Observation Preparation



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Atacama Large Millimeter/submillimeter Array
Expanded Very Large Array
Robert C. Byrd Green Bank Telescope
Very Long Baseline Array



Outline

- All-time, general observation preparation rules
- The EVLA is not the VLA!
 - Receivers and bandwidth
 - Spectral line correlator
 - Dynamic scheduling
- Detailed observation preparation hints
 - Dynamic observing considerations
 - Logistical considerations
 - Tactical considerations
 - Special topics
- Scheduling blocks in the OPT
 - New 2x 8 subband (2x | GHz maximum) wide OSRO resource catalogs
 - Check, check, check!
 - Quick run-through



All-time, general observation preparation rules:

• We receive lots of general questions which usually must be answered specifically for the science pursued, boundary conditions imposed and resources used (and which typically are not given along with the question ..)

Many questions don't have a straight general answer or solution!

- Remember during this talk
 - Everything is a <u>trade-off</u> and a zero sum game. Answers, solutions or tricks depend on <u>your</u> science goal, available flexibility to achieve that goal, and sensitivity to changing conditions. What may work for one may not work for another similar case, vice versa.
 - Simpler observations (setups, strategies, ..) are much easier to design and process
 - Choose so that you can <u>calibrate</u> your data versus extending on your on-source time;
 if you can't calibrate the data, your on-source time will most likely be lost anyway
 - There are tools to help planning planning increases the ability to do the right thing when observing and this planning preferably starts <u>before</u> proposing!
 - Planning tools may not accurately reflect the conditions during your observations



The EVLA is not the VLA!

- The EVLA has increased frequency coverage
 - Sensitivity to radio frequency interference (RFI) and the art (i.e. pain) of editing
- The EVLA has increased instantaneous bandwidth
 - Wide fractional bandwidth effects
 - Calibration (pro 's and con's)
 - Imaging
 - Subband boundaries, inhomogeneous spectral response over the observed bandwidth
- The EVLA has a complex correlator
 - Large data set sizes, always spectral line, possibly with complex data structures
 - May need to process in subsets or on computing clusters
- The EVLA primarily uses dynamic scheduling for flexibility
 - Each scheduling block must include a range of scheduling constraints
 - LST start range
 - Weather (atmospheric phase noise and wind limits)
 - Anticipate and protect against unknown startup slew time and cable wrap
 - Anticipate and protect against availability of calibrators/slew for different starting times

Detailed observation preparation hints

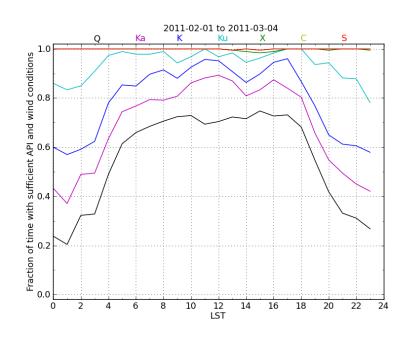
- Dynamic observing considerations:
 - Things to realize and anticipate when you don't know when your observation starts
- Logistical considerations:
 - When to do what in your SB to make your observation succeed
- Tactical considerations:
 - How to get your low priority SB more likely to be observed (no guarantees though!)
- Special topics:
 - Calibration and calibrators
 - Wideband issues
 - Spectral line issues
 - Shadowing

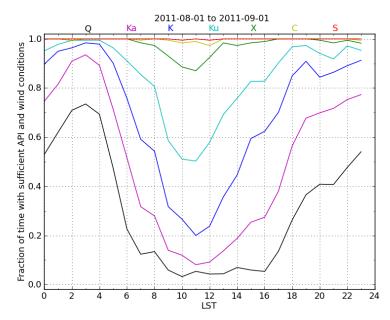
Dynamic observing considerations

- Before proposing consider
 - Where is my source on the sky
 - Popular LST ranges are more competitive (see pressure diagrams of previous cycles)
 - Can your science be done with a different source in a different, less popular LST range?
 - What observing frequencies does my science require
 - Higher frequencies (above ~15 GHz) are harder to observe and require more careful planning at every stage
 - Can your science be done with lower frequencies, or with less time in high frequency observing?
 - Does my target occur during day/night (etc) during this proposal cycle
 - Observing conditions are different for day/night and in spring/fall, etc.
 - Maintenance reduces the available observing time during the day and is thus always more competitive
 - Can your science be deferred to a different season when better conditions are available (see later)?
 - Where in this cycle is the array configuration scheduled that my science needs
 - e.g., A-array configuration conditions in early summer and in late summer are different, and priority determines when it will run, but there is no advance indication when this will be (and if they all are in the same season)
 - Can your science be done in a different array or the same array in a different season?
 - When all above is less favorable for my case, do I still propose for this cycle
 - Even when approved with a very high science rating, your program may not complete, or not be observed at all
 - This depends on observatory conditions, boundary conditions you set, competition from others, ...
 - You may get a review that will help improve your proposal for a next proposal cycle (still without guarantees)



High frequency conditions (pre-proposing)

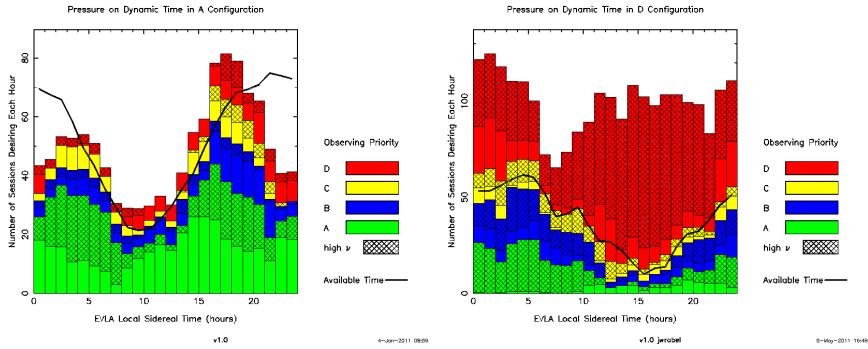




- February versus August high frequency observing conditions:
 - ✓ Observing LST 06-18h good in February, even daytime LSTs
 - ✓ Conditions LST 08-16 bad in August (daytime ~2 Q-days, maint!)
 - ✓ Below ~15 GHz almost always possible



LST pressure diagrams (post-proposing)



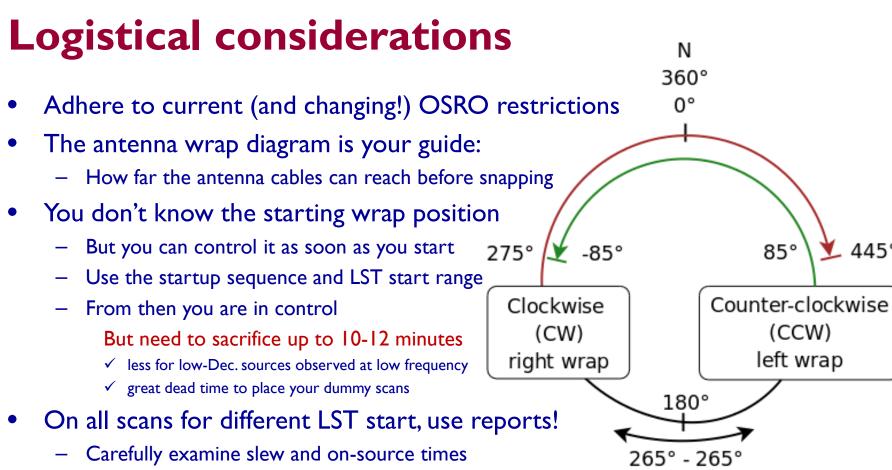
- IIA (A configuration) versus IIB (D configuration):
 - ✓ Daytime impact on available observing time
 - ✓ Popular versus unpopular LST ranges
 - ✓ Competition for time and weather conditions



Dynamic observing considerations

- When making the schedule consider
 - Where are my calibrators and target sources on the sky (if not done before proposing)
 - Strong, point like calibrators are good for bandpass and delay calibration
 - Nearby calibrators are good for phase calibration
 - Pointing calibrators should be in the same general region of the sky (Az, El)
 - Near Declination +34° is unfortunate for wraps choose calibrators on the same side as target
 - What is the limited choice of flux calibrators for this project
 - Can certain calibrations be combined with fewer calibrator sources
 - What trade-offs can be made, and at what cost/gain
 - What competition from other projects can I expect
 - Lower priority projects are doomed to fill up the gaps left by higher priority projects and/or use less favorable weather conditions, and therefore must be very flexible in scheduling constraints. If, e.g., your priority is a low C and needs 50h on a single target to do the science, it is not a shame not to submit anything (other than an improved proposal with more realistic goals at the appropriate next deadline).
 - Short scheduling blocks are good, but inefficient ,and impact uv-coverage and parallactic angle range
 - Less tight weather constraints are good, but may yield lesser quality data
 - Plan for observations at the end of the cycle, but submit scheduling blocks at the start
 - How do I keep data reduction simple
 - EVLA may produce huge data sets, which may be separated by months in different array configurations
 - Where do I place my subbands to minimize RFI (continuum) or maximize use of spectral coverage (line), etc.
 - Please follow guidelines for "standard observing modes" to enable application of pipeline calibration



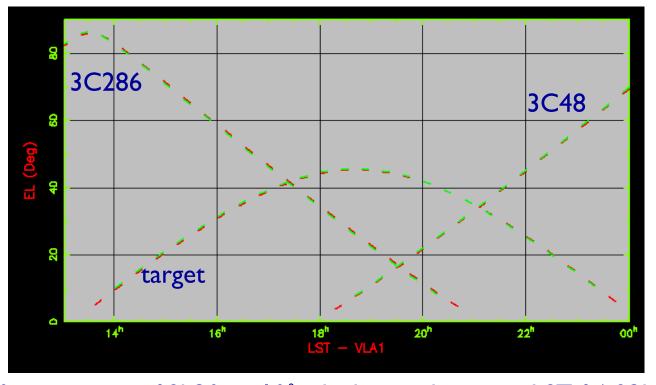


- Carefully examine slew and on-source times
- Carefully examine each azimuth and elevation, shadowing
- Play around with placing flux (etc) at start or end

Especially catching 3C286 and 3C48 at rise/set (or zenith) causes trouble (Dec. ~34°)

You may consider submitting separate scheduling blocks, with caution!

Example



- A 2-hour block for a target at 18h30m, -10° which is up between LST 14-23h
 - Early LST: flux calibrator 3C286, but cannot be used at late LST
 - Late LST: use flux calibrator 3C48, but cannot be used at early LST
- Possible solutions:
 - Choose to <u>start</u> with 3C286 and limit LST range to 14-19 (or less for a minimum elevation)
 - Choose to <u>end</u> with 3C48 and limit LST range 18-21 (or less for a minimum elevation)
 - Submit either block, or submit both! (but specify a note to operator; limited time allocation)

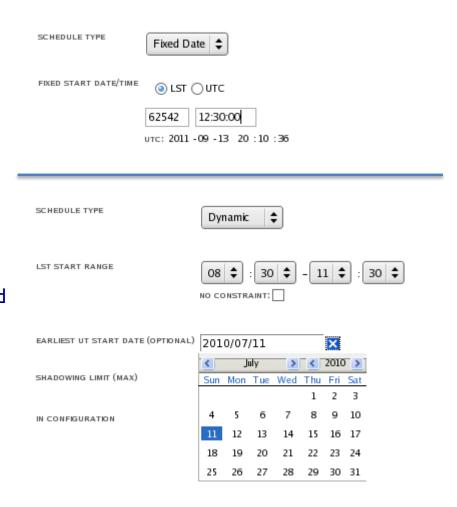


Tactical considerations

- If you have very high grades (low numbers) from the <u>science review panel</u> and a high scheduling priority (A) from the <u>time allocation committee</u>:
 - You can schedule whatever, whenever (within reason), mostly only care about logistics
- If you have priority A or B but are in a popular LST range or require good high frequency observing weather you have to be more flexible:
 - Make shortish blocks (up to ~3h) perhaps with non-overlapping LST ranges
 - Relax the weather conditions somewhat, especially if you can use self-calibration
 - More flexible blocks are more likely to fill a gap in the "triple-A" rated scheduling blocks
- If you have low grades and low priority (**C**), or intermediate priority (**B**) and require popular LST ranges in addition to high frequency conditions
 - Make your blocks as short as your science allows, perhaps only 30 minutes long
 - Consider observing the extreme LST ranges/elevations and weather
 - Make the block 30 seconds shorter than the nominal integer times 30 minutes
 - Submit early to perhaps jump the queue while others are still creating their blocks
 - Consider observing in between different array configurations (move time, ask schedsoc)
 - Pray, or if your science can't be done with partial observations, reconsider or repropose

Observing constraints in the OPT

- Fixed date observations
 - Only use for this purpose!
- LST start range
 - Has to be specified if any of your sources is below Declination +42°
 - The range when your sources are above the elevation limit: from the (start LST time) till the (end LST time minus scheduling block length)
 - Note that the choice of calibrators and where they are placed in the scheduling lock may significantly shorten the LST start range for an observation; consider breaking it up and supply two different scheduling blocks (use caution!)
 - Make blocks a multiple of 30 minutes





Observing constraints in the OPT

- Default wind and API setting
 - Will get you mostly good data for that frequency band (note: empirical limits)
 - Choose highest frequency in the block
- Wind limits
 - If pointing errors aren't your main concern (e.g., flux is less important) you could consider accepting higher winds than normal
- Atmospheric phase constraints
 - If you can self-calibrate your target,
 you don't need the best phase-stability
- Comments to the operator
 - Let them know about linked blocks,
 and when to choose between them
 - Operators cannot modify the scans!

SCHEDULING CONSTRAINTS/CONDITIONS

Description	Wind	Atmospheric Phase Limit
Lowest Frequencies (4, P, and L)	Any	Any
2.0GHz - 4.0GHz (S)	Any	60.0 degrees
4.0GHz - 8.0GHz (C)	Any	45.0 degrees
8.0GHz - 12.0GHz (X)	15.0 m/s	30.0 degrees
12.0GHz - 18.0GHz (Ku)	10.0 m/s	15.0 degrees
18.0GHz - 26.5GHz (K)	7.0 m/s	10.0 degrees
26.5GHz - 40.0GHz (Ka)	6.0 m/s	7.0 degrees
40.0GHz - 50.0GHz (Q)	5.0 m/s	5.0 degrees
Specified Constraints	m/s	degrees

COMMENTS TO THE OPERATOR

This (ID 4948641) is linked to scheduling block ID 1658644 for the same observation using a different LST start range. Please only observe only one of these blocks.



Special topics: calibration

- If you cannot calibrate, your data may end up to be useless!
 - Be conservative, better to spend a bit more observing time on your calibrator
 - Example for a 2-hour run on a single target at 5 GHz (2GHz bandwidth)

 Phase cal. scan 3 min every 30 min: on-source time is 81 min, rms is 4.9 uJy/beam

 Phase cal. scan 4 min every 15 min: on-source time is 66 min, rms is 5.1 uJy/beam
 - Better calibration allows for easier/smoother data reduction
- Calibration is derived for *individual* antennas, baselines and polarizations
 - Look up a current flux density for the calibrator (if available)
 - Calculate for 2 antennas, single polarization and robust weighing in the exposure calculator
 - Use appropriate bandwidth: channel for delay/bandpass, subband for flux/phase
 - Total bandwidth for weak phase cal, provided you have a strong one too (multiband corrected)
 - If you plan a particular calibrator scan length, calculate for 10 seconds less (QUACK)
- You want a good calibration solution; aim for 10-sigma or better
 - Bad data and data editing reduce the anticipated S/N, i.e. don't aim for 5-sigma and get noise
 - The calibrator flux density may be less than anticipated (they typically are outdated)



Special topics: calibrators

- How often, how bright, how far?
 - Remember: everything is a trade-off and depends on your science goal!
- Electronics (instrument) calibration:
 - Should not be (heavily) time dependent and independent of target source
 - (multiband) delay, flux density and bandpass calibration
 - Typically one scan sufficient, though a second may protect against e.g. RFI, missed scan, etc.
 - Bright enough to calibrate, see previous slide.
 - Direction does not matter much (apart for slewing/wraps/pointing)
- Delay calibration allows to use the full bandwidth for weaker phase calibrators
- Flux density bootstrapping
 - Scaling factor to get correlator output into Jansky's per beam
 - Flux doesn't need all antennas; solutions may be better to obtain using standard models (spex)
 - At high frequency maybe correct for opacity by only using scans taken near the same elevation
 - Be realistic in the error propagation, it is very hard to get better than 5% absolute accuracy
 - Attempting 5% or better accuracy requires dedicated flux density calibration and very careful data reduction
 - At high frequencies even typical careful standard data reduction may easily be off by 10-20% or more



Special topics: calibrators

- Atmosphere (gain) calibration:
 - Assumes that atmosphere for target and calibrator behave similarly and that solutions for calibrator (known properties) can be applied to target (unknown properties)
 - Derive solution (how bright?) in a similar patch of sky (how far?) often enough (how often?)
 - How bright? See before; 10 sigma in bandwidth on a single baseline, single polarization
 - Determines the calibrator scan length and impacts (minimum) cycle time
 - How often? Need to track phase in coherence time range 2-60 minutes
 - Depends on frequency, array configuration, elevation, weather: see documentation
 - How far? Need to see same atmosphere range 0-15 degrees, within coherence time!
 - Depends on frequency, array configuration, elevation, weather: see documentation
 - High frequency coherence times usually limit the cycle time to a few minutes, i.e. few degrees
 - Remember that you have to detect your calibrator within the cycle/scan, and you better make sure you do!
 - Cycle time includes: slew, calibration, slew back and target scan, within coherence time!
 - A weaker calibrator may be much better than a brighter distant calibrator: trade-offs!
 - Sometimes there is no nearby choice;
 - Consider to tighten the weather constraints or elevation/LST constraints: trade-offs!
 - Note: L band in A array during TEC activity may need few-minute loops (if no self-cal)



Special topics: pointing

- Pointing calibration:
 - Only needed for high frequencies (~15 GHz and up)
 - Use the predefined pointing setups (in NRAO default resource catalog)
 - Brightness: solution derived for
 - 10 second subscans (but pointing procedure requires 2.5 minutes on-source!)
 - 128 MHz bandwidth (i.e. a single subband in wideband mode, for now)
 - a single polarization (RR and LL determined separately, then averaged)
 - X or C band

so typically your pointing source should (currently) be 100 mJy or brighter (for 10-sigma)

- Pointing should be done in the general Azimut/Elevation direction of the target (~20-30°)
- Movement of the target toward the pointing direction is better than moving away from it
- Do a new pointing if
 - Observing Az/El has changed considerably (tens of degrees) from the last pointing (typically each hour)
 - Temperature or gravitational pull is changing rapidly (observing at sunset/sunrise or near zenith)
 - Moving to or away from your flux (etc.) calibrator (unless it is near your target)
- If your first scans need pointing, the startup time is 2.5 minutes more (~12 minutes)!
 - Must allow for most unfortunate slew to get to your pointing source first, then add 2.5 minutes on source



Special topics: frequencies, wideband

- NRAO default frequencies should be the best trade-off between widest contiguous continuum in the two basebands and the least RFI presence
 - Not guaranteed RFI free: you must be prepared to inspect and edit your data
- At lower frequency bands, the bandwidth ratio is 1:2!
 - Calibrator sources may (will) have a considerable spectral energy distribution (index),
 which is only modeled for the standard calibrators
 - You may be able to derive the spectral index in your observation for free
 - The primary beam and field-of-view for the high-frequency end and the low-frequency end of the observing band will differ up to a factor 2
 - A very strong source in the (first) null at higher frequencies may be prominent and complicate your calibration and imaging at lower frequencies
 - There will be more effects, mostly in imaging, that will not be mentioned here
- We will see more of this when 8 GHz bandwidth observing is enabled



Special topics: frequencies, spectral line

- Spectral line observations need a newly user created resource (see later)
 - Depends on your science, and trade-offs you can/want to make (e.g. easy data reduction)
 - Fundamentally restricted to have a regular (7x) 128 MHz suck-out comb (1GHz baseband)
 - Over 128 MHz per baseband there is thus no contiguous homogeneous spectral response
 - Use interleaving/feathering if your science allows (i.e. trade-off the second baseband)
 - Have to be careful to set up, and check frequency resource reports!
 - Doppler setting tunes the baseband center as required and keeps the comb "frozen" relative to the baseband
- Note: post-processing Hanning smoothing or (diurnal LSR Galactic) Doppler tracking corrections reduce the effective spectral resolution (by a factor 2)
 - Must realize, act and prevent before submitting the scheduling block for observations
 - Use more channels than your science requires (if possible; use half the original bandwidth)
 - Use short blocks so that Doppler setting on different days set the fixed sky frequency; if short enough, the Doppler velocity doesn't vary too much over the block duration and thus the effect over the block can be ignored (not always negilible for narrow channels)
- If you change your resource (e.g. frequency), use bulk edit in the OPT
 - And check, check your reports if it applied the change that you expected



Special topics: shadowing

- Shadowing: part of the collecting area of the 'far' antenna is blocked by the 'near' antenna of the baseline, and thus less radiation will be collected on that particular baseline, resulting in reduced sensitivity on that baseline.
- The OPT reports per scan the maximum shadowing if it occurs
 - Mostly an issue in D array configuration (antennas packed close together)
 - May be an issue for low elevation in the azimuth direction of the array arms (i.e. also in C)
 - For "Any" array configuration, the OPT calculates the worst case scenario (i.e. for D)
 - You can set the limit for which you want to see/ignore the warnings, e.g., 15 meters

Do not panic!

Is not linear

 But anticipate loss for your calibration scans (i.e. use longer scans)

Shadowing	Fraction of area blocked	Baseline sensitivity loss
1 meter	0.01	0.5%
5 meters	0.10	5%
12.5 meters	0.39	22%
18 meters	0.64	40%
25 meters	1	100%



Scheduling blocks and the OPT

- Creating scheduling blocks in the OPT
 - Your source lists should be transferred from the PST, but check the coordinates!
 - Submitted scheduling blocks will be checked by staff for <u>logistics and overall sanity</u>;
 scheduling blocks <u>will **not** be checked for science</u> content, intent and sanity!
 (i.e. get your Doppler frequencies right for all different target sources!)
- Noticed that most of you have used the OPT, so what is new?
 - New 2x 8 subband 2 GHz wide OSRO resource catalogs/capability available to everyone
 - Need to more carefully check reports of resources, and scans
 - Adhere to <u>current</u> (IIB semester) OSRO restrictions
 - They change constantly, but only less restrictive (old blocks remain valid)
 - OPT manual needs updating (especially RCT), but is still a great source of information
 - Also check the general OPT information page (next to the OSRO restrictions page)
 - Quick run-through (details in documentation or the 2010 synthesis imaging workshop lecture)

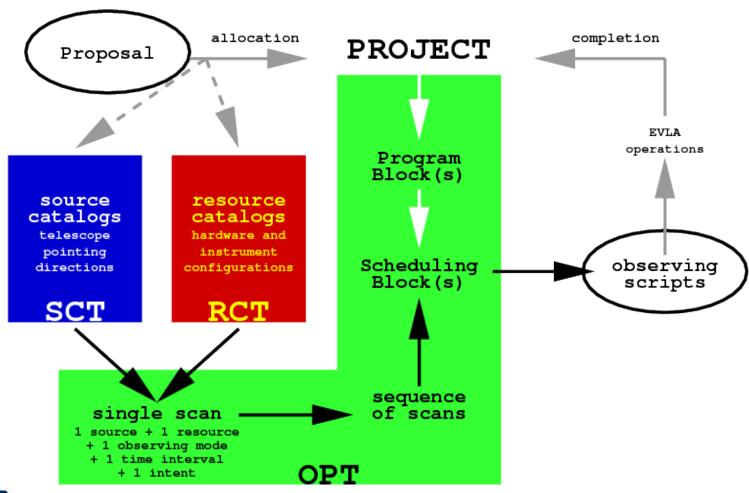


Set up a 2012 OSRO resource in the RCT

- Previous (pre-D 2011) OSRO was similar to common VLA setups
 - 2 IF/baseband/subband tunings, up to 128 MHz, fixed number of channels (polarization)
- Current OSRO has
 - 2 basebands, each covering 1024 MHz and independently tunable
 - There are some restrictions for the higher frequency bands
 - Up to 8 subbands per baseband, each up to 128 MHz
 - All subbands must have a fixed number of channels (depending on polarization products)
 - All subbands in a baseband must be contiguous in frequency (introducing bandpass effects)
 - No subband may not cross any of the 128 MHz suck-out frequencies in a baseband
 - Doppler setting is applied to the baseband; the suck-out comb pattern will move with it
 - Can use the second baseband to interleave the first for homogeneous frequency response (or OSRO2)
 - Be more careful in designing the resource setup (but the RCT tries to keep it simple)
- Later, in 2013, gradually more flexibility will be allowed
 - Earlier opportunities for the RSRO program visitors
- But first a quick refresher of the OPT software (using 2010 SIW slides)

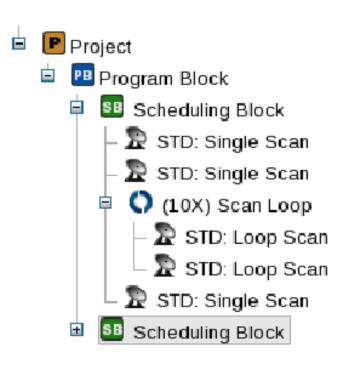


Flow diagram





Example project "tree"



- Project has Program Blocks (PB)
- PB has <u>Scheduling Blocks</u> (SB)
 - Is "observing run" script
 - Sequence of scans and/or (loops of) loops of scans
 - Includes <u>calibration</u>!



Designing a schedule

- Designing a schedule: determine which/what
 - Sources to observe, and for how long (with each setup)
 - Setups ("resources") to use
 - Receivers, (baseband/subband) signals, correlator
 - Calibrations to perform
 - Extra sources/resources? (e.g. pointing)
 - Strategies and tactics to apply (short, long, less quality acceptable?)
 - Optimization and constraints to choose (slew, wrap, weather..)
- Everything is part of a <u>trade-off</u> and depends on the <u>science goals</u>
- Demonstration of how to make a scheduling block (SB) with the OPT



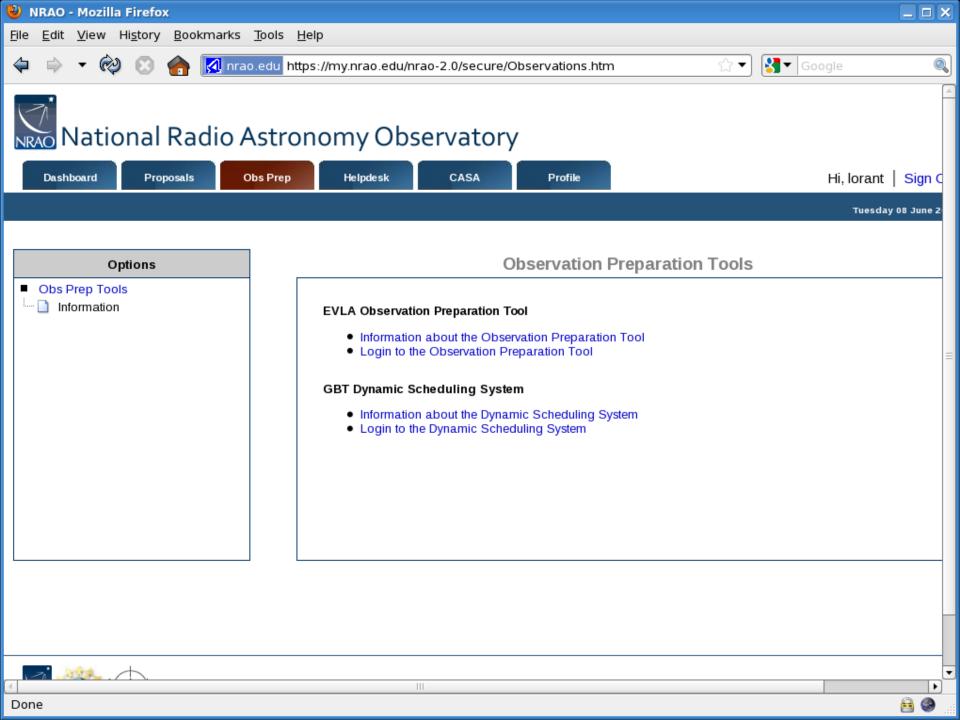
How to start the **OPT**

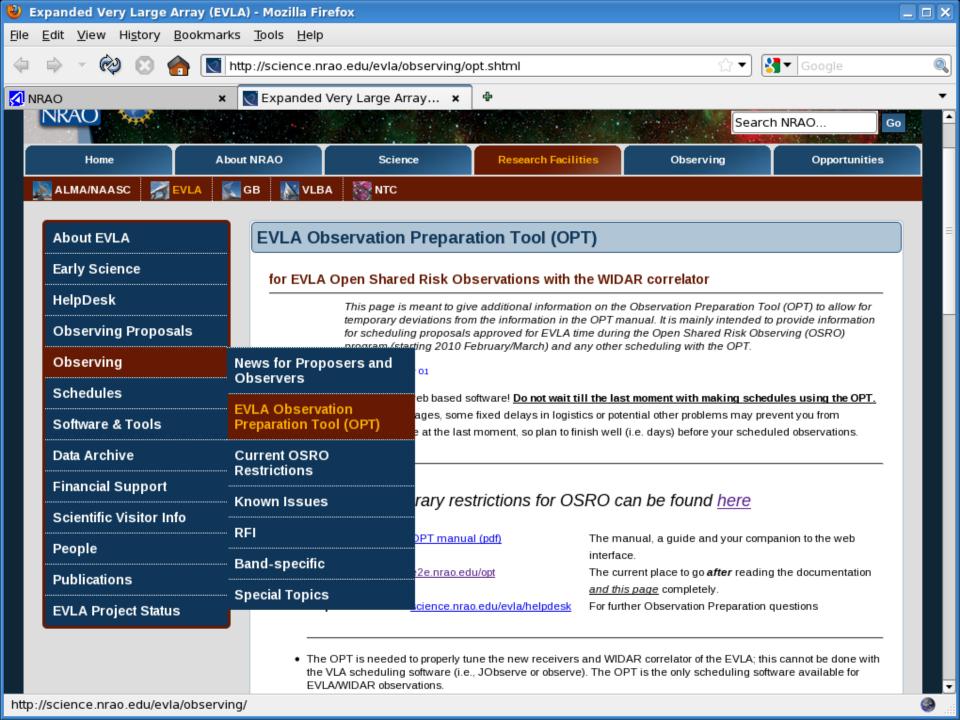
Register at

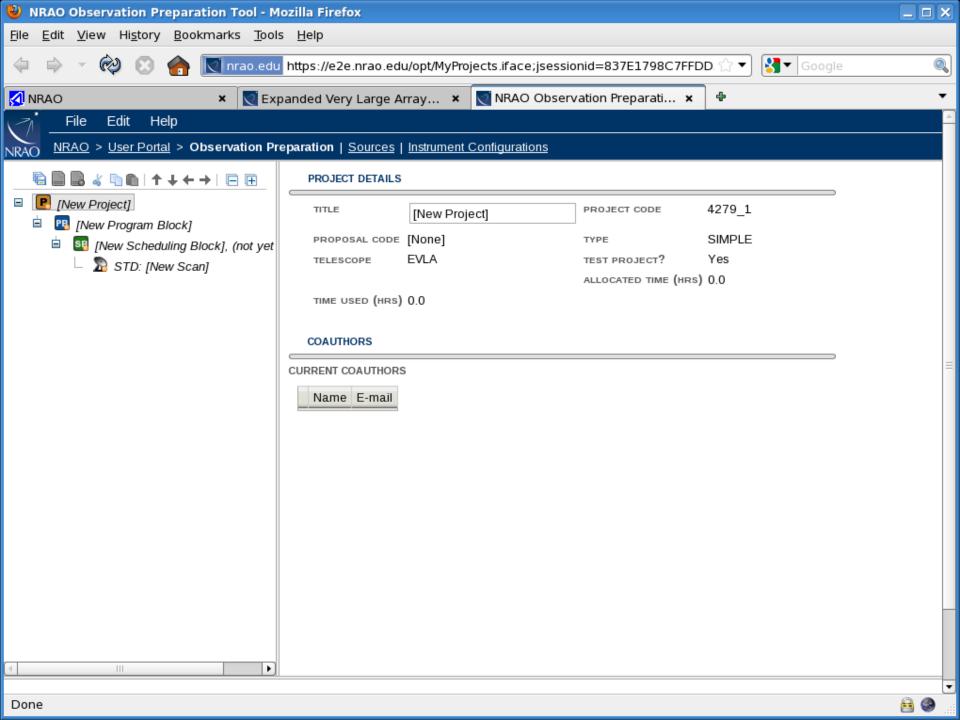
https://my.nrao.edu

- Registration/login/webaddress also good for
 - Proposal submission for EVLA, GBT, VLBA (HSA)
 - (preparing EVLA observations)
 - Proprietary data retrieval (archive.nrao.edu)
 - Obtaining user help (help.nrao.edu)
 - **–** ...

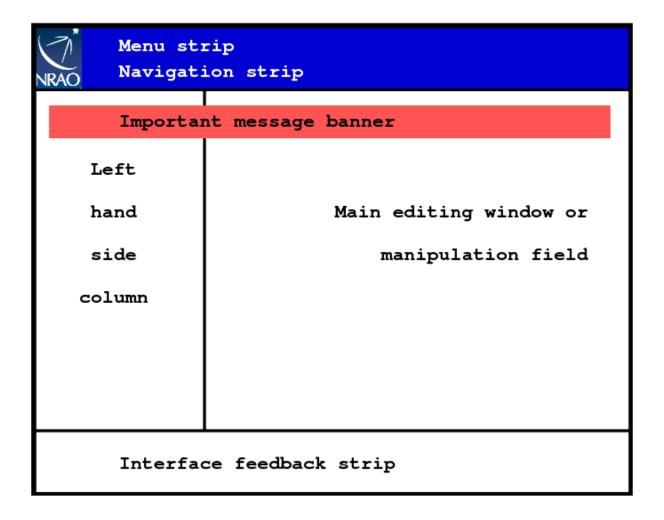








Layout of the OPT





Layout of the OPT

OPT

Icon menu

OPT:

List of projects (PBs/SBs) and scans SCT

Source Search

Tool

Icon menu

SCT:

List of source catalogs and groups RCT

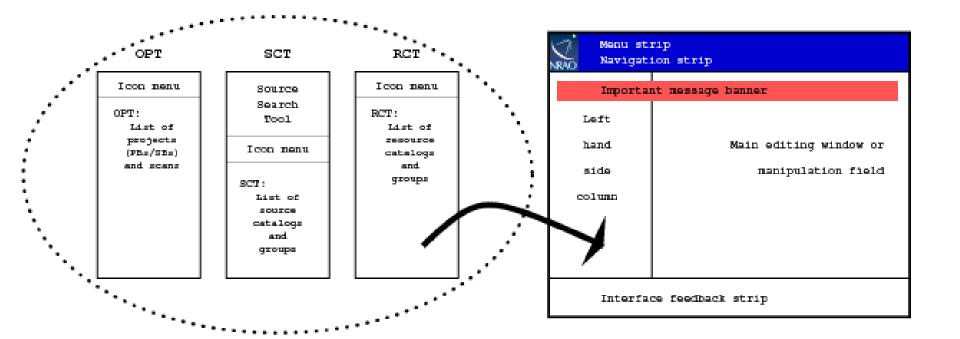
Icon menu

RCT:

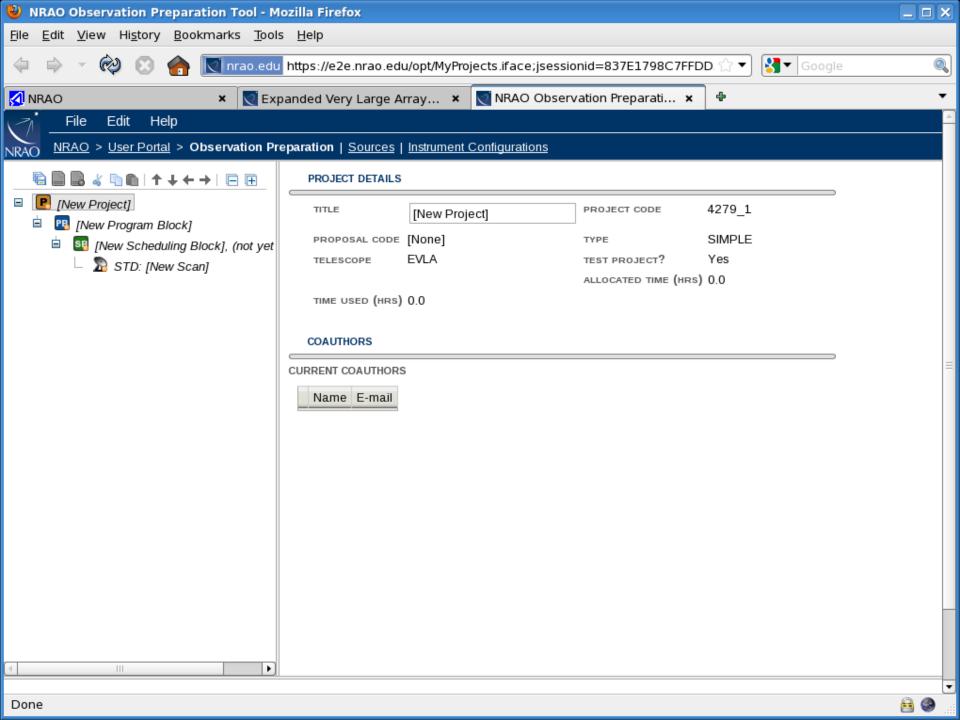
List of resource catalogs and groups

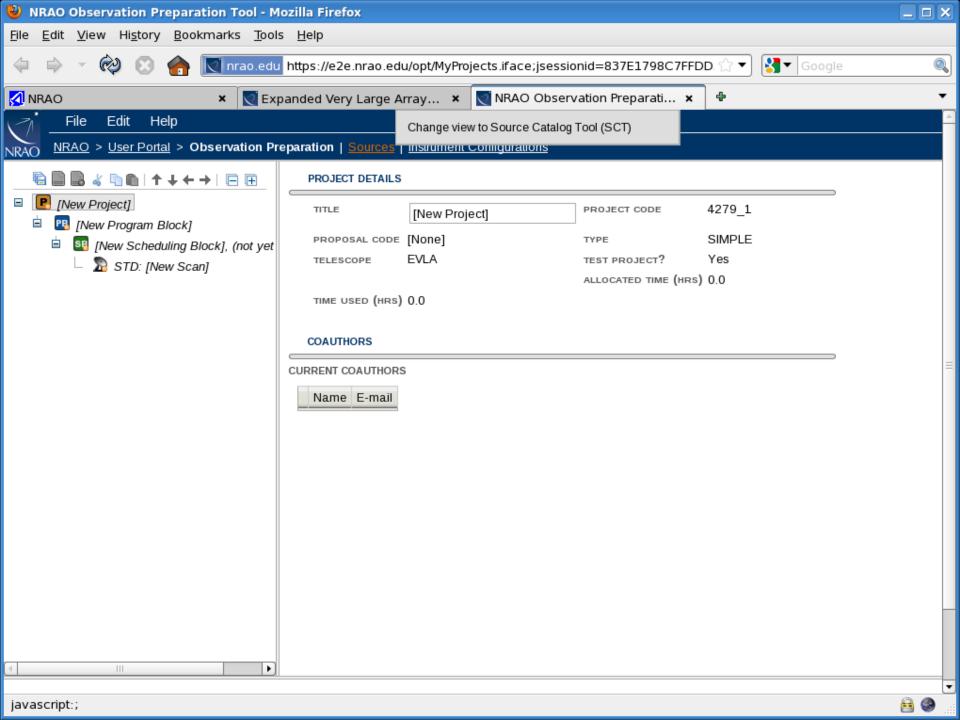


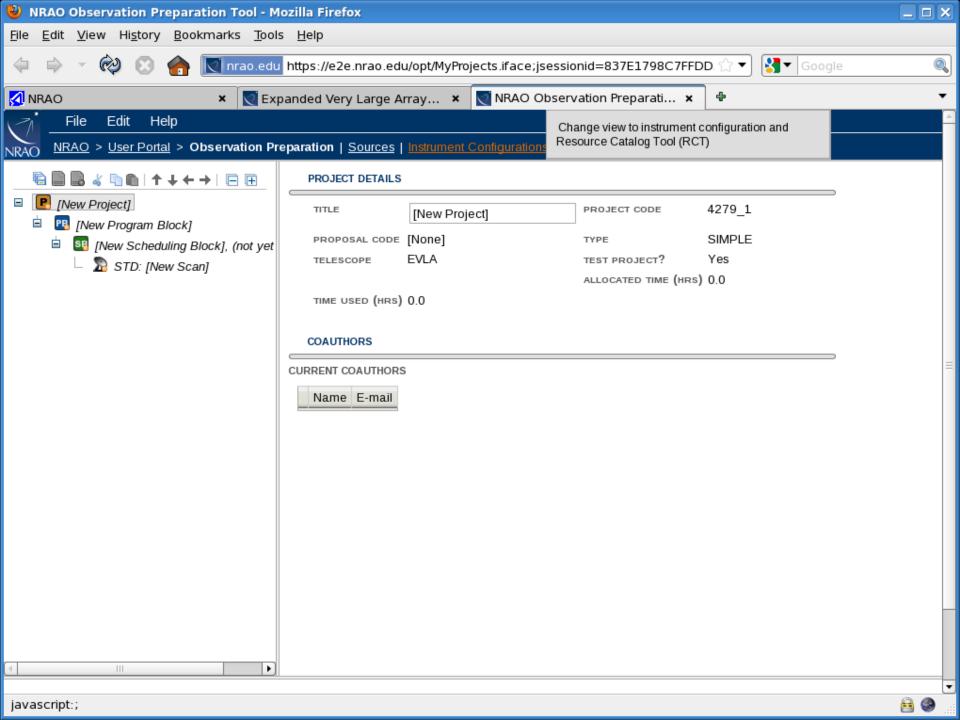
Layout of the OPT

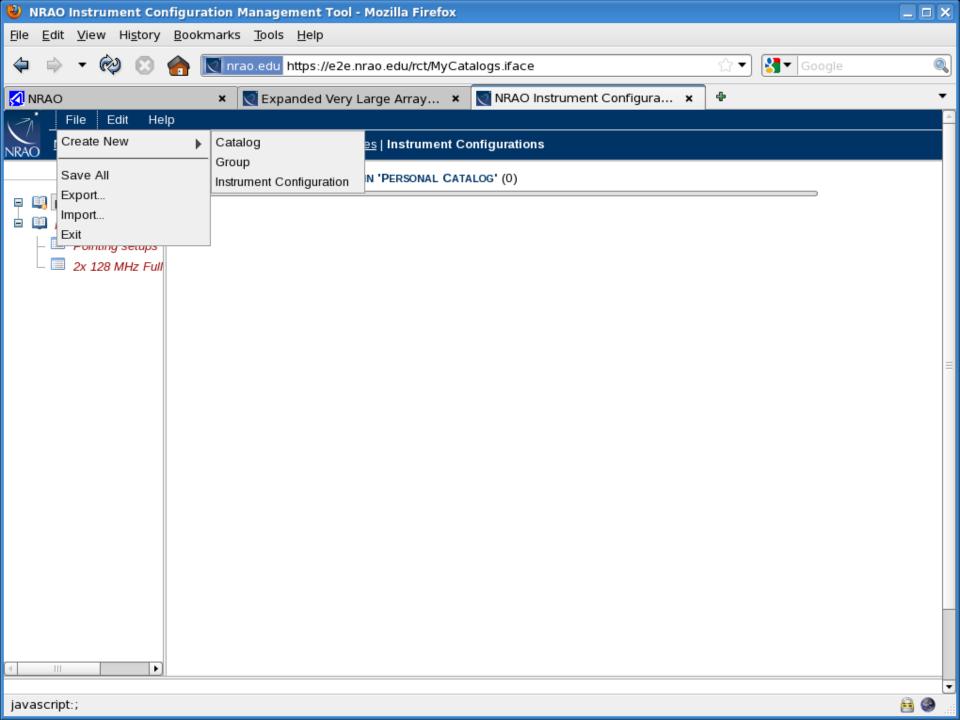


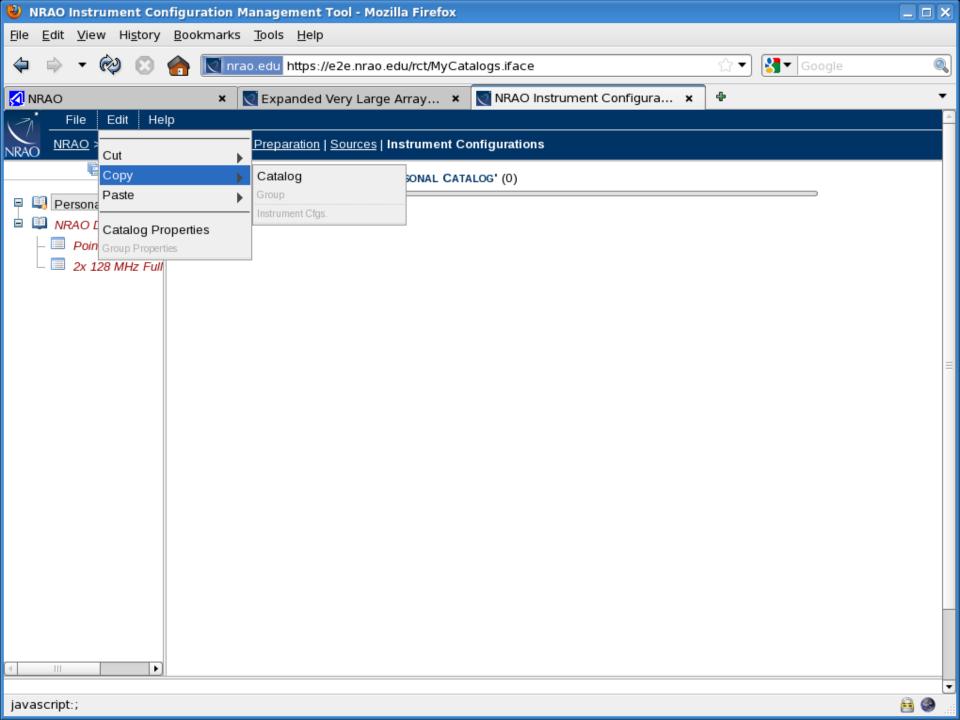


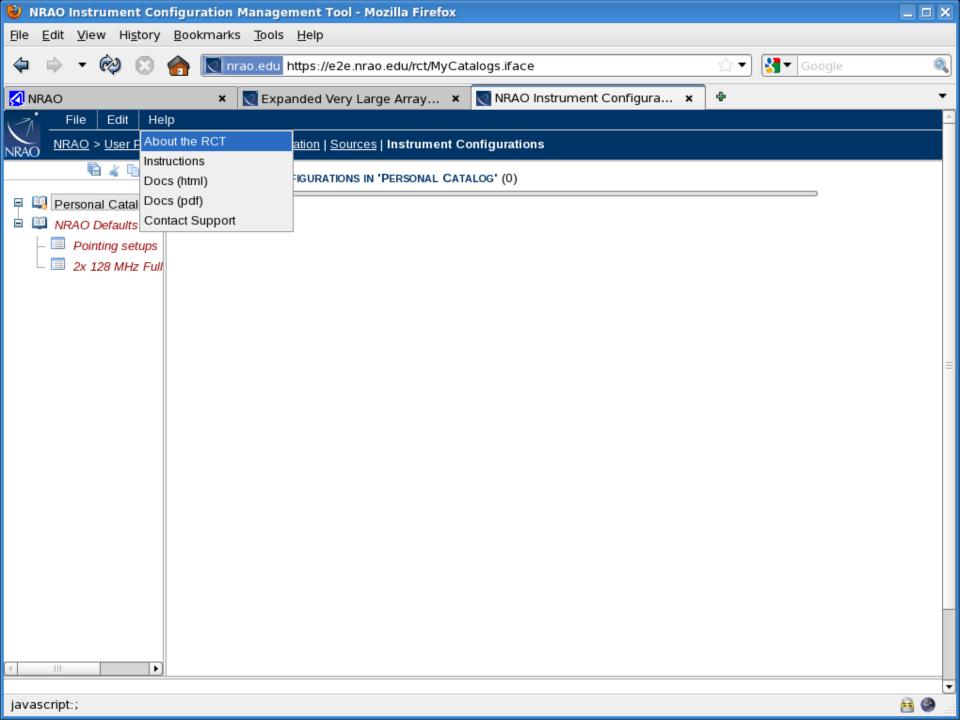


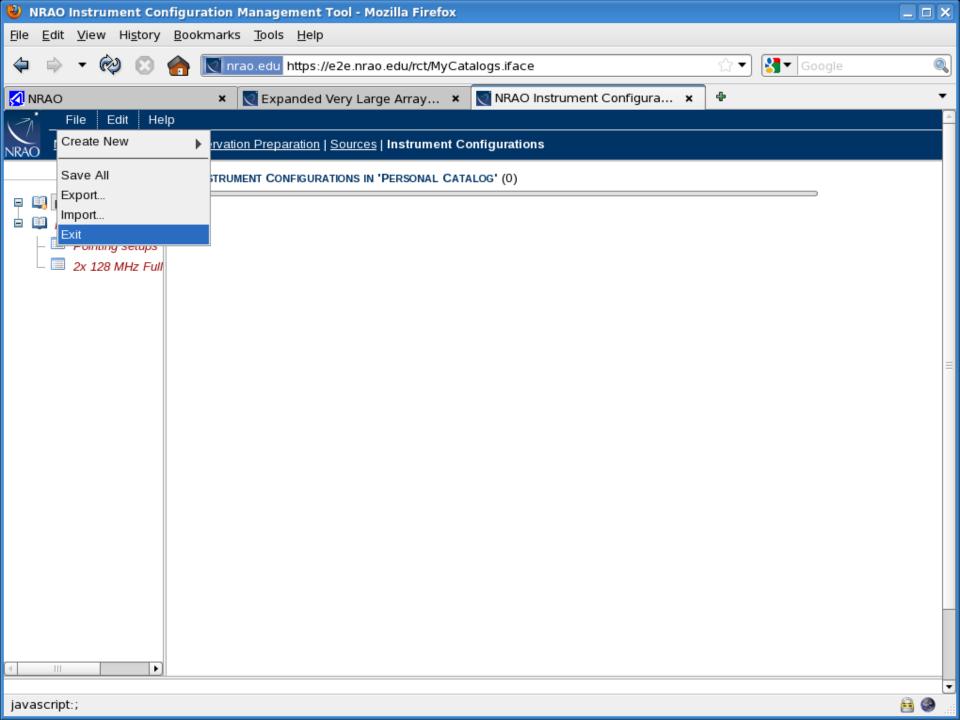








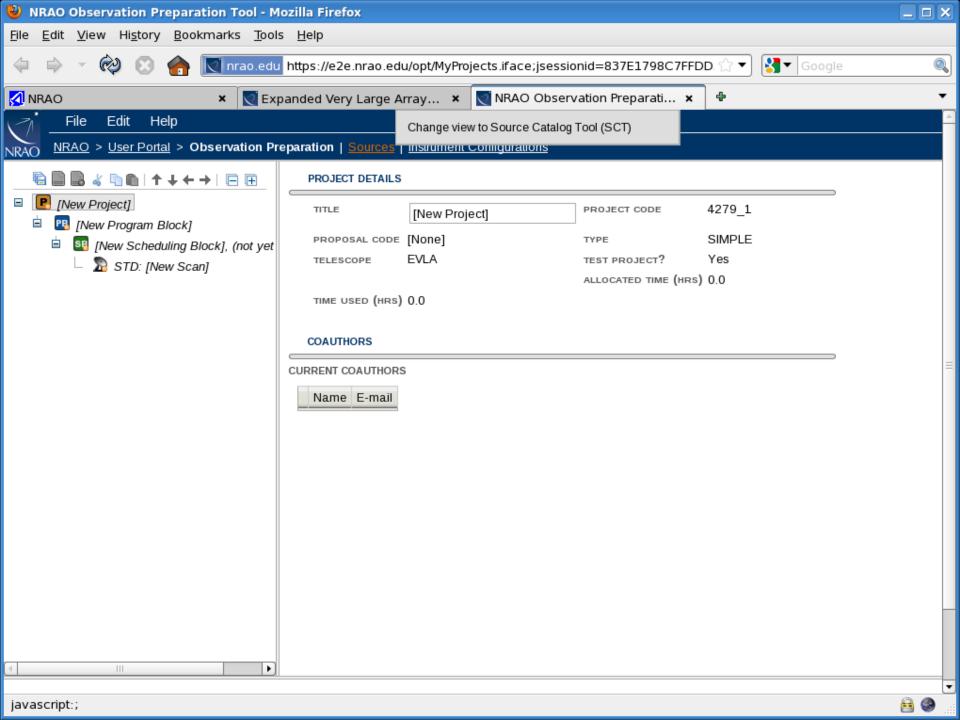


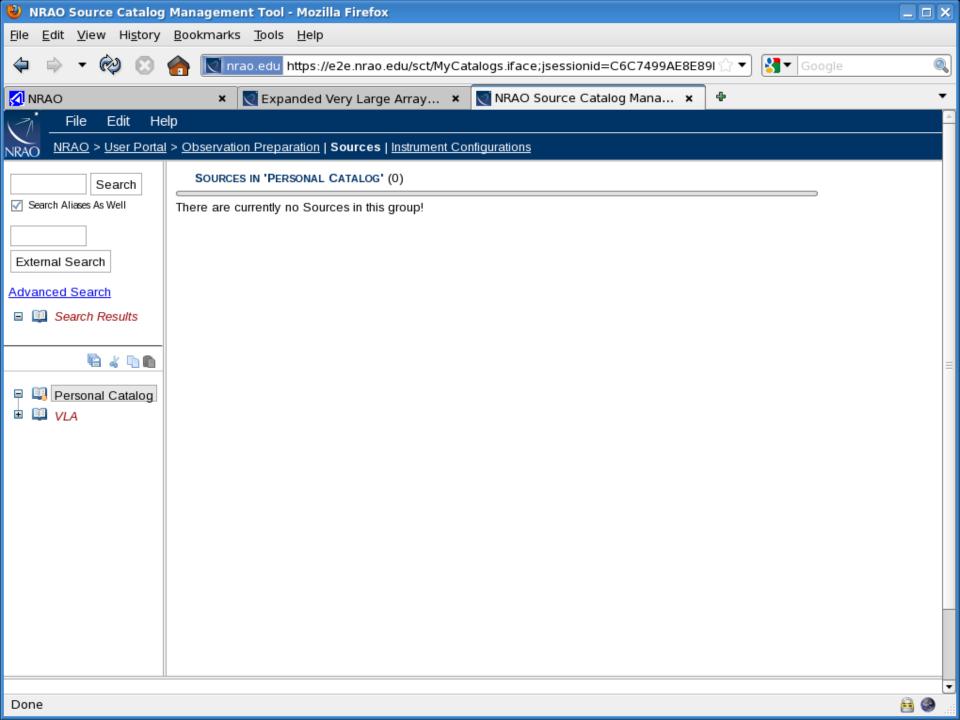


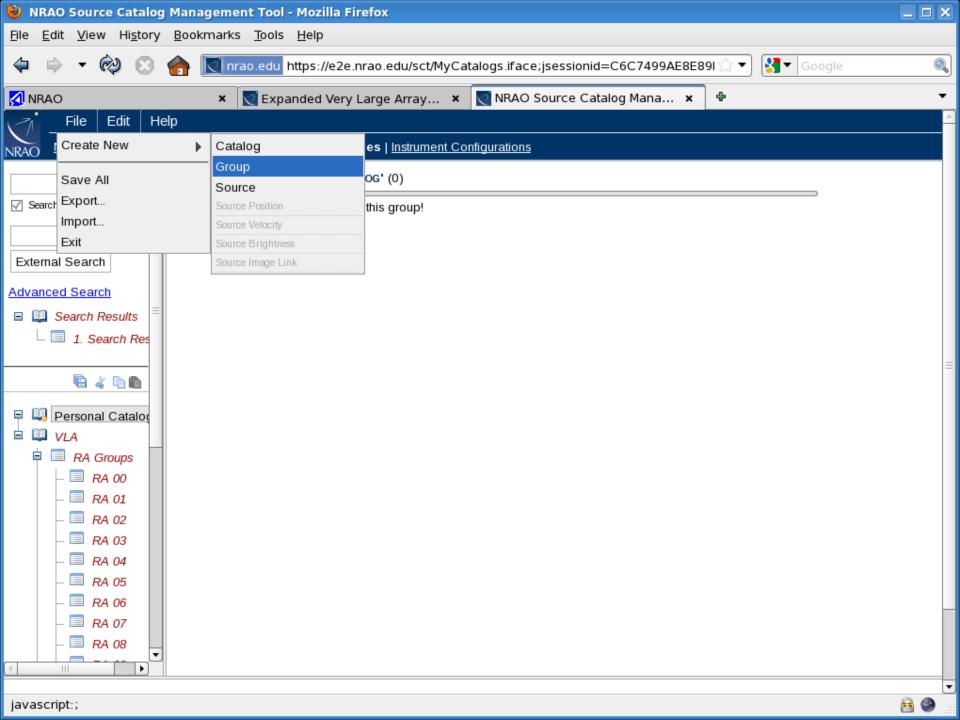
Creating source catalogs using the SCT

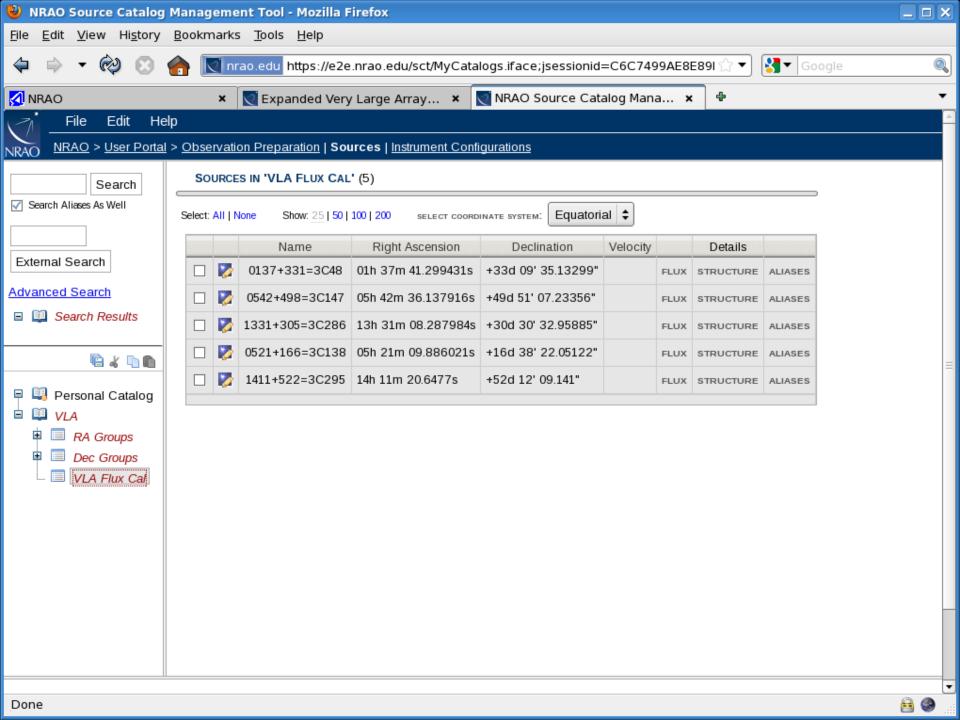
- Sources may have been transferred from the proposal information
 - If so, check the source positions, velocity definitions, ...
- Can define your own sources (telescope pointing directions)
- Can use sources predefined in (VLA) "calibrator" source list
 - Not every (calibrator) source in this list is useful for calibration in every observation;
 check its properties before using for calibration!
 - Some (calibrator) sources make great target sources...
- Useful to combine <u>all</u> sources to be used in an SB(/PB) in its own "group"
- Make sure all sources are defined before creating the scan list

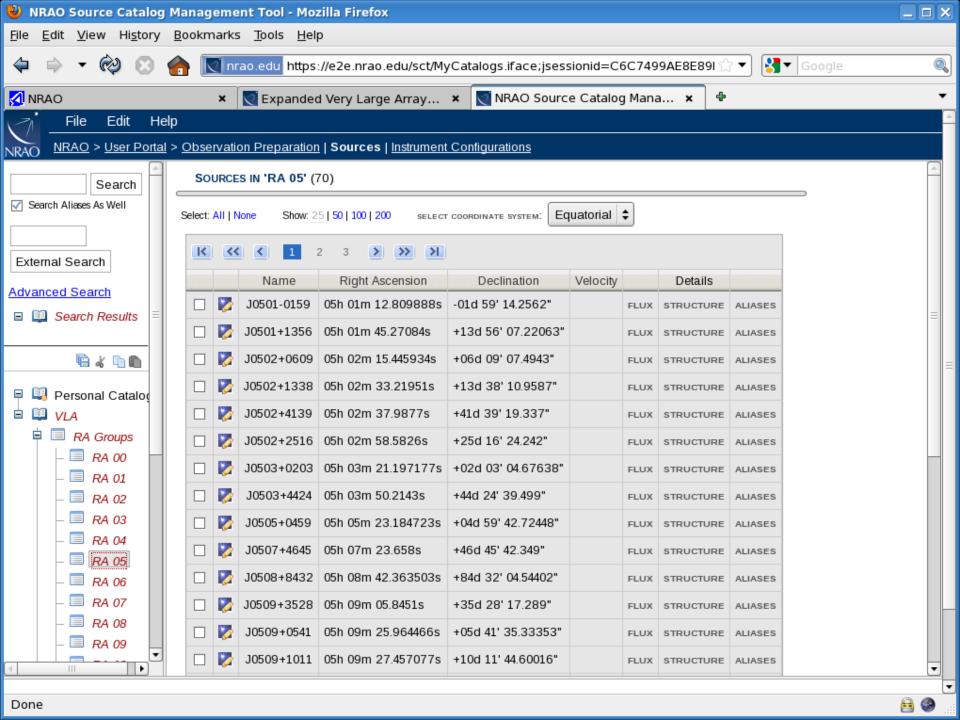


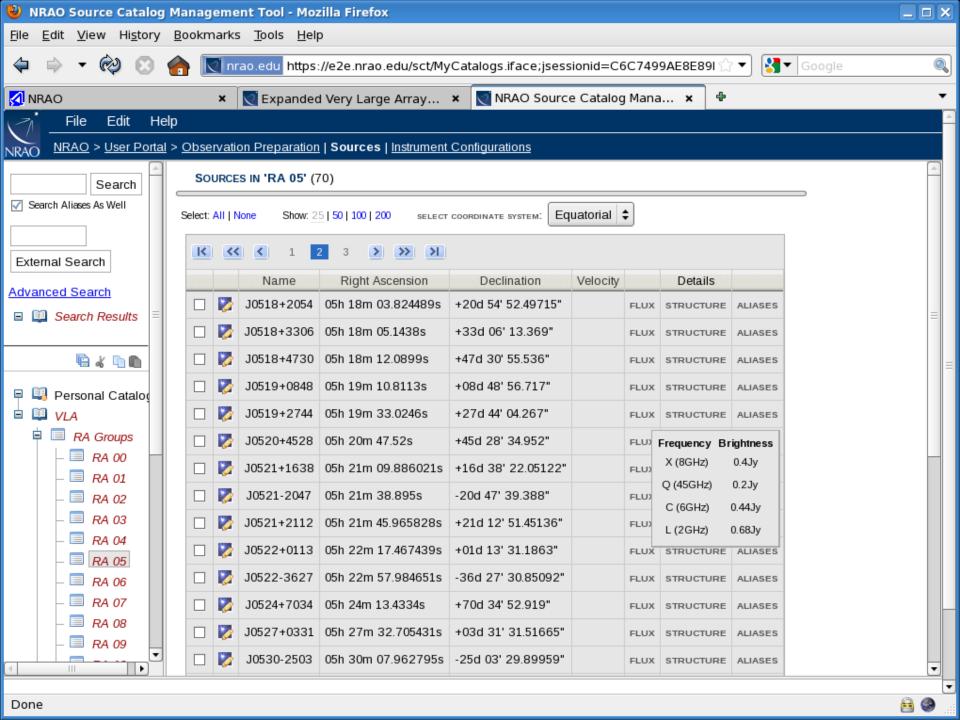


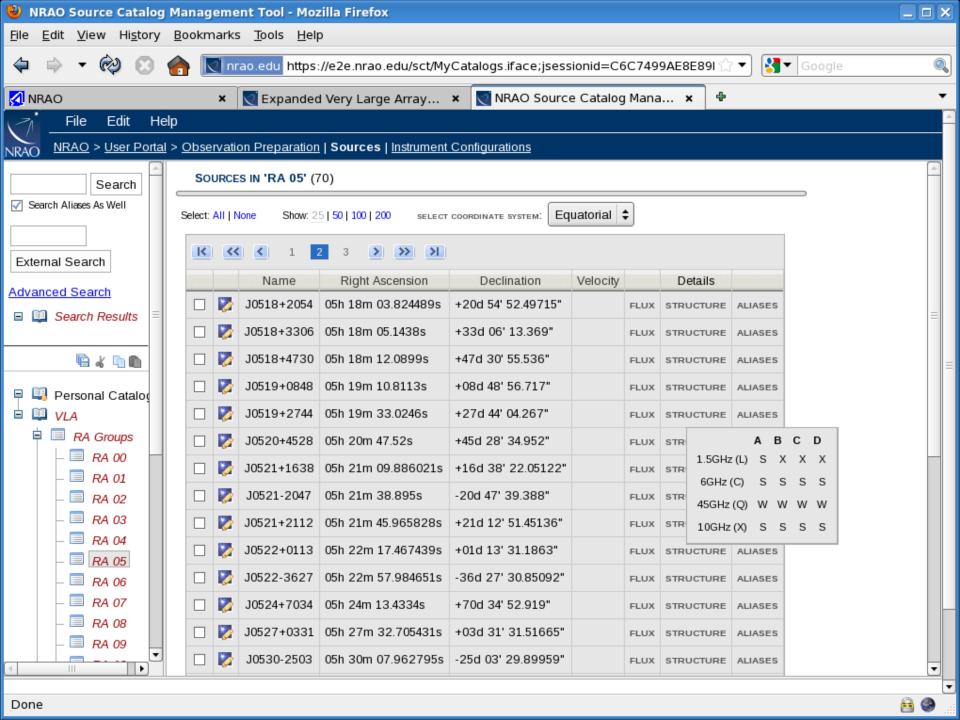


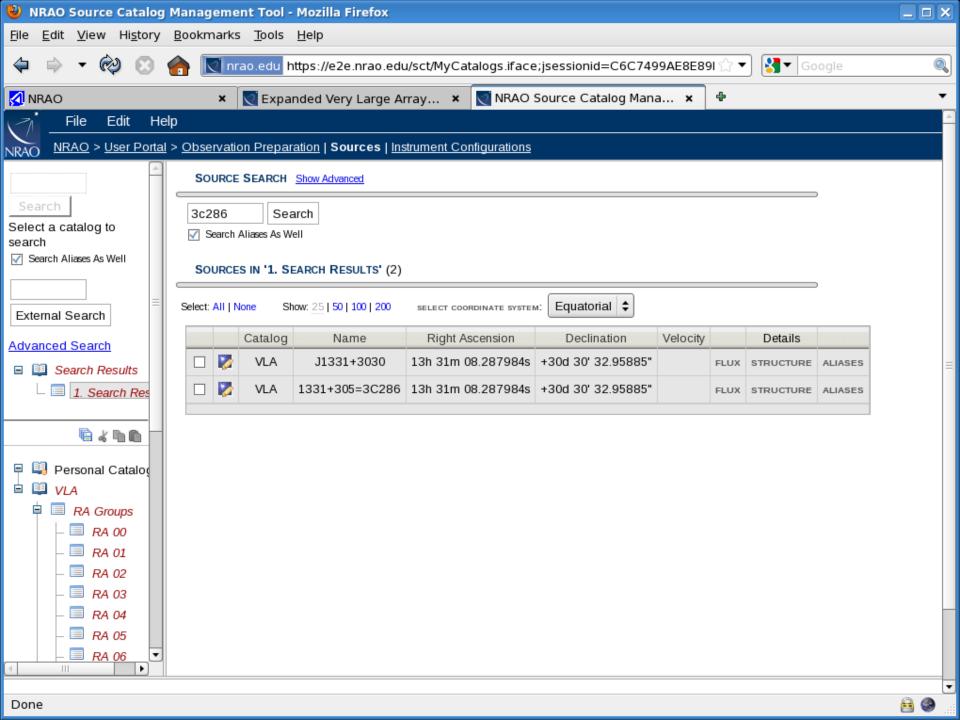


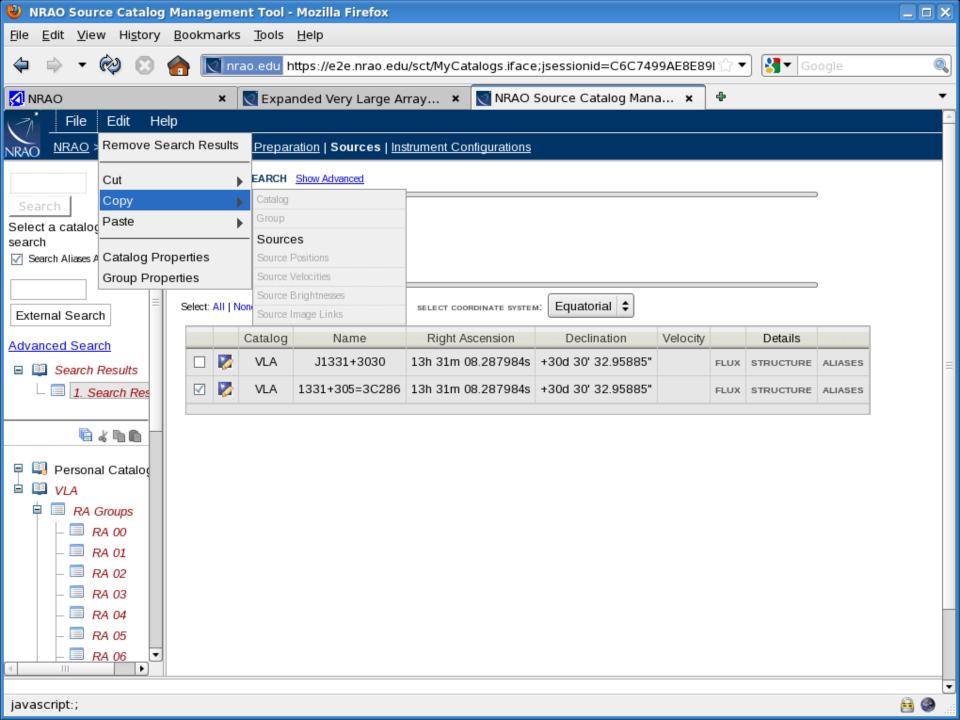


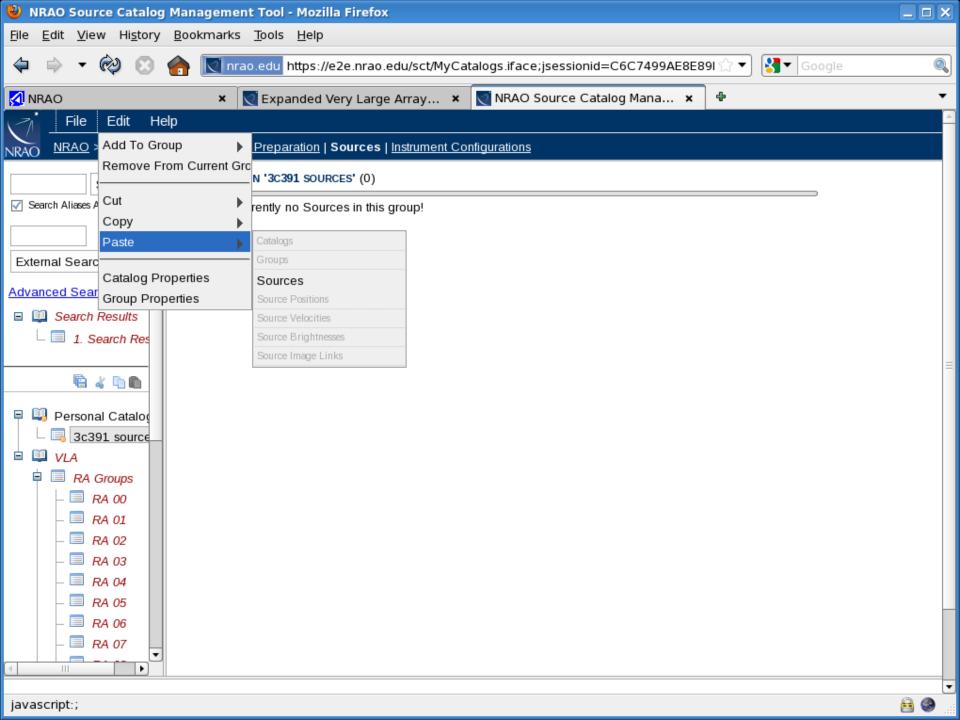


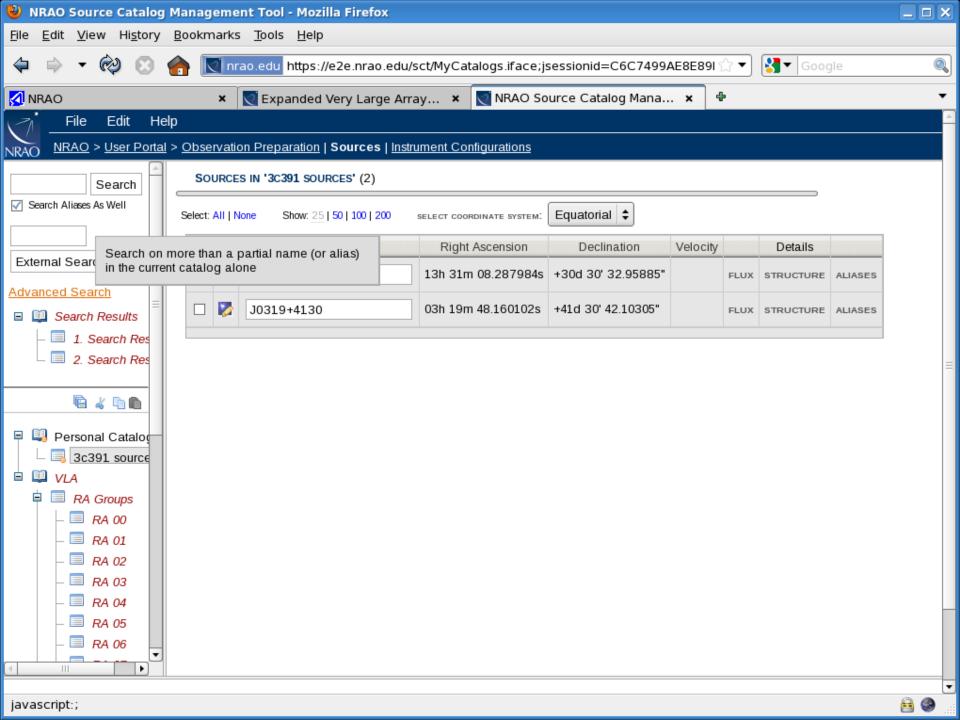


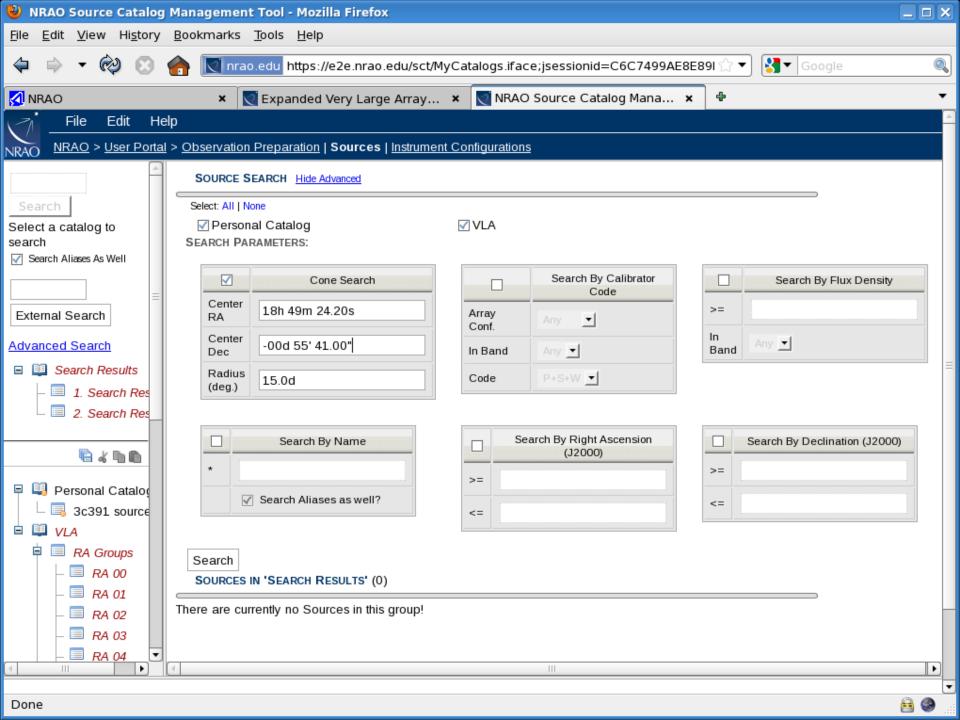


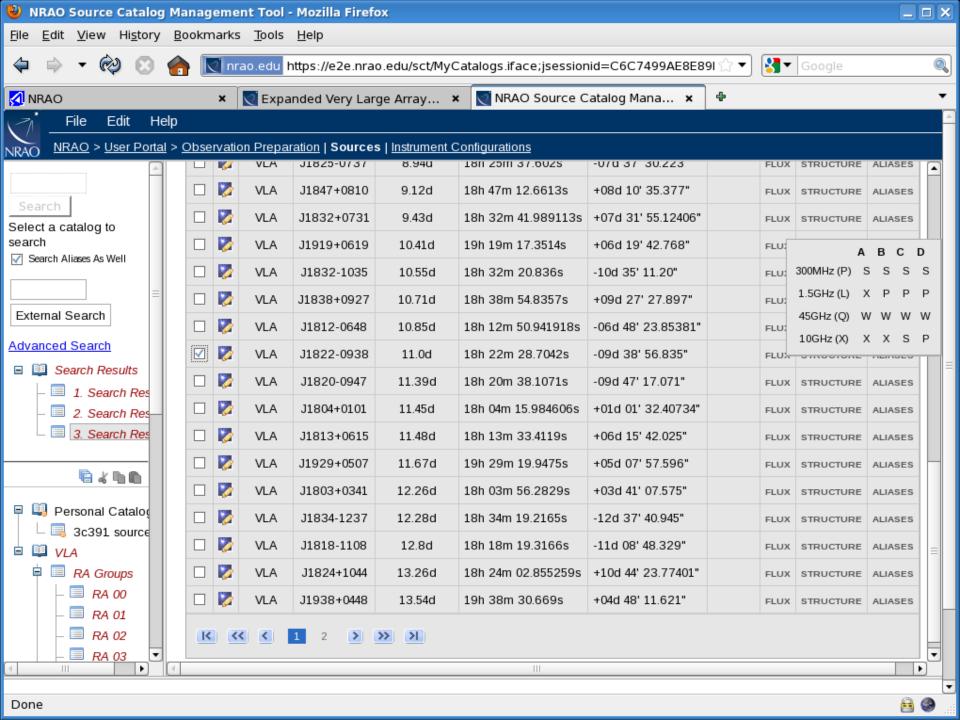








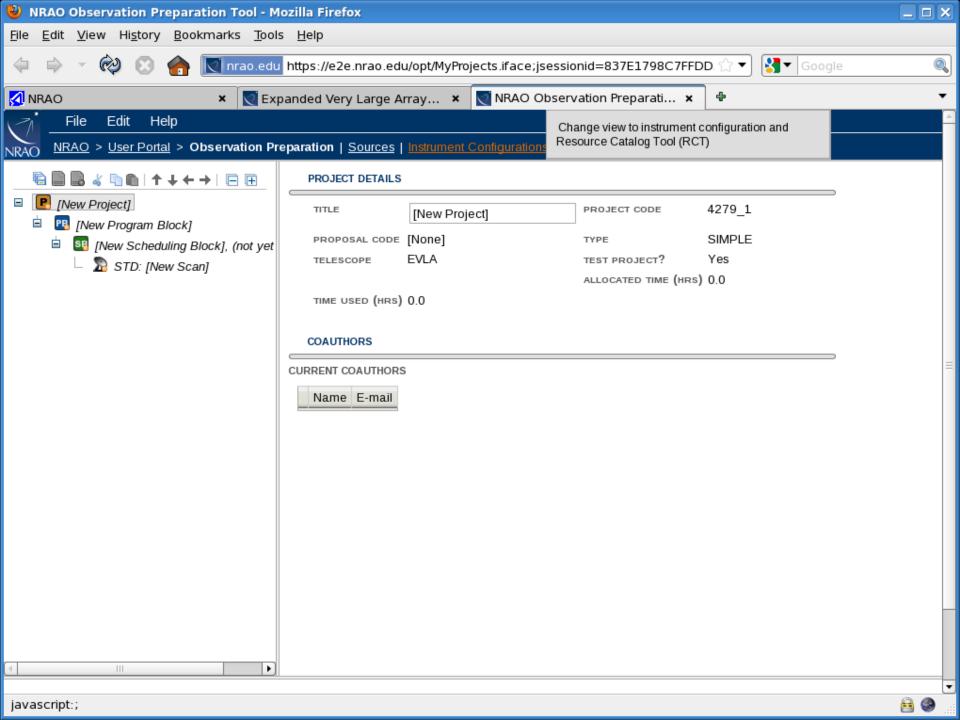


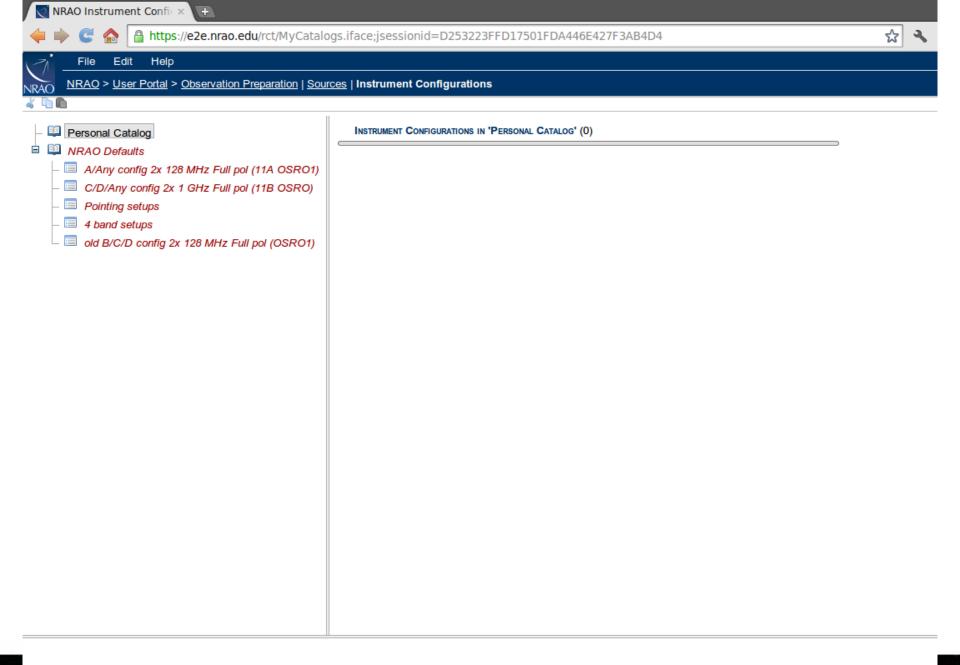


Creating resource catalogs in the RCT

- Can define your own resources (hardware/instrument configurations)
 - Different targets may need different Doppler setups/resources and extra dummy scans
- Can use resources predefined in default continuum "OSRO" resource list
 - Not every (continuum) resource in this list is useful for every continuum observation;
 check its properties before using!
 - More (default) resources will appear with advances in commissioning
 - We also kept the old OSRO2 available (used in last D array configuration, twice the channels in one subband)
 - But it is probably more advantageous to use 8 interleaved subbands to get 4x the bandwidth for calibration
- Useful to combine resources to be used in an SB(/PB) in its own "group"
- Make sure all resources are defined before creating the scan list















NRAO > User Portal > Observation Preparation | Sources | Instrument Configurations





A/Any config 2x 128 MHz Full pol (11A OSRO1) C/D/Any config 2x 1 GHz Full pol (11B OSRO)

Help

Pointing setups

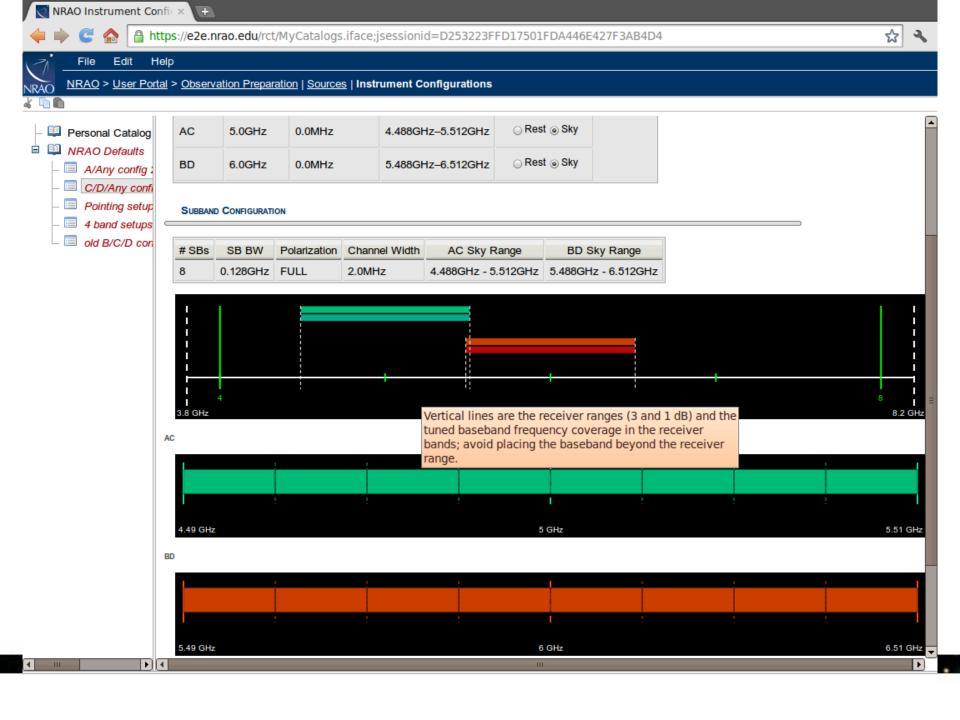
4 band setups

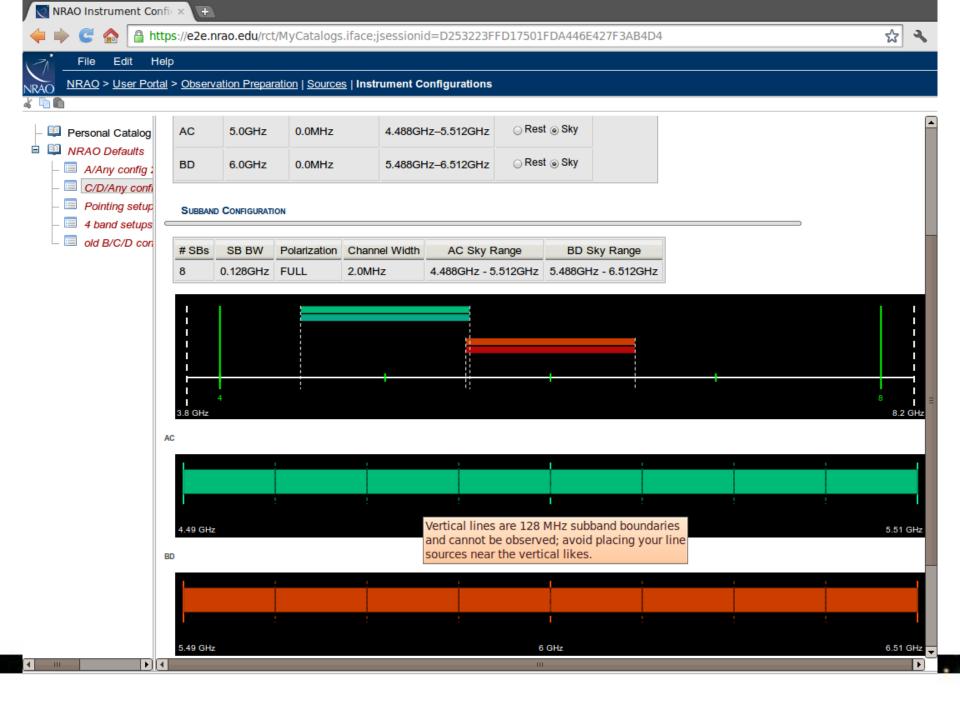
old B/C/D config 2x 128 MHz Full pol (OSRO1)

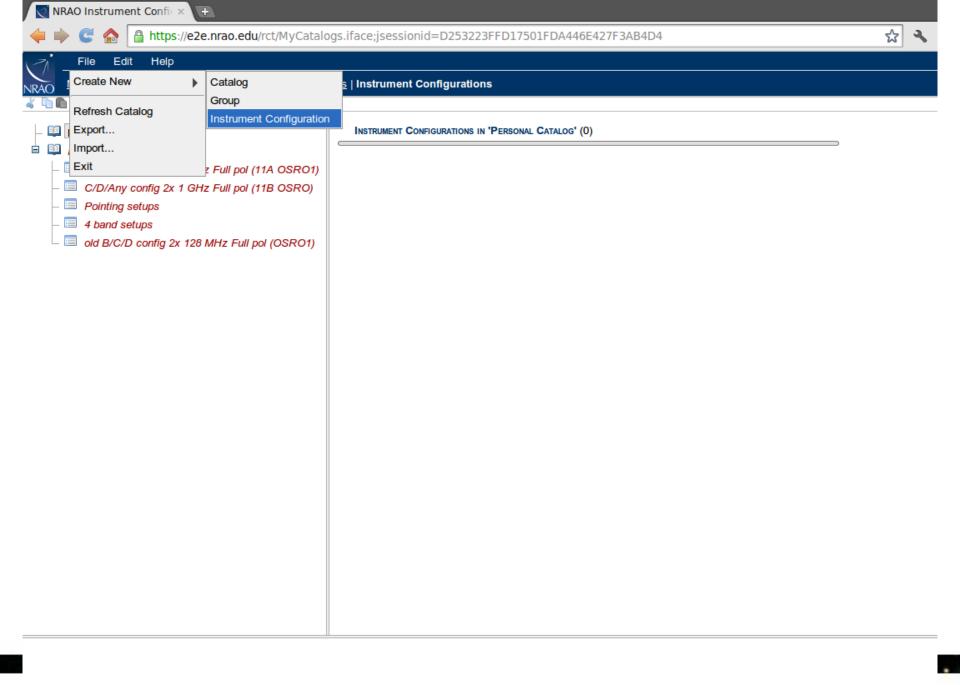
INSTRUMENT CONFIGURATIONS IN 'C/D/ANY CONFIG 2x 1 GHz Full Pol (11B OSRO)' (8)

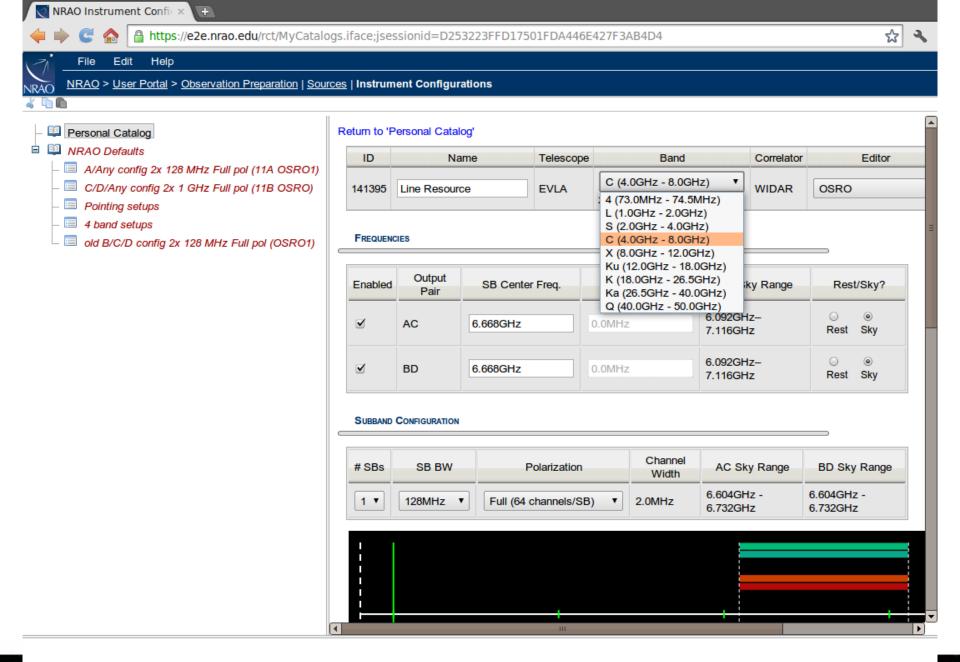
Show: 25 | 50 | 100 | 200 Select: All | None

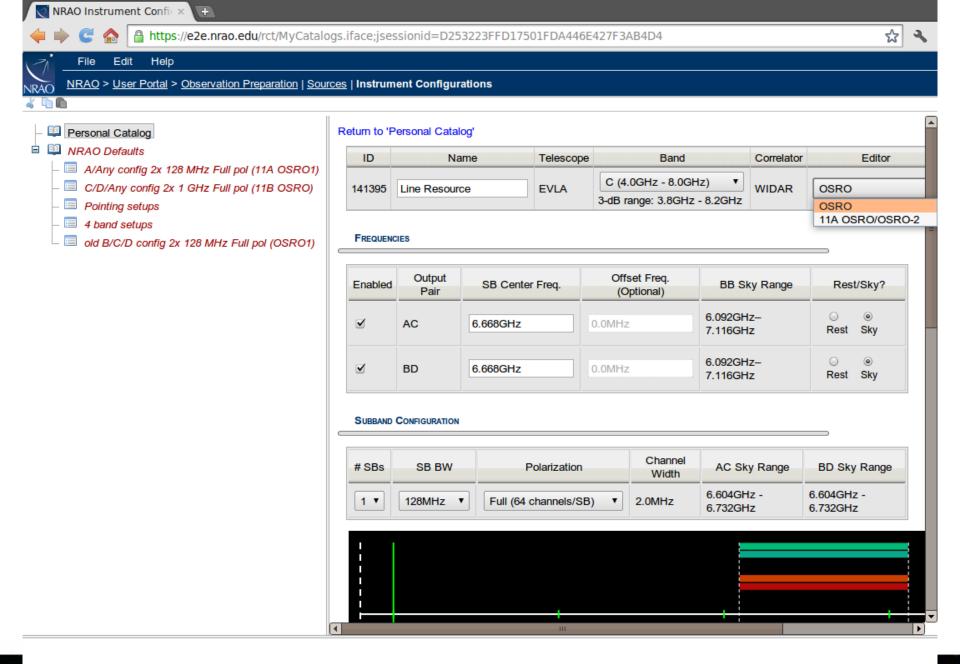
	Name	Band	Back End	T _{int}	AC BW BD BW	AC Freq BD Freq	Comments
>	L full width	L	WIDAR / OSRO	5	0.512GHz 0.512GHz	1.25GHz 1.75GHz	
>	S full width	s	WIDAR / OSRO	5	1.024GHz 1.024GHz	2.5GHz 3.5GHz	
>	C wide band	С	WIDAR / OSRO	5	1.024GHz 1.024GHz	5.0GHz 6.0GHz	
>	X band, mixed wide/narrow	x	WIDAR / OSRO	3	1.024GHz 1.024GHz	8.5GHz 9.5GHz	Uses all antennas, though some do not have the new wideband X band receivers. Good to obtain data on all baselines, but many will have no data in the second (BD) baseband.
>	Ku wide band	Ku	WIDAR / OSRO	3	1024.0MHz 1024.0MHz	13.5GHz 14.5GHz	
>	K wide band	к	WIDAR / OSRO	3	1024.0MHz 1024.0MHz	20.7GHz 21.7GHz	
>	Ka wide band	Ka	WIDAR / OSRO	3	1.024GHz 1.024GHz	32.52GHz 31.52GHz	
>	Q wide band	Q	WIDAR / OSRO	3	1024.0MHz 1024.0MHz	40.5GHz 41.5GHz	

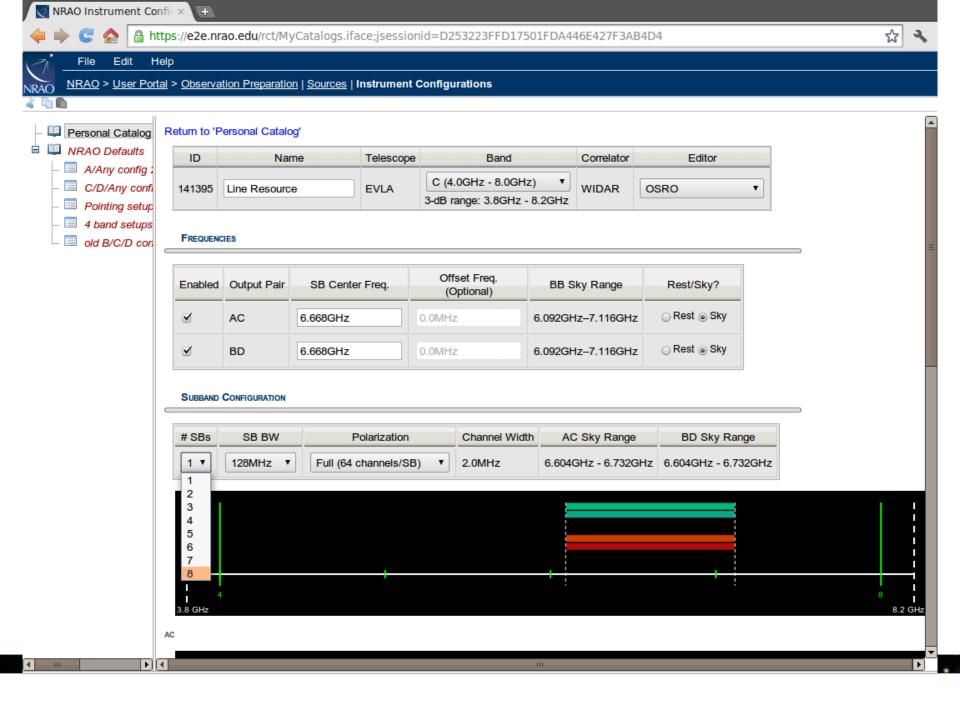


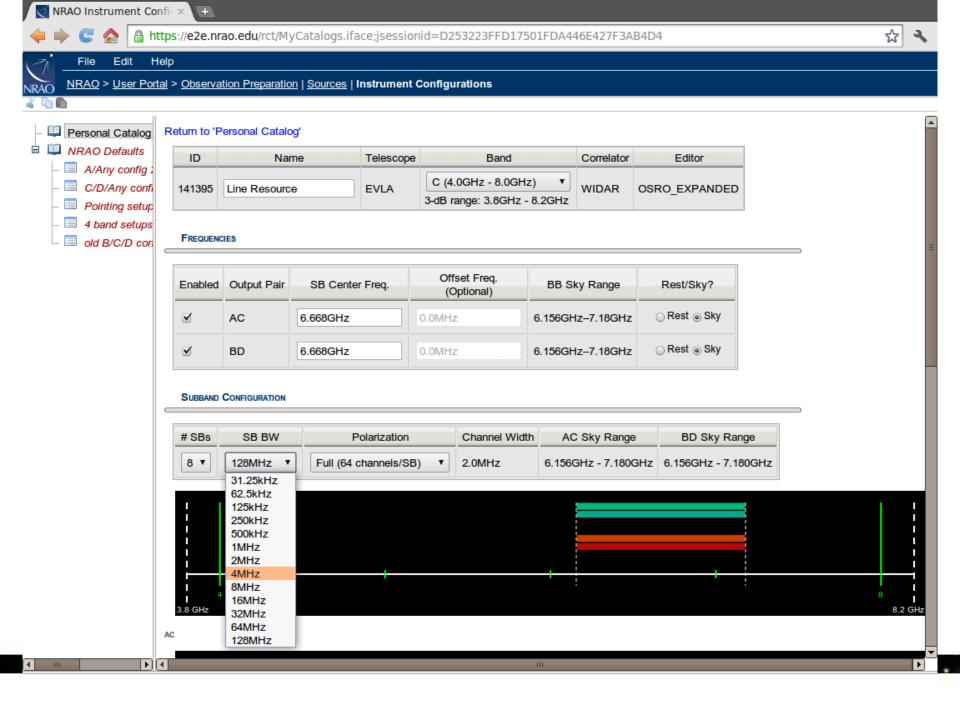


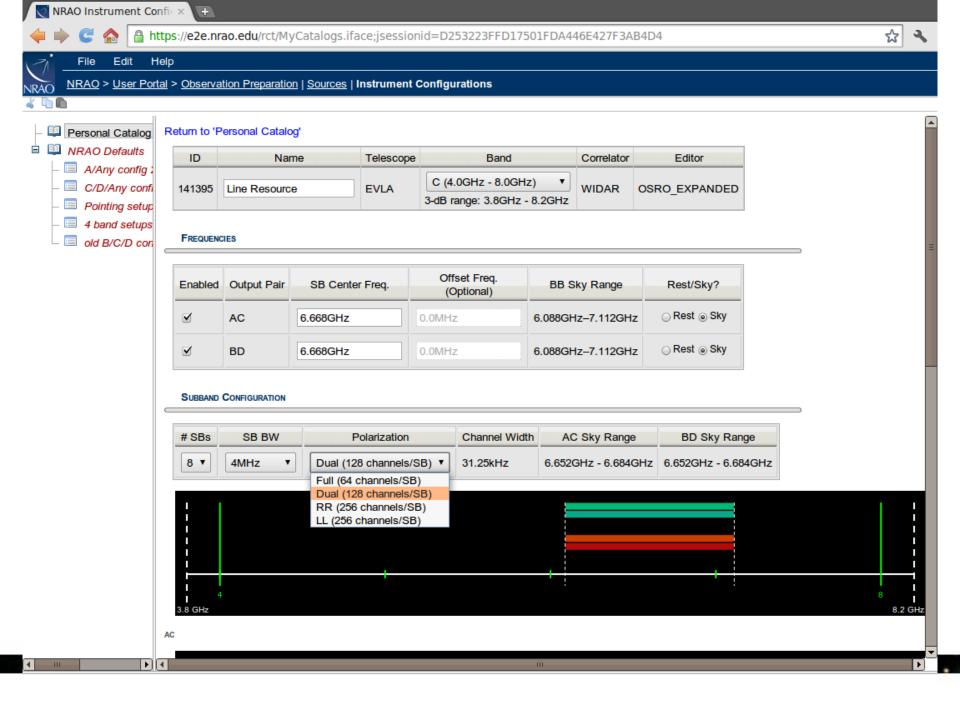


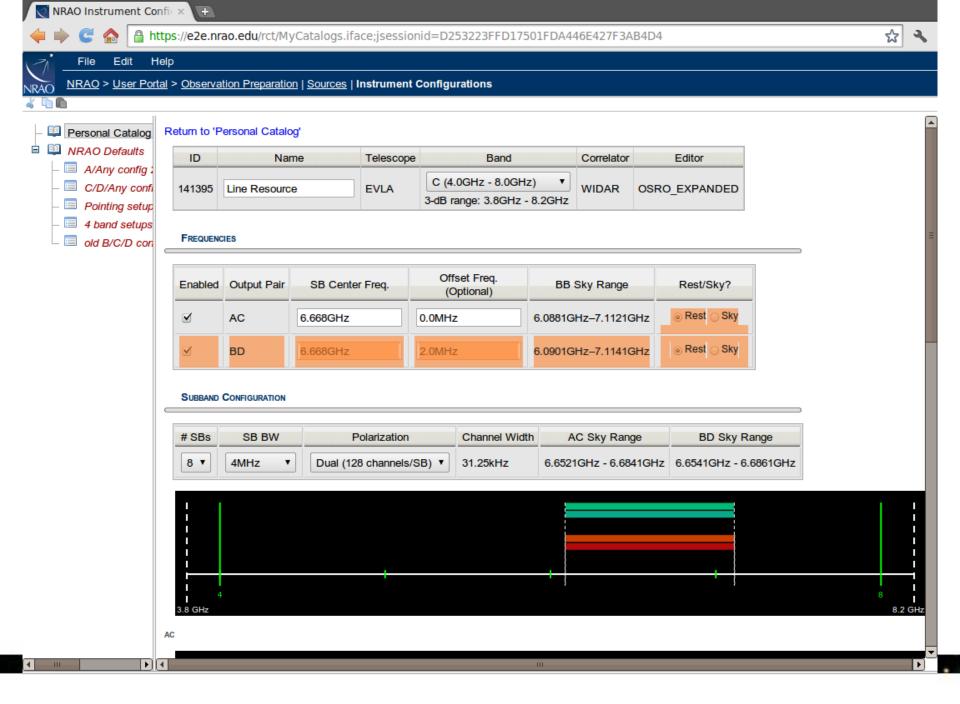


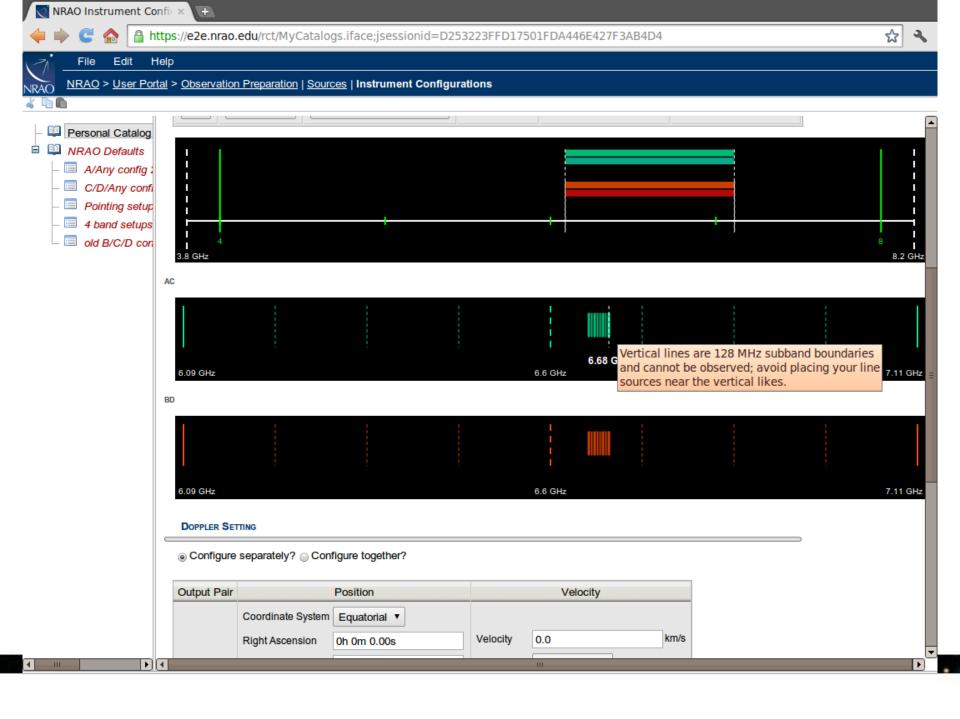


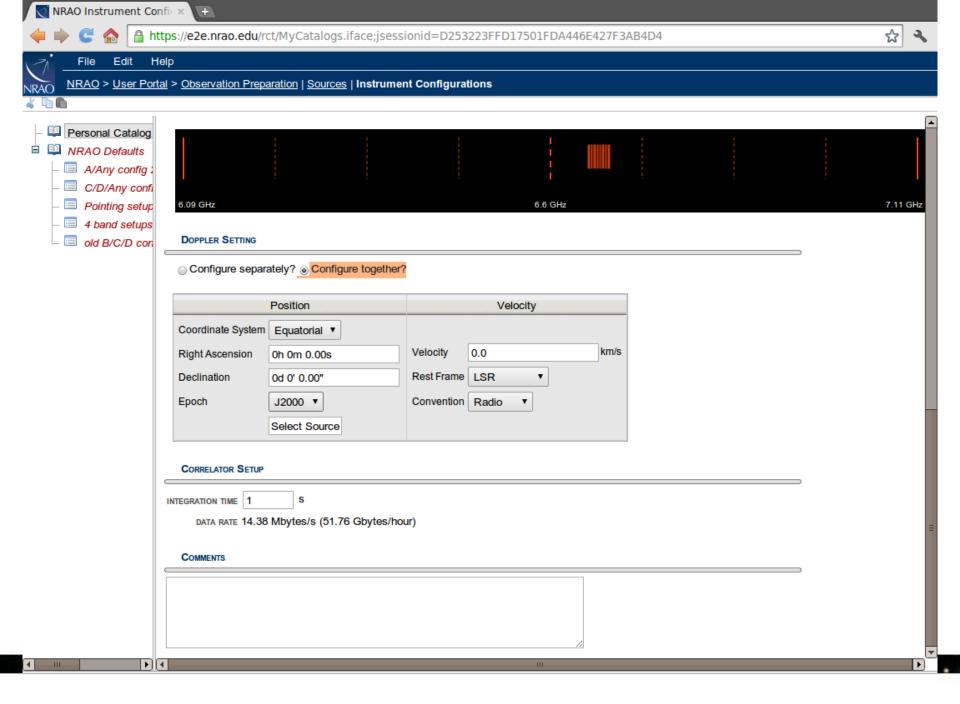


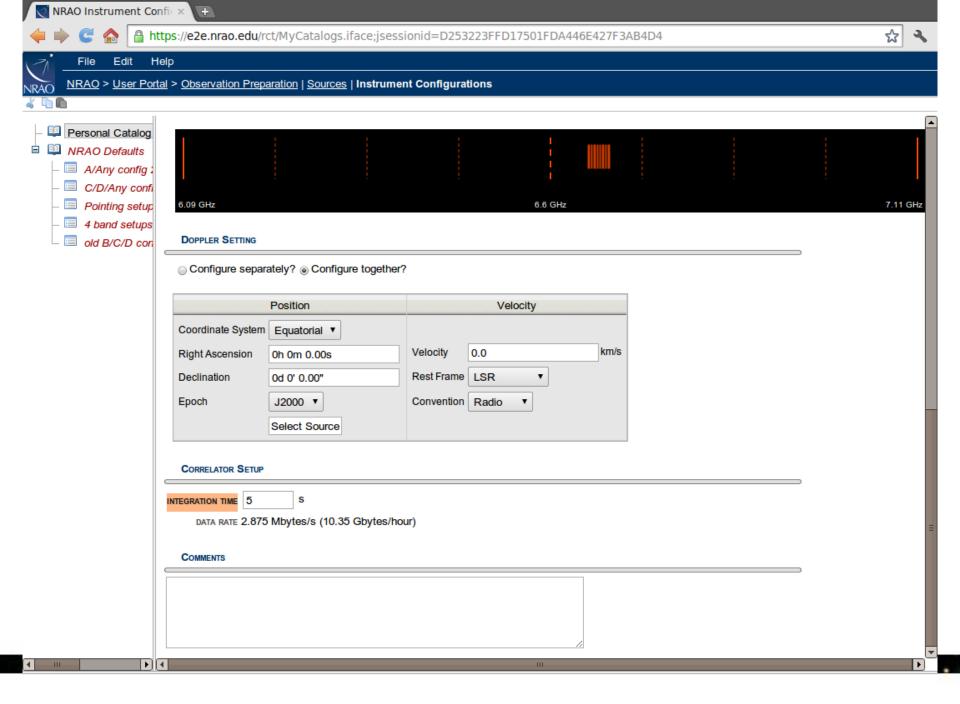


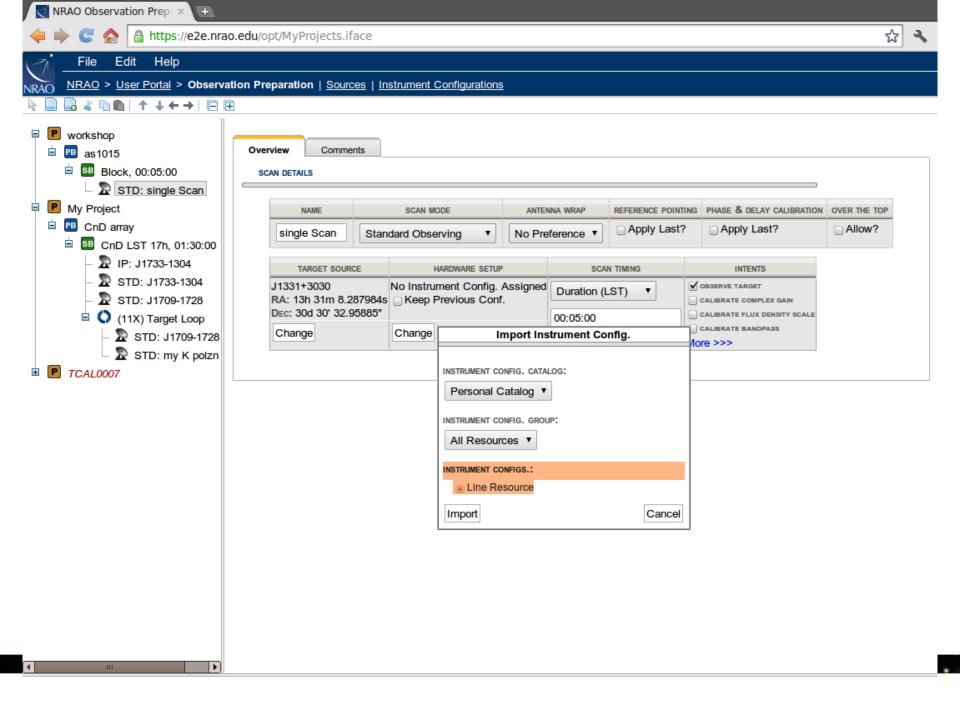


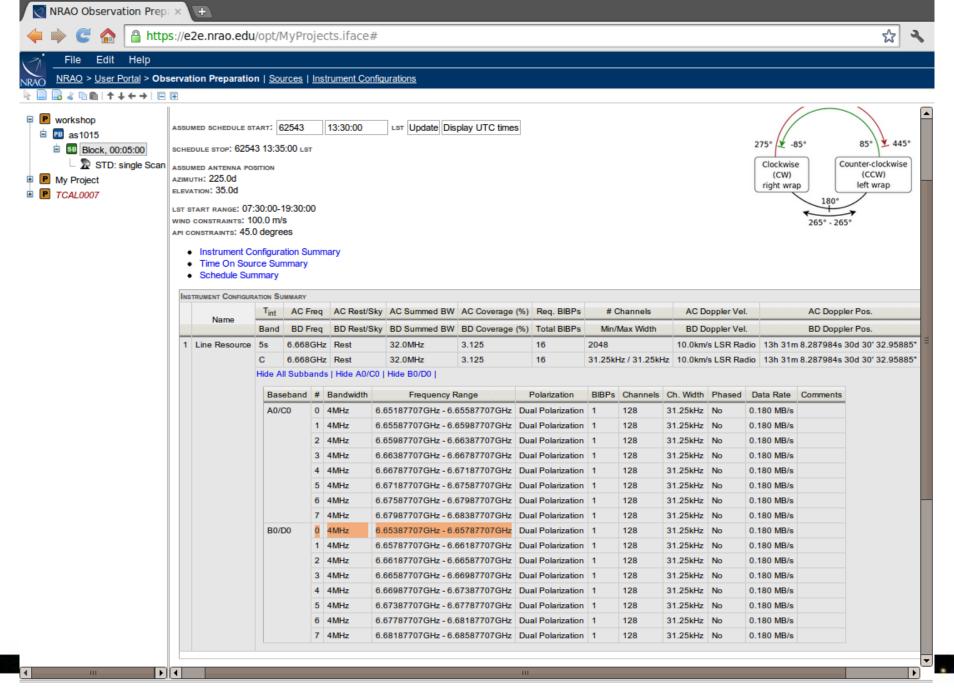












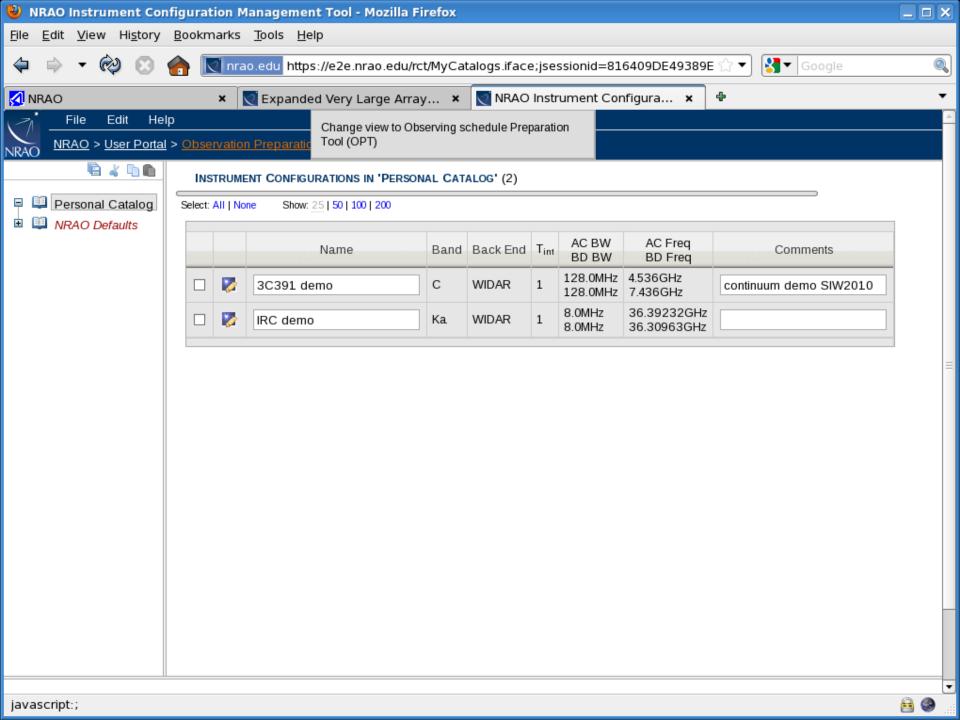
- The remaining ~30 slides on how to create a scheduling block are recycled from the SIW, synthesis imaging workshop held in 2010, using the demo data sets used there; they are useful to look at if you like, or have a look at the full, original SIW OPT talk posted on the SIW web pages. Also, the manual has a lot of useful information if you can see through the belated updates; be aware that the OPT software is still changing almost weekly..
- If you encounter problems that are not addressed in the documentation, create a ticket through the NRAO helpdesk: https://help.nrao.edu/ (use your my.nrao.edu account to log in or log in through my.nrao.edu)
- Questions?



Building Scheduling Blocks (SBs)

- All sources to be used are defined
 - Transferred from the proposal semi-automatically (i.e. check!)
 - Found in existing catalogs
 - Created new
- All sources to be used are conveniently grouped
 - More user friendly if there are many, or from different catalogs
- All resources to be used are defined see above
- All resources to be used are conveniently grouped see above
- The project Program Blocks are present
 - Proposal information and read-only details already entered
 - Pl and Co-ls, type, array config., time allocated, (priority,) ..., (...)





Designing the 3C391 observing run

- 7 field mozaic, 7 sources with roughly equal time on source
 - Enough signal to be able to use self-calibration
- Could use default C band continuum (OSROI) setup (but didn't)
- Flux: **3C286**, can also be used for EVPA calibration (answer: -66°)
- Bandpass: 3C84
- Unpolarized nearby point source: J1822-0938 (gain, pol.z leakage)
 - C band (i.e. about 5 GHz) thus cycle every ~30 minutes
- Do not break into small observing runs (polarization angle)
- Spread sources over hour angle for best (u,v)-coverage
 - Set up sequence and loop several times
- Spread polarization leakage calibrator: include in loop



