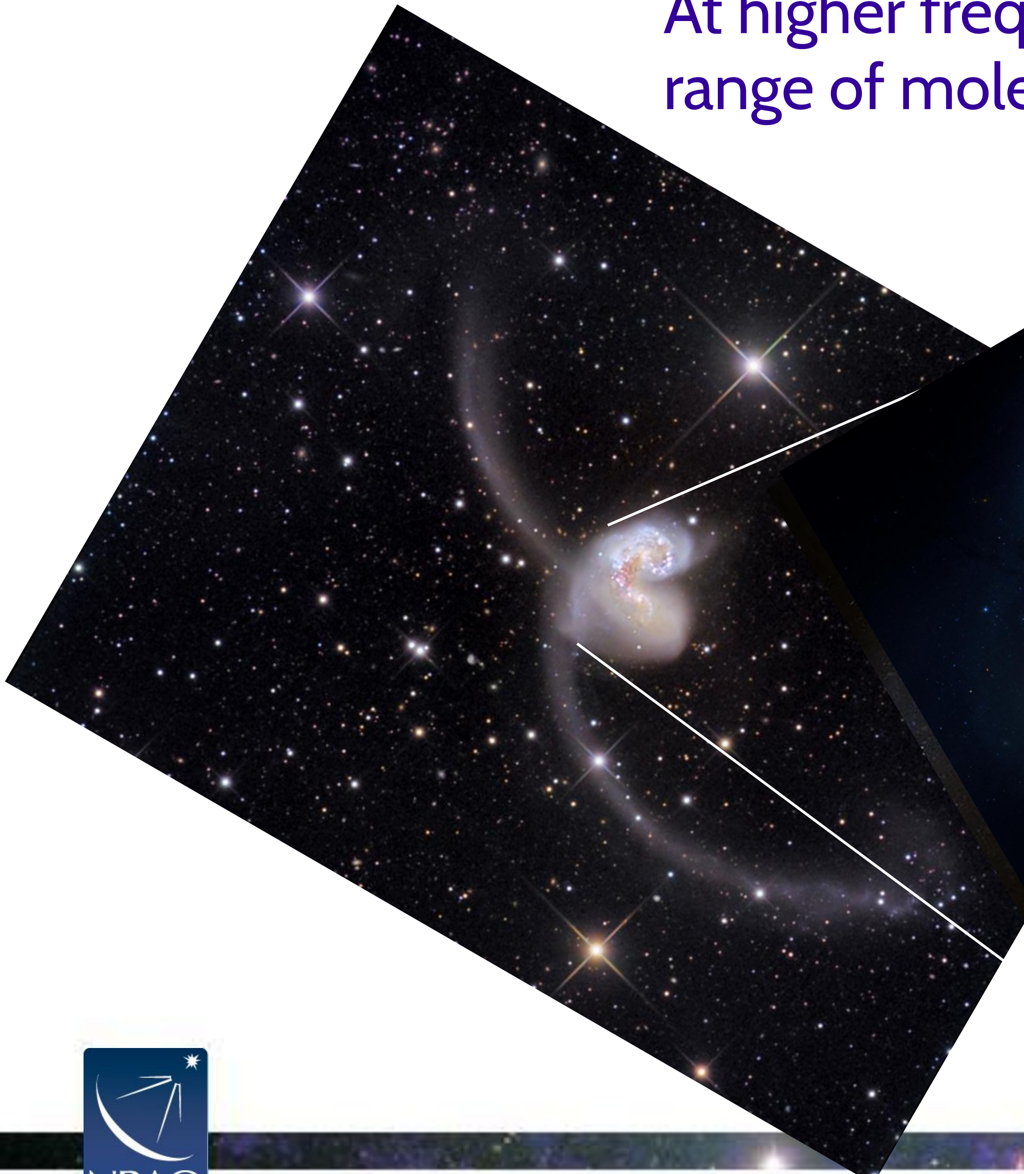

HI emission and absorption, free-free absorption in galaxies

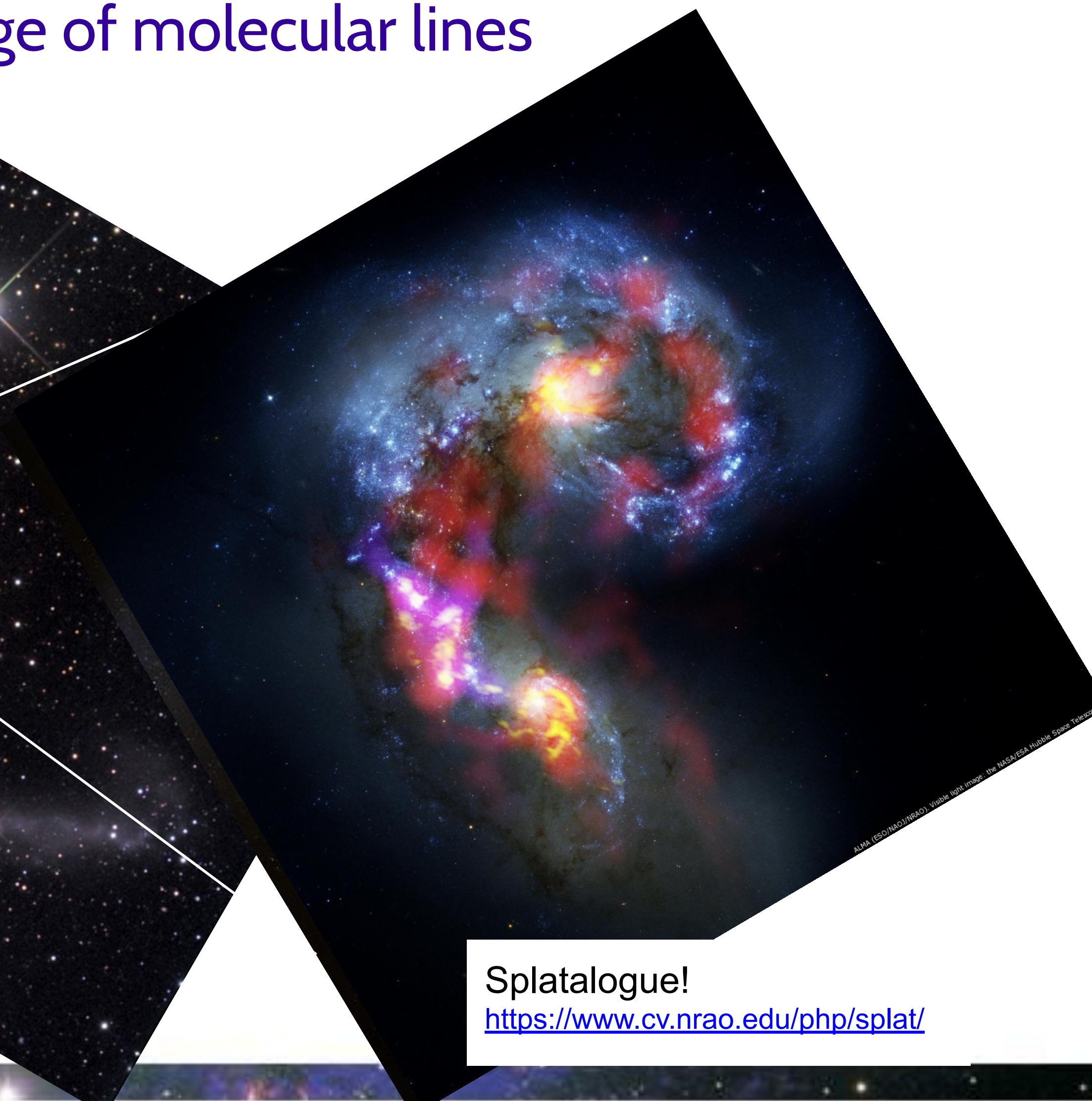


# What can we observe?

At higher frequencies we can observe a broad range of molecular lines



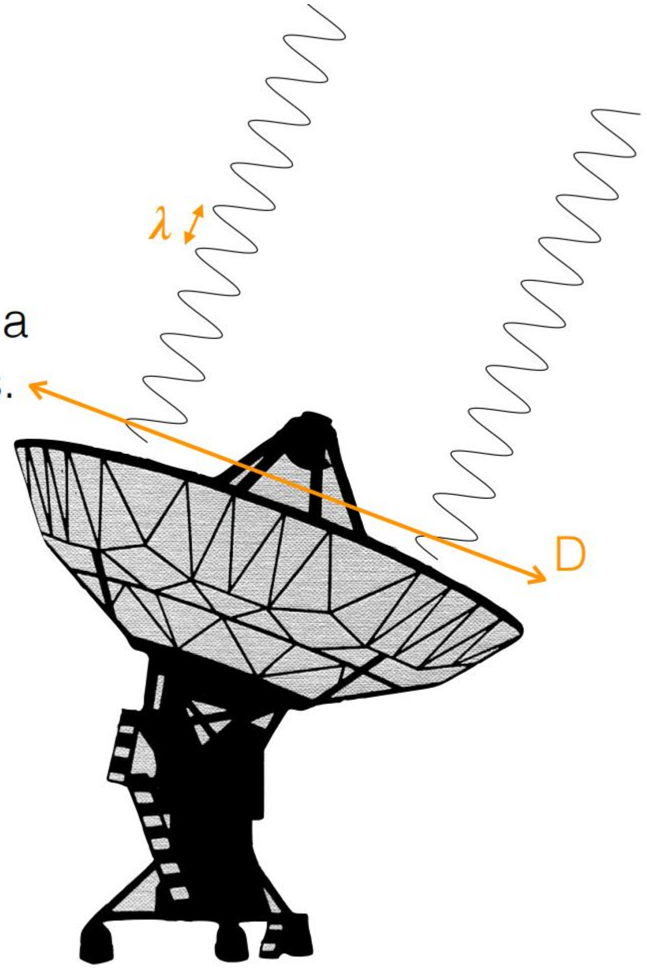



Splatalogue!
https://www.cv.nrao.edu/php/splat/

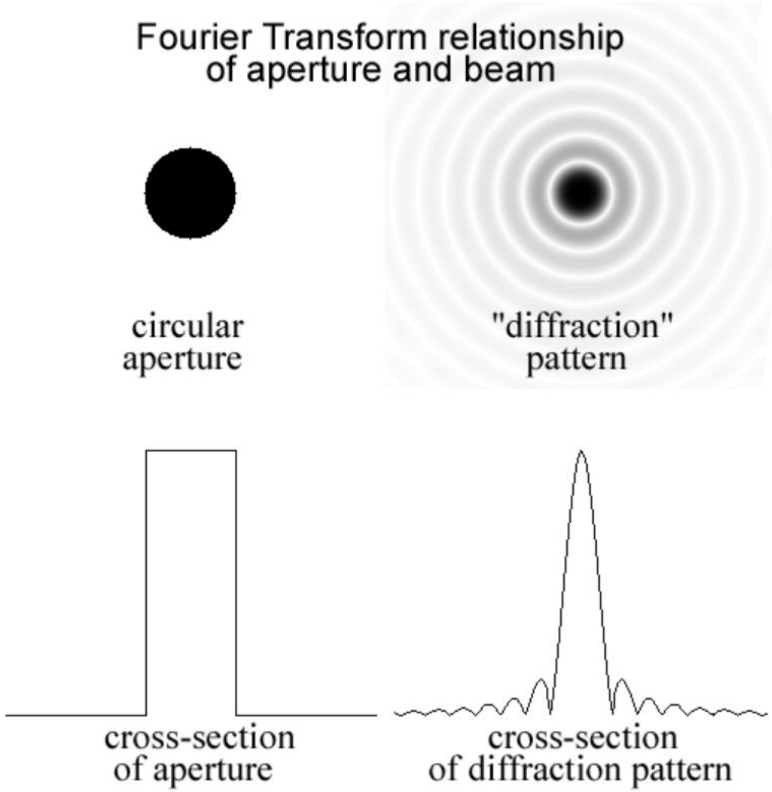
# Long wavelength means no glass mirrors

# Single Dish

Diffraction theory: this telescope (by itself) has a resolution ~ $\lambda$/D radians.



Associated with the collecting area (effective area) is the *beam pattern*, also called the *primary beam*, which is just the *Fourier Transform* of the aperture, as shown in the figure below.



# Resolution of Observations

**Angular resolution for most telescopes is ~ λ/D**

D is the diameter of the telescope and λ is the wavelength of observation

**For the Hubble Space Telescope:**

*λ ~ 1um / D of 2.4m = resolution ~ 0.13"*

**To reach that resolution at λ ~1mm, we would need a 2 km-diameter dish!**

Instead, we use arrays of smaller dishes to achieve the same high angular resolution at radio frequencies

**This is interferometry!**

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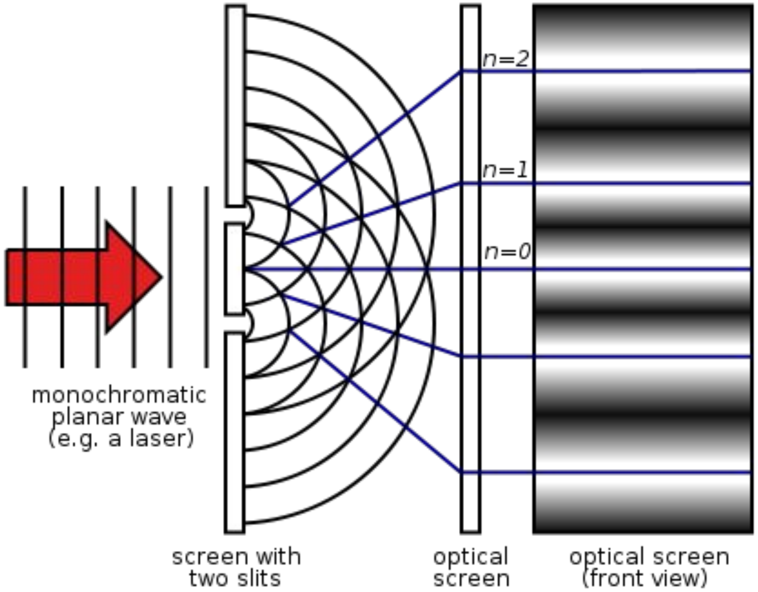
# What is an interferometer?

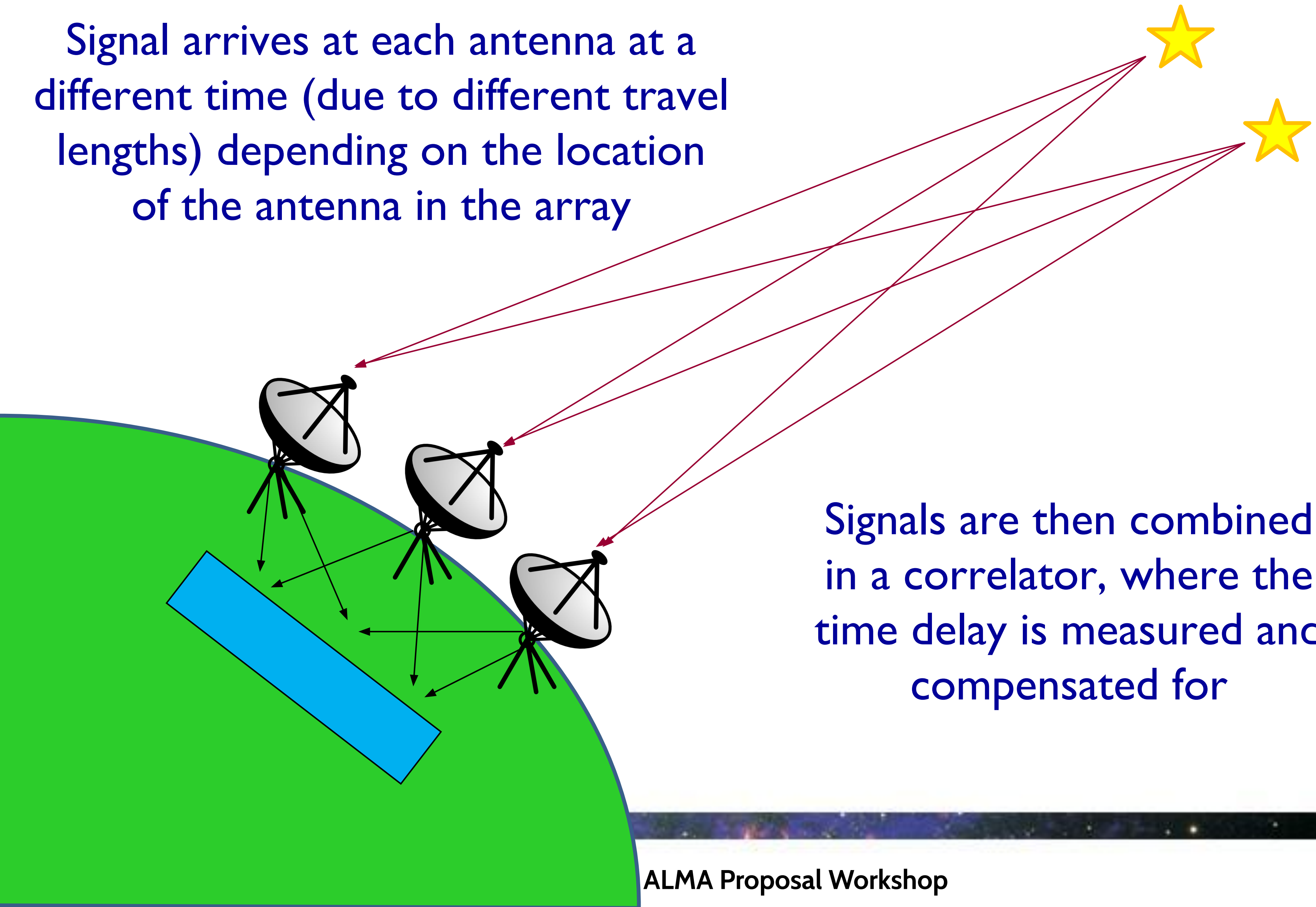
An *interferometer* measures the interference pattern produced by multiple apertures, much like a 2-slit experiment




******However, the interference patterns measured by radio telescopes are produced by* ***multiplying*** *- not adding - the wave signals measured at the different telescopes (i.e. apertures)*

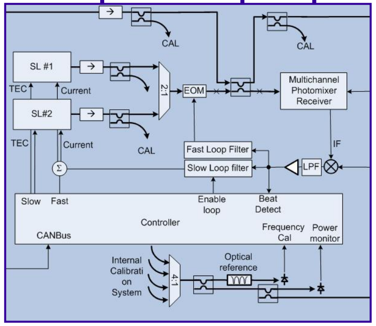
# How Do We Use Interferometry?

Signal arrives at each antenna at a different time (due to different travel lengths) depending on the location of the antenna in the array



Signals are then combined in a correlator, where the time delay is measured and compensated for

# Some Instrument Details




## To precisely measure arrival times we need very accurate clocks

At Band 10 one wavelength error = 1 picosecond = $10^{-12}$ s (!!)

Need << 1 wavelength timing precision, so each antenna has an on-board clock with high sampling rates

Once determined, the reference time is distributed to all antennas

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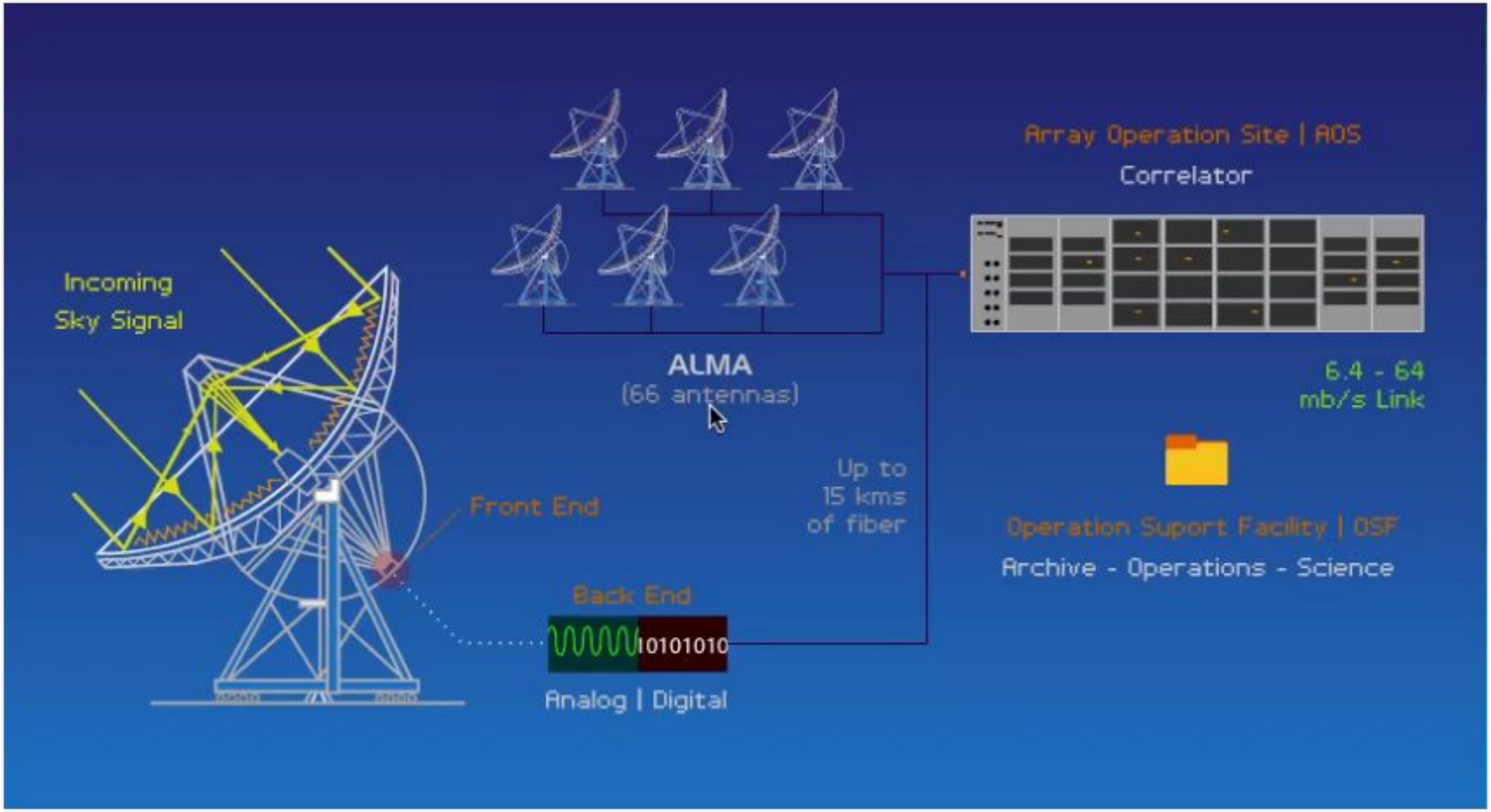
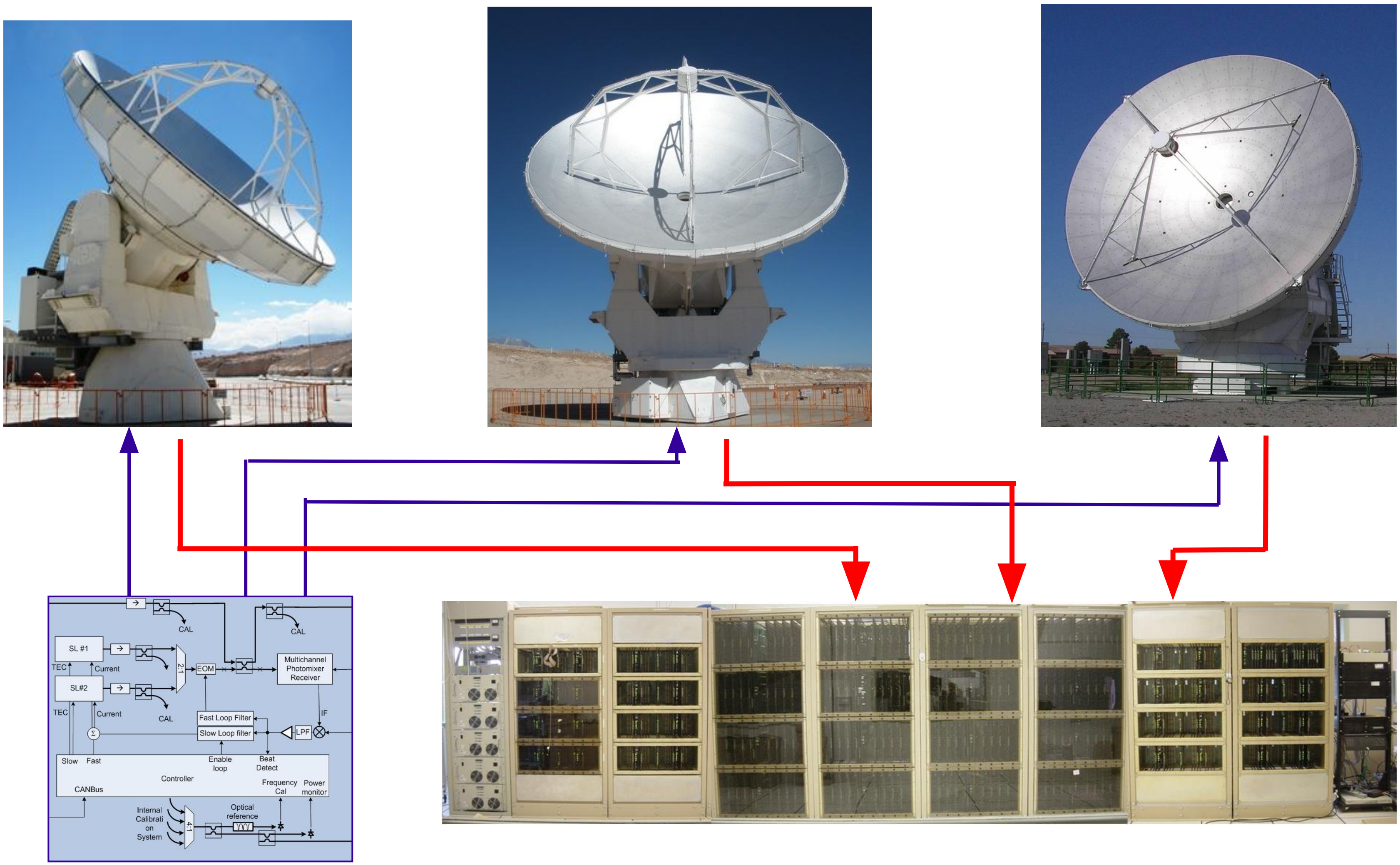
# An Interferometer In Action



# Some Instrument Details

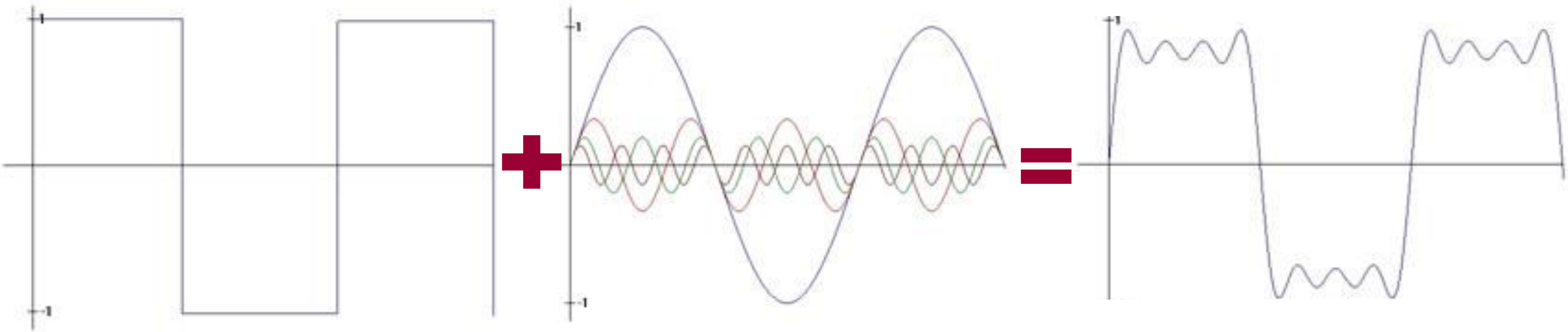



Signals from each antenna are digitized and sent to the correlator for multiplication & averaging

For ~50 antennas, the data rate is 600 GB/sec for the correlator to process

# Introducing the Fourier Transform

**Fourier theory states that any well behaved signal (including images) can be expressed as the sum of sinusoids**



Reference signal

4 sinusoids

Sum of sinusoids & signal

The Fourier transform is the mathematical tool that decomposes a signal into its sinusoidal components

The Fourier transform contains *all* of the information of the original signal

- Fourier Transform Wikipedia Page: https://en.wikipedia.org/wiki/Fourier_transform
- Here is a nifty simple interactive demonstration of Fourier Transforms. http://www.jezzamon.com/fourier/

# Visibility and Sky Brightness

*The van Cittert-Zernike theorem*

- **Visibility as a function of baseline coordinates (u,v) is the Fourier transform of the sky brightness distribution as a function of the sky coordinates (x,y)**

$$V(u,v) \xrightarrow{\text{FT}} T(x,y)$$

V(u,v) = the complex visibility function $= \iint T(x,y)e^{2\pi i(ux+vy)}dxdy$

T(x,y) = the sky brightness distribution $= \iint V(u,v)e^{-2\pi i(ux+vy)}dudv$

# The Fourier Transform relates the measured interference pattern to the radio intensity on the sky

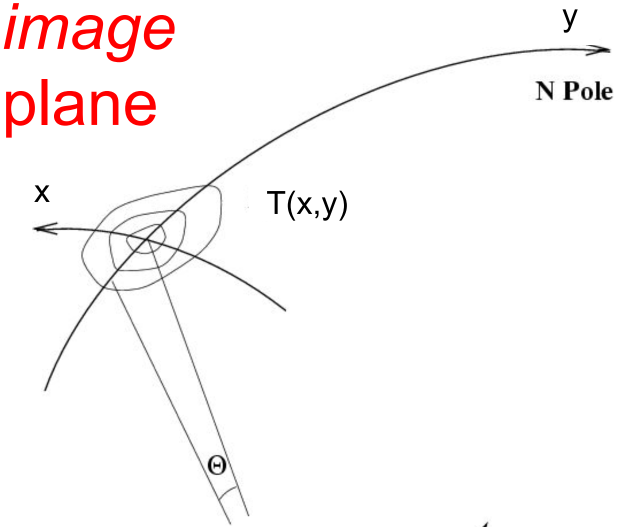
Fourier

$$V(u,v) = \int\int T(x,y)e^{2\pi i(ux+vy)}dxdy$$

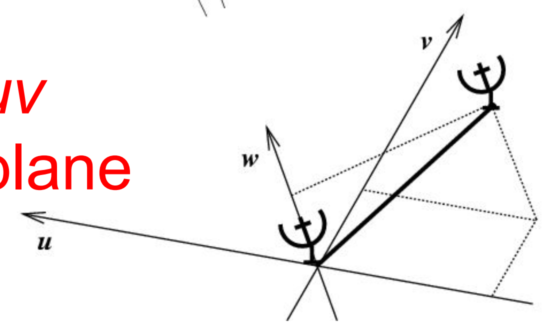
Image

$$T(x,y) = \int\int V(u,v)e^{-2\pi i(ux+vy)}dudv$$

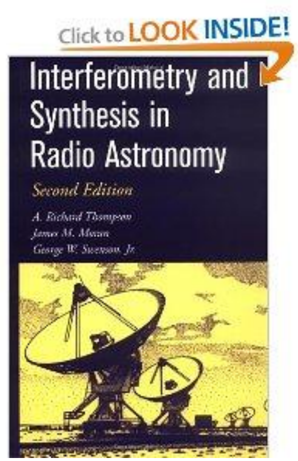
*image* plane




*uv* plane

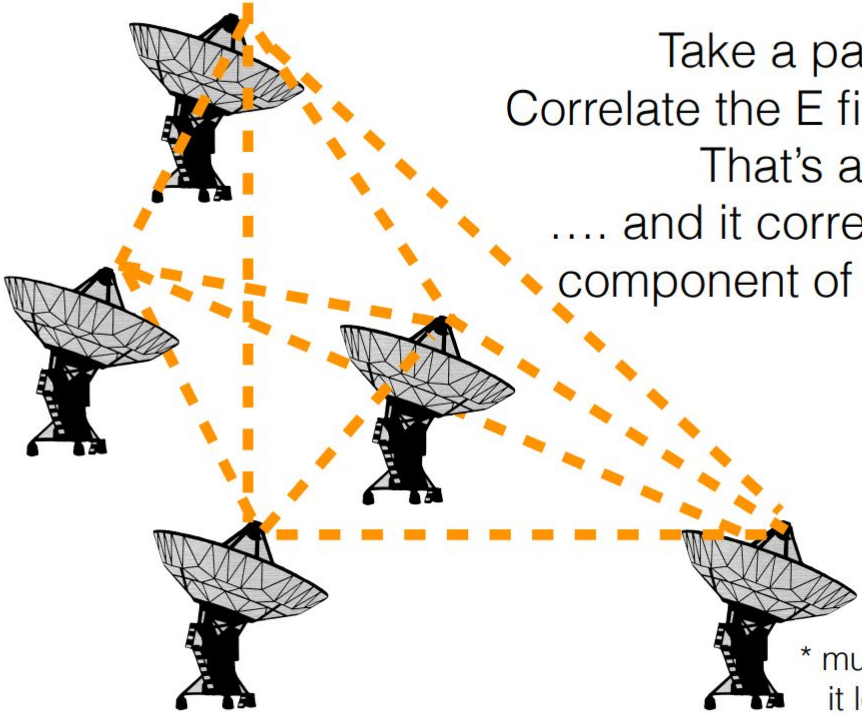



(for more info, see e.g. Thompson, Moran & Swenson)



## Interferometer theory, very loosely.

correlation* of E field at Earth = FT of brightness distribution of the sky.



Take a pair of antennas.
Correlate the E fields at each antenna.
That's a "visibility"…
…. and it corresponds to a Fourier component of the sky brightness.

* mutual coherence function, but it looks a lot like a correlation

# What Are Visibilities?

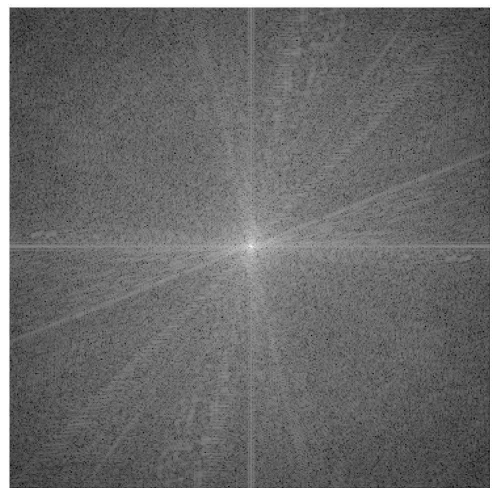
Each V(u,v) contains information on T(x,y) everywhere

Each V(u,v) is a complex quantity

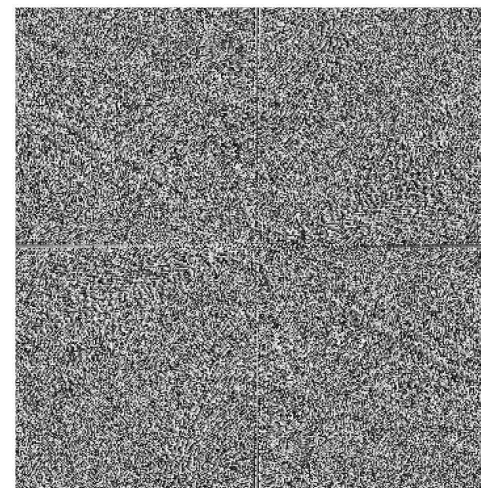
Expressed as (real, imaginary) or (amplitude, phase)

T(x,y)

FT

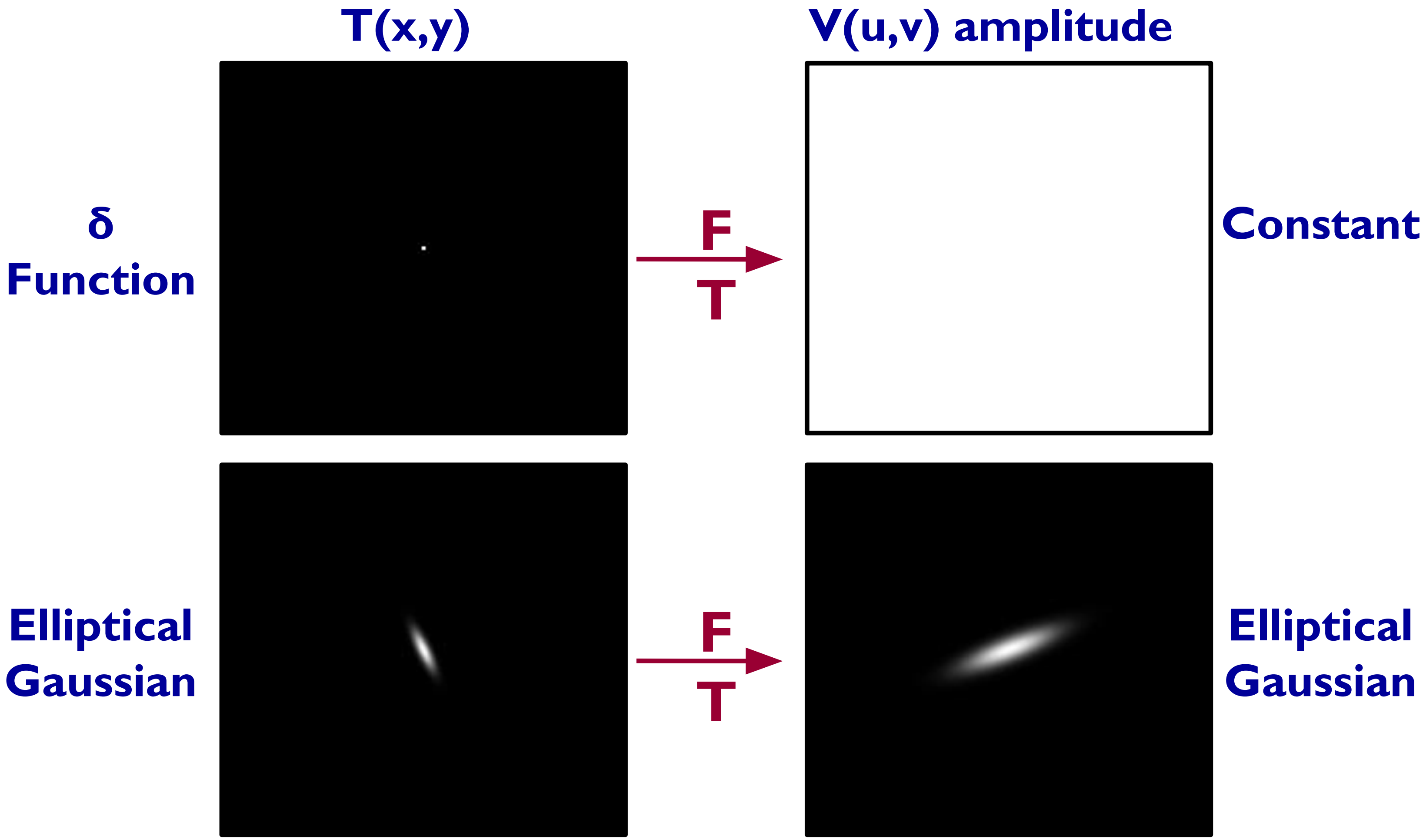


V(u,v) amplitude



V(u,v) phase

# Examples of 2D Fourier Transforms

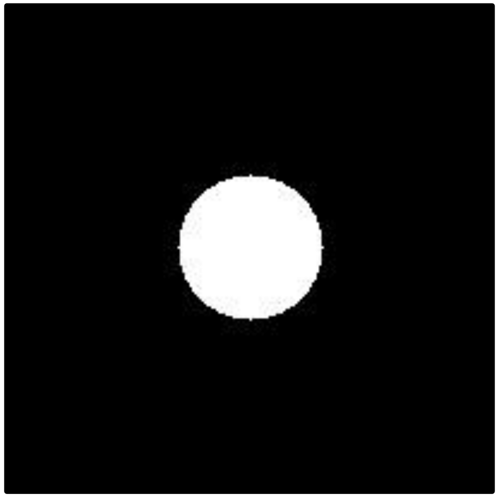



**Rules of the Fourier Transform:**
Narrow features transform to wide features (and vice versa)

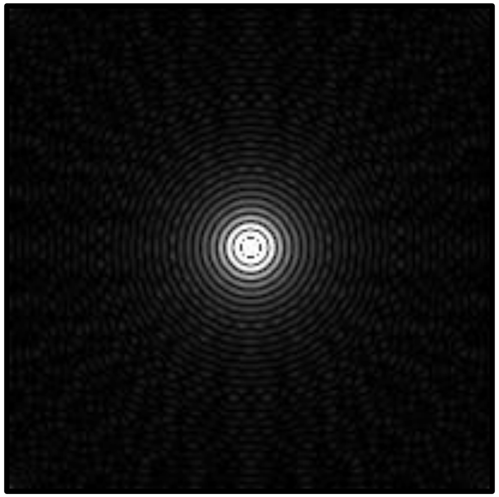
# Examples of 2D Fourier Transforms

**T(x,y)** | **V(u,v) amplitude**

**Uniform Disk**

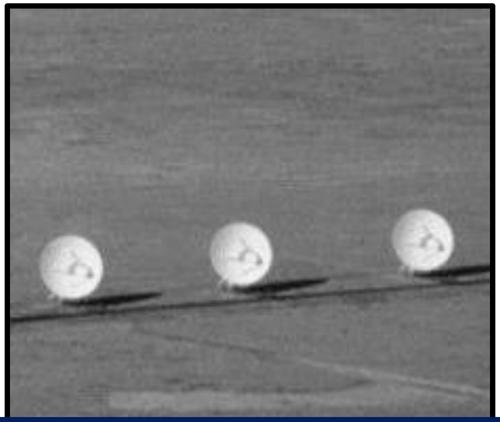


**FT** →



**Bessel Function**

**VLA Antennas**



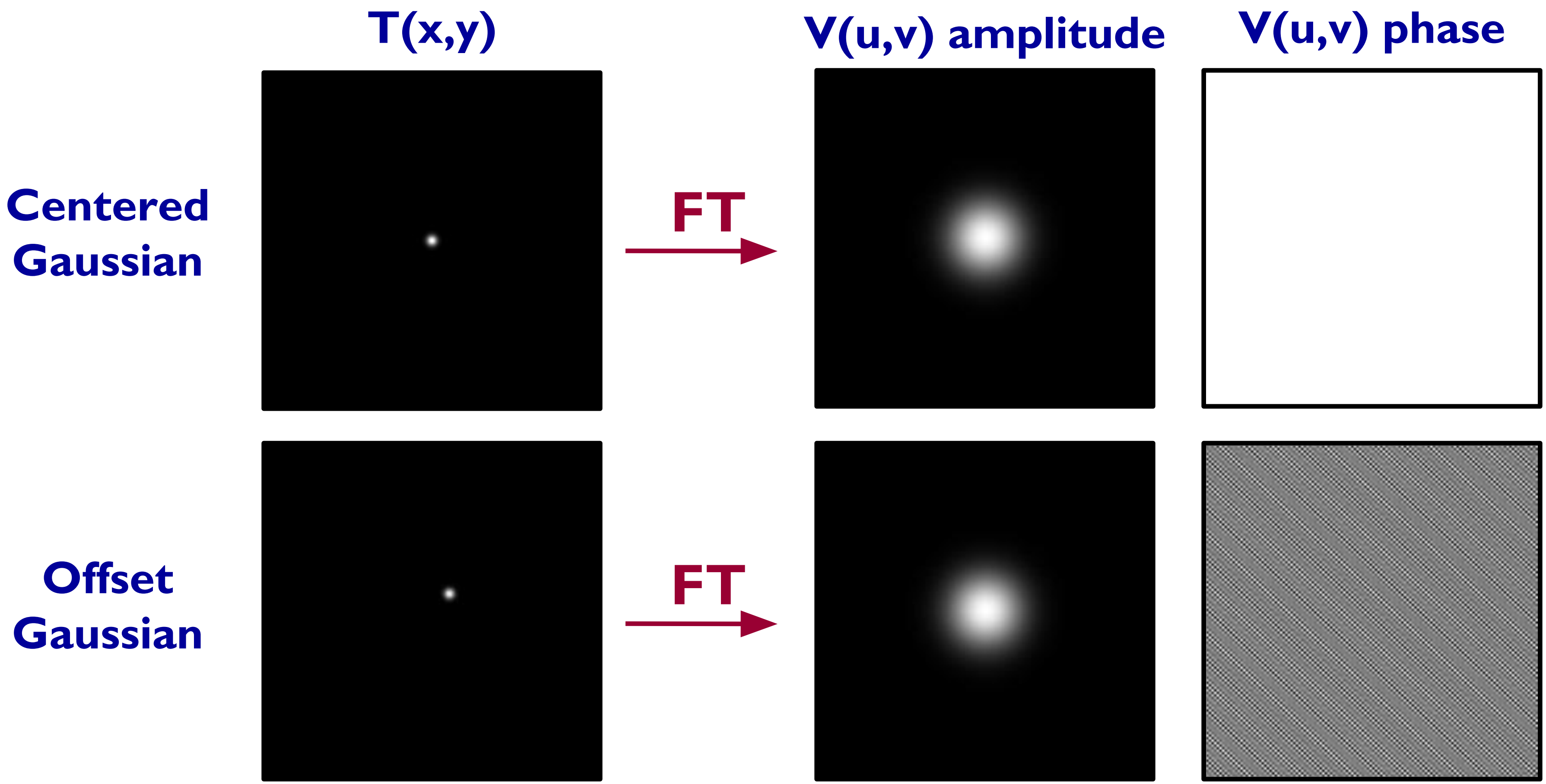
**FT** →

**Bessel Function!**

**Rules of the Fourier Transform:**

Sharp features (edges) result in many high spatial features

# Examples of 2D Fourier Transforms




**Rules of the Fourier Transform:**
Amplitude tells you 'how much' of a spatial frequency
Phase tells you 'where' the spatial frequency is

# Some Comments on Visibilities

- The Visibility is a unique function of the source brightness.
- The two functions are related through a Fourier transform. $V_{\upsilon}(u,v) \Leftrightarrow I(l,m)$
- An interferometer, at any one time, makes one measure of the visibility, at baseline coordinate (u,v).
- `Sufficient knowledge' of the visibility function (as derived from an interferometer) will provide us a `reasonable estimate' of the source brightness.
- How many is 'sufficient', and how good is 'reasonable'?
- These simple questions do not have easy answers…

# Basics of Aperture Synthesis

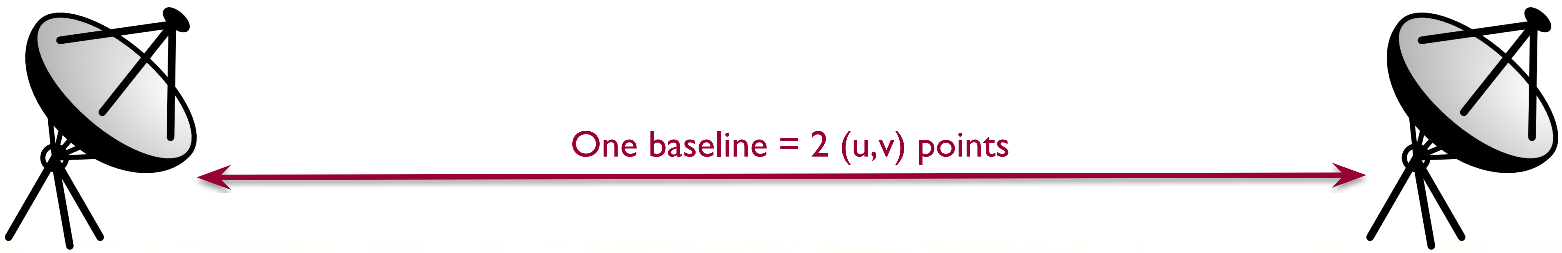
**Idea:** Sample V(u,v) at an enough (u,v) points using distributed small aperture antennas to synthesize a large aperture antenna of size ($u_{max}$,$v_{max}$)

One pair of antennas = one baseline

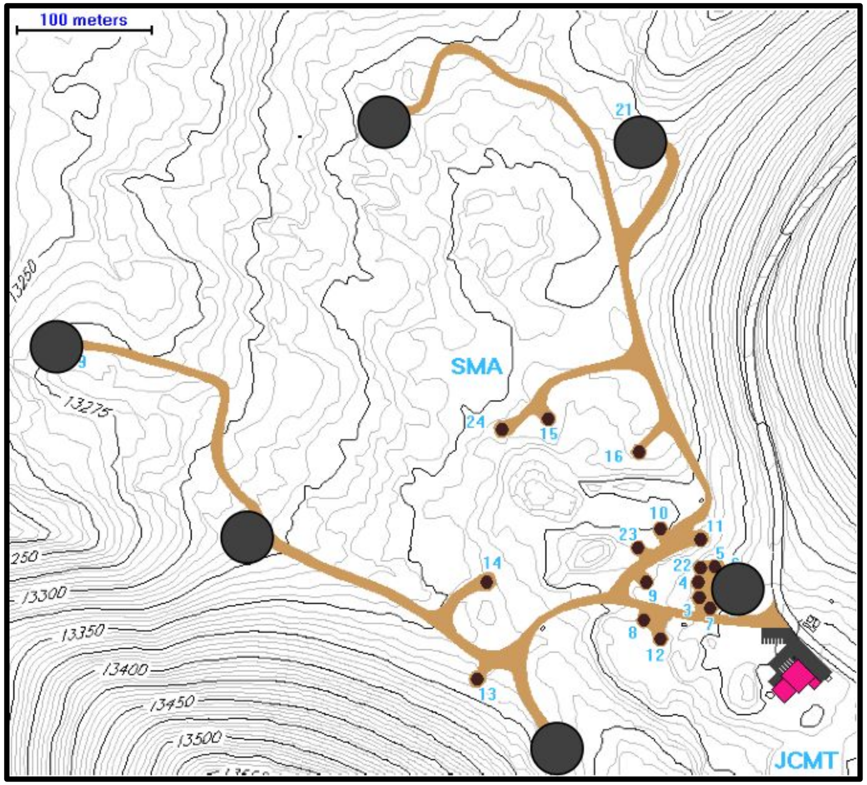
For **N antennas**, we get **N(N-1) samples** at a time

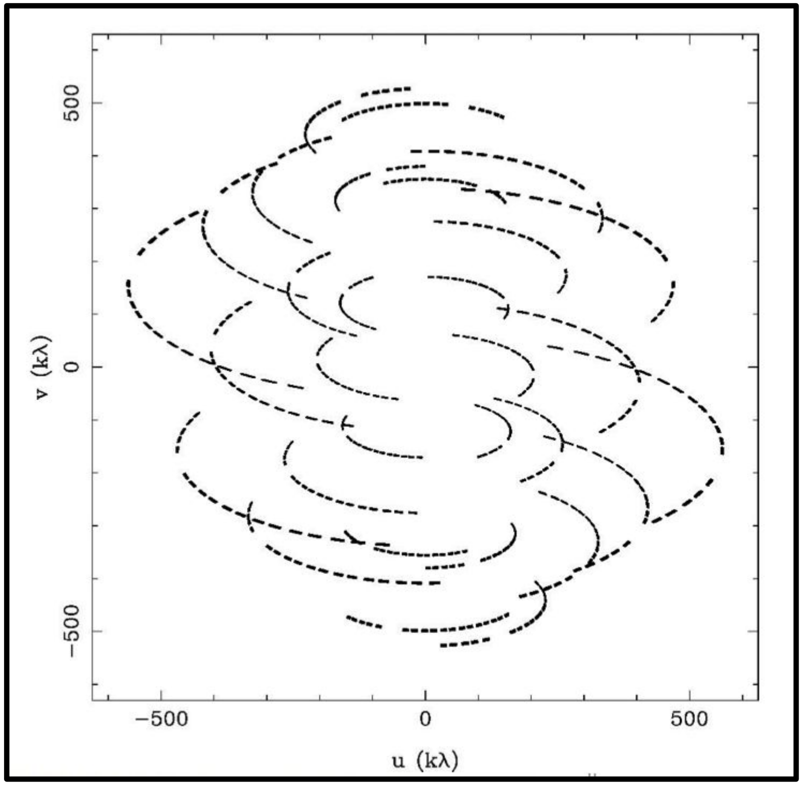
**How do we fill in the rest of the (u,v) plane?**

1. Earth's rotation
2. Reconfigure physical layout of N antennas



# (u,v) Plane Sampling



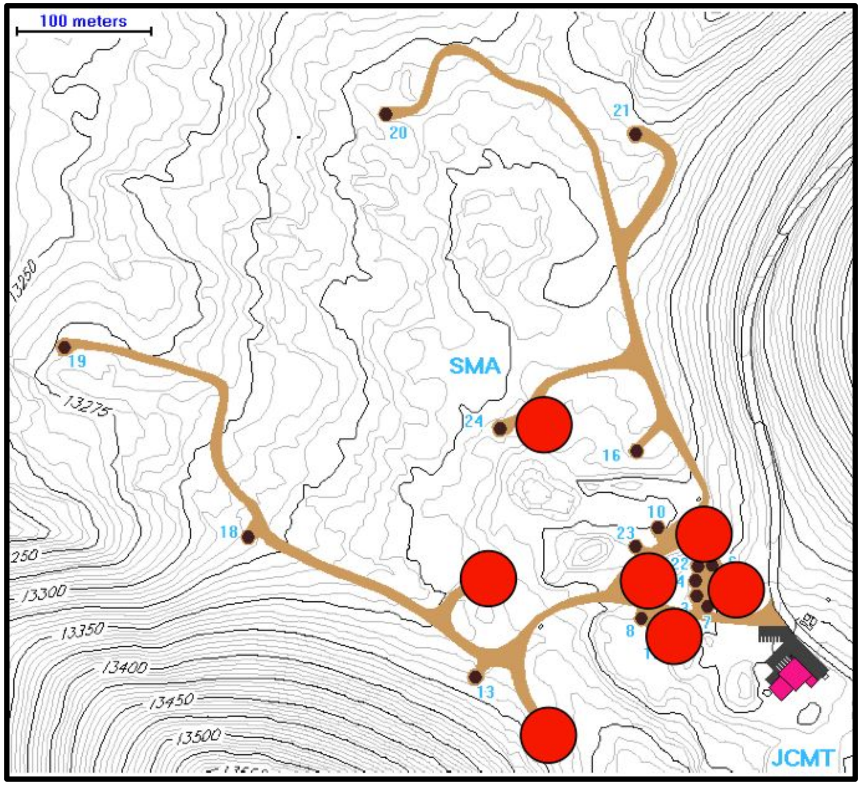



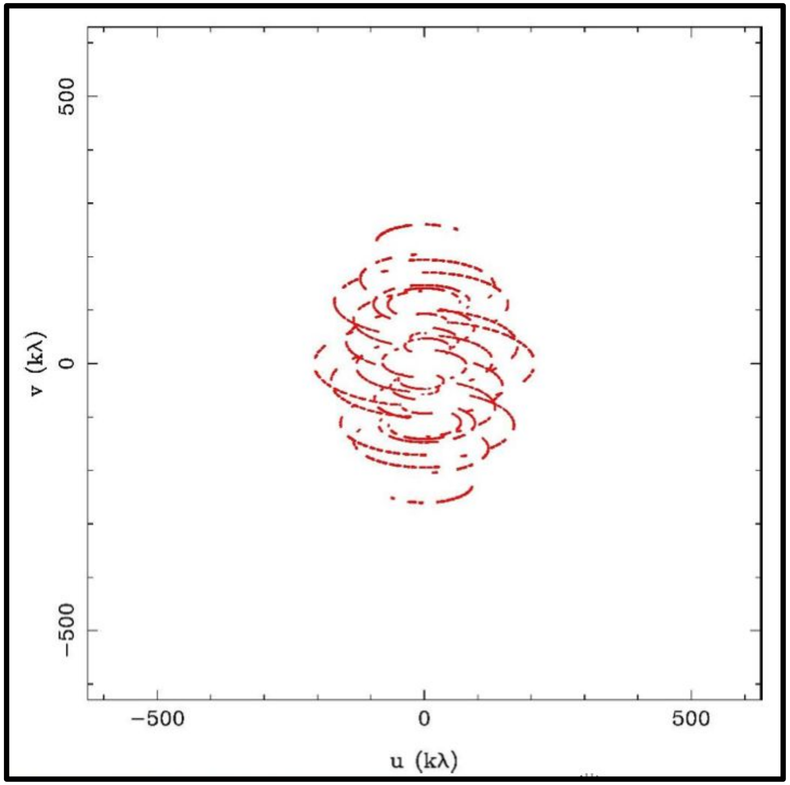

**Very Extended SMA configuration**

(most extended baselines)

345 GHz, DEC = +22

# (u,v) Plane Sampling



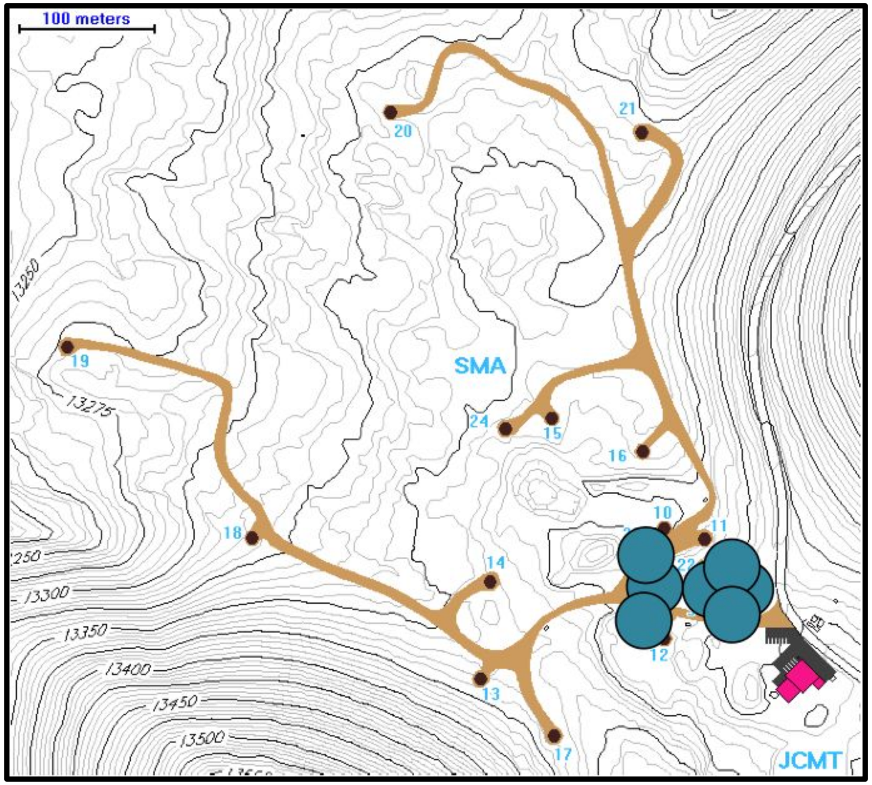



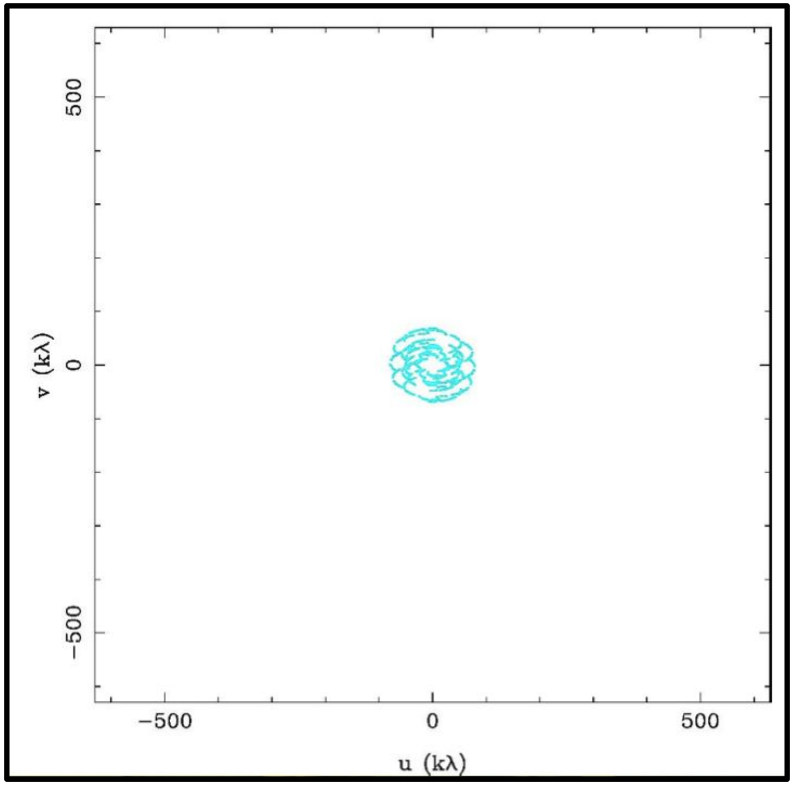

**Extended SMA configuration**

(extended baselines)

345 GHz, DEC = +22

# (u,v) Plane Sampling



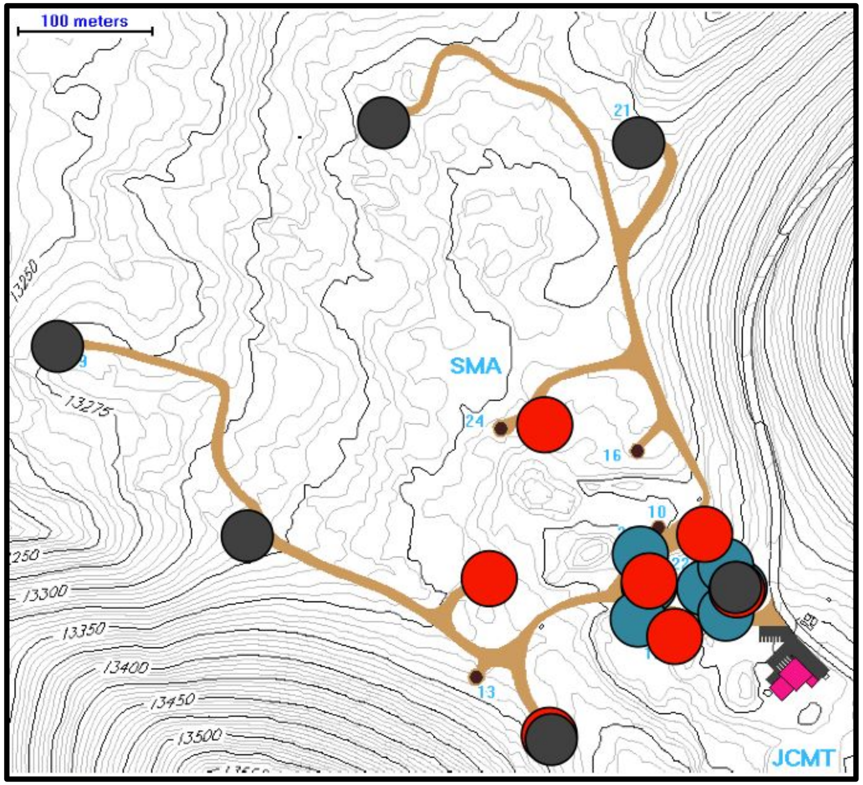



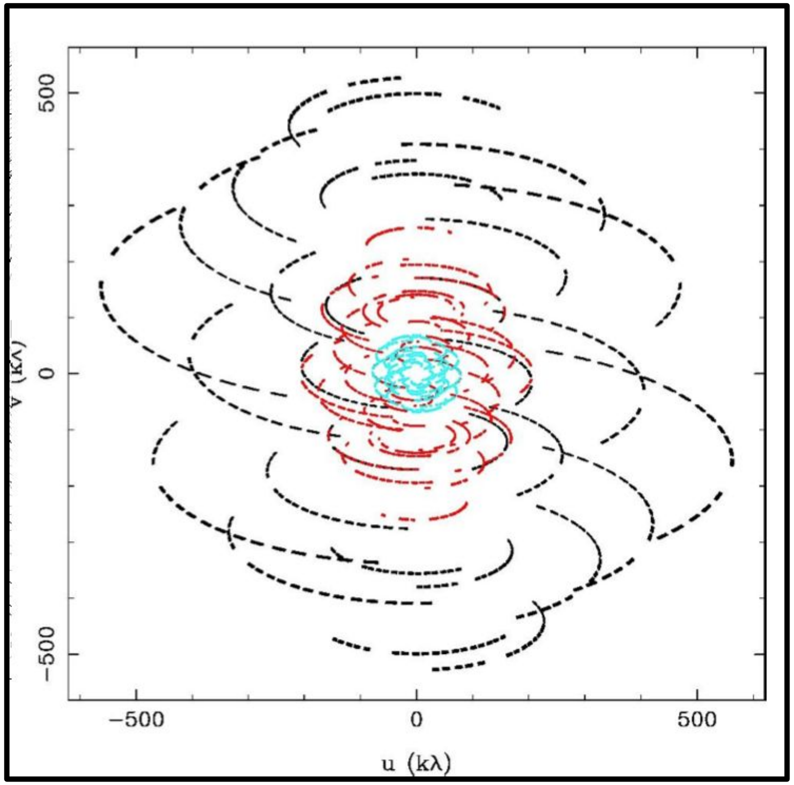

**Compact SMA configuration**

(compact baselines)

345 GHz, DEC = +22

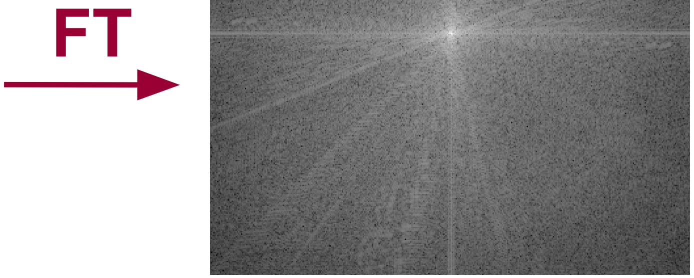
# (u,v) Plane Sampling



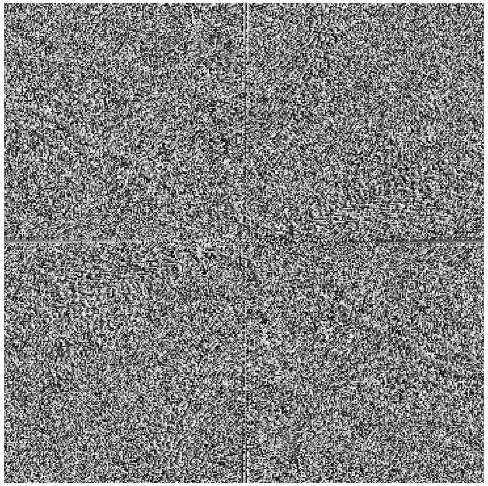




**Combine multiple configurations to get the most complete coverage of the (u,v) plane**

T(x,y)



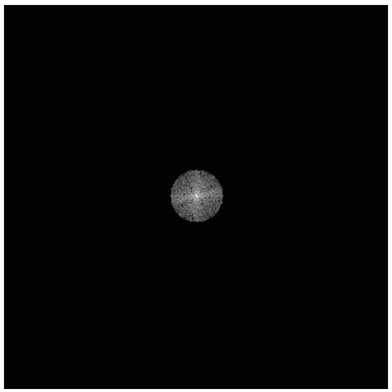
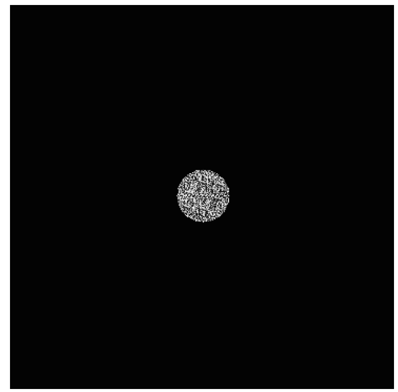
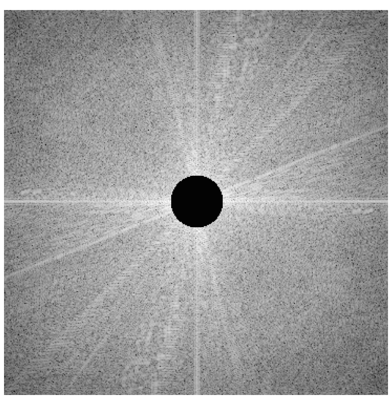
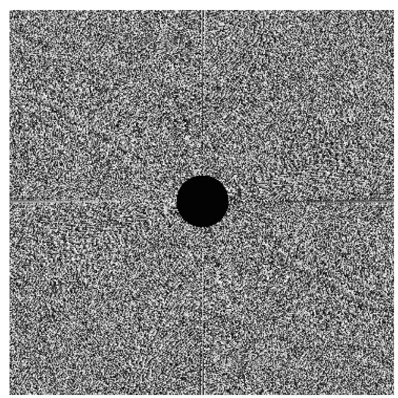

V(u,v) amplitude



V(u,v) phase

# Implications of (u,v) Coverage

## What does it mean if our (u,v) coverage is not complete?

| | V(u,v) amplitude | V(u,v) phase | | T(x,y) |
|---|---|---|---|---|
| Missing High Spatial Frequencies |  |  | FT → |  |
| Missing Low Spatial Frequencies |  |  | FT → |  |

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# Characteristic Angular Scales

**Angular resolution of telescope array:**

$\sim \lambda/B_{max}$ ($B_{max}$ = longest baseline)

**Maximum angular scale:**

$\sim \lambda/B_{min}$ ($B_{min}$ = shortest distance between antennas)

**Field of view (FOV):**

$\sim \lambda/D$ (D = antenna diameter)

*Sources more extended than the FOV can be observed using multiple pointing centers in a mosaic

**An interferometer is sensitive to a range of angular sizes: $\lambda/B_{max} < \theta < \lambda/B_{min}$**

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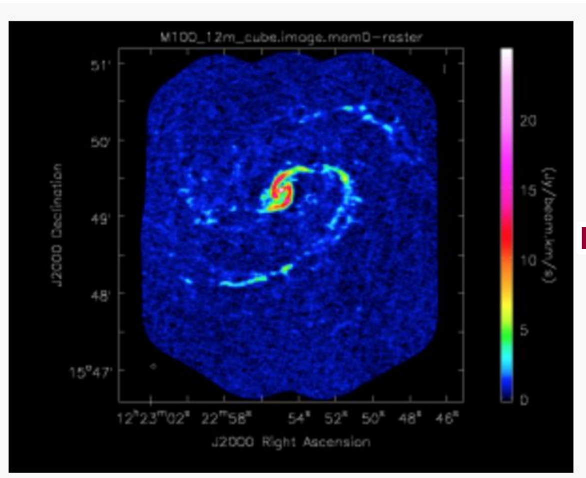
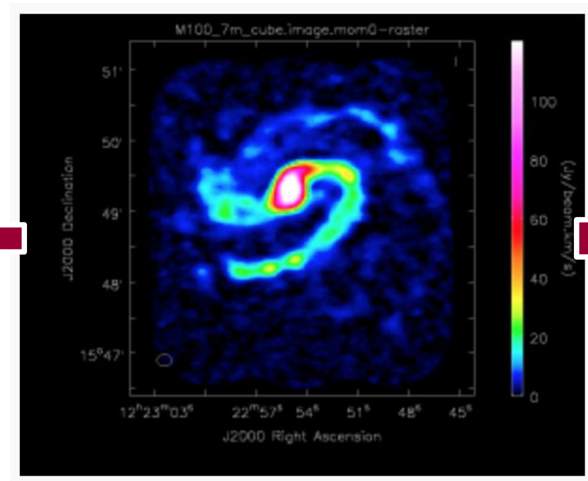
# Characteristic Angular Scales: M100

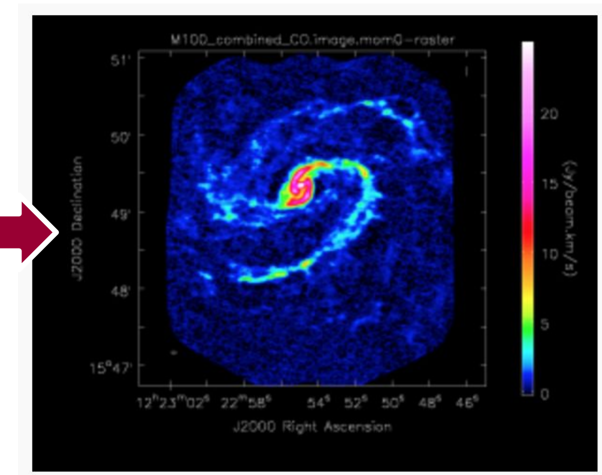
ALMA 12m




ACA 7m




Combined




ALMA 12m shows smaller spatial scales (denser, clumpier emission)
ACA 7m data shows larger spatial scales (diffuse, extended emission)

**To get both — you need a combined image!**

# Interferometry: Spatial Scales

- The **sensitivity** is given by the number of antennas times their area.
- The **field of view** is given by the beam of a single antenna (corresponding to the resolution for a single dish telescope or the primary beam).
- The **resolution** is given by the largest distance between antennas (called the synthesized beam).
- The **largest achievable angular scale** that can be imaged is given by the shortest distance between antennas.

# Angular Scales — A Proposal Tip!

**Interferometers act as spatial filters - shorter baselines are sensitive to larger targets, so remember …**

Spatial scales larger than the smallest baseline cannot be imaged.

Spatial scales smaller than the largest baseline cannot be resolved.

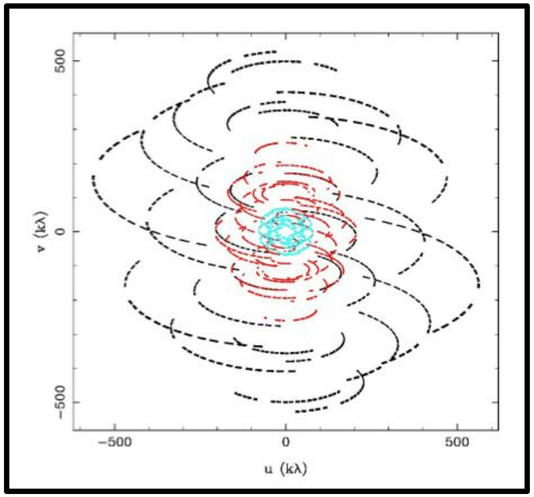
Table A-1: Angular Resolutions (AR) and Maximum Recoverable Scales (MRS) for the Cycle 8 configurations

| Config | Lmax | | Band 3 | Band 4 | Band 5 | Band 6 | Band 7 | Band 8 | Band 9 | Band 10 |
|---|---|---|---|---|---|---|---|---|---|---|
| | Lmin | | 100 GHz | 150 GHz | 185 GHz | 230 GHz | 345 GHz | 460 GHz | 650 GHz | 870 GHz |
| 7-m | 45 m | AR | 12.5" | 8.4" | 6.8" | 5.5" | 3.6" | 2.7" | 1.9" | 1.4" |
| | 9 m | MRS | 66.7" | 44.5" | 36.1" | 29.0" | 19.3" | 14.5" | 10.3" | 7.7" |
| C-1 | 161 m | AR | 3.4" | 2.3" | 1.8" | 1.5" | 1.0" | 0.74" | 0.52" | 0.39" |
| | 15 m | MRS | 28.5" | 19.0" | 15.4" | 12.4" | 8.3" | 6.2" | 4.4" | 3.3" |
| C-2 | 314 m | AR | 2.3" | 1.5" | 1.2" | 1.0" | 0.67" | 0.50" | 0.35" | 0.26" |
| | 15 m | MRS | 22.6" | 15.0" | 12.2" | 9.8" | 6.5" | 4.9" | 3.5" | 2.6" |
| C-3 | 500 m | AR | 1.4" | 0.94" | 0.77" | 0.62" | 0.41" | 0.31" | 0.22" | 0.16" |
| | 15 m | MRS | 16.2" | 10.8" | 8.7" | 7.0" | 4.7" | 3.5" | 2.5" | 1.9" |
| C-4 | 784 m | AR | 0.92" | 0.61" | 0.50" | 0.40" | 0.27" | 0.20" | 0.14" | 0.11" |
| | 15 m | MRS | 11.2" | 7.5" | 6.1" | 4.9" | 3.3" | 2.4" | 1.7" | 1.3" |
| C-5 | 1.4 km | AR | 0.54" | 0.36" | 0.30" | 0.24" | 0.16" | 0.12" | 0.084" | 0.063" |
| | 15 m | MRS | 6.7" | 4.5" | 3.6" | 2.9" | 1.9" | 1.5" | 1.0" | 0.77" |
| C-6 | 2.5 km | AR | 0.31" | 0.20" | 0.17" | 0.13" | 0.089" | 0.067" | 0.047" | 0.035" |
| | 15 m | MRS | 4.1" | 2.7" | 2.2" | 1.8" | 1.2" | 0.89" | 0.63" | 0.47" |
| C-7 | 3.6 km | AR | 0.21" | 0.14" | 0.11" | 0.092" | 0.061" | 0.046" | 0.033" | 0.024" |
| | 64 m | MRS | 2.6" | 1.7" | 1.4" | 1.1" | 0.75" | 0.56" | 0.40" | 0.30" |
| C-8 | 8.5 km | AR | 0.096" | 0.064" | 0.052" | 0.042" | 0.028" | N/A | N/A | N/A |
| | 110 m | MRS | 1.4" | 0.95" | 0.77" | 0.62" | 0.41" | N/A | N/A | N/A |

From the ALMA Cycle 6 Proposal Guide

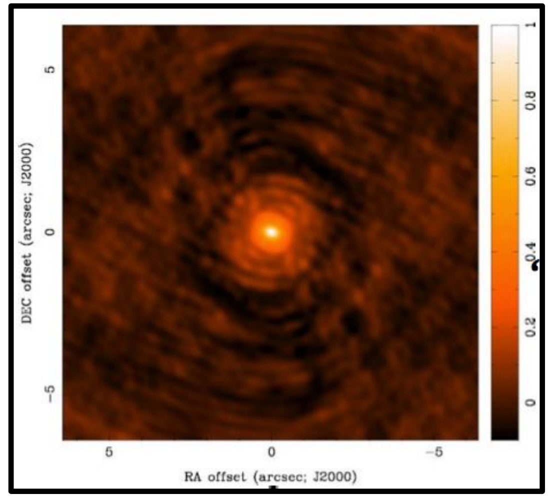
# The Dirty Beam

**S(u,v)**

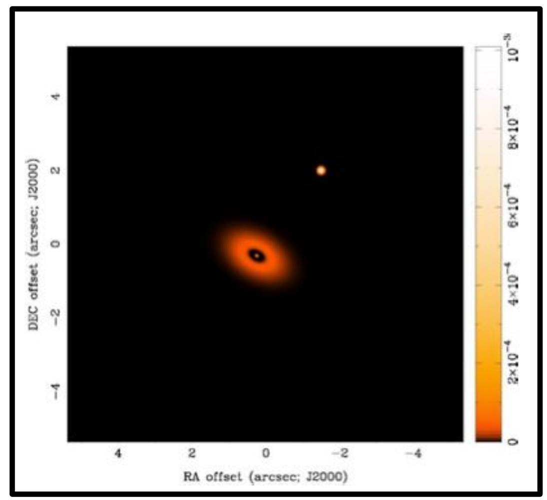



**FT**

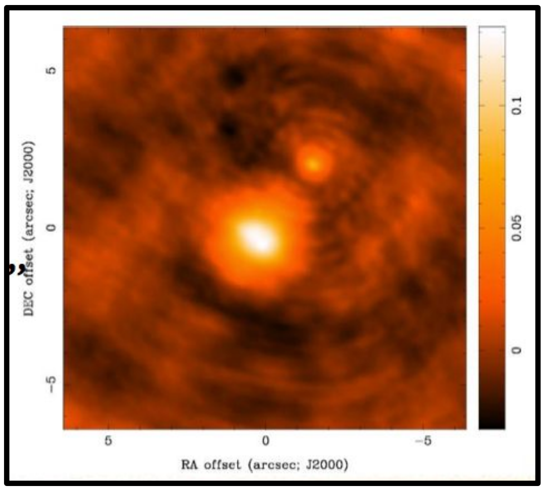
**s(x,y)**
**“Dirty Beam”**




*** (Convolution)**

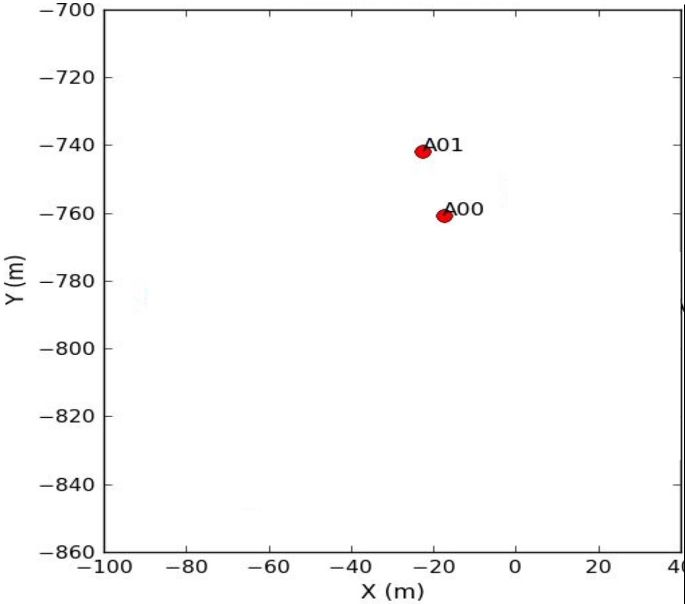


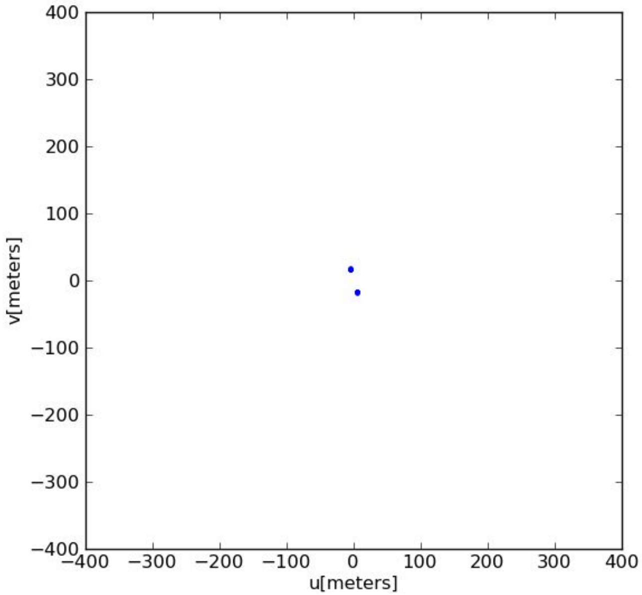

**T(x,y)**

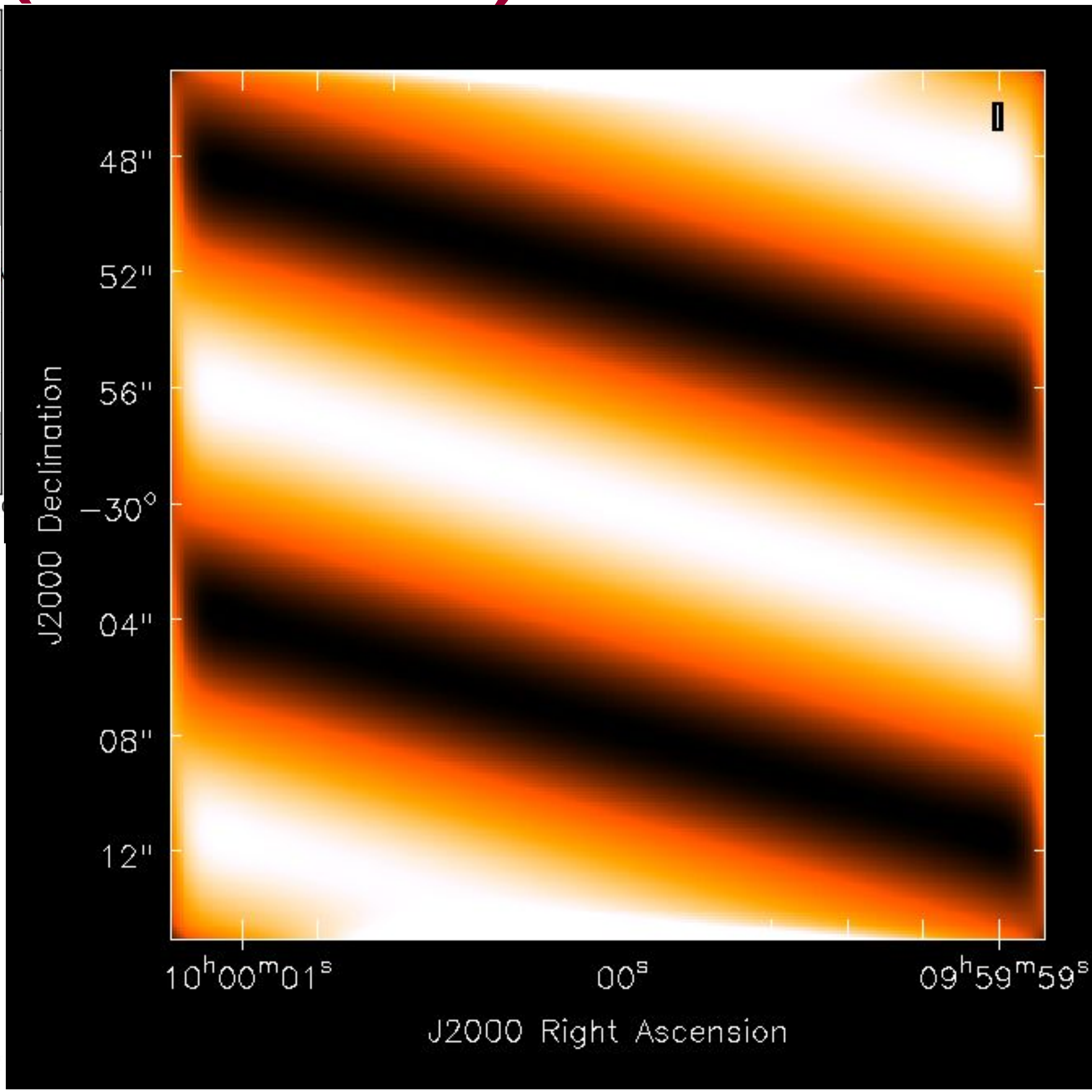



**$T_D(x,y)$ “Dirty Image”**

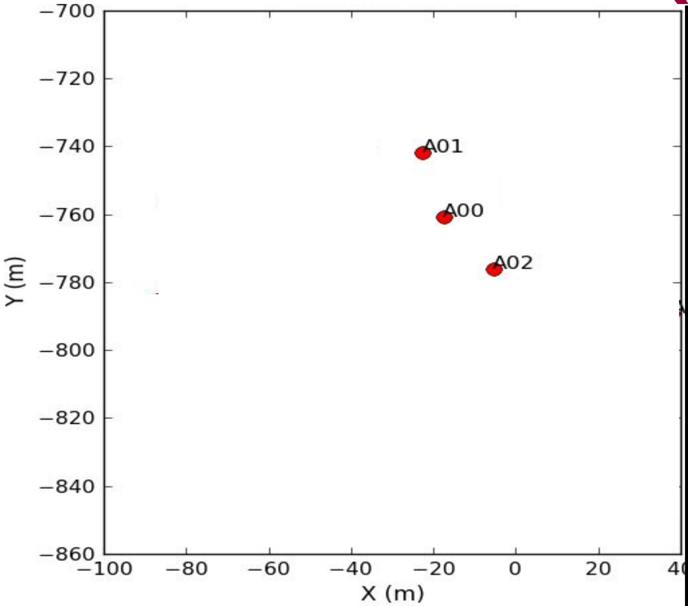
# Example: Fringe pattern with 2 Antennas (one baseline)

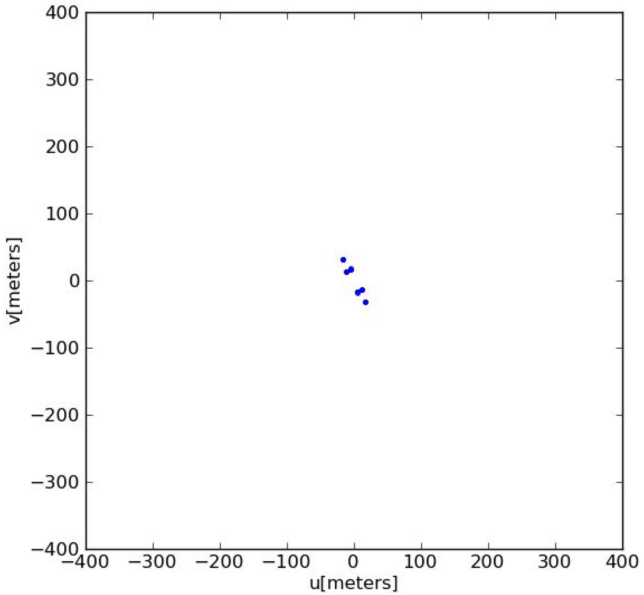


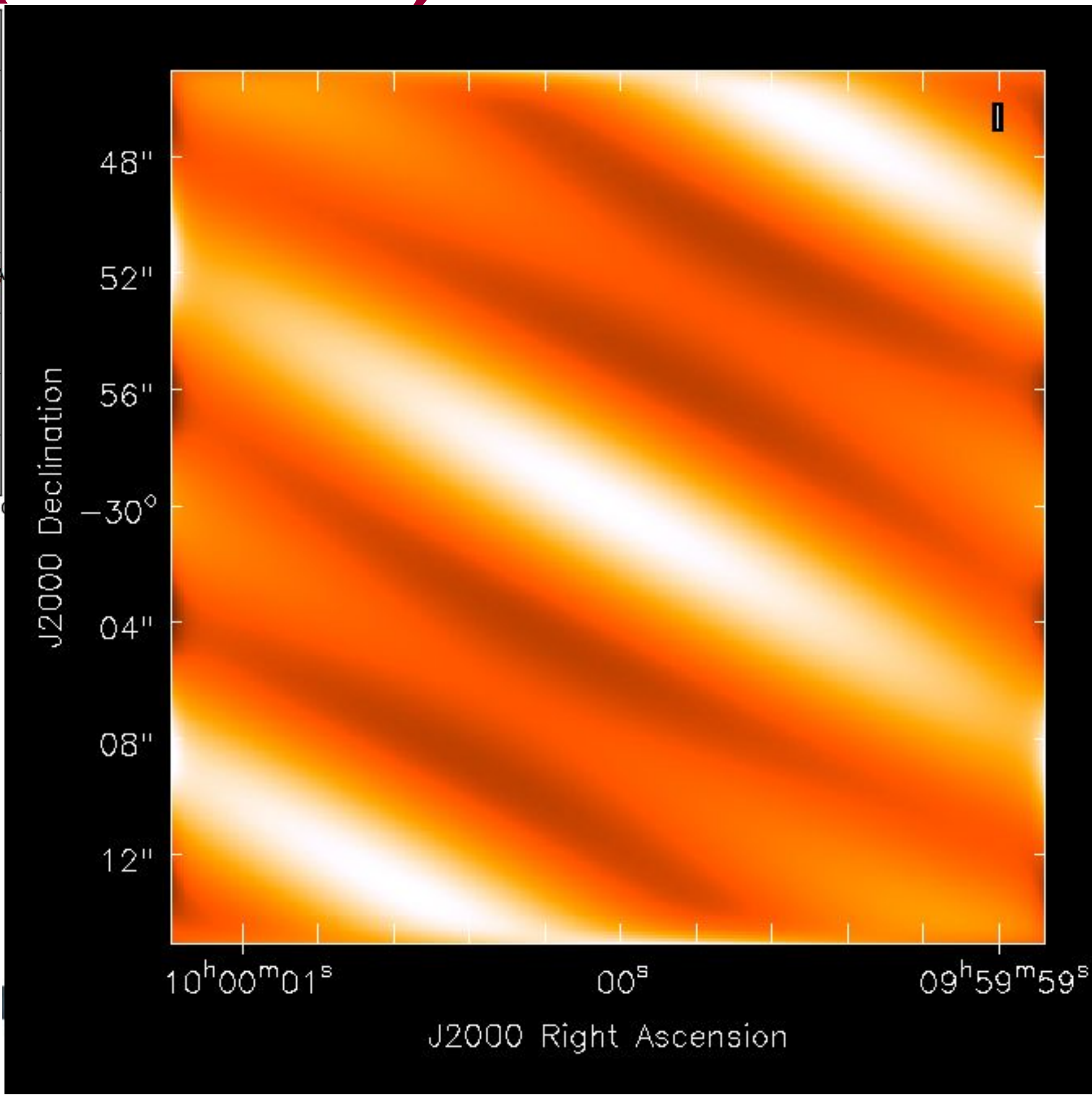



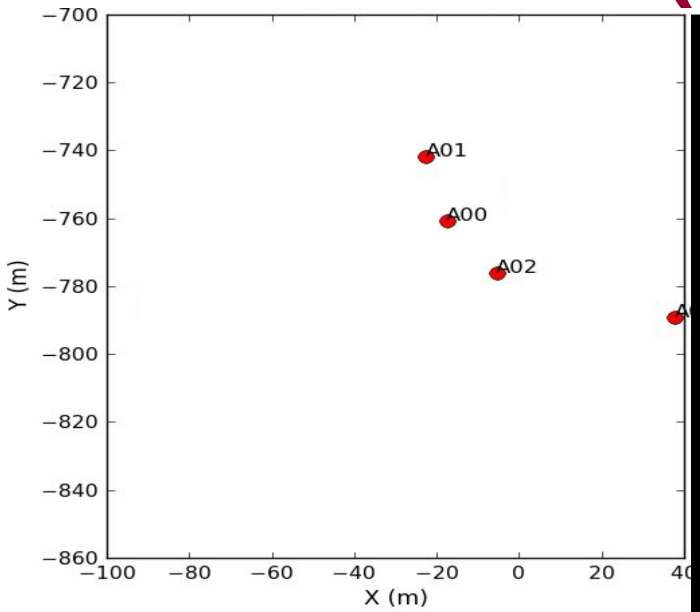
# Example: Fringe pattern with 3 Antennas (3 baselines)

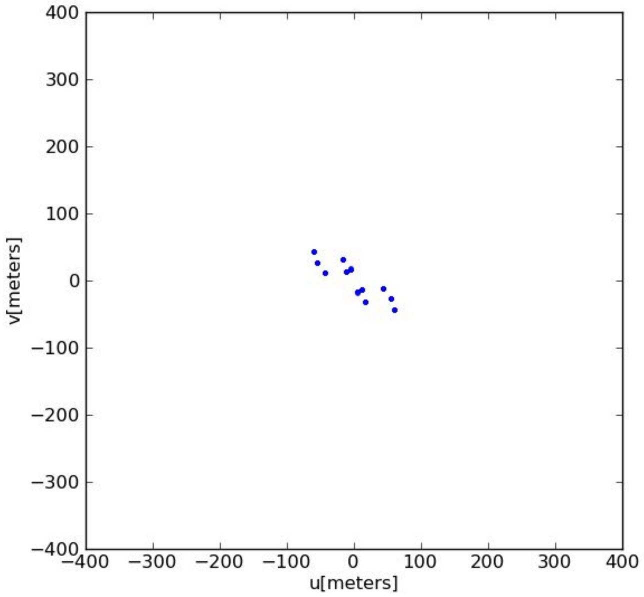


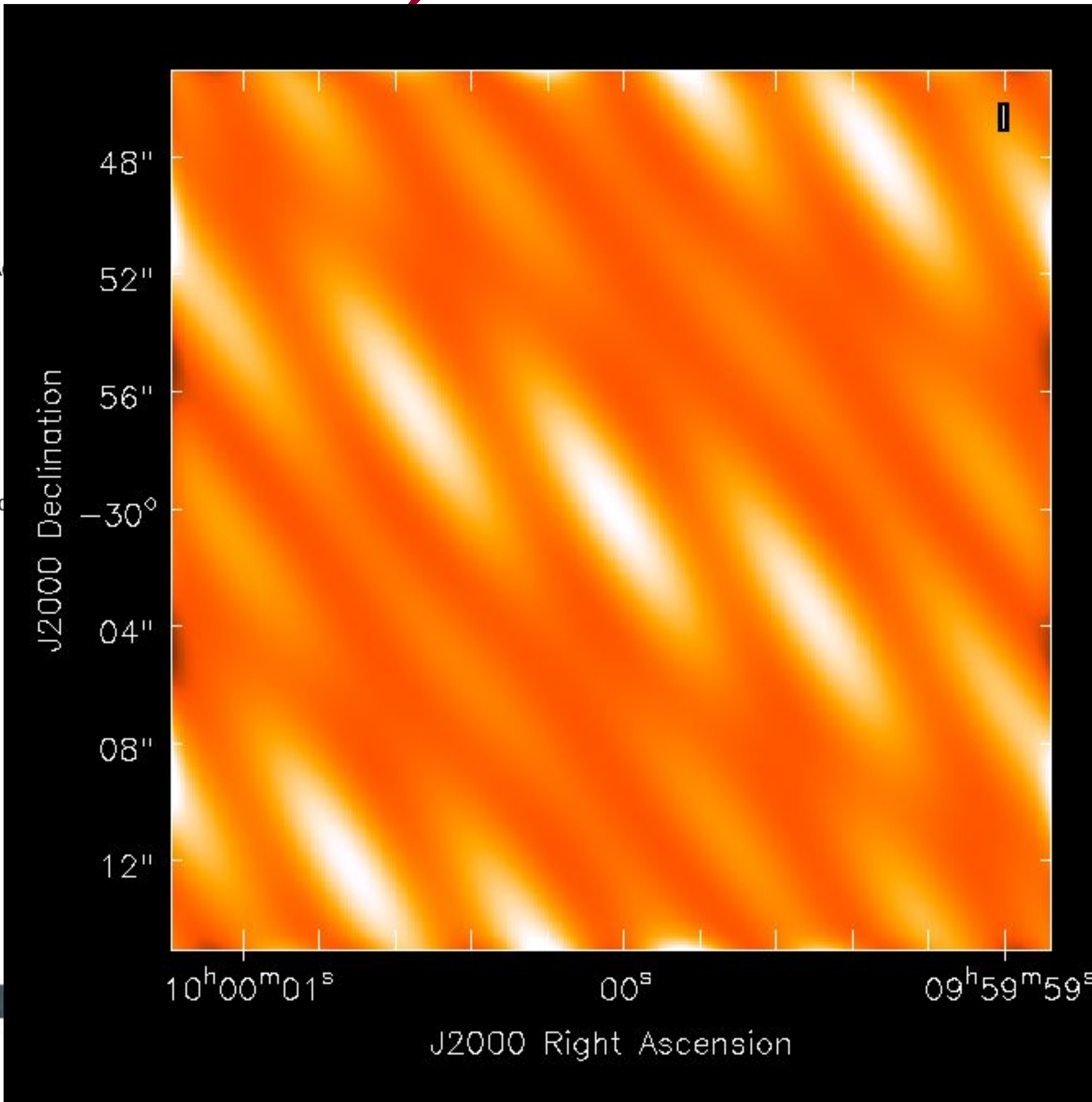



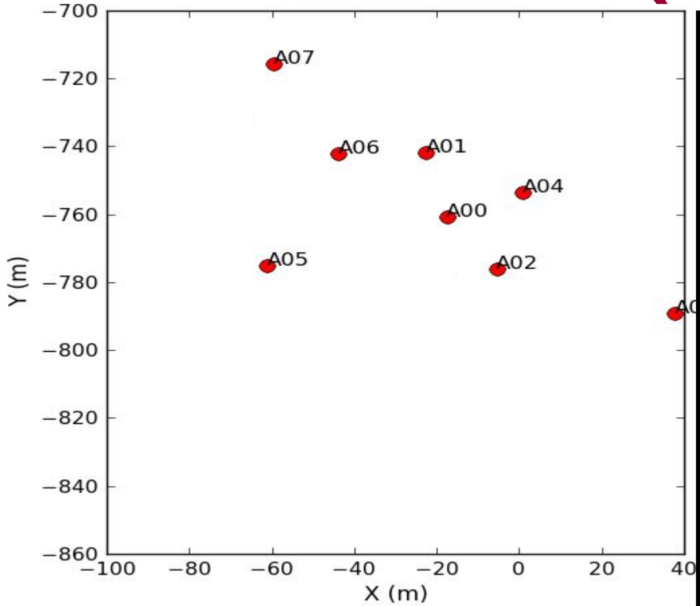
# Example: Fringe pattern with 4 Antennas (6 baselines)

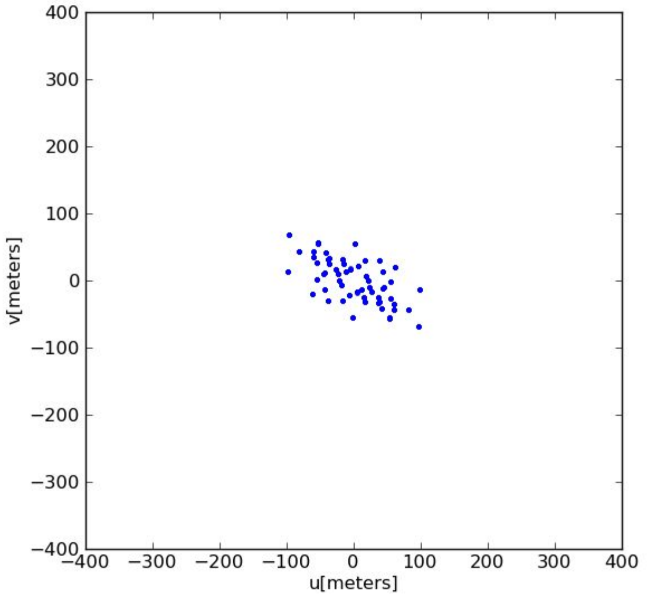


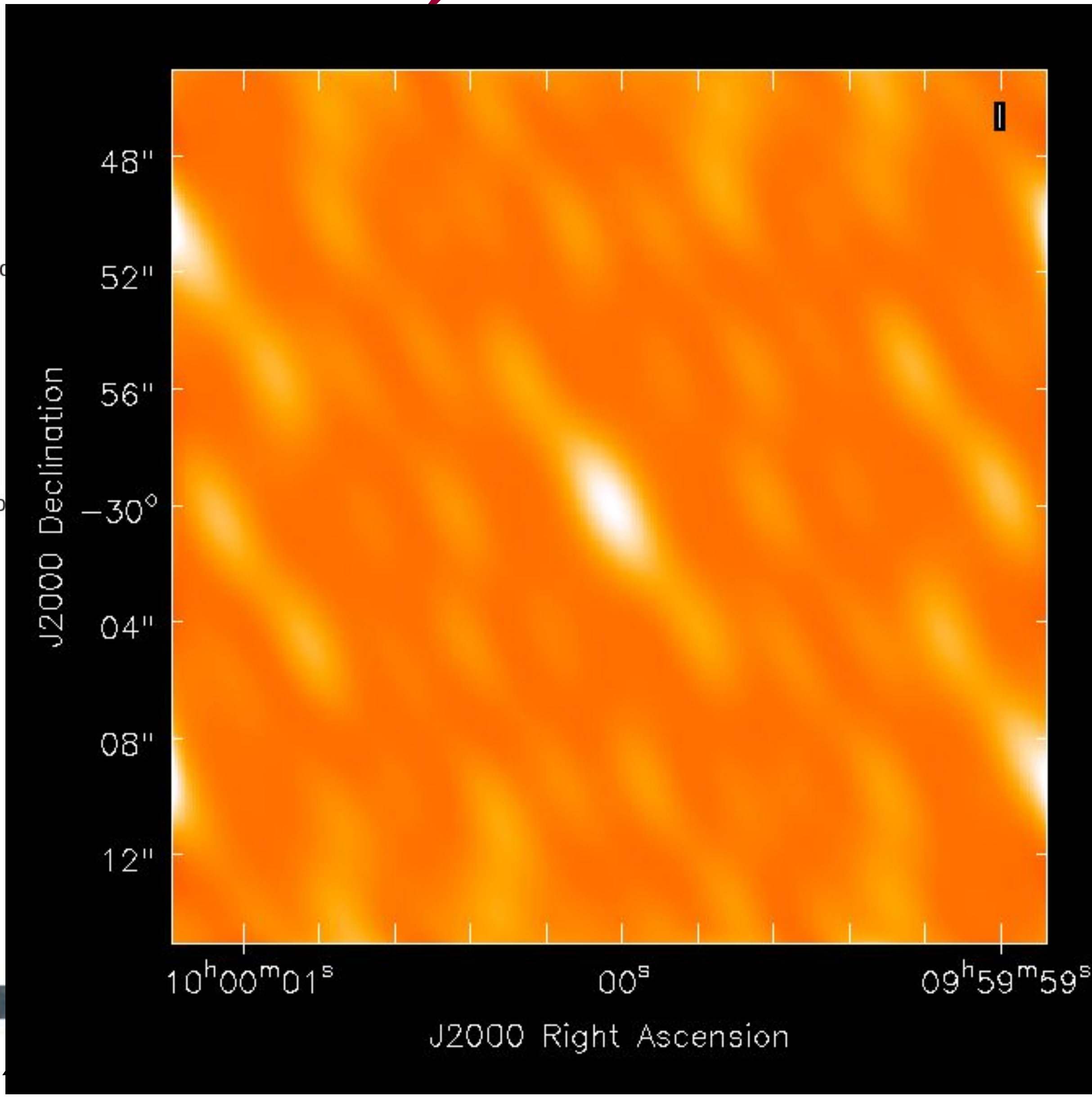



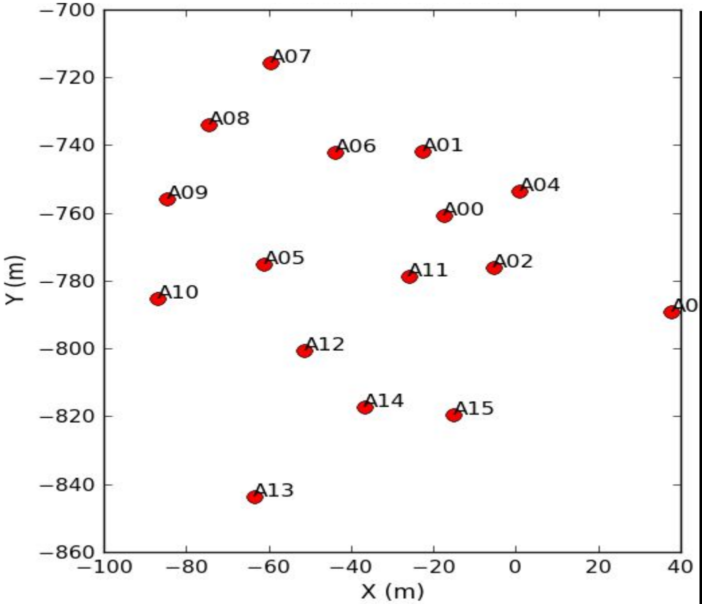
# Example: Fringe pattern with 8 Antennas (28 baselines)

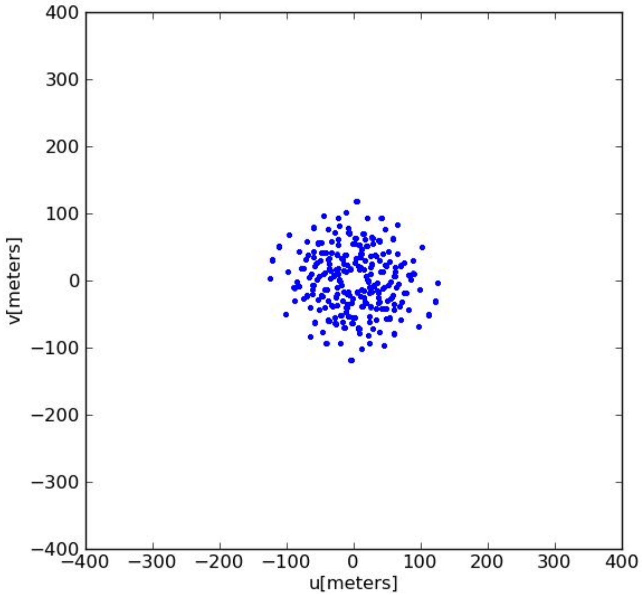


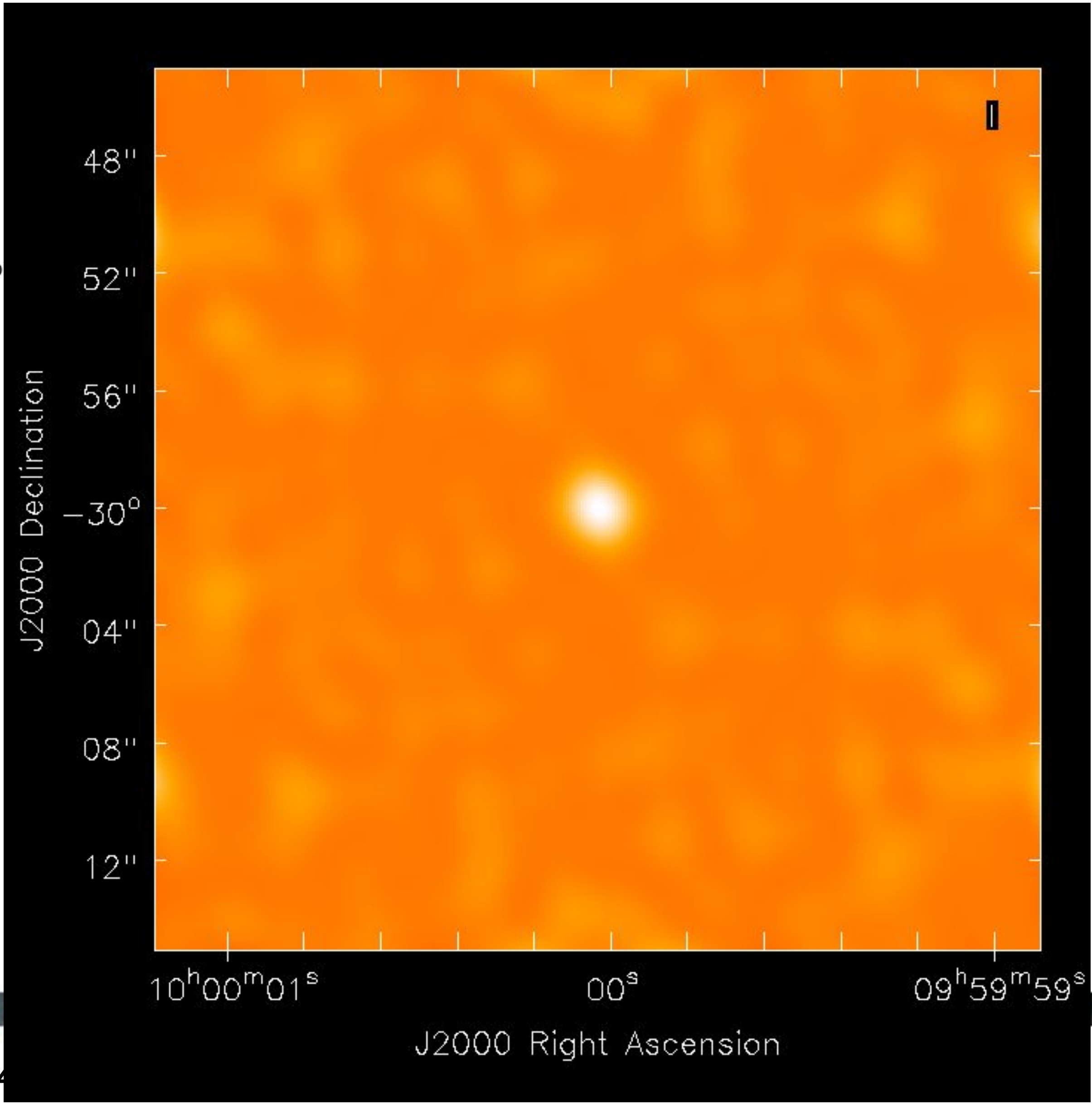



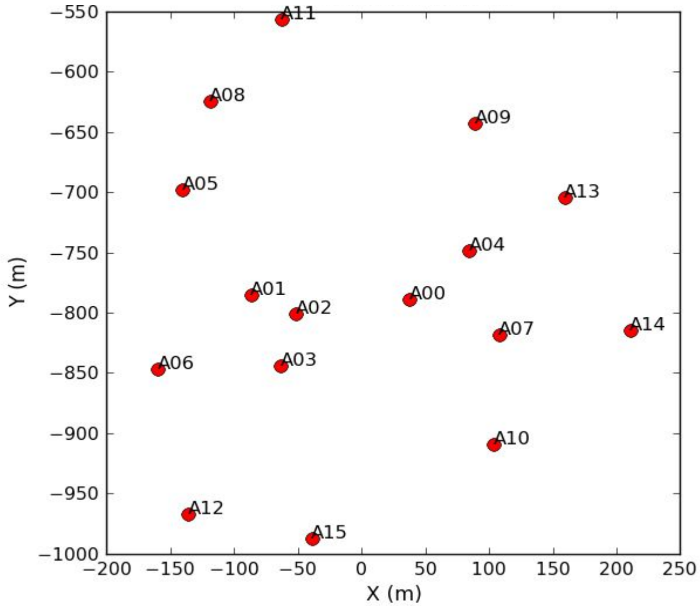
# 16 Antennas – Compact Configuration

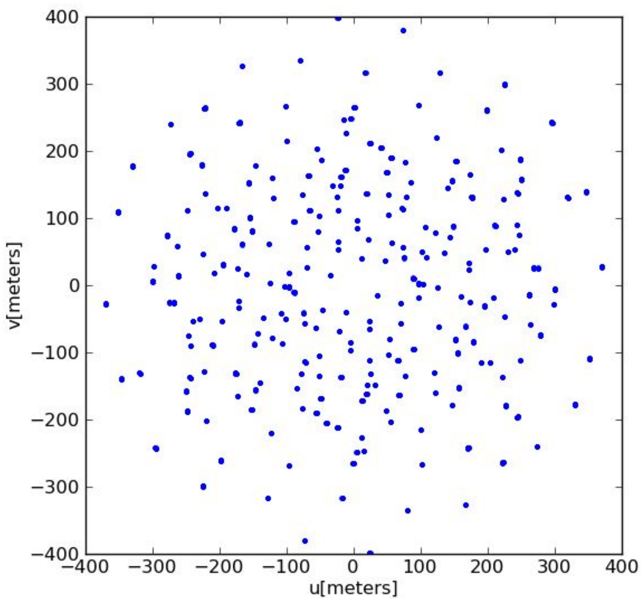


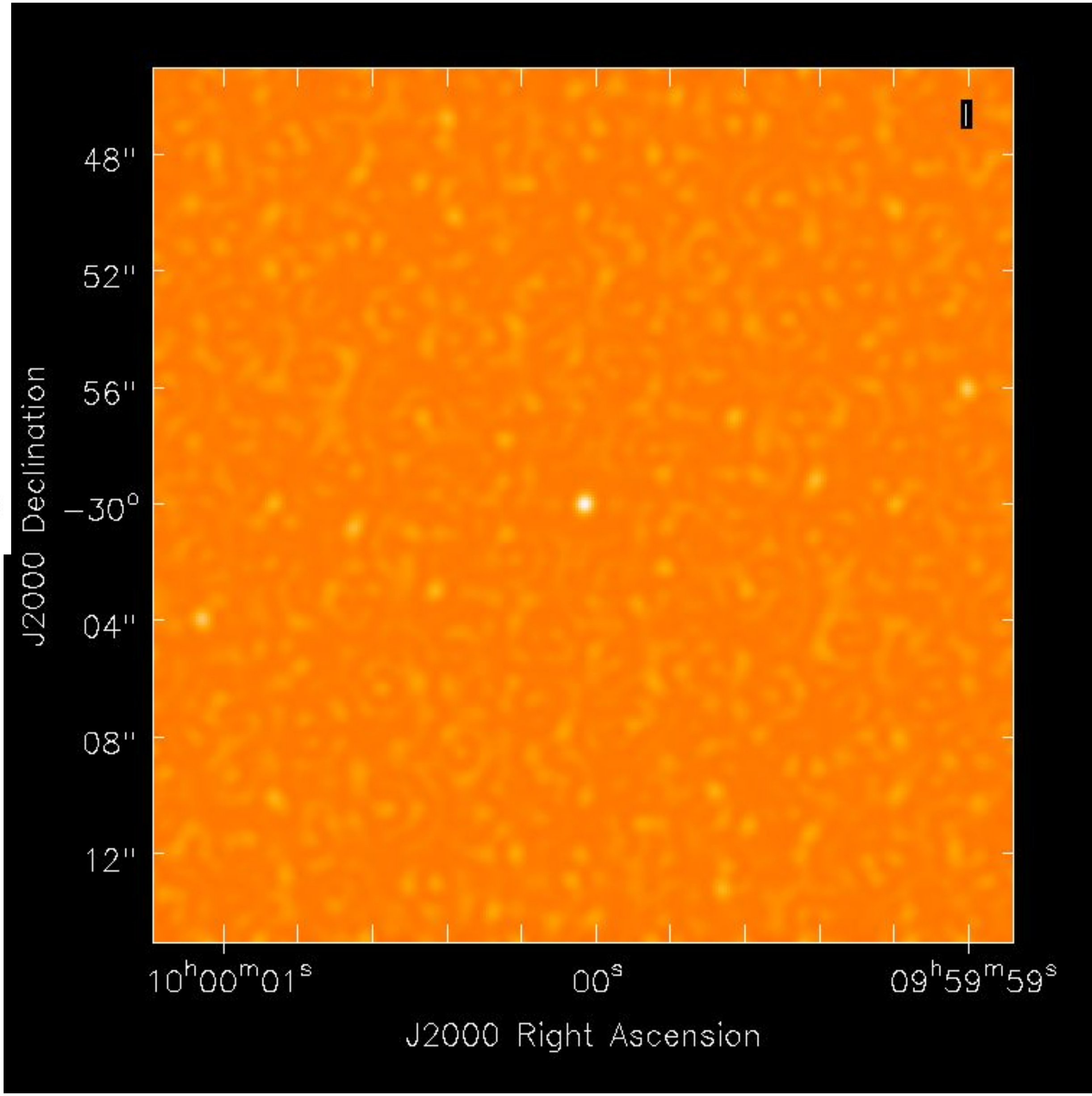



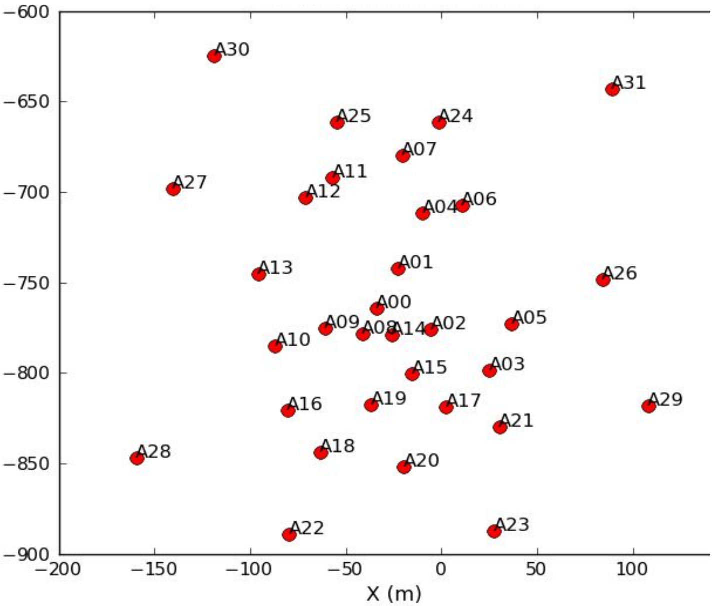
# 16 Antennas – Extended Configuration

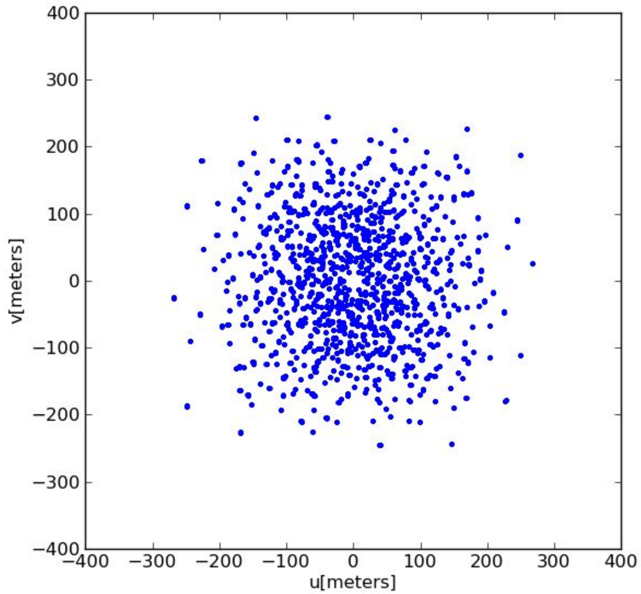


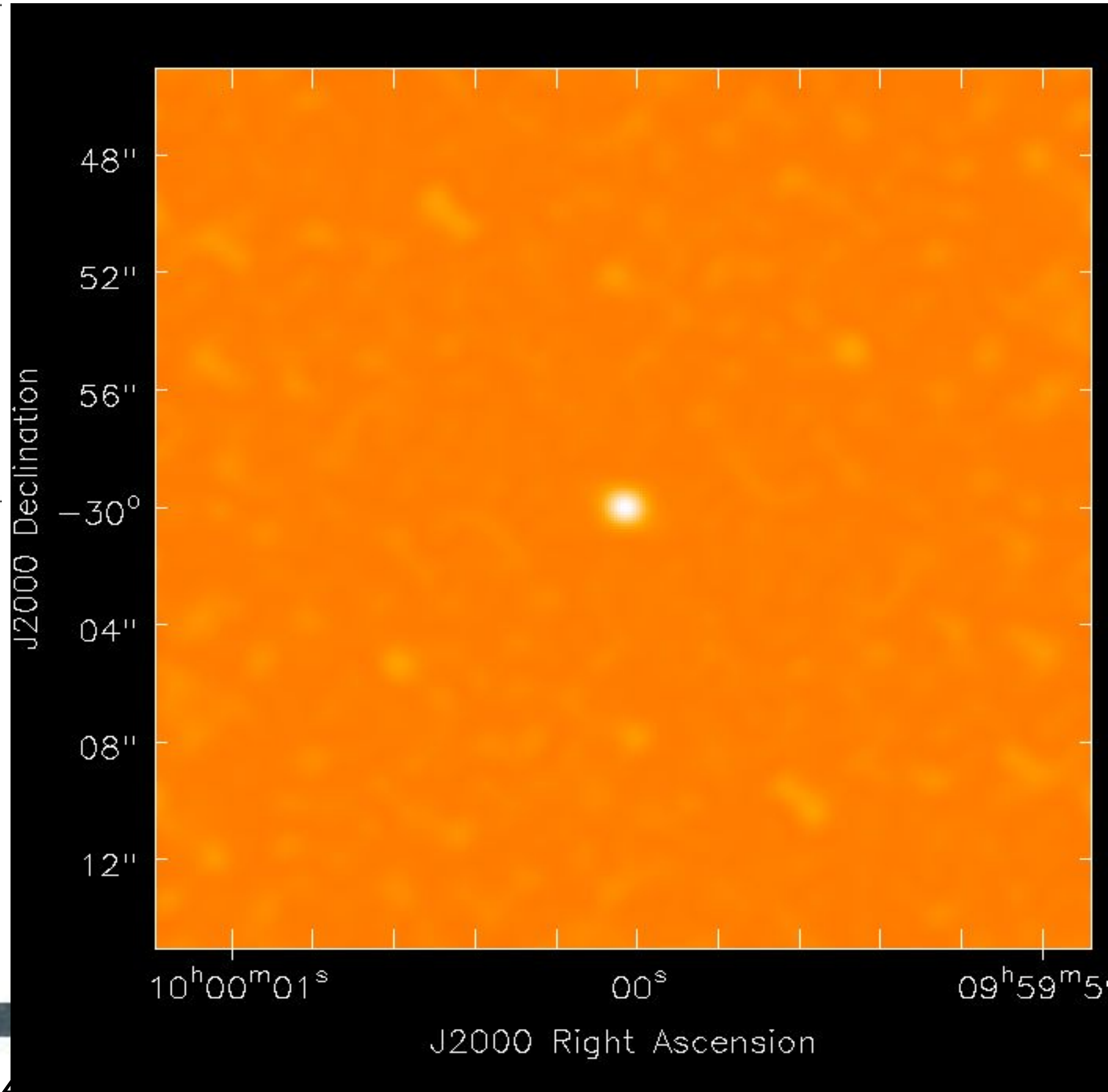



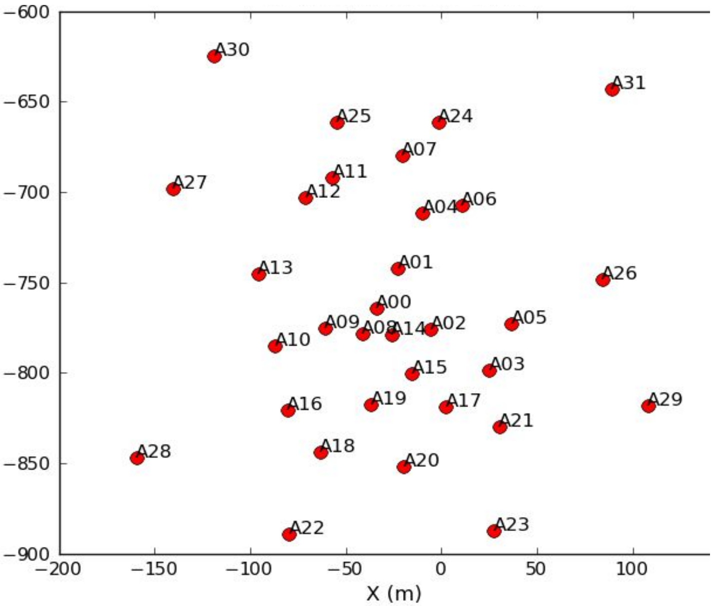
# 32 Antennas – Instantaneous

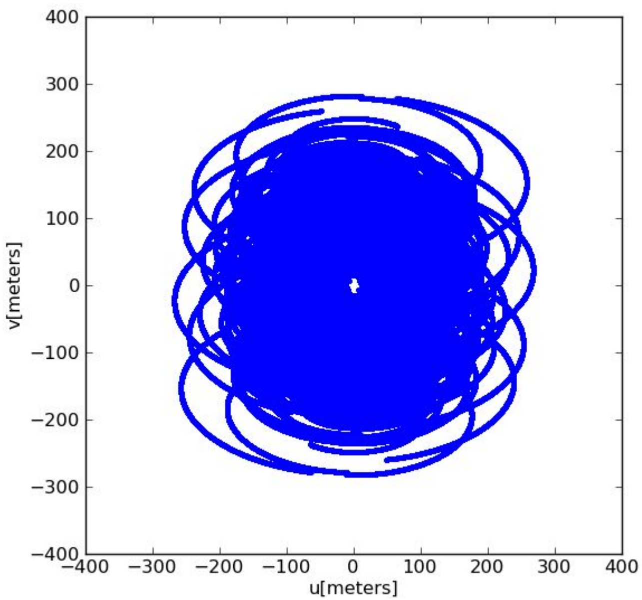


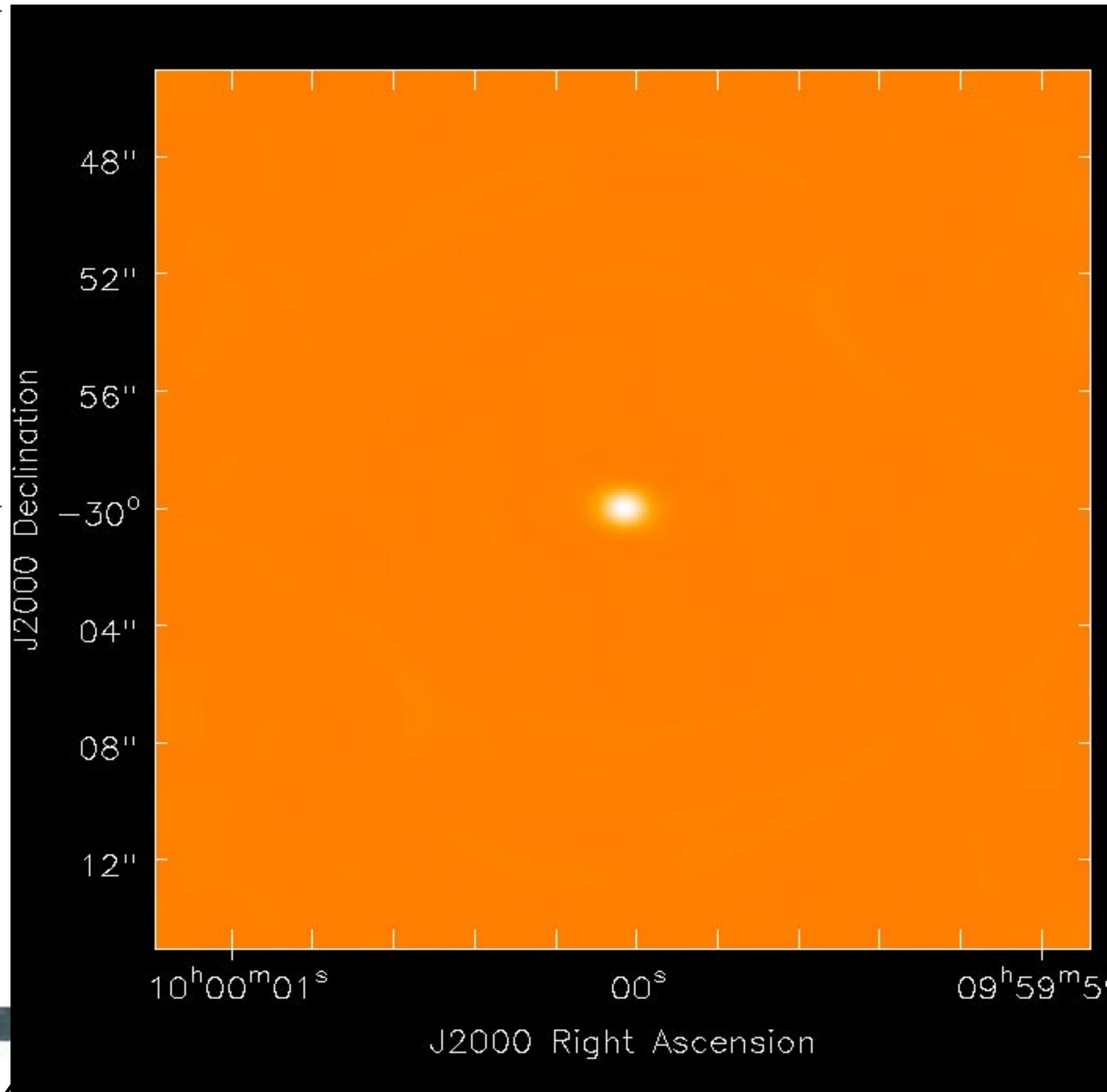



# 32 Antennas – 8 hours

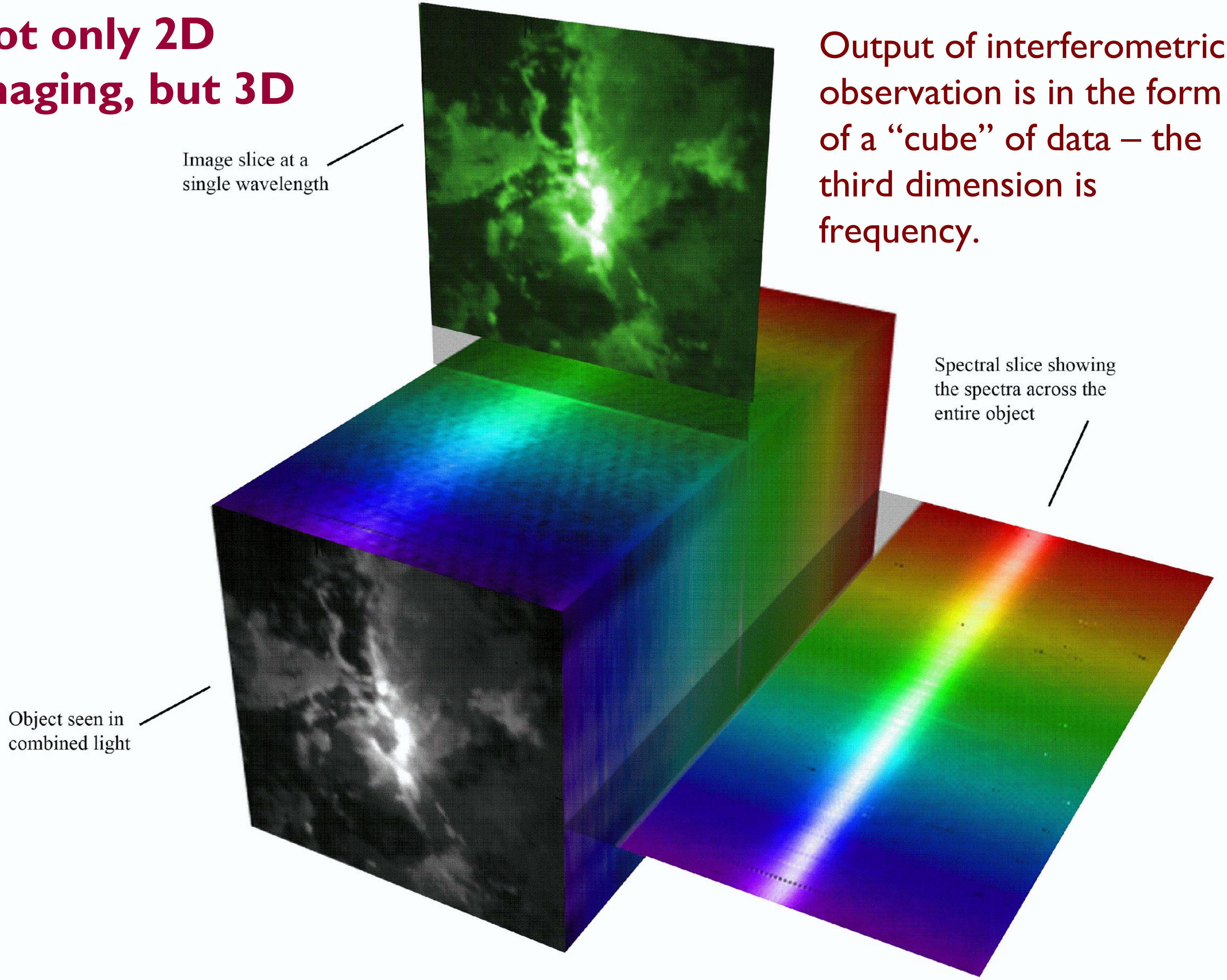


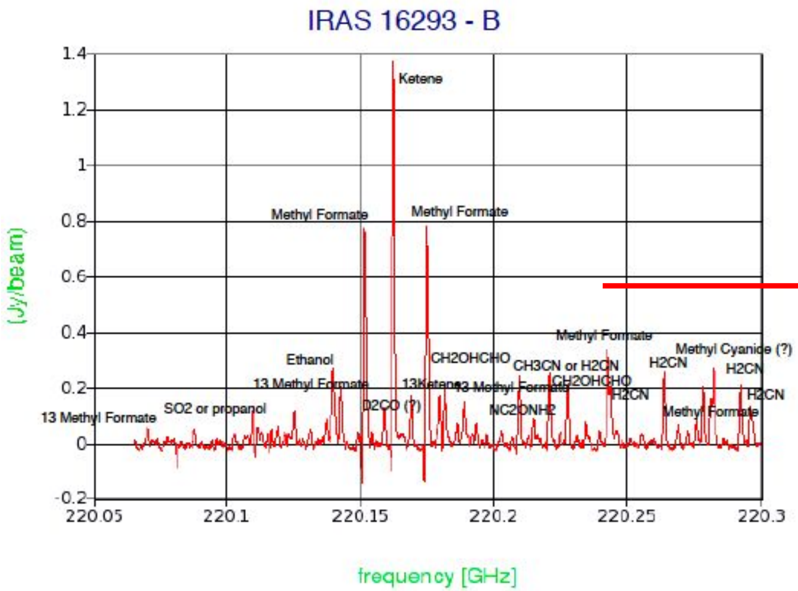



# Not only 2D imaging, but 3D

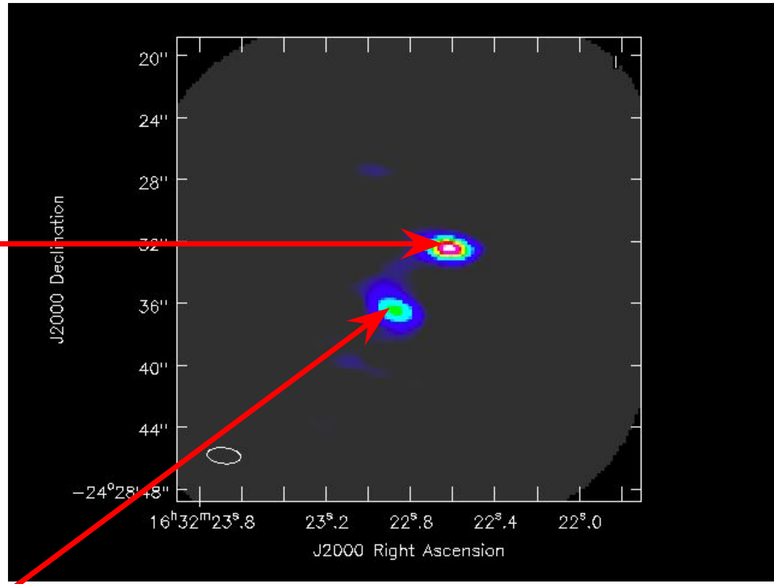
Output of interferometric observation is in the form of a "cube" of data – the third dimension is frequency.



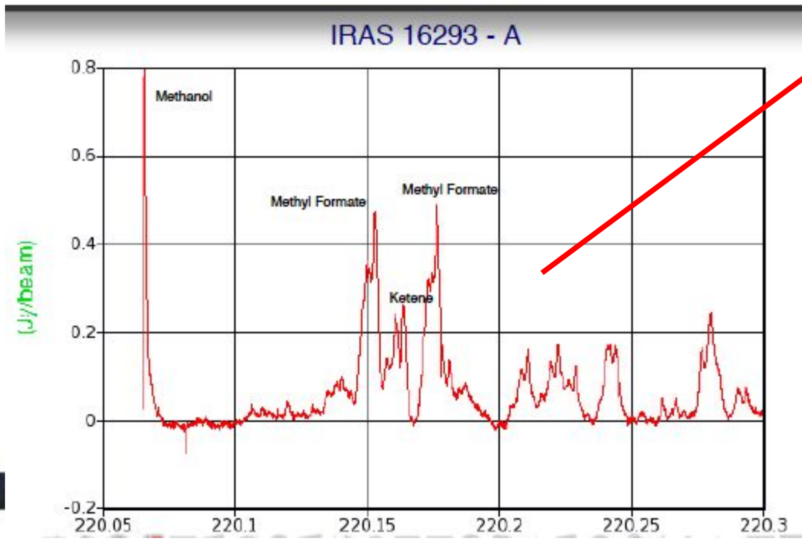
# Sometimes the most interesting science lies in the third dimension




Band 6




J. Turner & ALMA CSV team

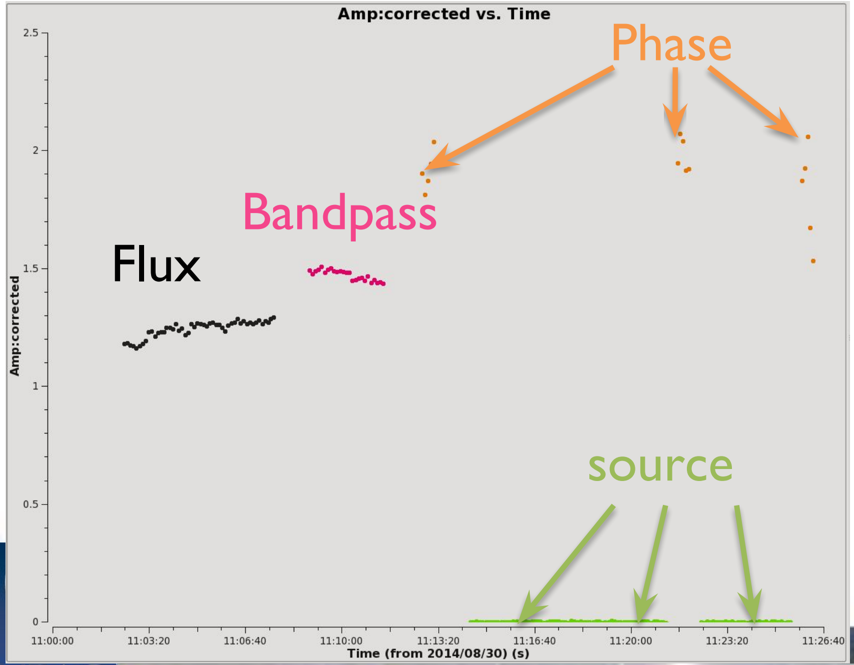


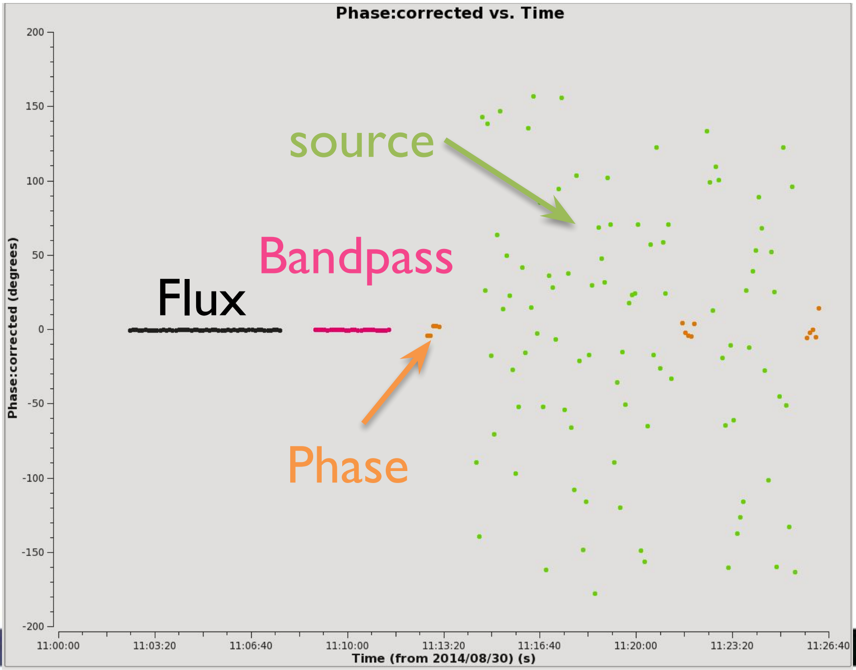

## Young Low Mass Stars: IRAS16293

- Note narrow lines toward pre-protostellar core B (top) with infall apparent in methyl formate and ketene lines.

# A Brief Word on Calibration

- Interferometers measure visibilities, i.e., the amplitude and phase of the cross-correlated signals between pairs of antennas, as a function of time and frequency.
- We calibrate these data by determining the complex gains (amplitude and phase), the frequency response (bandpass) and flux scale for each antenna.

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# Calibration Process

Calibration is the effort to measure and remove the time-dependent and frequency-dependent atmospheric and instrumental variations.

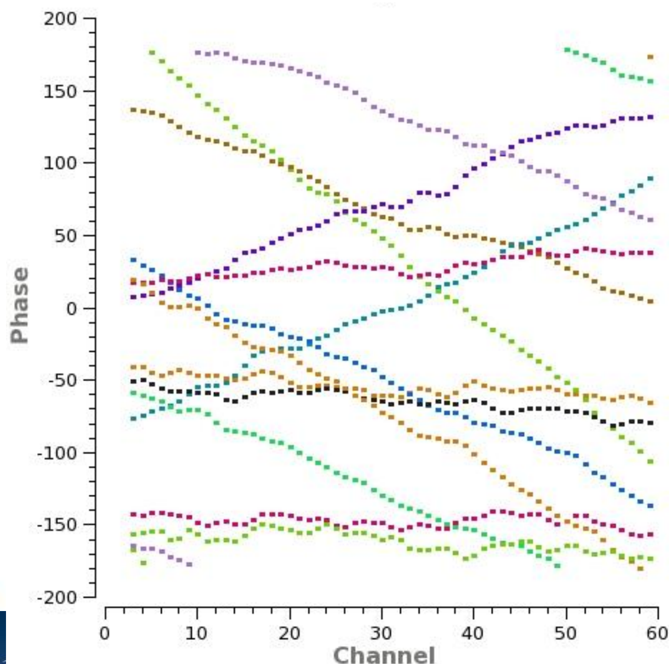
Steps in calibrating interferometric data:
(Note: You don't have to worry about these in your observational set up!)

- Bandpass calibration (correct frequency-dependent telescope response)
- Phase and amplitude gain calibration (remove effects of atmospheric water vapor and correct time-varying phases/amplitudes)
- Set absolute flux scale

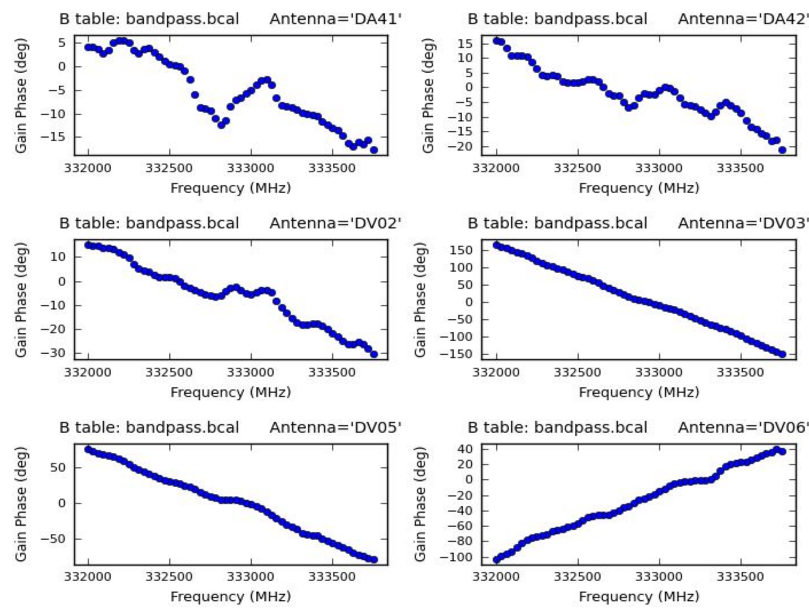
# Bandpass Calibration: Phase

* Analogous to optical “flat fielding” + bias subtraction for each antenna.
* Primarily correcting for frequency dependent telescope response (i.e. in the correlator/spectral windows)
* Done once in an SB, uses bright point sources like quasars
* Typically, baseline responses are inverted to antenna-based correction

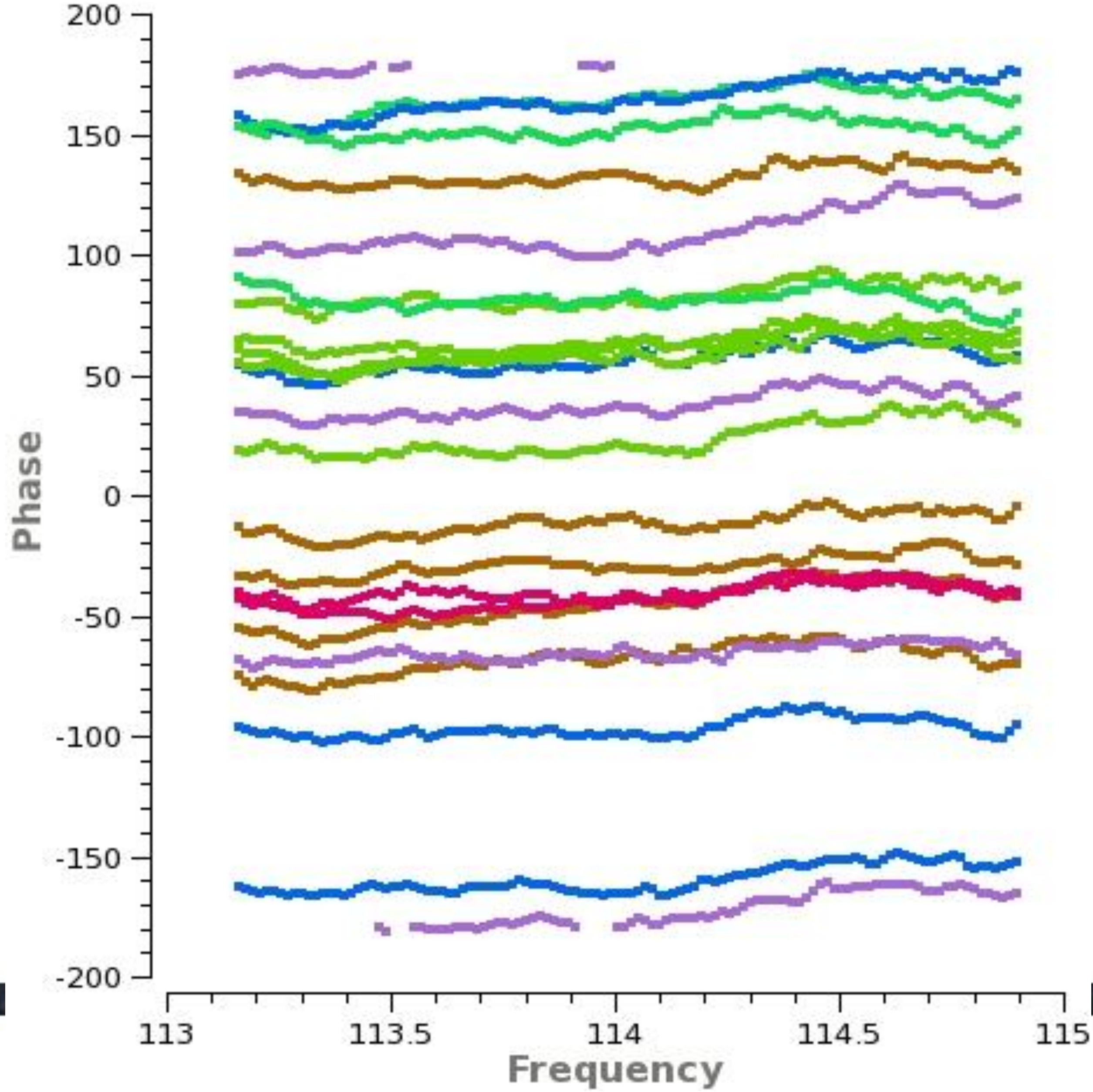
Baselines to one antenna



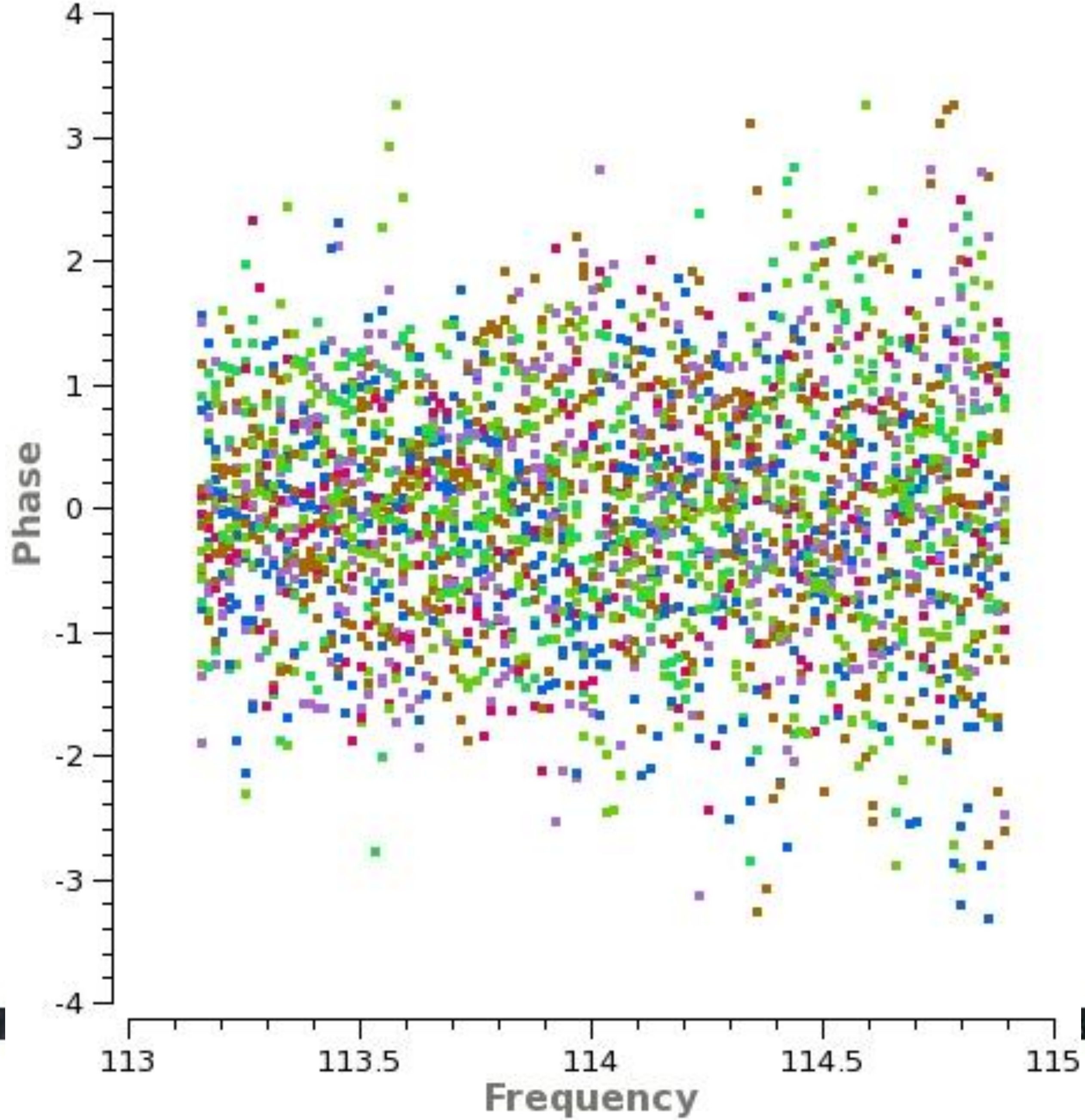

Antenna-based Bandpass Solutions



# Bandpass Phase vs. Frequency (Before)

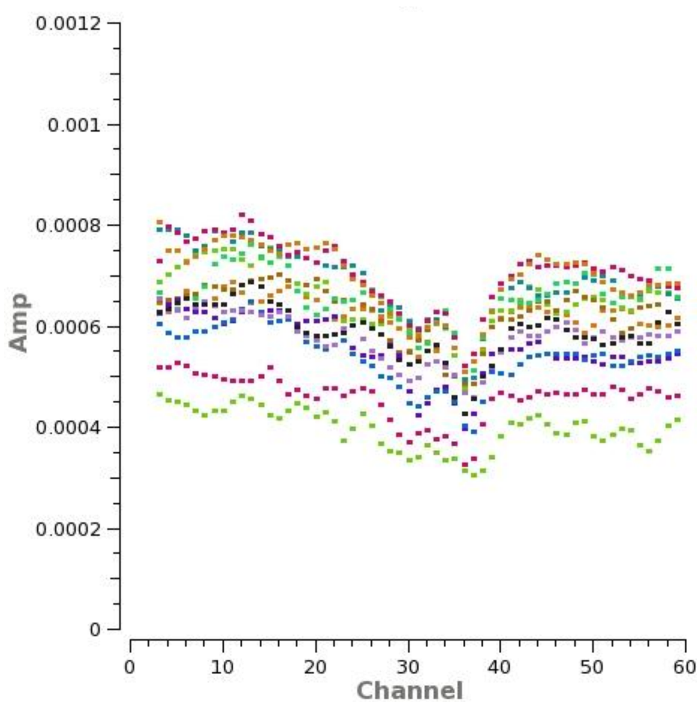


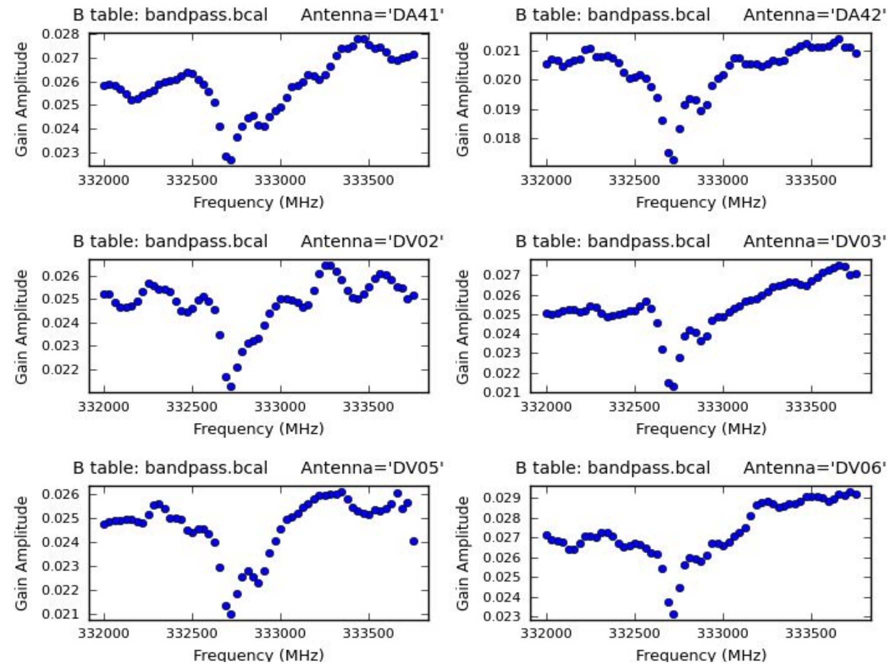
# Bandpass Phase vs. Frequency (After)



# Bandpass Calibration: Amplitude

Baselines to one antenna




Amplitude Before Bandpass Calibration

Bandpass solutions for individual antennas

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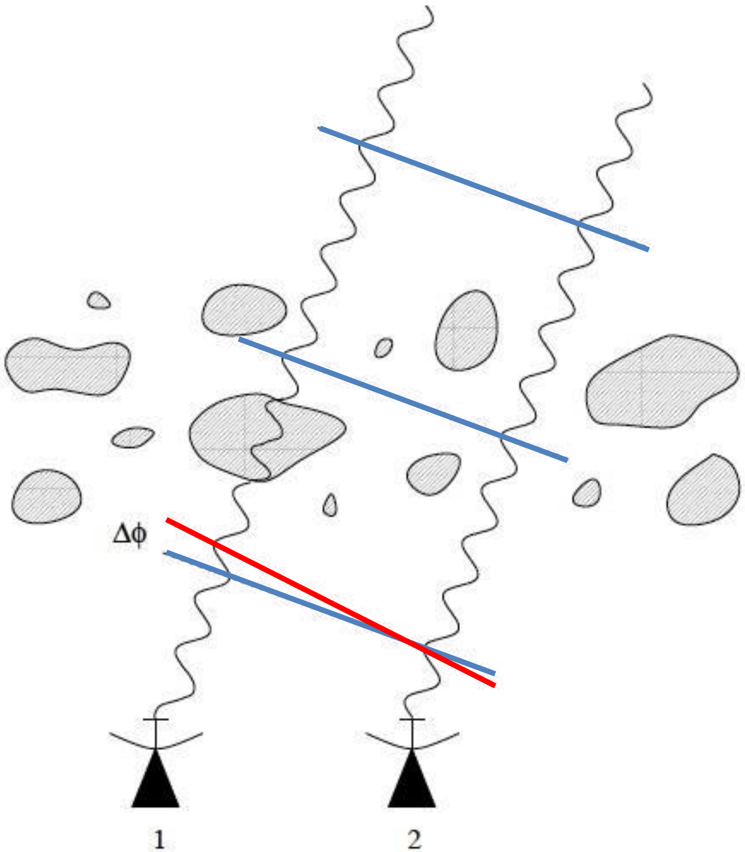
# Atmospheric Phase Correction

- Variations in the amount of precipitable water vapor cause phase fluctuations that result in:
  - Low coherence (loss of sensitivity)
  - Radio "seeing" of 1arcsec at 1mm
  - Anomalous pointing offsets
  - Anomalous delay offsets

Patches of air with different water vapor content (and hence index of refraction) affect the incoming wave front differently.



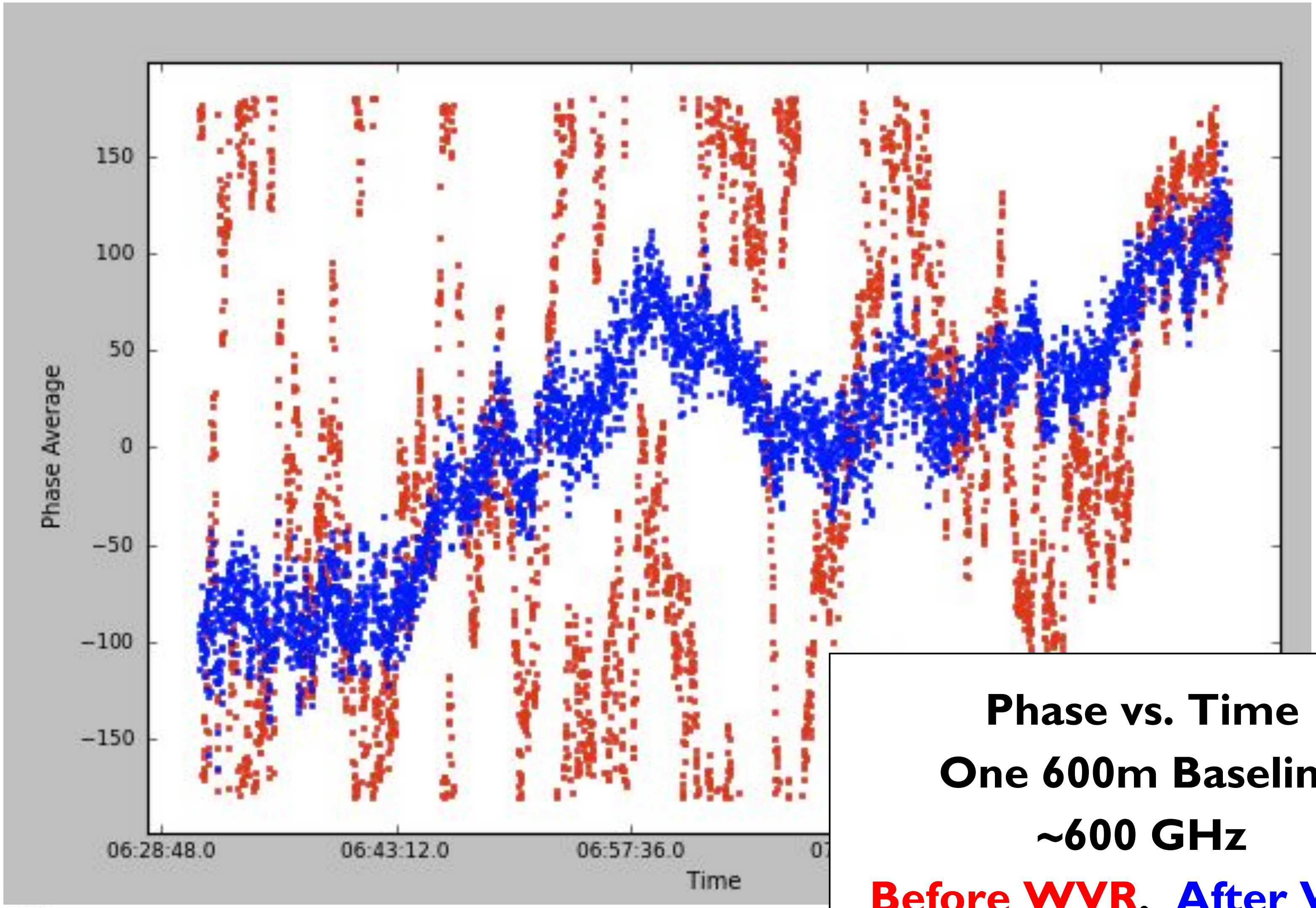
# Phase & Amplitude Gain Calibration

Determines the variations of phase and amplitude over time

- First pass is atmospheric correction from Water Vapor Radiometers readings.

- Final correction from gain calibrator (point source near to target) that is observed every few minutes throughout the observation (analogous to repeat trips to a standard star).

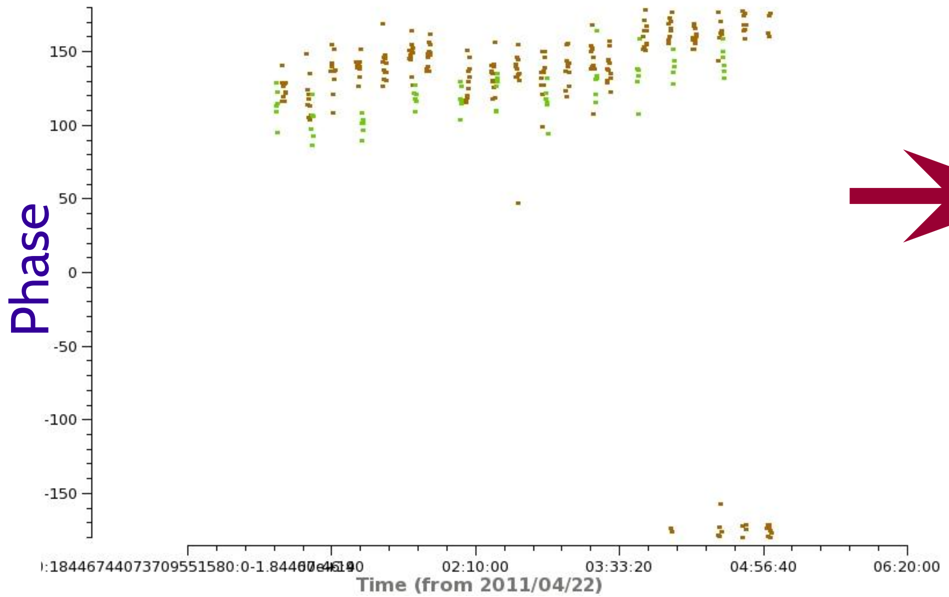
# Water Vapor Correction on ALMA

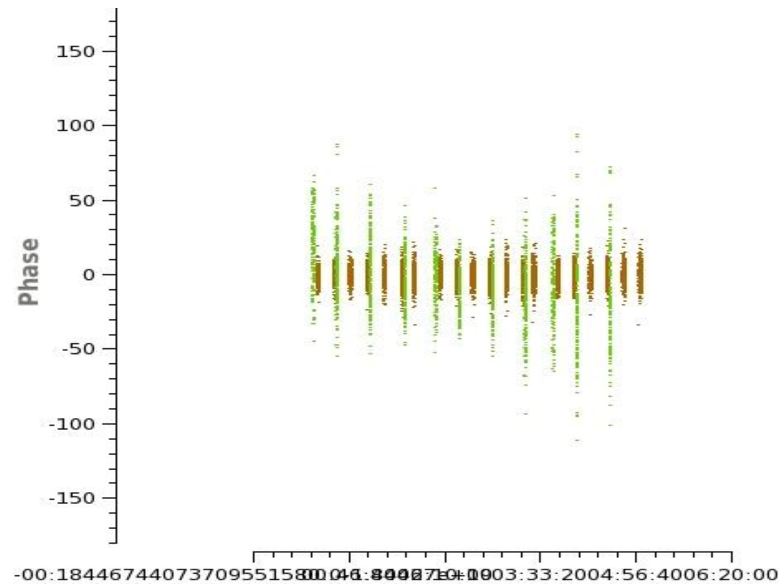



**Phase vs. Time**
**One 600m Baseline**
**~600 GHz**
**Before WVR, After WVR**

# Phase Calibration

The phase calibrator must be a point source close to the science target and must be observed frequently.  This provides a model of atmospheric phase change along the line of sight to the science target that can be compensated for in the data.




Time

Corrected using point source model

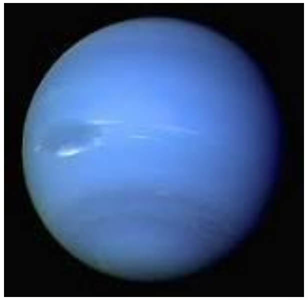
# Flux (or Amplitude) Calibration

Two Steps:

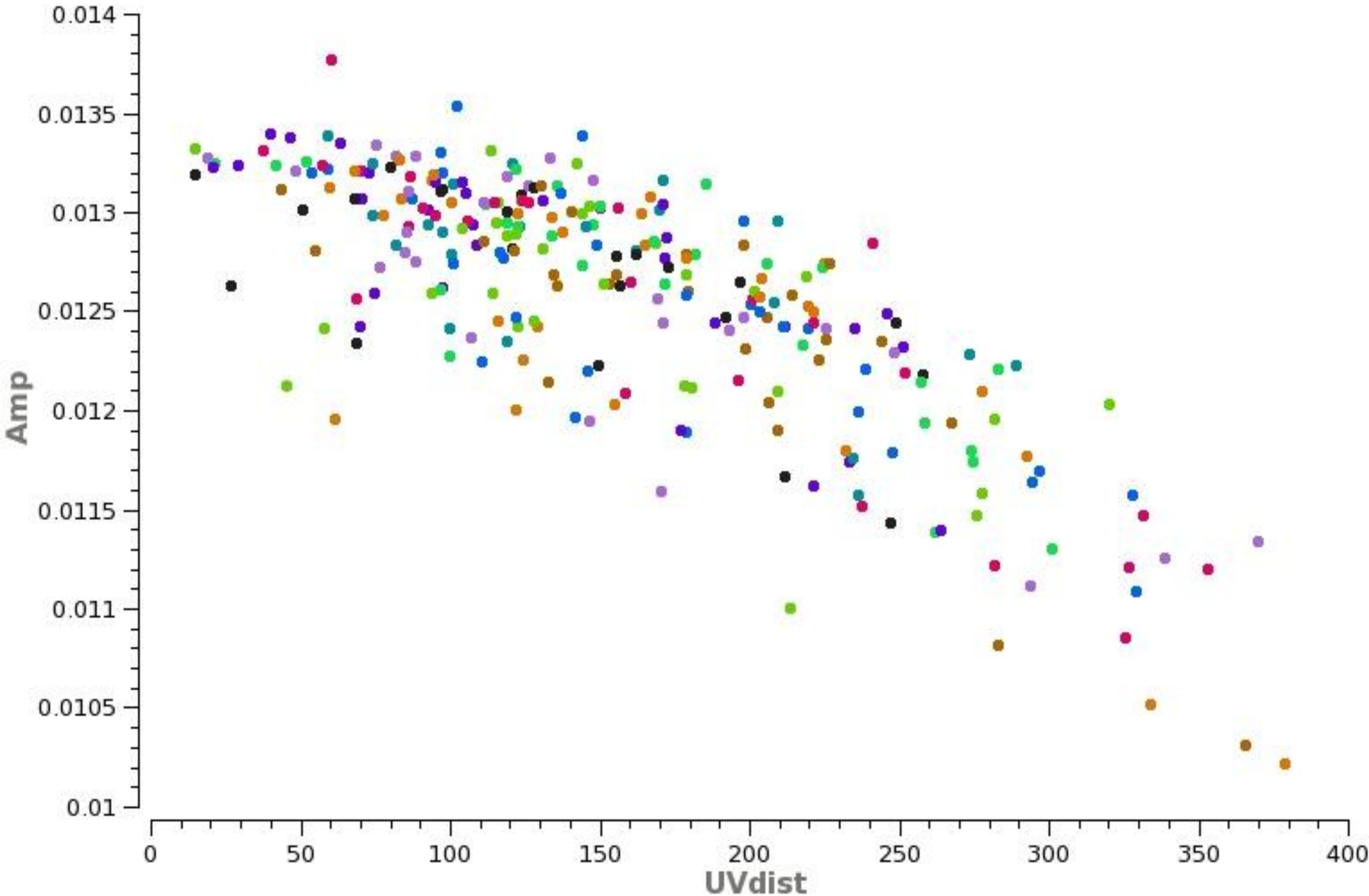
1. Use calibration devices with known temperatures (hotload and ambient load) to measure System Temperature frequently.

2. Use a source of known flux to convert the signal measured at the antenna to common unit (Janskys). If the source is resolved, or has spectral lines, it must be modeled very well.

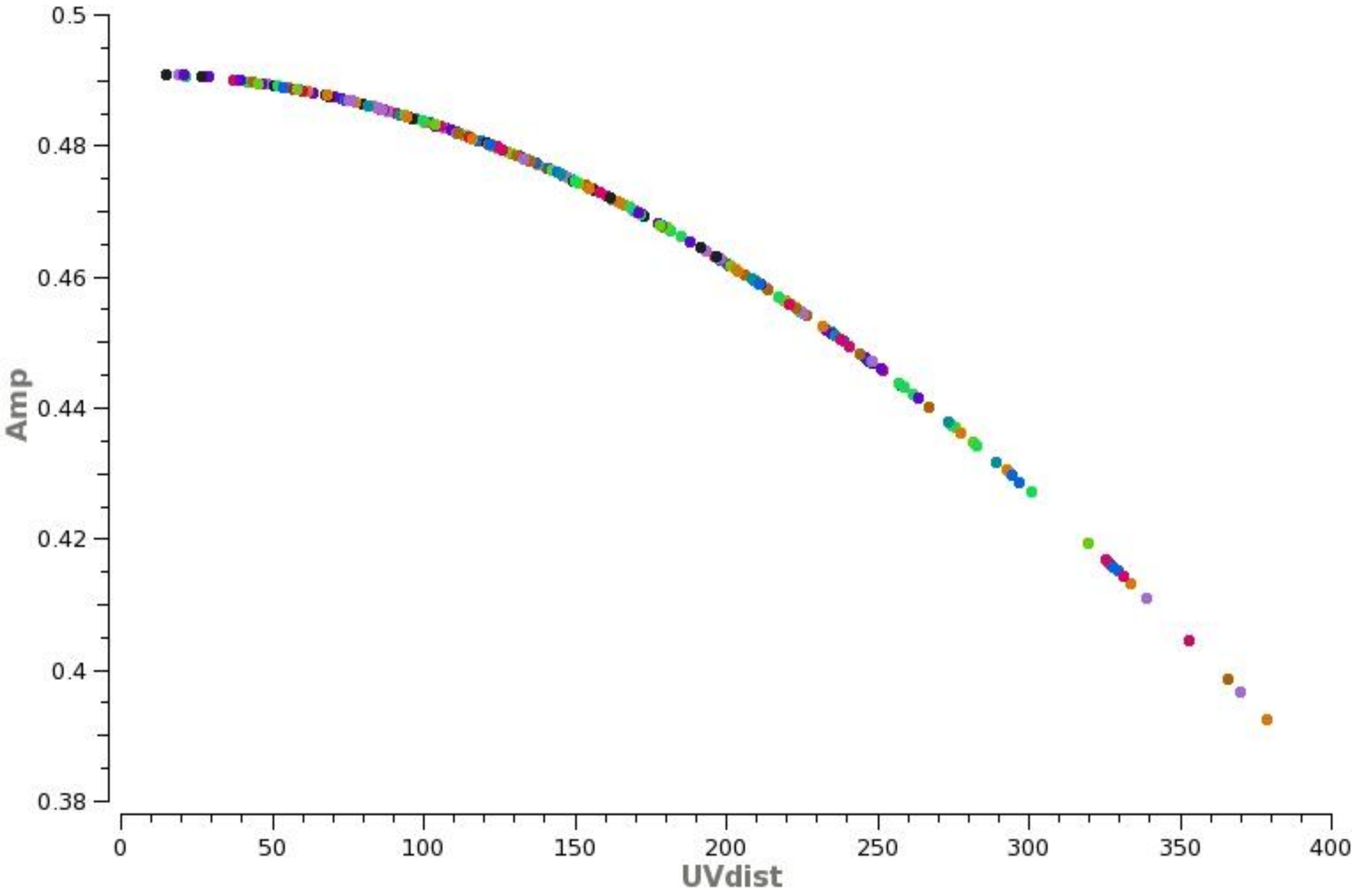
The derived amplitude vs. time corrections for the flux calibrator are then applied to the science target.



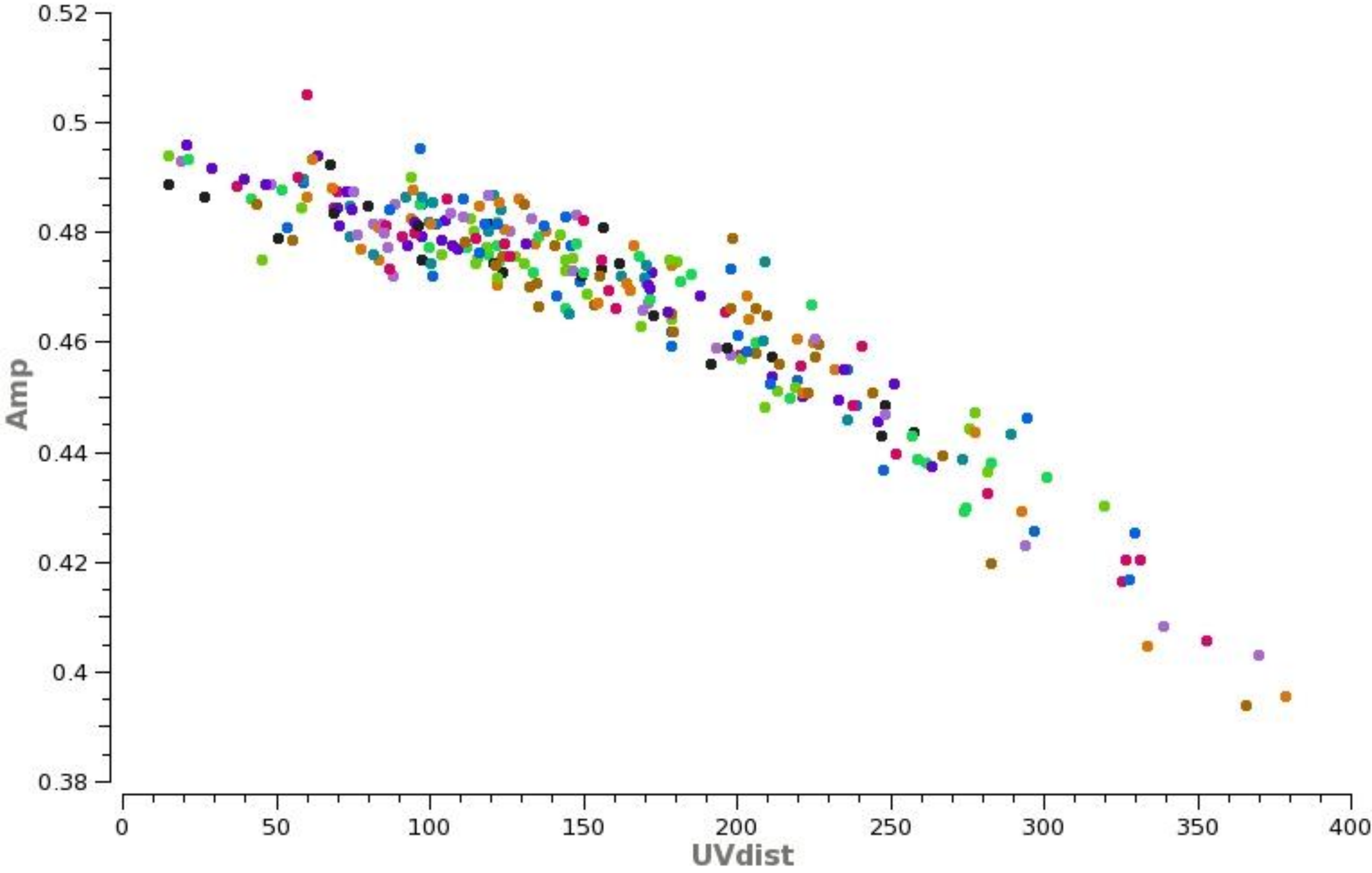
# Amp-Calibrators Amp vs. uv-distance (Before)



# Amp-Calibrators Amp vs. uv-distance (Model)

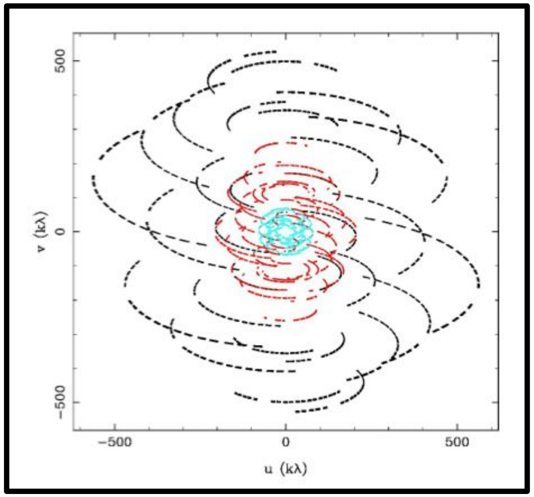


# Amp-Calibrators Amp vs. uv-distance (After)



# The Dirty Beam

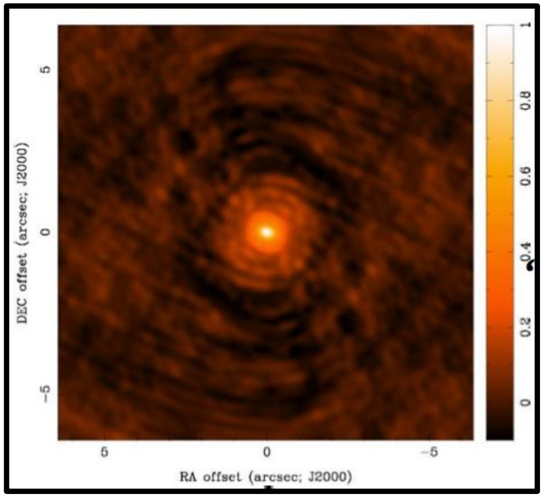
S(u,v)



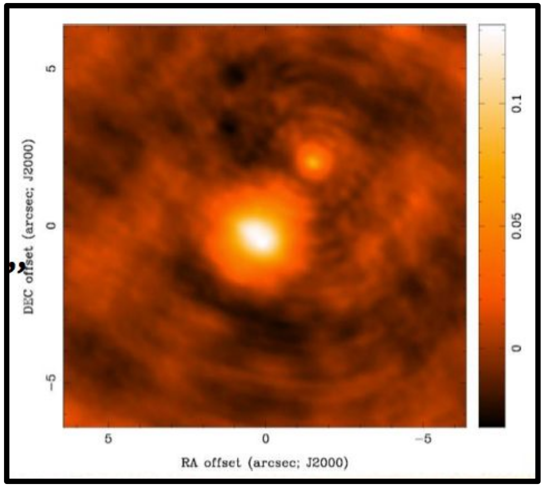

FT

s(x,y)

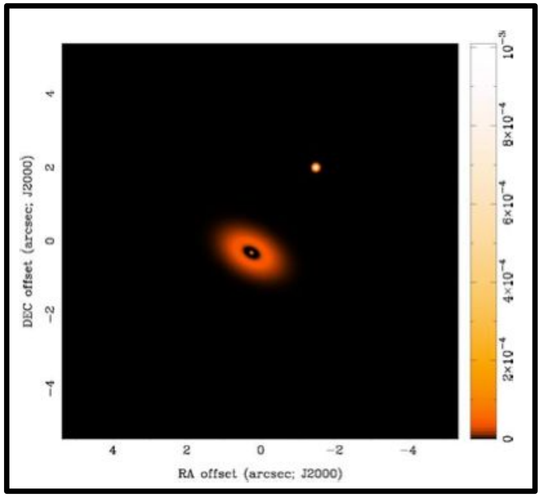
"Dirty Beam"




* (Convolution)




$T_D(x,y)$ "Dirty Image"




T(x,y)

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# Deconvolution Philosophy

- use non-linear techniques to interpolate/extrapolate samples of *V(u,v)* into unsampled regions of the *(u,v)* plane

(remove sidelobes of the dirty beam from the image)

- aim to find a sensible model of *T(l,m)* compatible with data
- requires *a priori* assumptions about *T(l.m)* to pick plausible "invisible" distributions to fill unsampled parts of *(u,v)* plane
- main assumption: real sky does not look like typical dirty beam
- "clean" deconvolution algorithm (and its variants) by far dominant in radio astronomy, though there are others in use
- a very active research area, e.g. compressed sensing

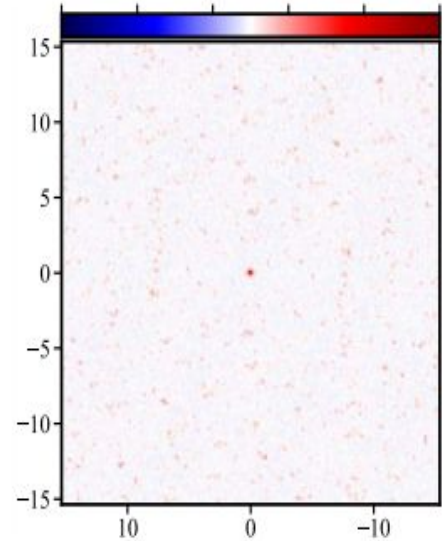
# Classic Högbom (1974) clean Algorithm

- *a priori* assumption: *T(l,m)* is a collection of point sources

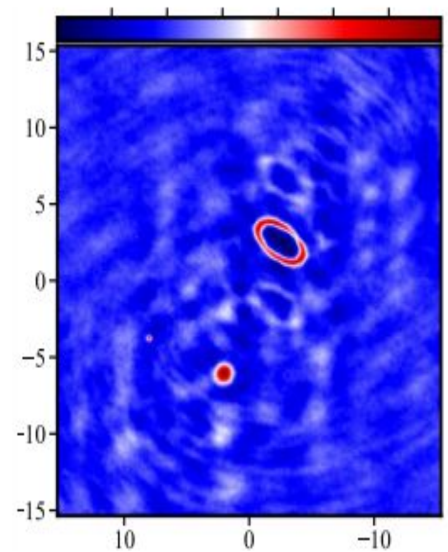
initialize a *clean component* list

initialize a *residual image* = dirty image

1. identify the highest peak in the *residual image* as a point source
2. subtract a scaled dirty beam *s(l,m)* x "loop gain" from this peak
3. add this point source location and amplitude to the *clean component* list
4. goto step 1 (an iteration) unless stopping criterion reached




*s(l,m)*




$T^D(l,m)$

# Classic Högbom (1974) clean Algorithm

- stopping criterion
  - *residual map* maximum < threshold = multiple of rms , e.g. 2 x rms (if noise limited)
  - *residual map* maximum < threshold = fraction of dirty map maximum (if dynamic range limited)

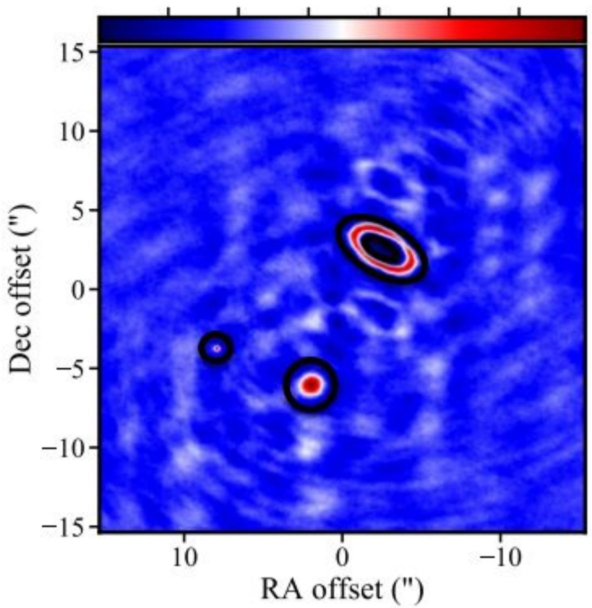
- loop gain parameter
  - good results for g=0.1 (CASA `tclean` default)
  - lower values can work better for smooth and extended emission

- don't "overclean" to artificially low noise level
  - generally a problem only when *(u,v)* coverage is sparse

# Classic Högbom (1974) clean Algorithm

- finite support
  - easy to include *a priori* information about where in the dirty map to search for *clean components*
  - implemented as image masks or clean boxes; CASA `tclean` "mask"
  - very useful, often essential for best results, but potentially dangerous
  - use with care

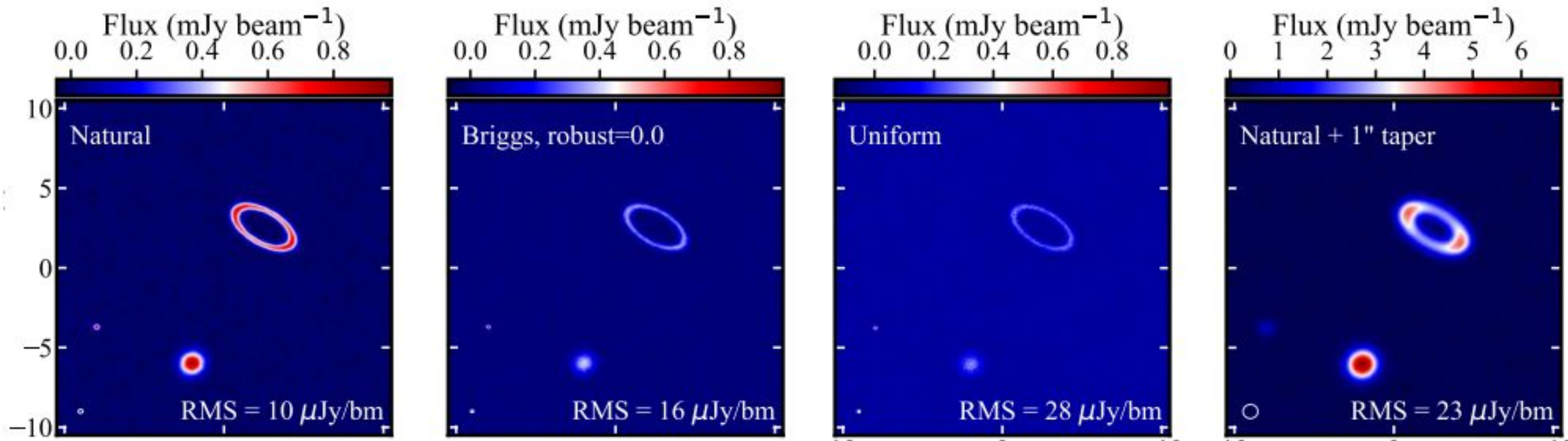
  - can be an arduous manual process; automasking under development



# Classic Högbom (1974) clean Algorithm

- last step is to create a final "restored" image
  - make a model image with all point source *clean components*
  - convolve point source model image with a "clean beam", an elliptical Gaussian fit to the main lobe of the dirty beam (avoids super-resolution of the point source component model)
  - add back *residual map* with noise and structure below the threshold
- restored image is an estimate of the true sky brightness *T(l,m)*
  - units of the restored image are (mostly) Jy per clean beam area = intensity, or brightness temperature

# Results from Different Weighting Schemes




| natural | robust=0 | uniform | natural + 1” taper |
| --- | --- | --- | --- |
| 0.29x0.25 p.a. -81 | 0.19x0.17 p.a. -78 | 0.17x0.15 p.a. -87 | 0.93x0.88 p.a. -86 |

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# Good Future References

Thompson, A.R., Moran, J.M., Swensen, G.W. 2017 “Interferometry and Synthesis in Radio Astronomy”, 3rd edition (Springer)
http://www.springer.com/us/book/9783319444291

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# Good Future References

| |
|---|
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| Examples of UV coverage from Ian Czekala<br>https://drive.google.com/file/d/1fy3edrJNATo175WopB49-3mZ7QeZPK5O/view |

NRAO

**www.nrao.edu**
**science.nrao.edu**

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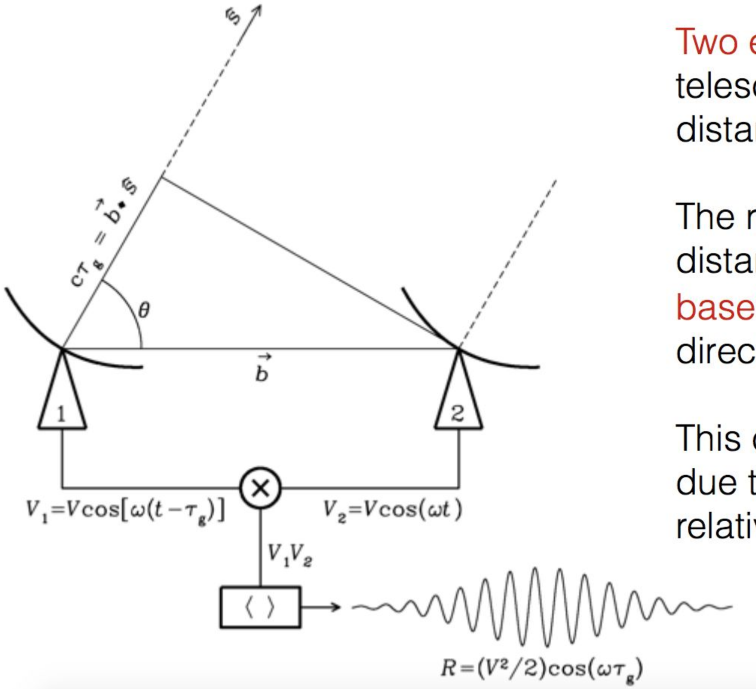



Two element interferometer: Two identical telescopes observe the electric field of some distant source (c.f. Young's double slit).

The radiation to antenna 1 travels an extra distance $\vec{b} \cdot \hat{s} = b \cos \theta$, where $\vec{b}$ is the vector baseline length and $\hat{s}$ a unit vector in the direction of the source.

This can be expressed as a geometric delay due to the projected position of the source, relative to the baseline of the antennas.

$$\tau_g = \vec{b} \cdot \hat{s}/c$$

For a quasi-monochromatic interferometer (responds to a narrow frequency range ν = 2π / λ), the output voltages over time *t* from the two antennas are,

$$V_1 = V\cos[\omega(t - \tau_g)] \quad \text{and} \quad V_2 = V\cos(\omega t)$$

The correlator multiples the voltages from the two antennas together to give,

$$V_1 V_2 = V^2 \cos[\omega(t - \tau_g)] \cos(\omega t) = \left(\frac{V^2}{2}\right) [\cos[2\omega t - \omega\tau_g] + \cos(\omega\tau_g)]$$

and then a time average [$\Delta t \gg (2\omega)^{-1}$] to remove the high frequency component to give,

$$R = \langle V_1 V_2 \rangle = \left(\frac{V^2}{2}\right) \cos(\omega\tau_g)$$

Uncorrelated noise from gain variations within the receivers, the atmosphere and radio frequency interference does not correlate (advantage over single dish measurements).

The output voltage *R* varies sinusoidally with the change of the source direction in the interferometer frame, i.e. the delay changes. These sinusoids are called fringes, and we can define the fringe phase as,

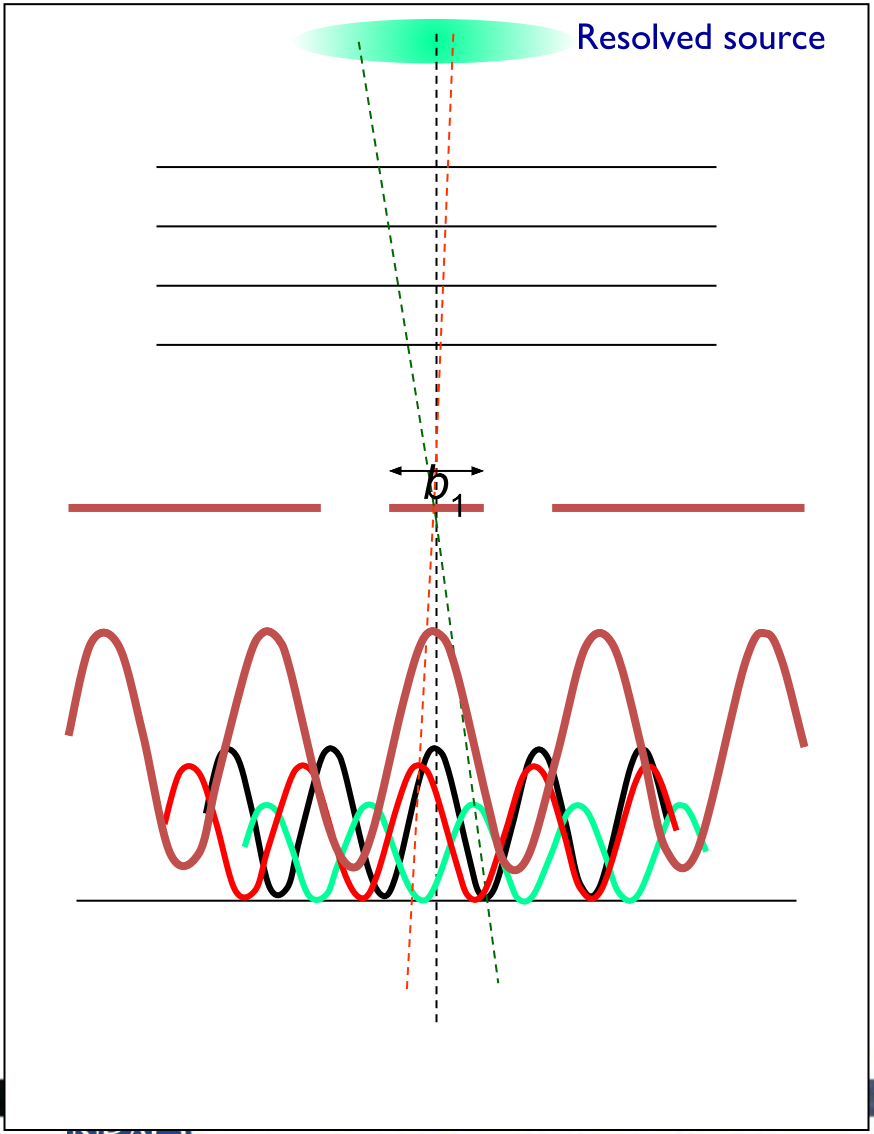
$$\phi = \omega\tau_g = \frac{\omega}{c} b \cos\theta \quad \text{and} \quad \frac{d\phi}{d\theta} = \frac{\omega}{c} b \sin\theta = 2\pi \left(\frac{b \sin\theta}{\lambda}\right)$$

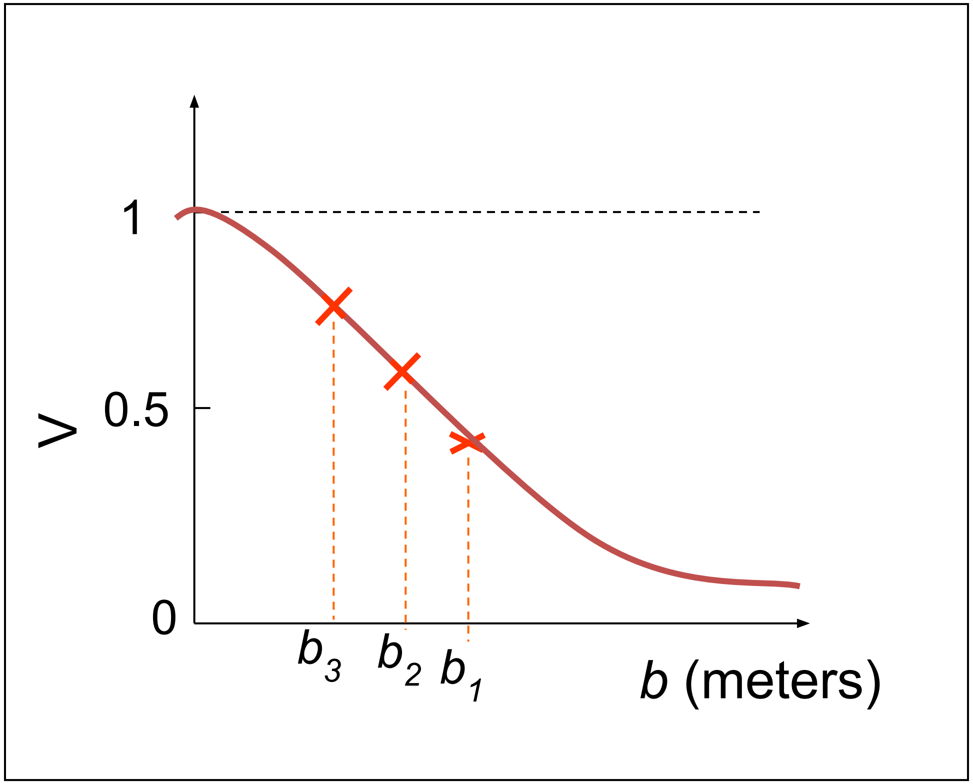
The fringe period ($\Delta\phi = 2\pi$) corresponds to an angular change of $\Delta\theta = \lambda / (b \sin\theta)$, and so, for large *b*, interferometers can measure very accurate positions of sources (typically $\sigma_\theta \sim 10^{-3}$ arcsec).

As the source(s) moves across the sky, the response of the interferometer changes because the geometric delay changes. The maximum in the fringe pattern occurs when $\tau_g c$ is an integral number of wavelengths (similar to the Young's double slit).

# Visibility and Sky Brightness

Graphic courtesy Andrea Isella

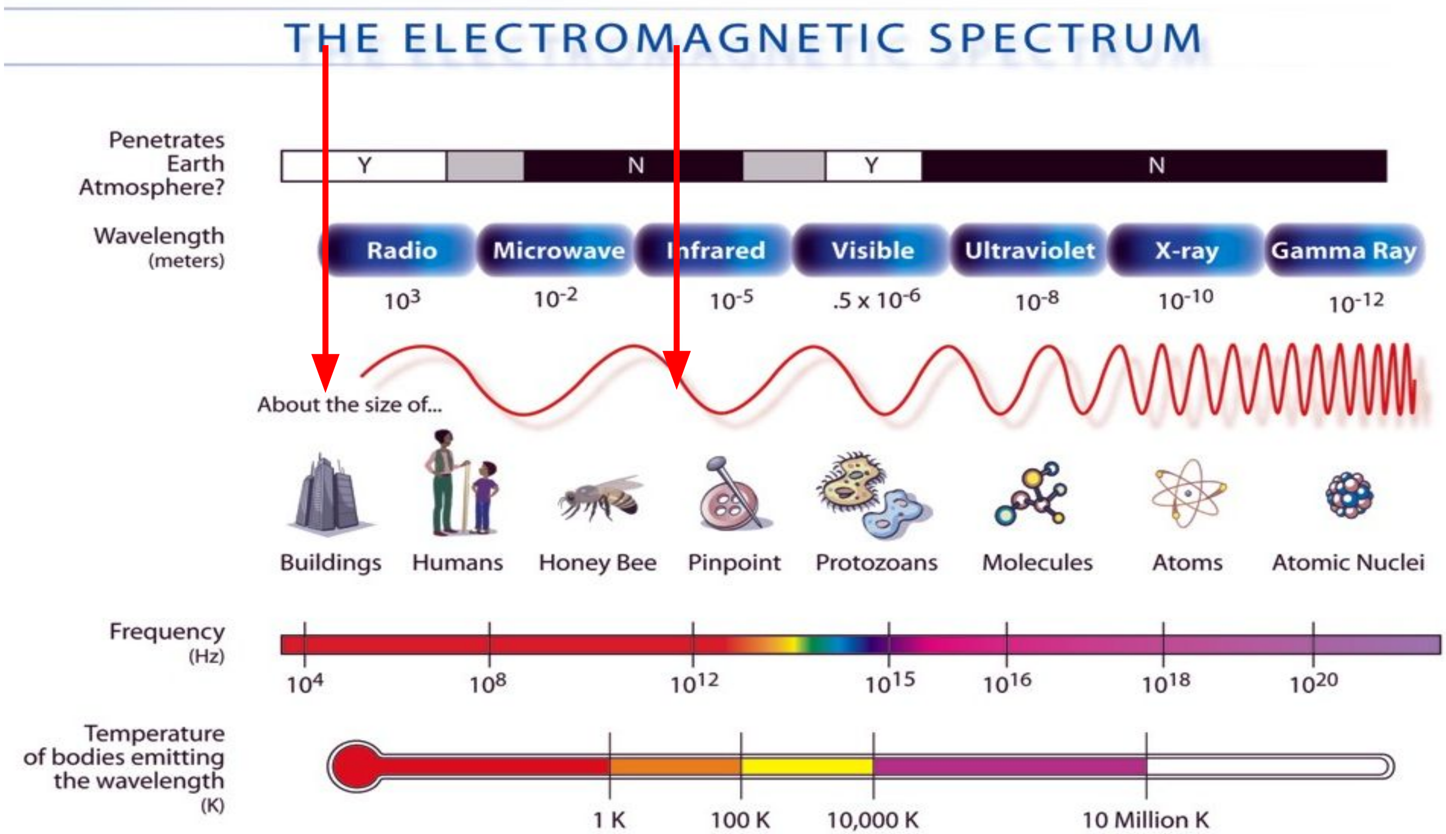



$$|V| = \frac{I_{max} - I_{min}}{I_{max} + I_{min}} = \frac{Fringe\ Amplitude}{Average\ Intensity}$$

# Radio Astronomy

Now used to refer to most telescopes using heterodyne technology



# What is heterodyne?

In a heterodyne receiver, observed sky frequencies are converted to lower frequency signals by mixing with a signal artificially created by a Local Oscillator. The output can then be amplified and analyzed more easily while retaining the original phase and amplitude information.

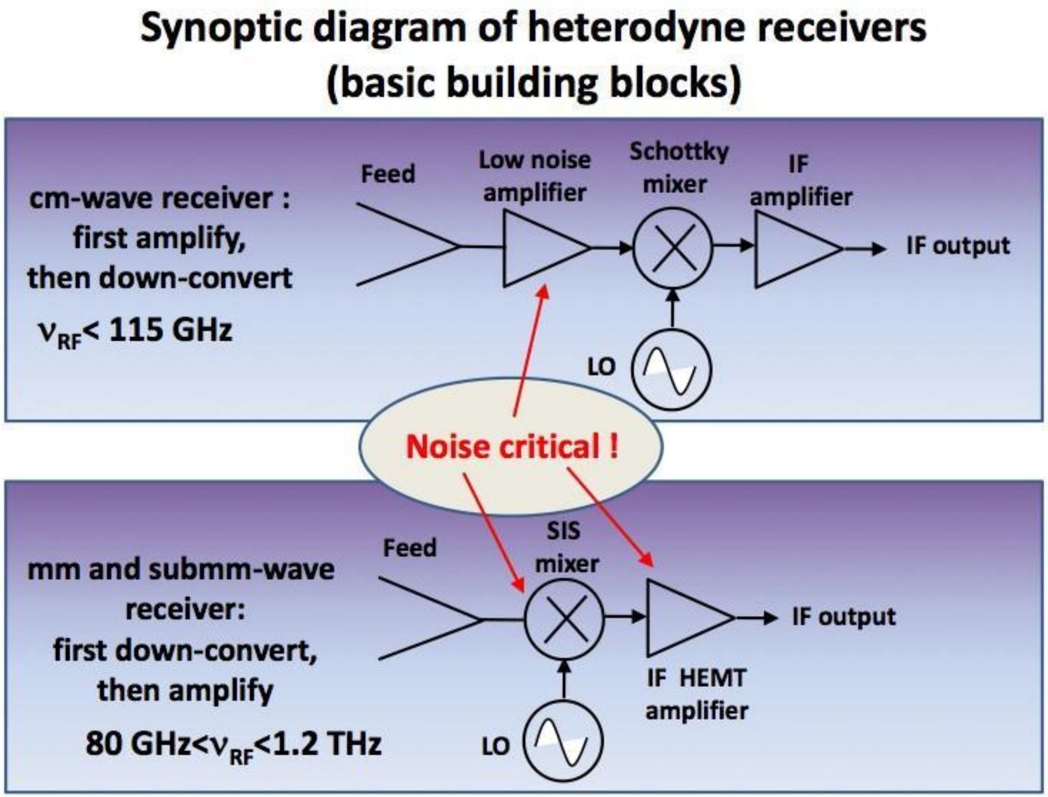
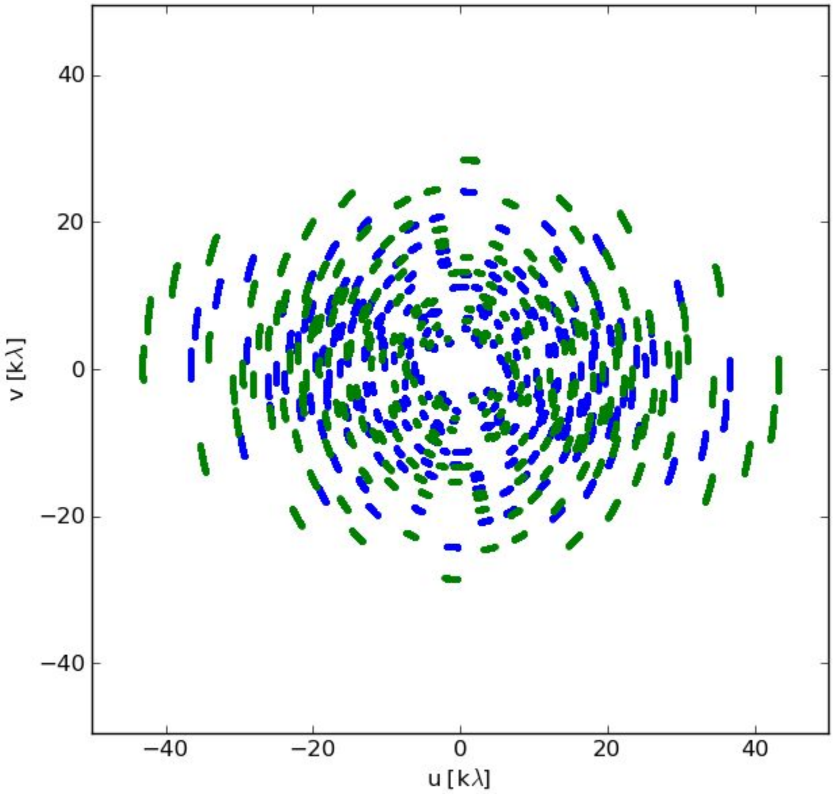
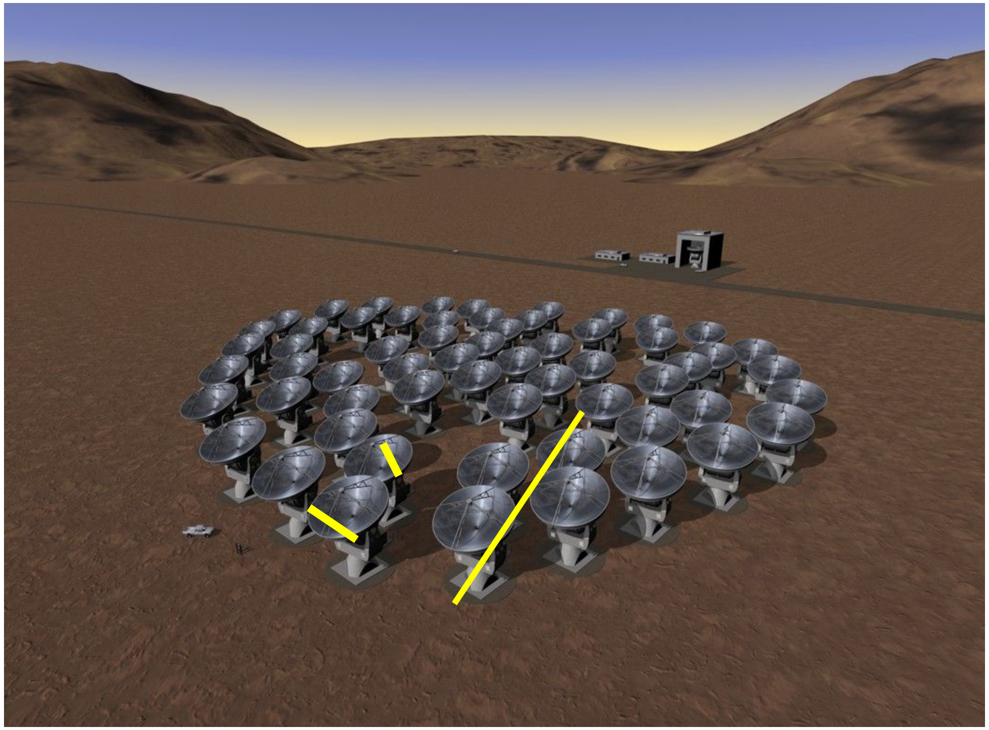



Image from Alessandro Navarrini (IRAM)

# Sampling Function

Each antenna pair samples only one spot; the array cannot sample the entire Fourier/uv domain resulting in an **imperfect image**



Small uv-distance: short baselines (measure extended emission)
Long uv-distance: long baselines (measure small scale emission)
Orientation of baseline also determines orientation in the uv-plane

# uv coverage: why the central hole?

- The central hole in the sampling of the uv plane arises due to **short baselines.**
- The largest angular scale that an interferometer is sensitive to is given by the shortest distance between 2 antennas.
- The field of view is given by the beam of a single antenna.
- A single antenna diameter will always be < the shortest distance between two antennas.
- So the field of view is always > the largest angular scale
- If your source is extended, you will always have some flux at short spacings (i.e. extended emission) that is not recovered.
- **Solutions:** We can extrapolate to these shorter spacings after our observations are taken or we can fill in the information with 7m observations or ultimately single dish data.

# A Brief Word on Calibration

## Calibration requirements (Handled by ALMA):

**Gain calibrator**
Bright quasar near science target
Solves for atmospheric and instrumental variations with time

**Bandpass calibrator**
Bright quasar
Fixes instrumental effects and variations vs frequency

**Absolute flux calibrator**
Solar system object or quasar
Used to scale relative amplitudes to absolute value