

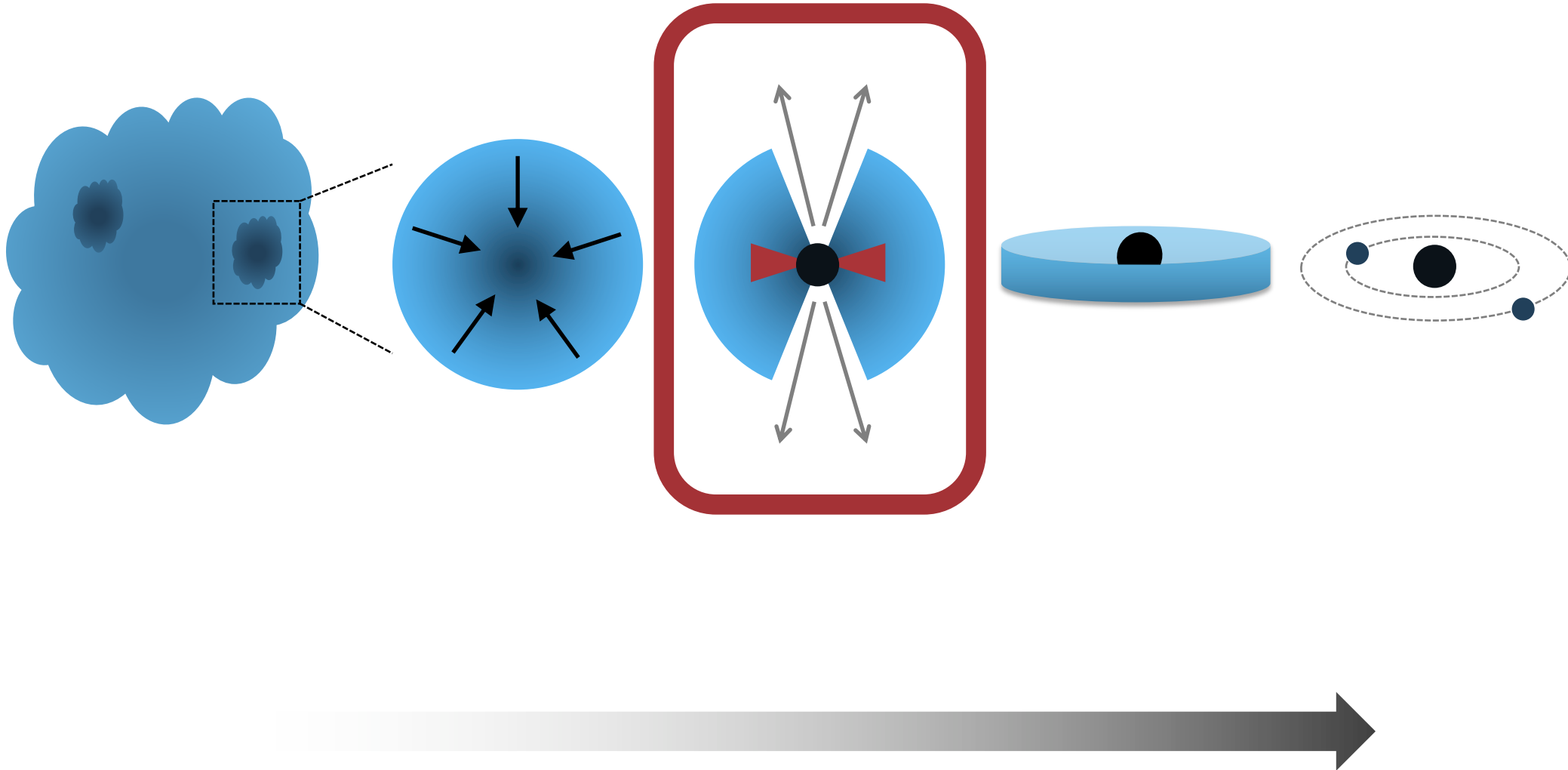
YOUNG EMBEDDED DISKS:

The ALMA (WSU) and JWST era

MEREL VAN 'T HOFF

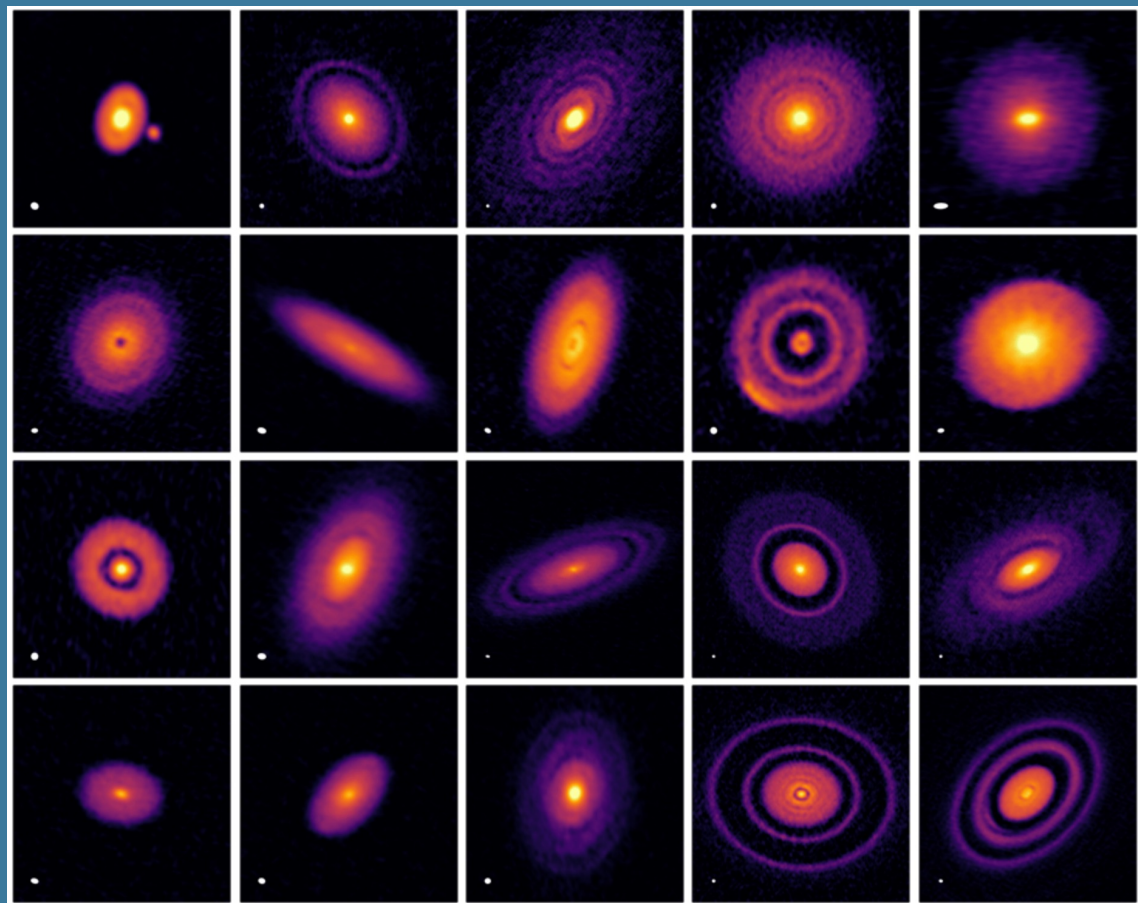
Purdue University

YOUNG EMBEDDED DISKS



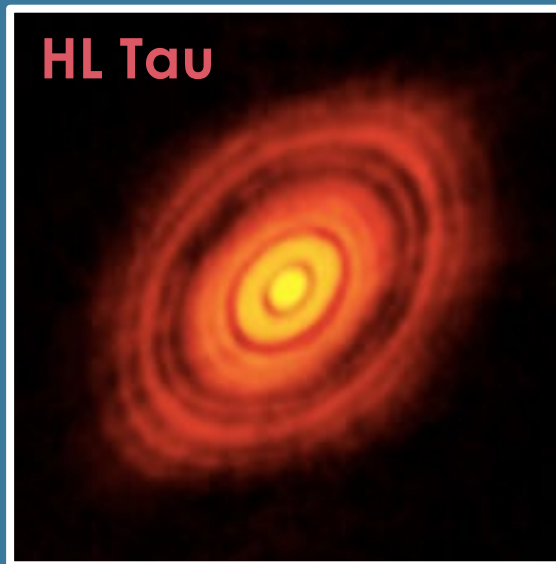
CONTINUUM

Are dust substructures common in young disks?

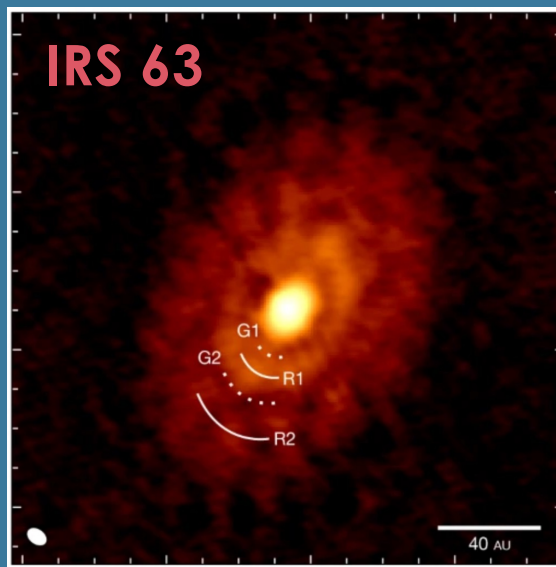


Mature protoplanetary disks (DSHARP)

Andrews et al. 2020



ALMA Partnership 2015

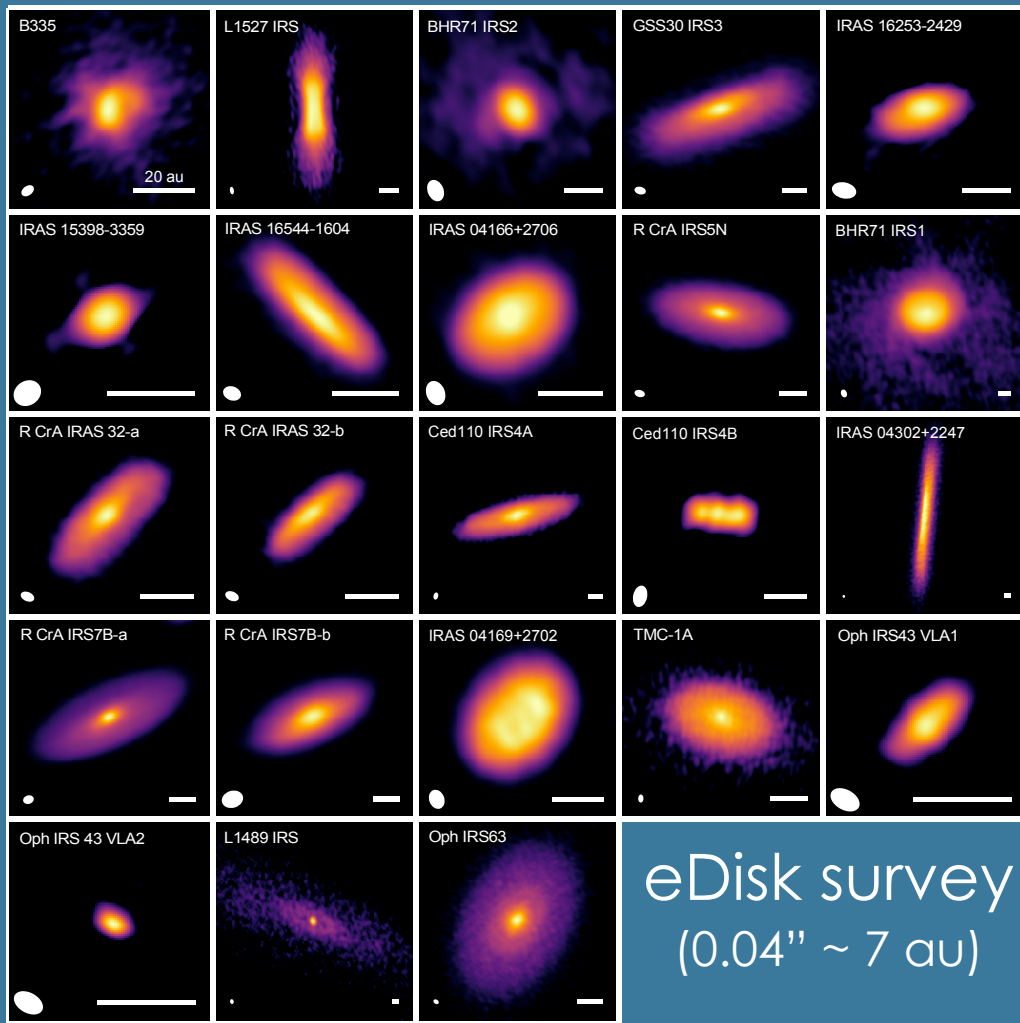


Segura-Cox et al. 2020

Young embedded disks

CONTINUUM

Are dust substructures common in young disks?



Maybe not...

... but dust may be optically thick

... but eDisk is biased toward edge-on disks

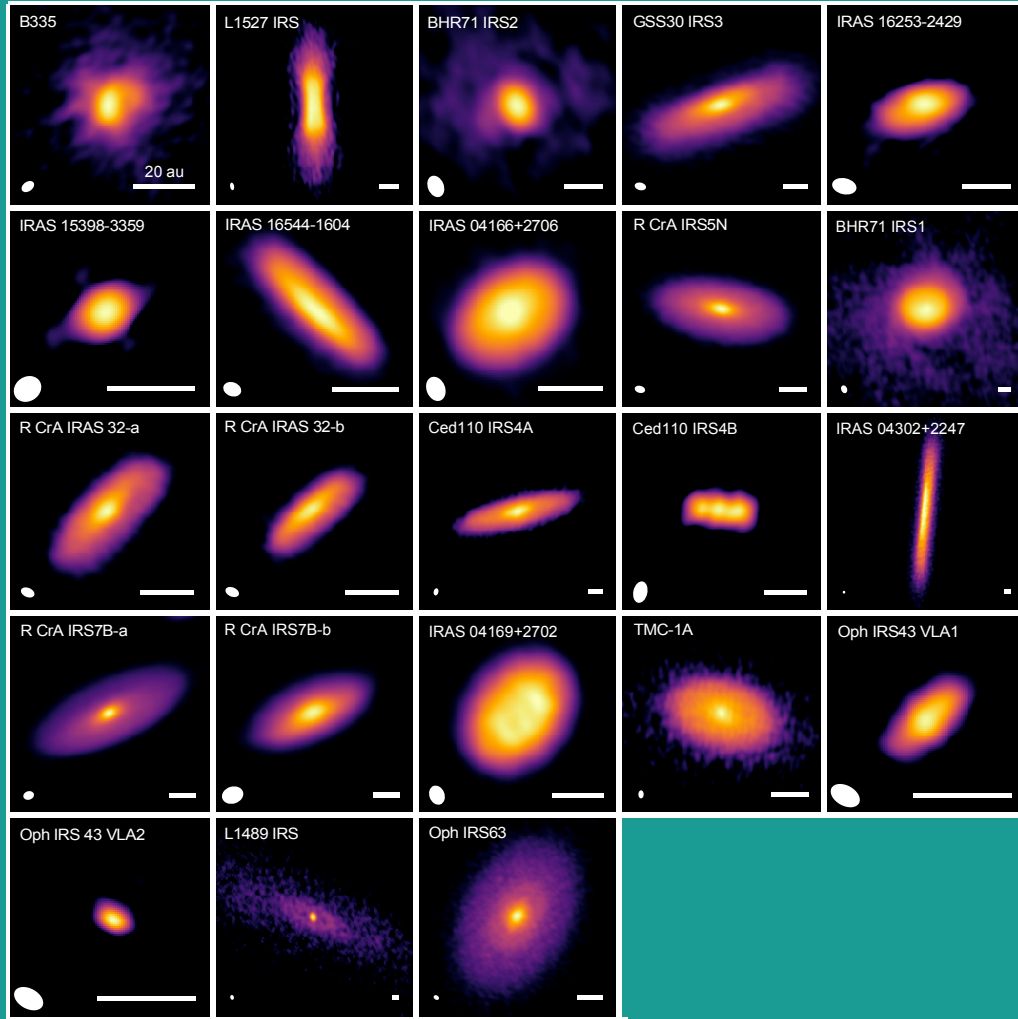
If not, substructures may form quickly once Class 0 and I protostars evolve into Class II.
(see also Nazari et al. 2024)

Most common structure:

Asymmetry along the disk minor axis,
Suggest optically thick vertically extended disk, so dust not settled toward midplane.

CONTINUUM

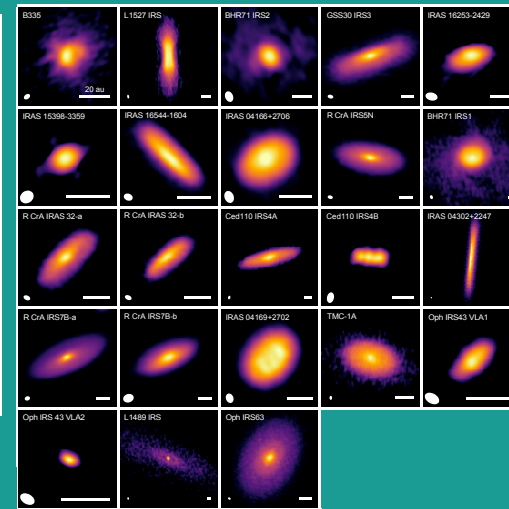
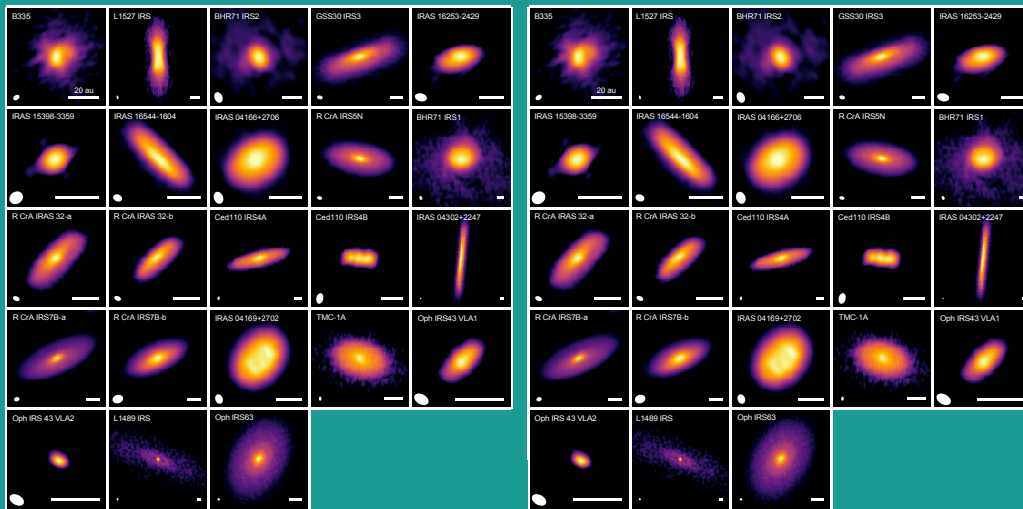
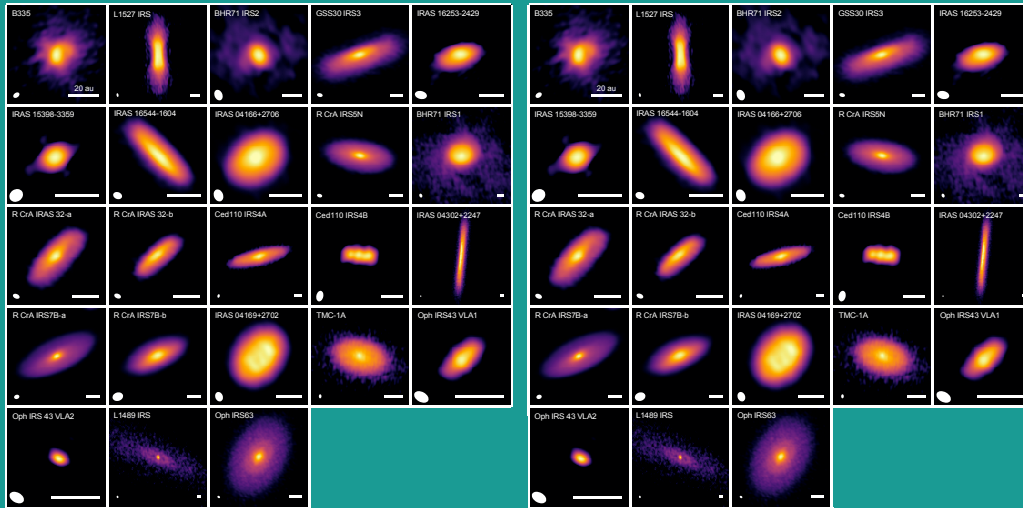
Are dust substructures common?



$\geq 4.8x$ (for 2x BW correlation)
increase in continuum imaging
speed

CONTINUUM

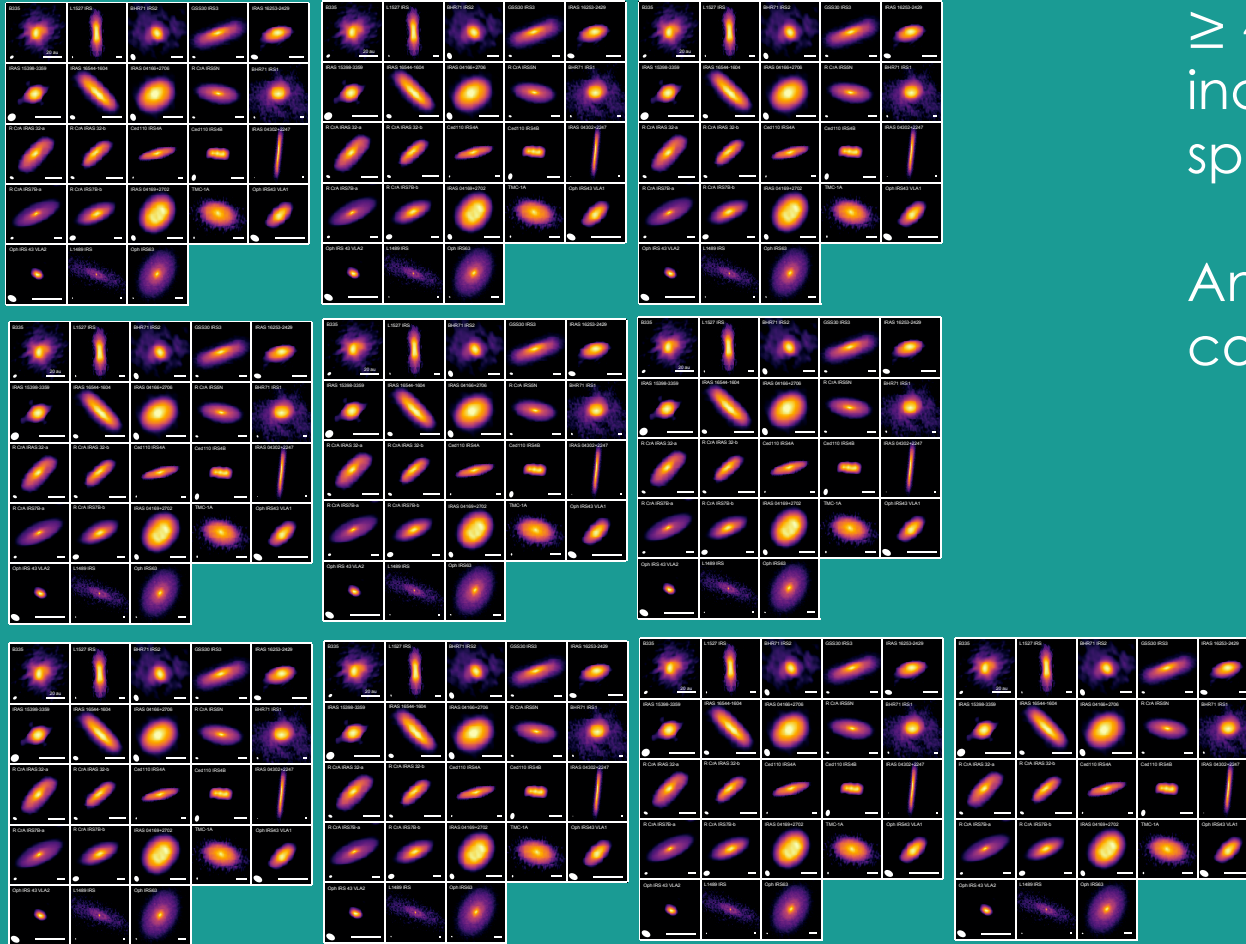
Are dust substructures common?



$\geq 4.8x$ (for 2x BW correlation)
increase in continuum imaging
speed

CONTINUUM

Are dust substructures common?



$\geq 4.8x$ (for 2x BW correlation)
increase in continuum imaging
speed

Another 2x increase for 4x BW
correlation

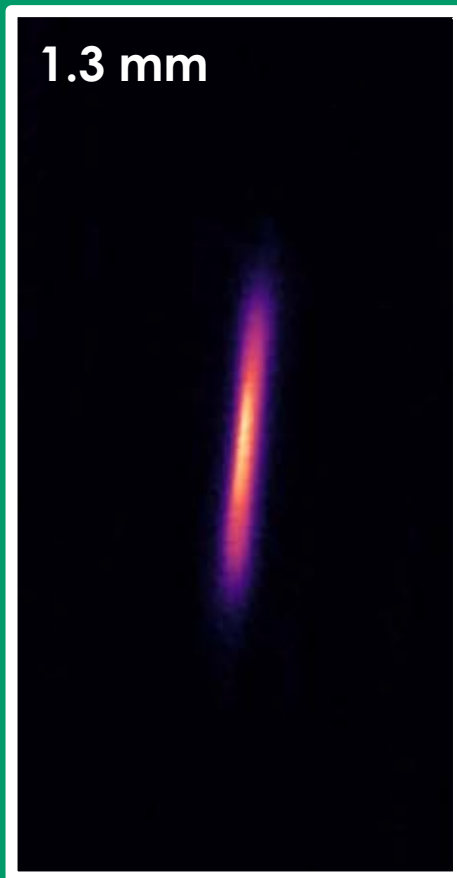
CONTINUUM

Disk structure

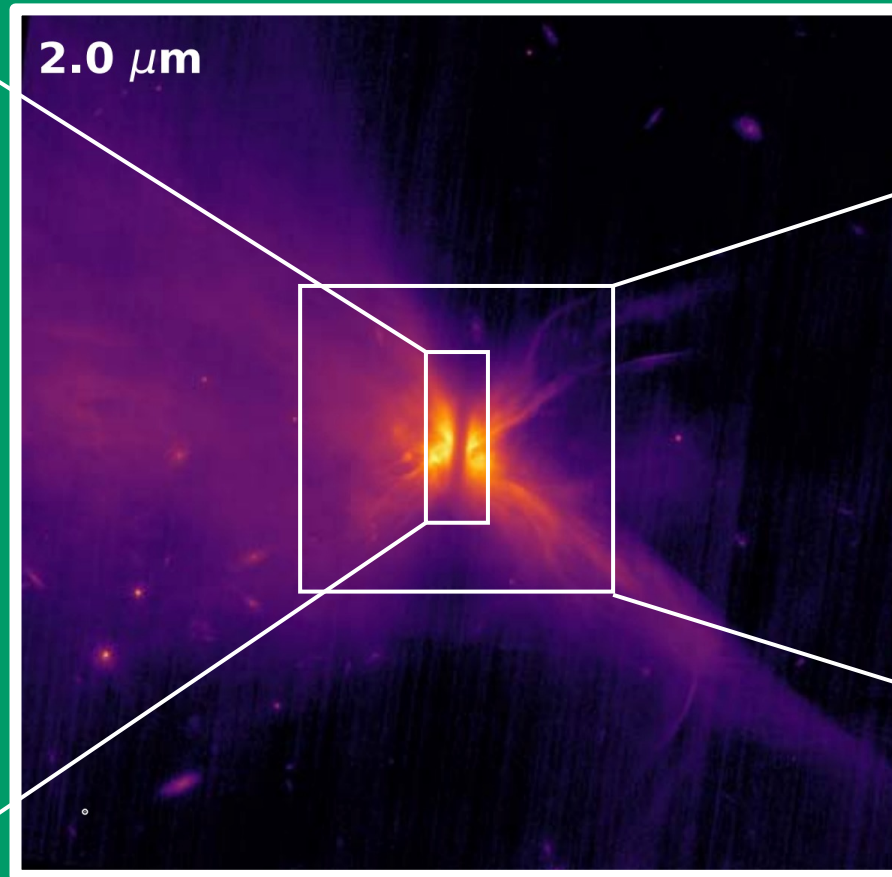
Asymmetry in scattered light suggest tilted inner disk

ALMA: thermal emission

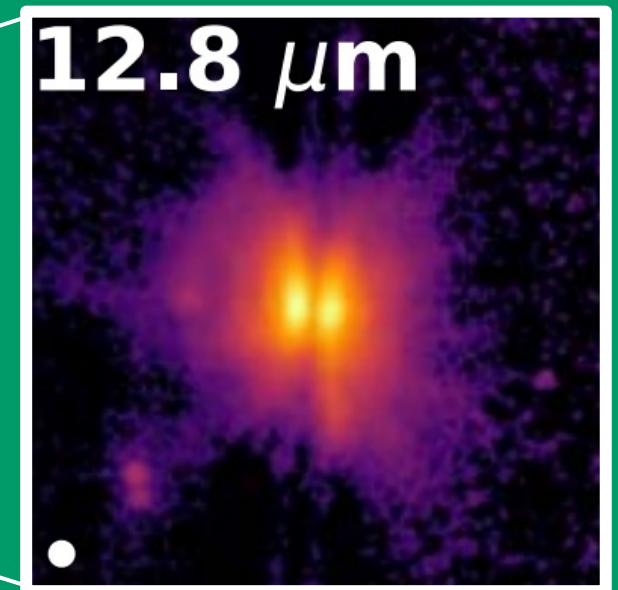
JWST: scattered light



eDisk: Lin et al. 2023



Villenave et al. 2024



CONTINUUM

Dust composition

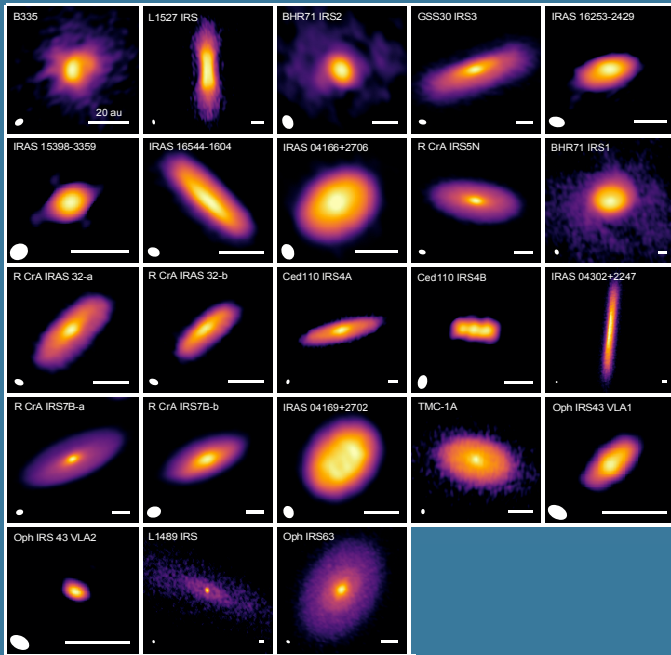
JWST MIRI spectrum can be decomposed into contributions from different minerals

CONTINUUM

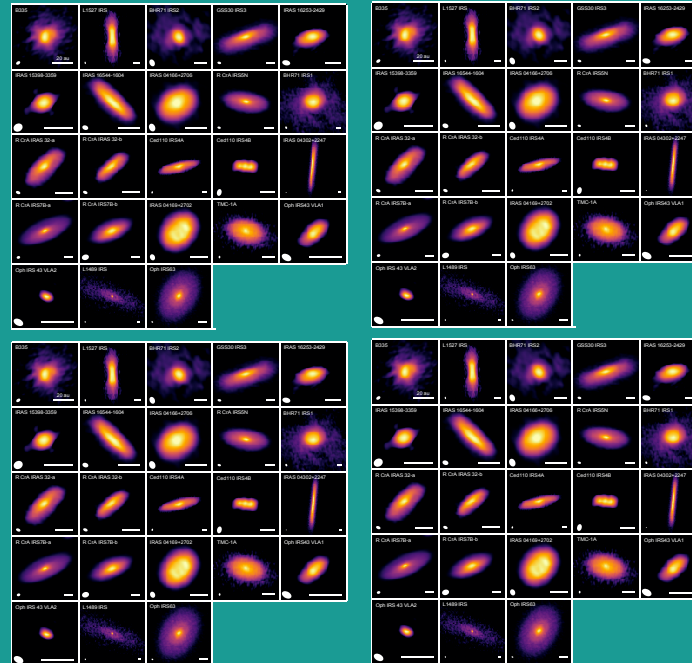
Dust in young disks

ALMA

Substructures not common
in young disks



Study much larger disk
sample



ALMA WSU

Complementary info on
structure and composition

JMST

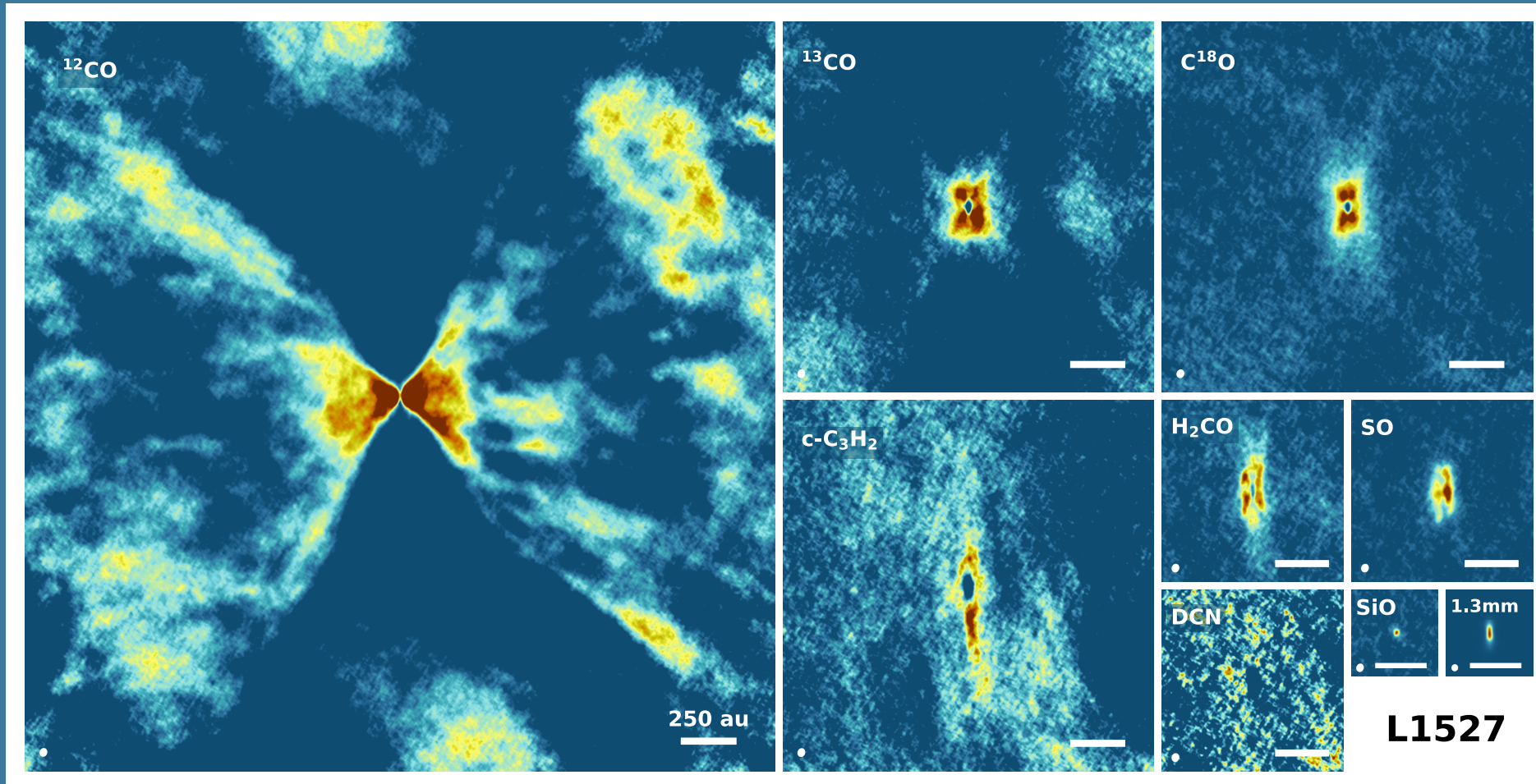
MOLECULAR LINES

Physical and chemical structure of young disks

ALMA

ALMA WSU

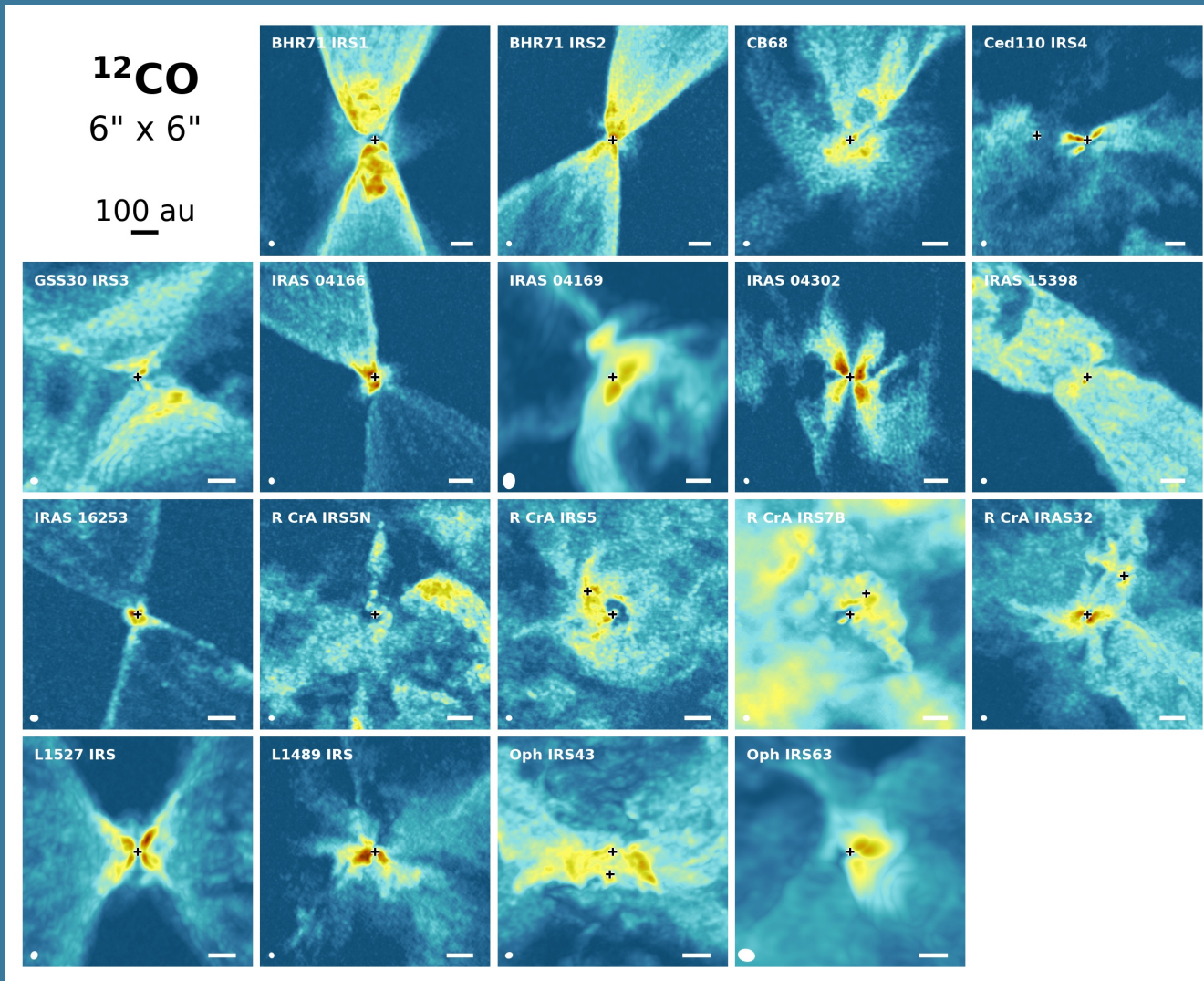
JWST



eDisk: van 't Hoff et al. 2023

MOLECULAR LINES

Physical and chemical structure of young disks

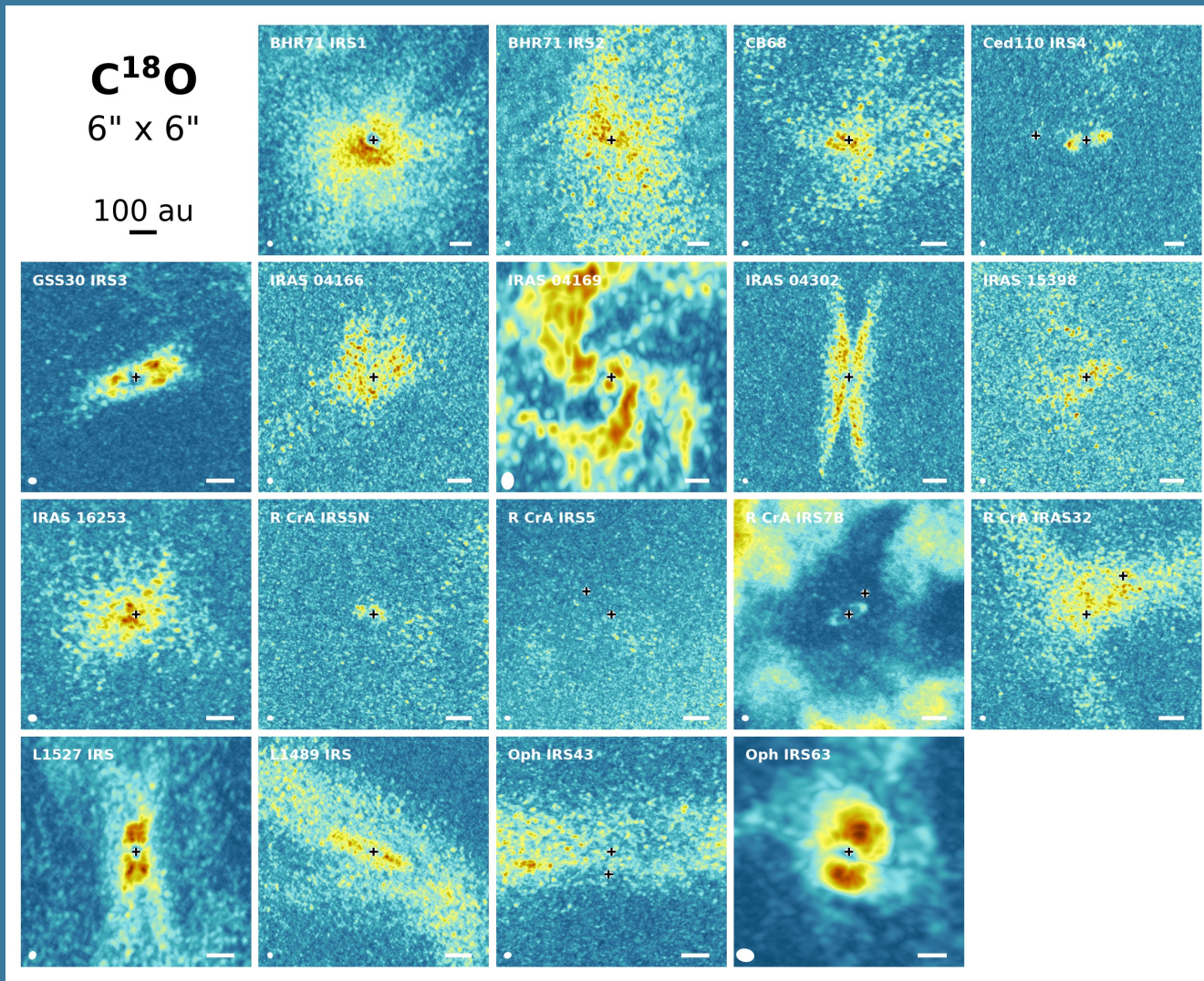


^{12}CO mainly traces outflows

eDisk: Sharma et al. subm

MOLECULAR LINES

Physical and chemical structure of young disks



C¹⁸O mainly traces disk
(and some envelope)

eDisk: Sharma et al. subm

MOLECULAR LINES

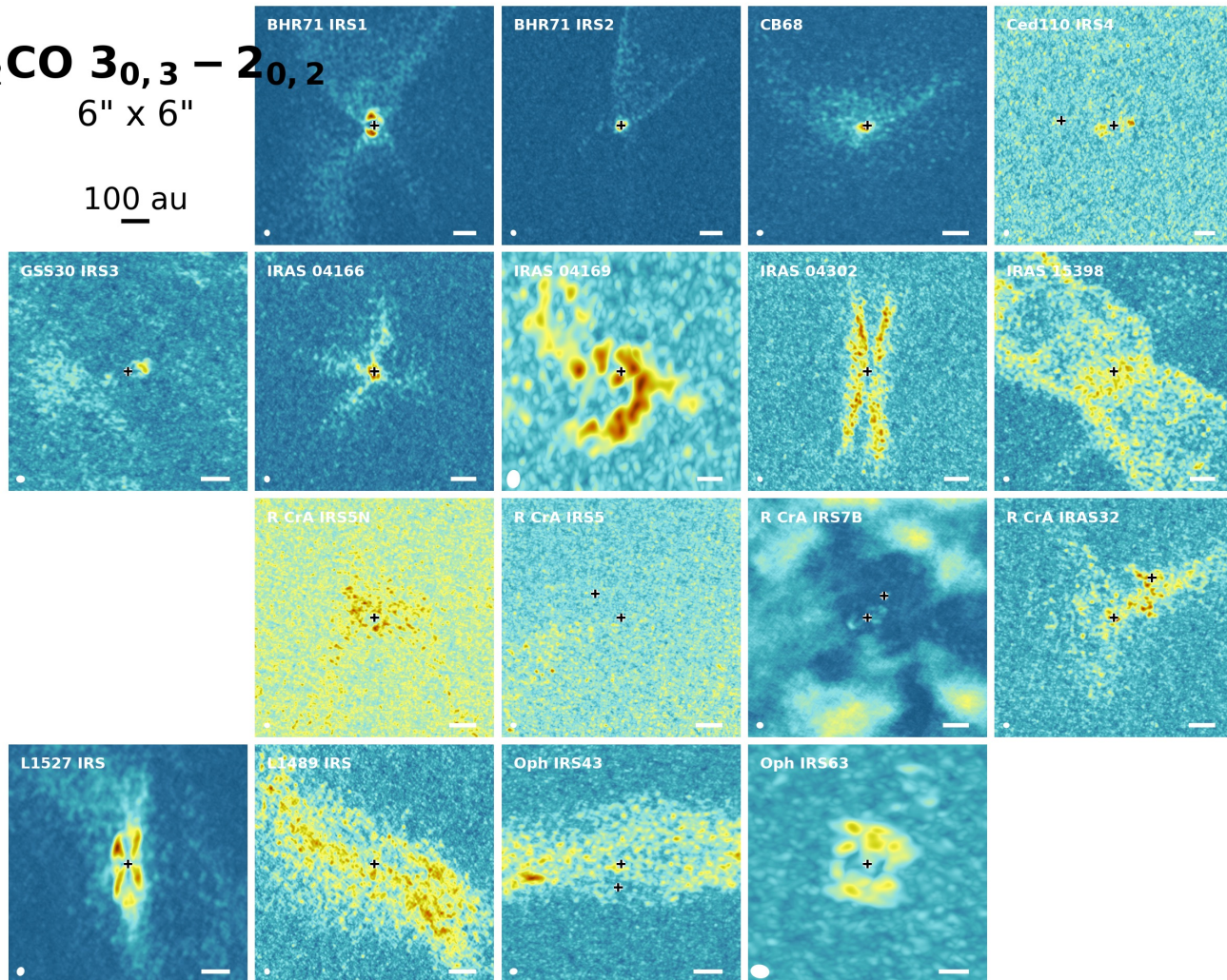
Physical and chemical structure of young disks

ALMA

ALMA WSU

JWST

$\text{H}_2\text{CO } 3_{0,3} - 2_{0,2}$
6" x 6"
100 au



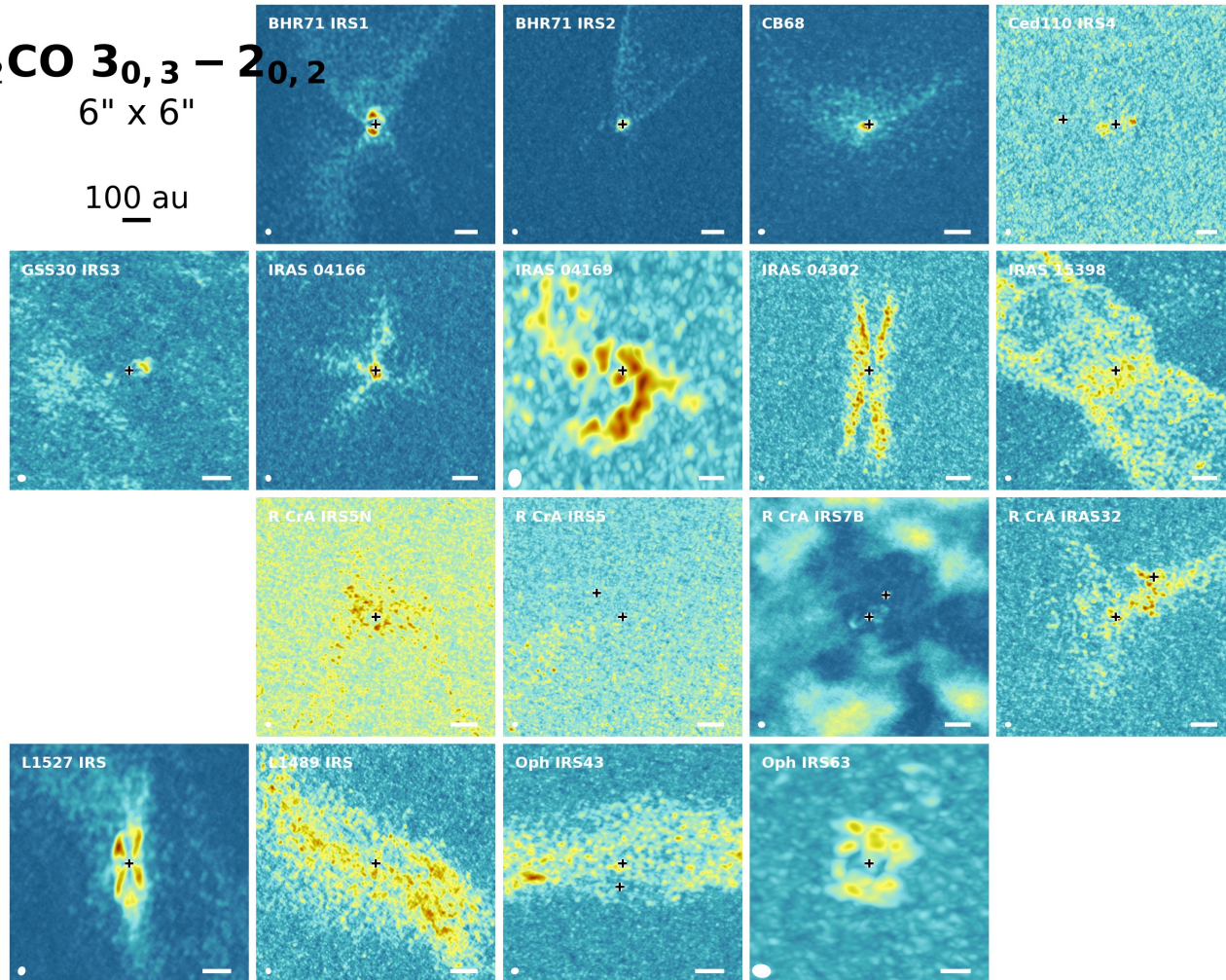
H_2CO can trace different components

eDisk: Sharma et al. subm

MOLECULAR LINES

Physical and chemical structure of young disks

H₂CO 3_{0,3} – 2_{0,2}
6" x 6"
100 au

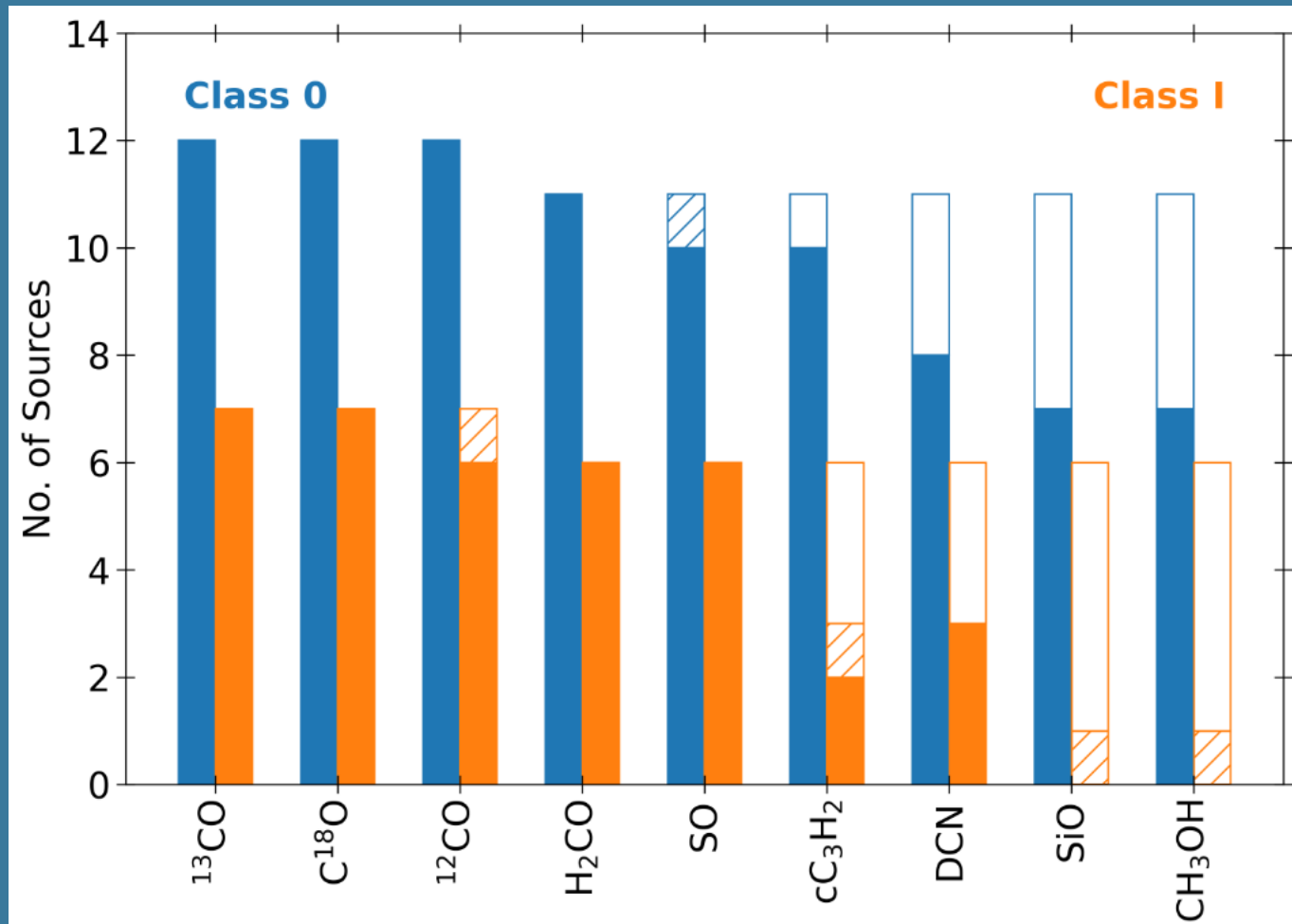


Up to a factor $\sim\sqrt{2}$ to $\sim\sqrt{3}$
deeper

eDisk: Sharma et al. subm

MOLECULAR LINES

Physical and chemical structure of young disks



eDisk: Sharma et al. subm

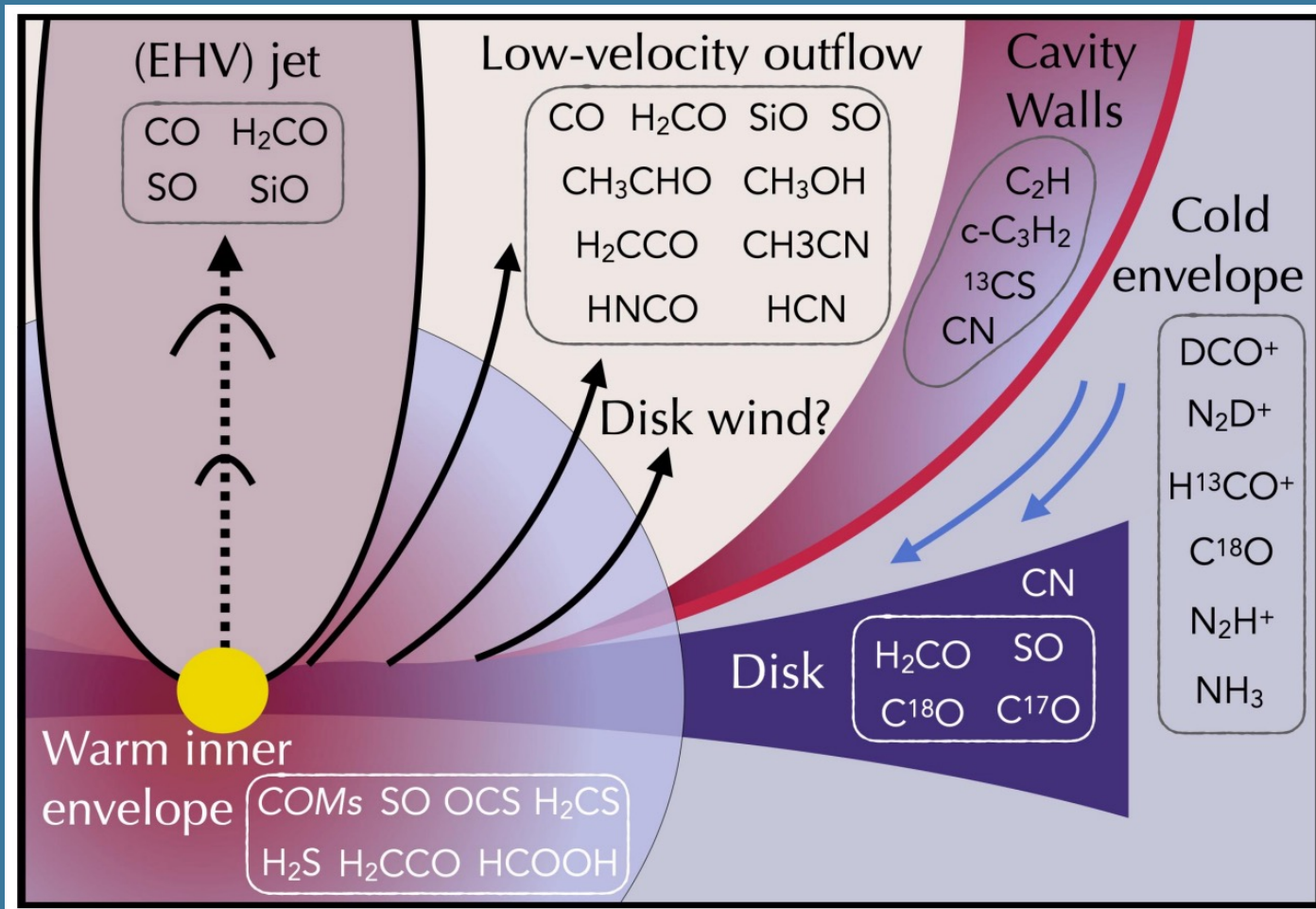
MOLECULAR LINES

Physical and chemical structure of young disks

ALMA

ALMA WSU

JWST

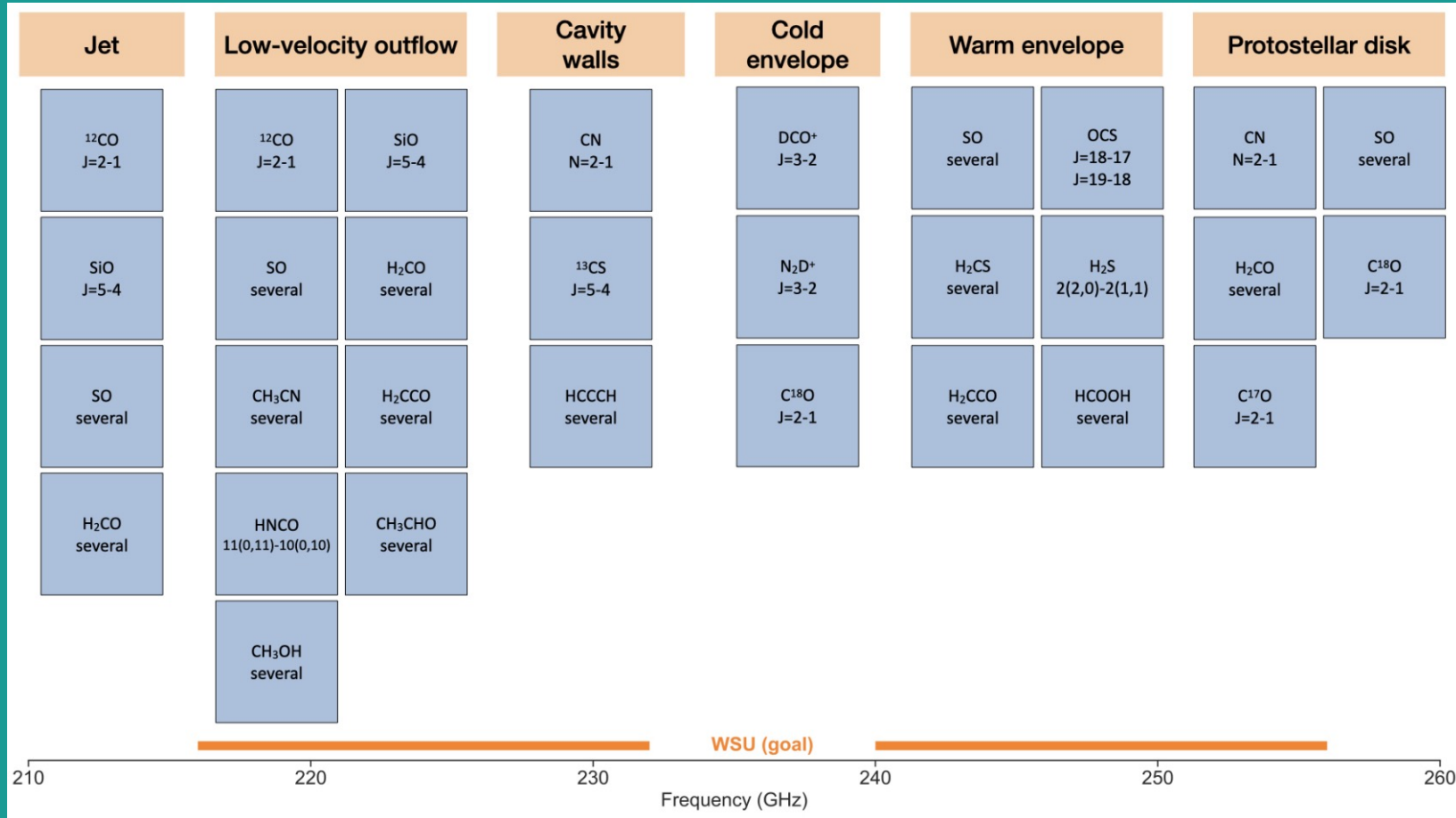


6 different programs:
27 species
(incl. isotopologues)

Tychoniec et al. 2021

MOLECULAR LINES

Physical and chemical structure of young disks

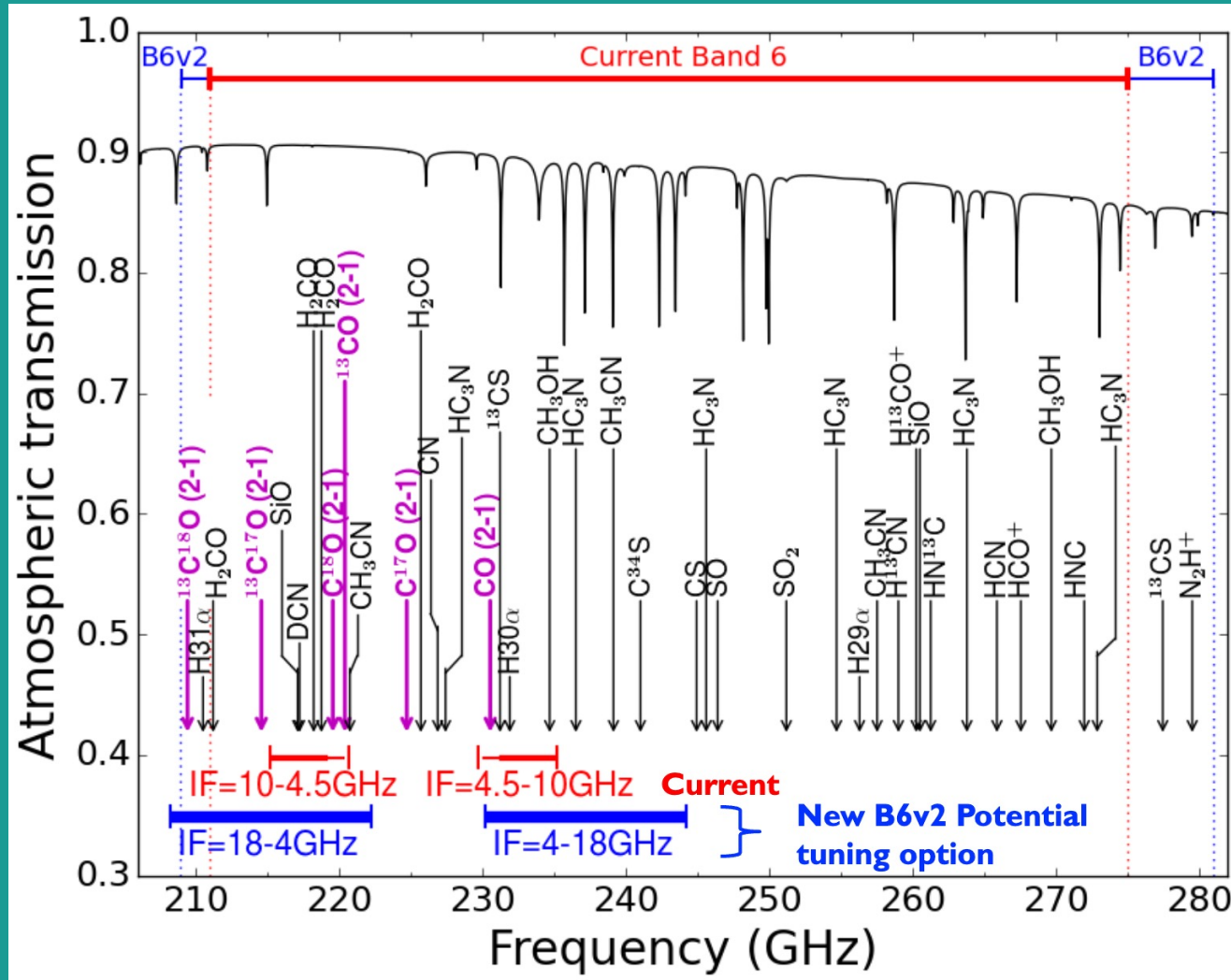


All these species
(and more) in one
tuning

Carpenter et al. 2022

MOLECULAR LINES

Physical and chemical structure of young disks



New Band 6 will cover $^{13}C^{18}O$ and N_2H^+

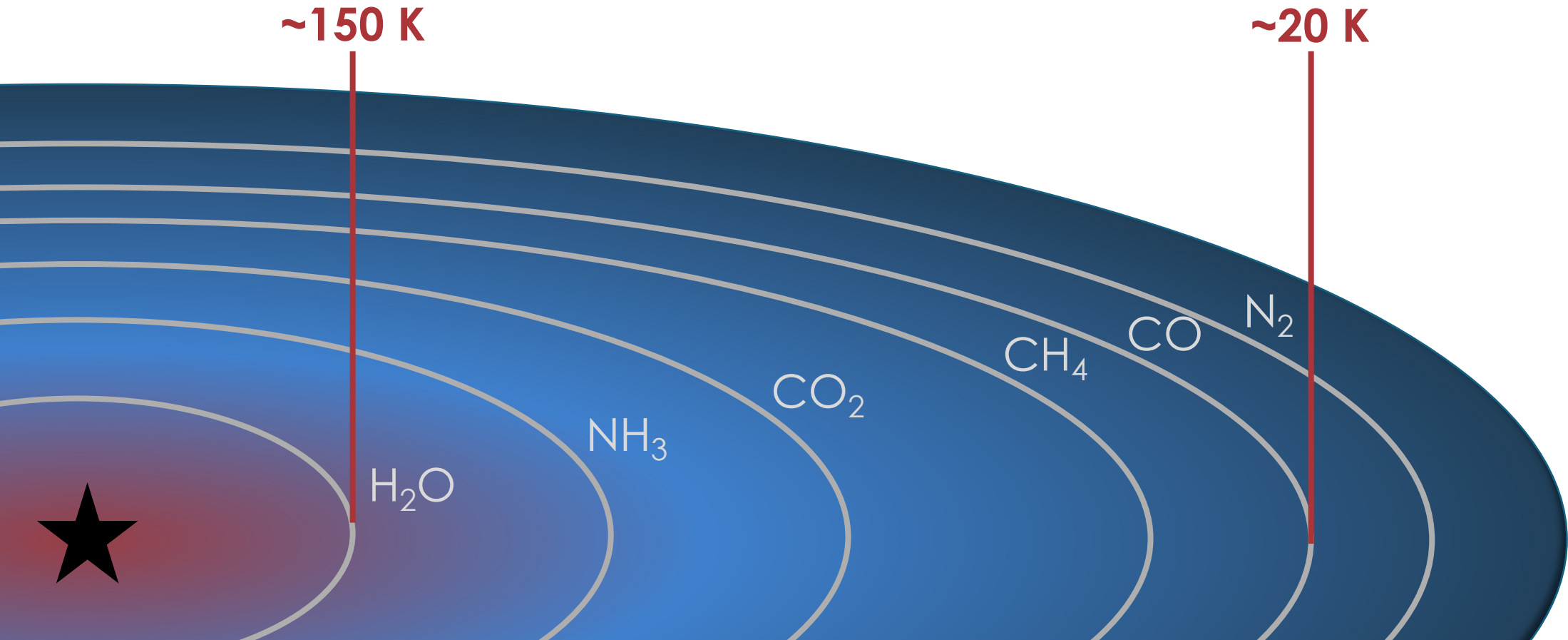
MOLECULAR LINES

Temperature structure

ALMA

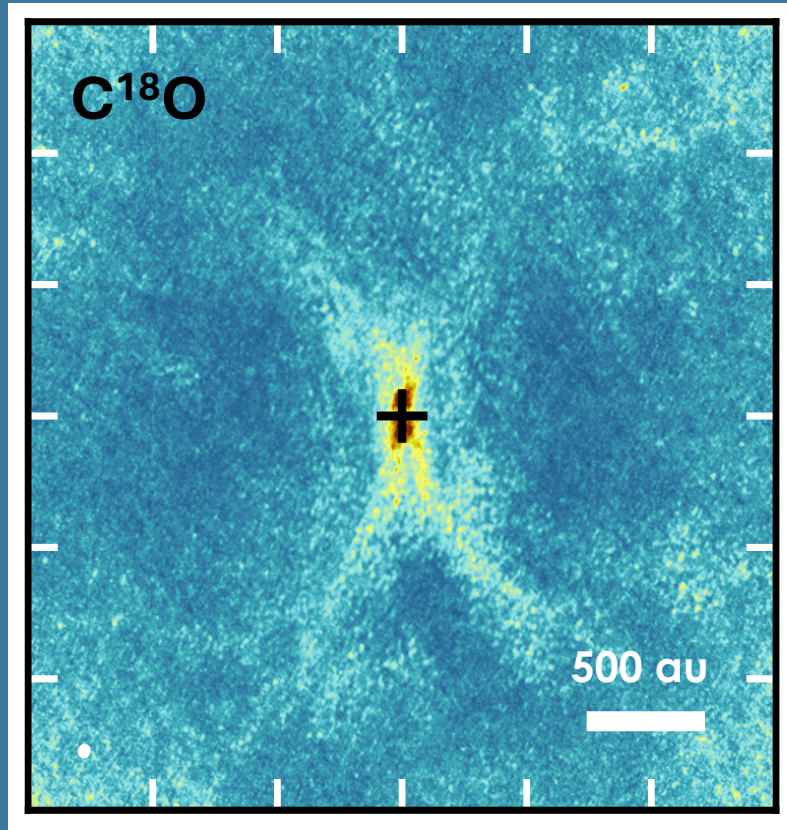
ALMA WSU

JWST



MOLECULAR LINES

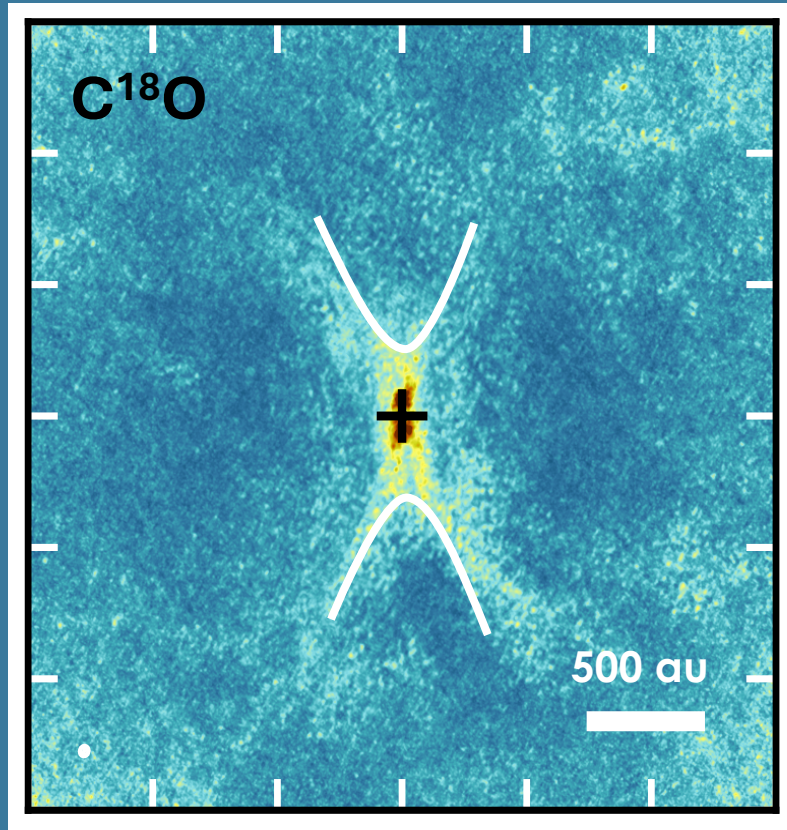
Temperature structure



eDisk: van 't Hoff et al. 2023

MOLECULAR LINES

Temperature structure



eDisk: van 't Hoff et al. 2023

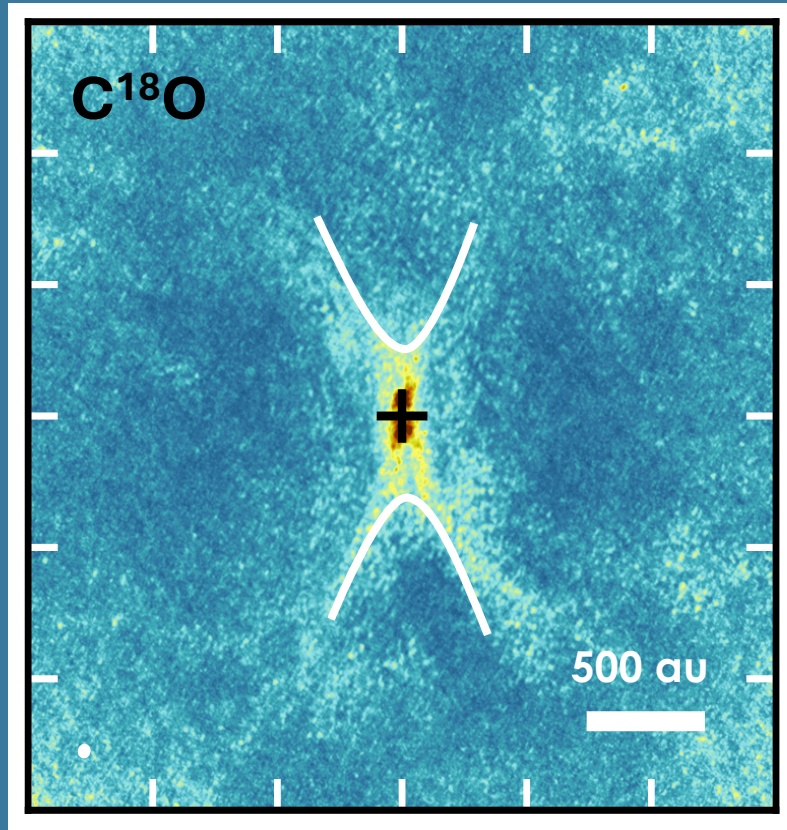
MOLECULAR LINES

Temperature structure

ALMA

ALMA WSU

JWST



CO snowline outside the disk:
Disk warmer than ~ 20 K

See also van 't Hoff et al. 2018, 2020

eDisk: van 't Hoff et al. 2023

MOLECULAR LINES

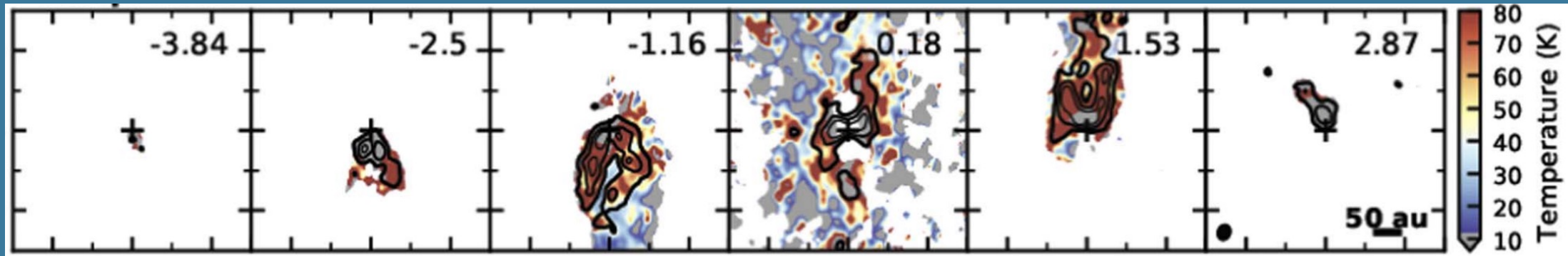
Temperature structure

ALMA

ALMA WSU

JWST

Temperature in velocity channels from the ratio of two H₂CO transitions



eDisk: van 't Hoff et al. 2023

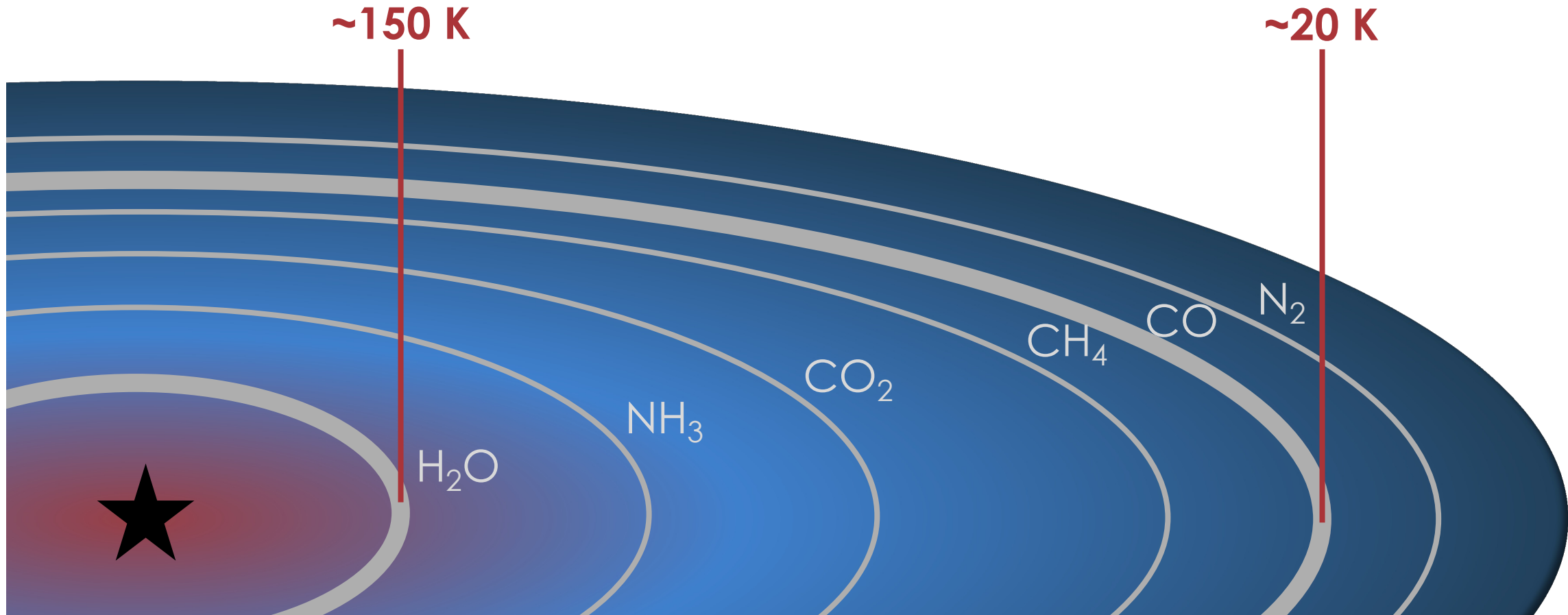
MOLECULAR LINES

Temperature structure

- Increased imaging speed ($\sim 2.2x$) will make it easier to do
 - Deeper observations
 - Higher resolution observations
 - Larger samples
 - Observations of different transitions from a molecule in different Bands
- Increased bandwidth will allow more temperature probes to be observed simultaneously

MOLECULAR LINES

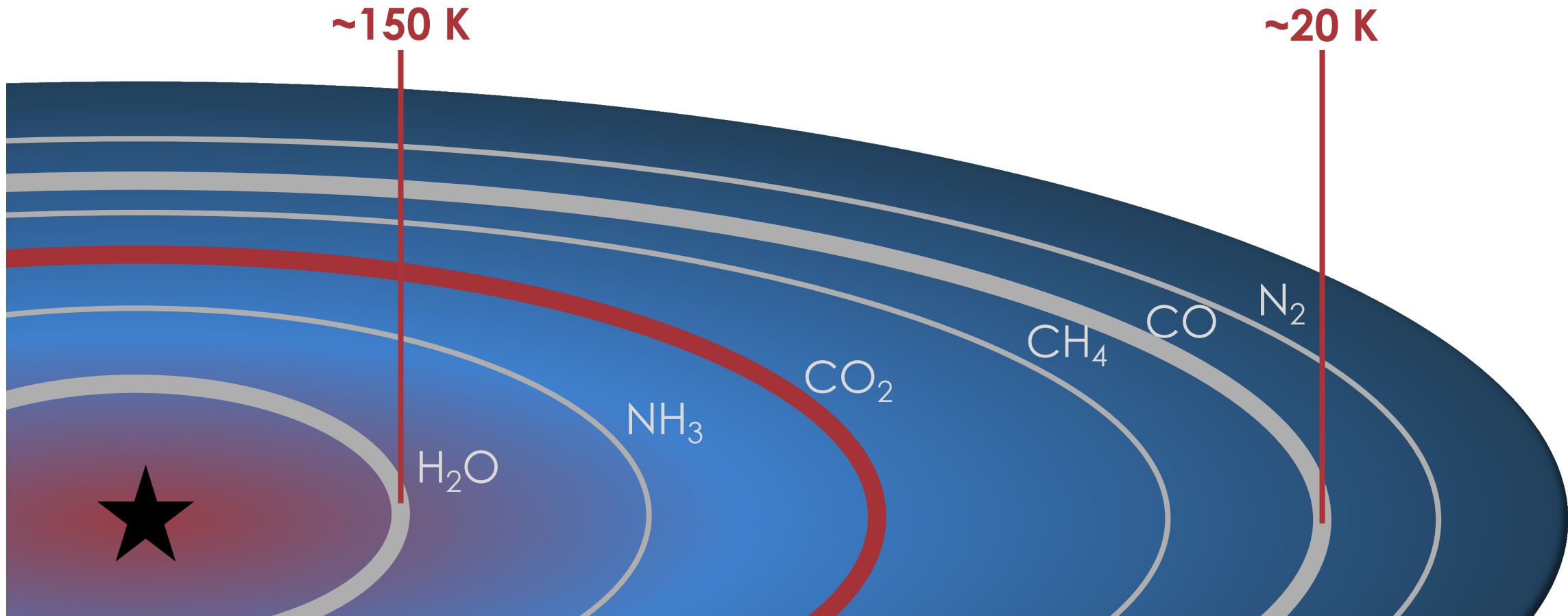
Temperature structure



MOLECULAR LINES

Temperature structure

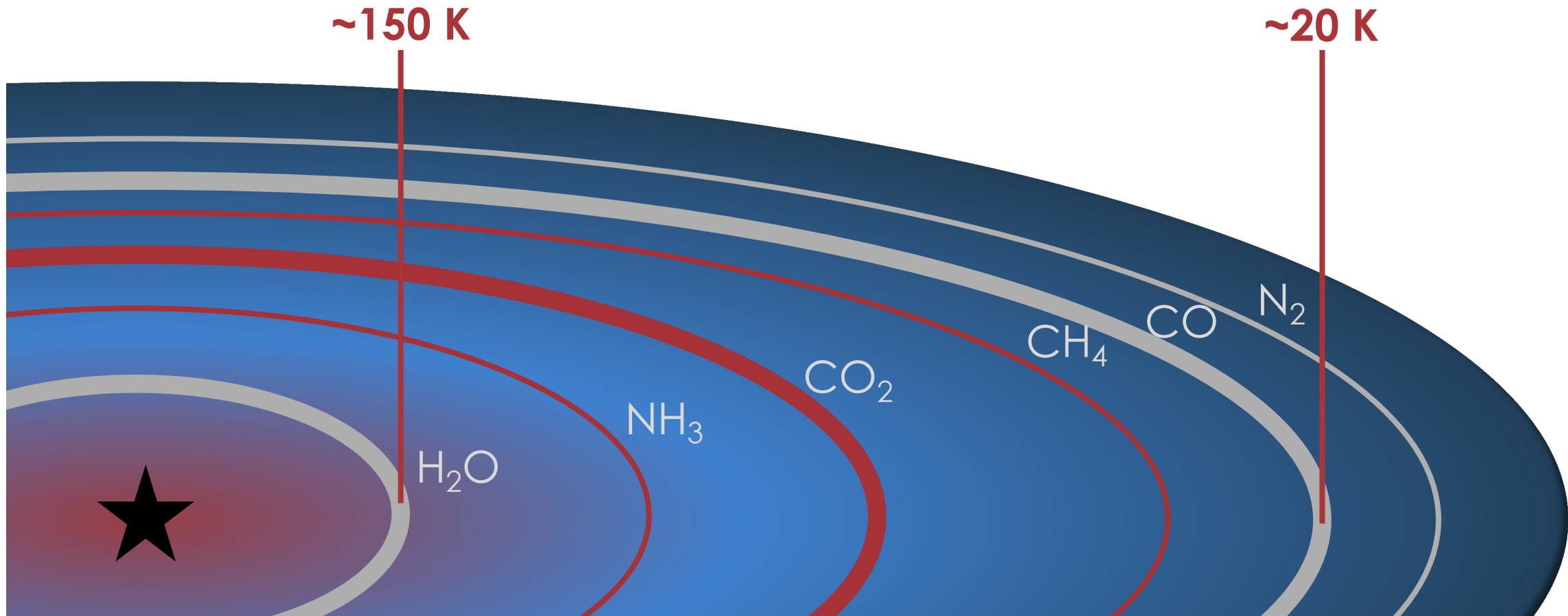
Kaçan et al. in prep.



MOLECULAR LINES

Temperature structure

Kaçan et al. in prep.

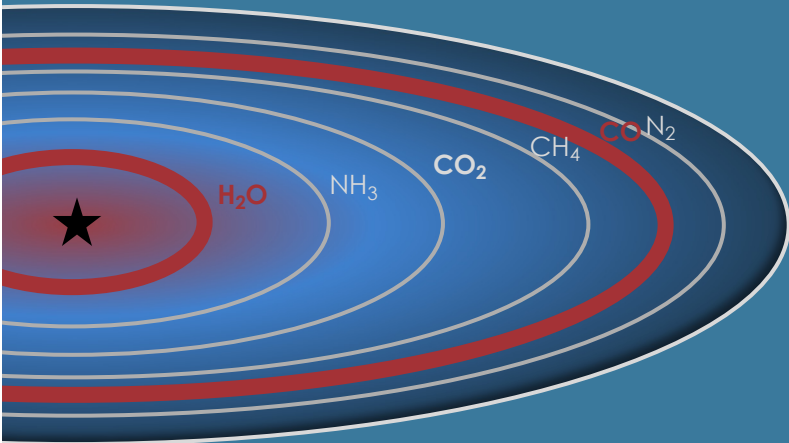


MOLECULAR LINES

Physical and chemical structure of young disks

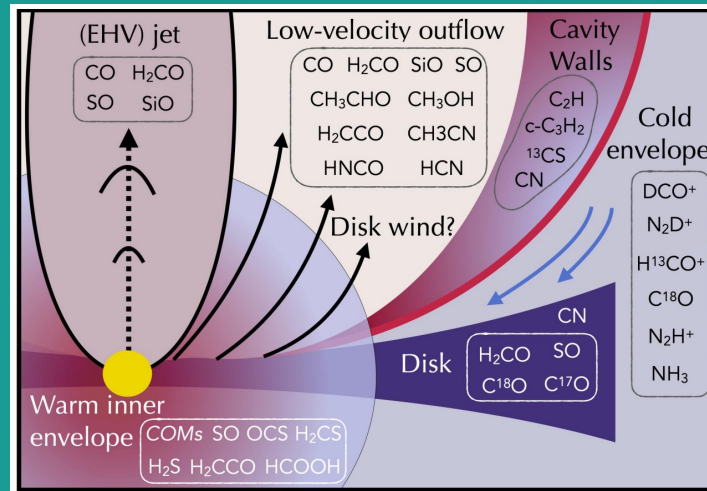
ALMA

Young disks traced by C¹⁸O;
Too warm for CO ice



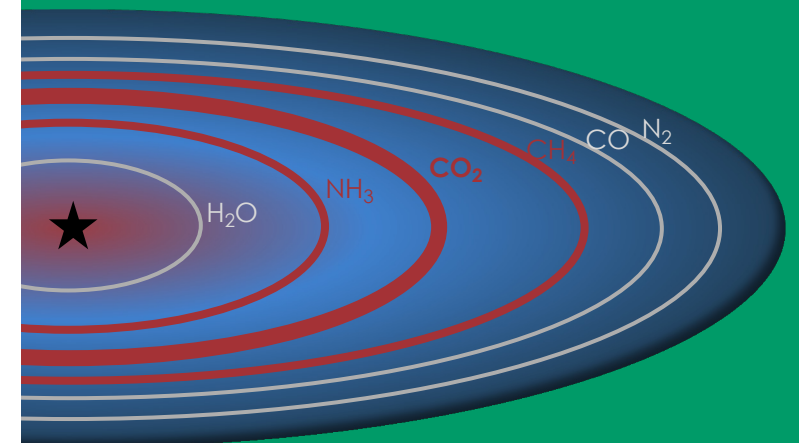
ALMA WSU

Study many tracers
simultaneously



JWST

Complementary info on
molecules without rotational
transitions



MOLECULAR LINES

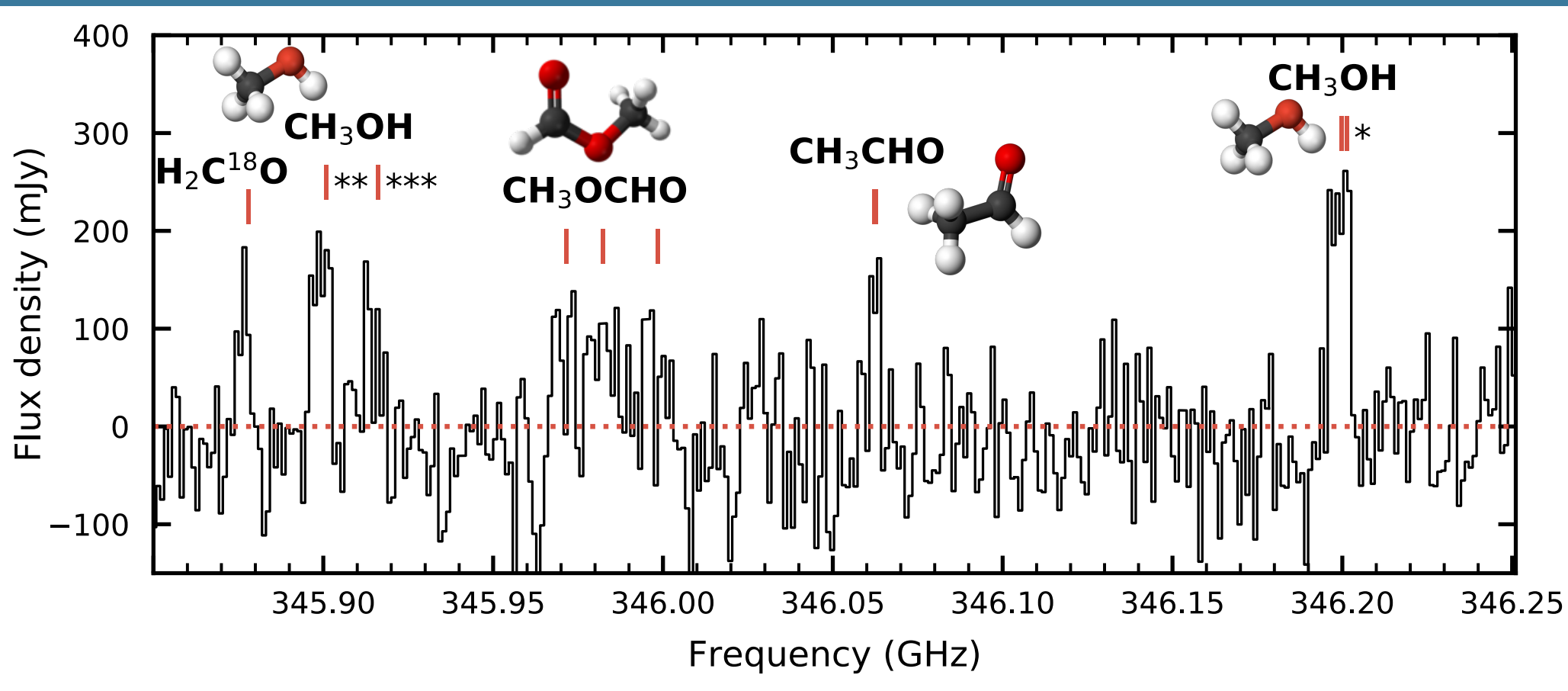
Chemical complexity

ALMA

ALMA WSU

JWST

V883 Ori (1 minute with ALMA)

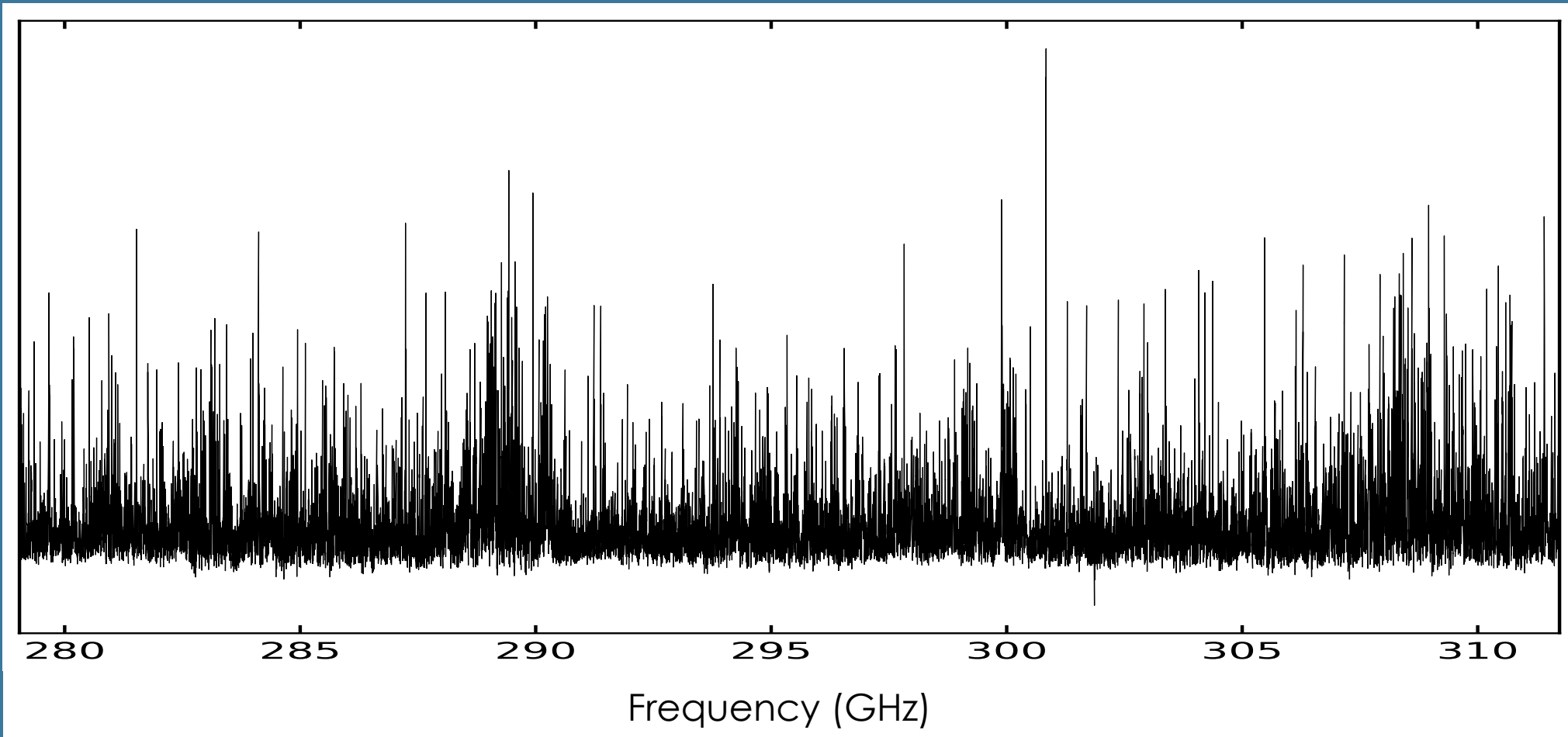


van 't Hoff et al. 2018 (see also Lee et al. 2019; Yamato et al. 2024; Jeong et al. 2024)

MOLECULAR LINES

Chemical complexity

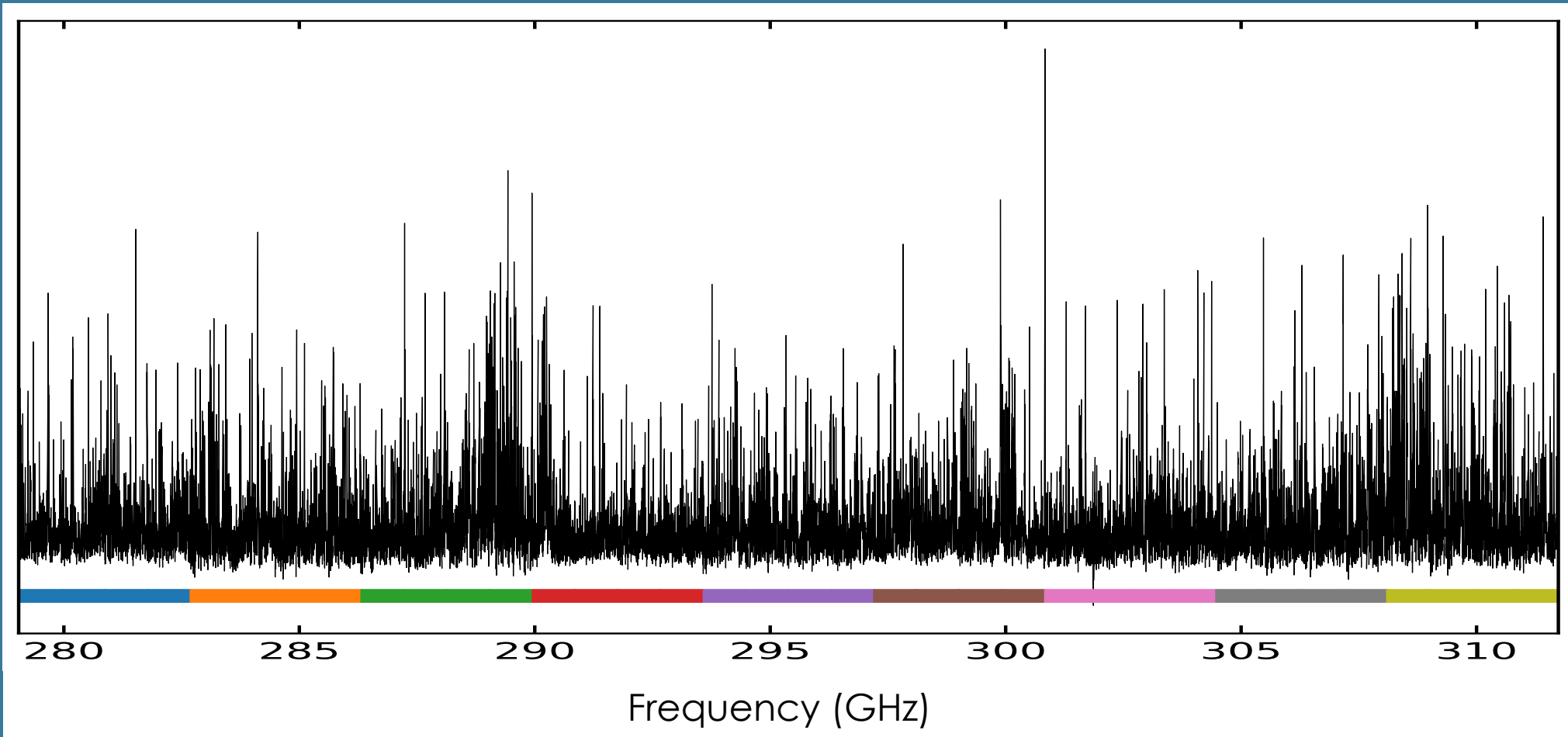
V883 Ori with COMPASS Large Program (PI: Jes Jørgensen)



MOLECULAR LINES

Chemical complexity

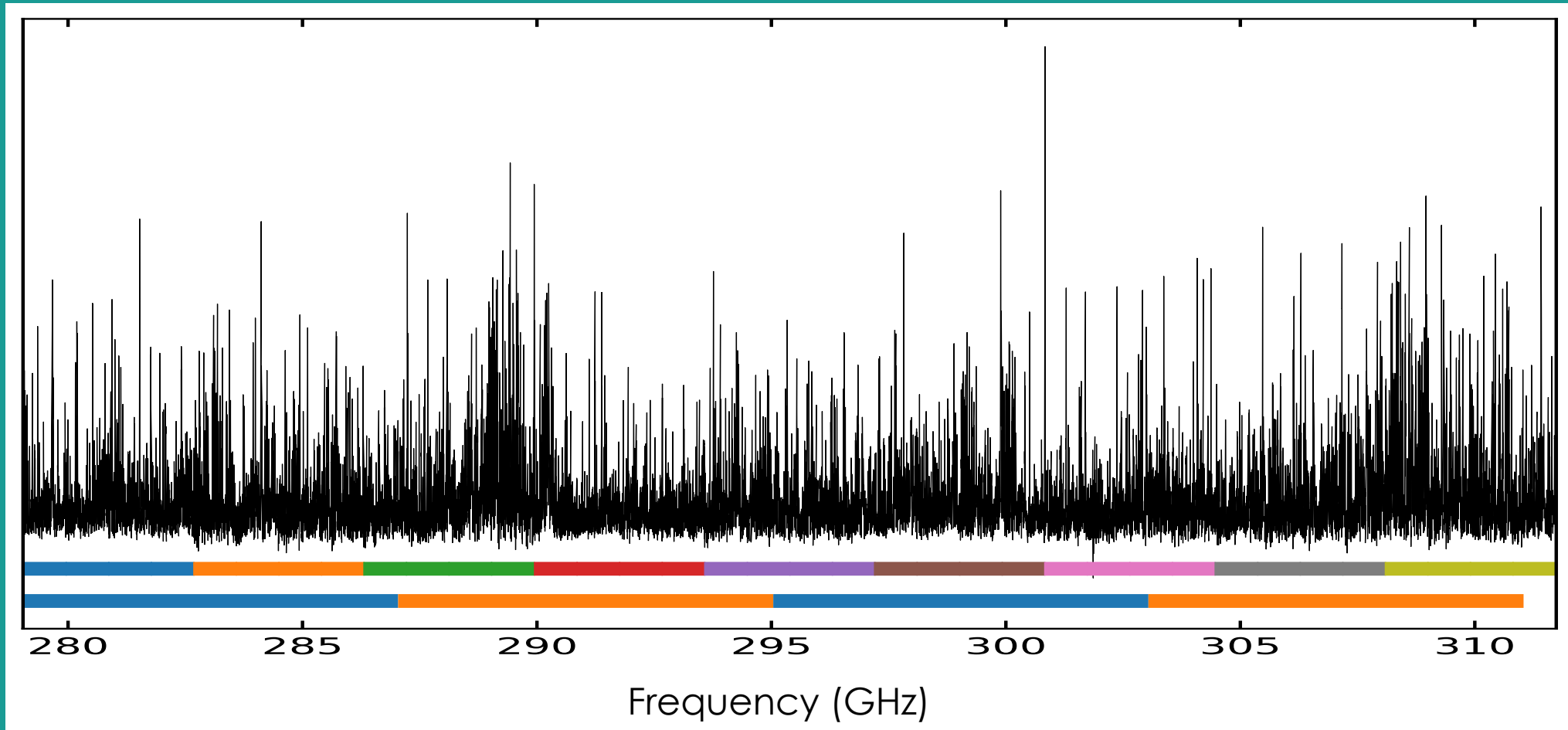
V883 Ori with COMPASS Large Program (PI: Jes Jørgensen)



MOLECULAR LINES

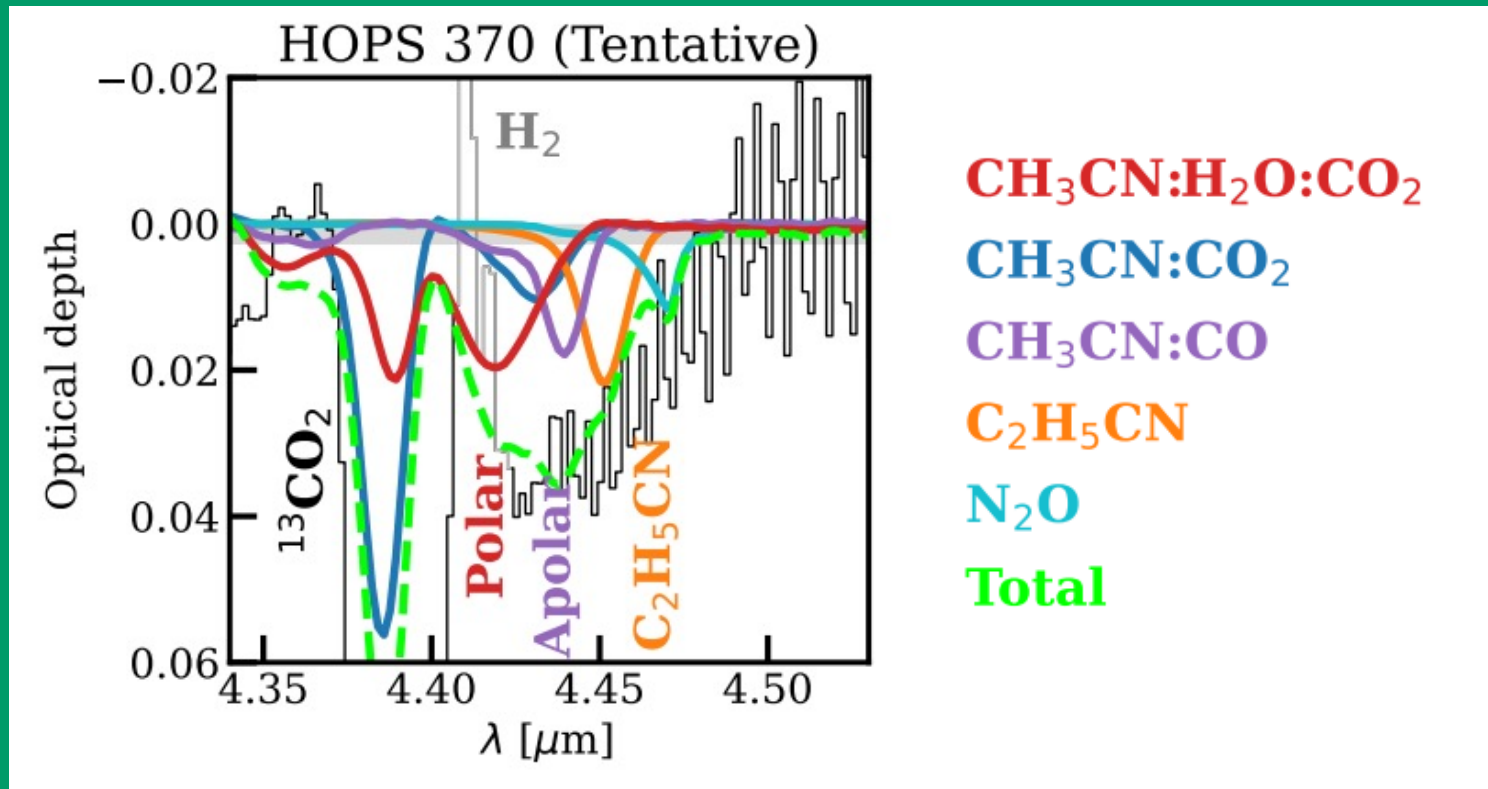
Chemical complexity

V883 Ori with COMPASS Large Program (PI: Jes Jørgensen)



MOLECULAR LINES

Chemical complexity



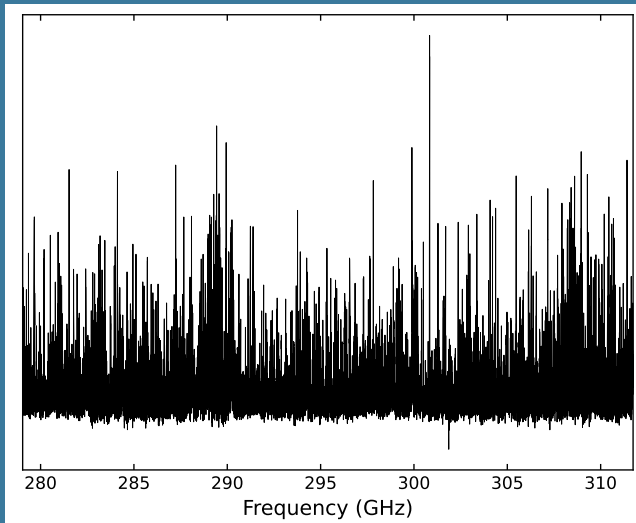
Complementary
COMPASS JWST
program to
study ice
composition in
the envelope

MOLECULAR LINES

Chemical complexity in young disks

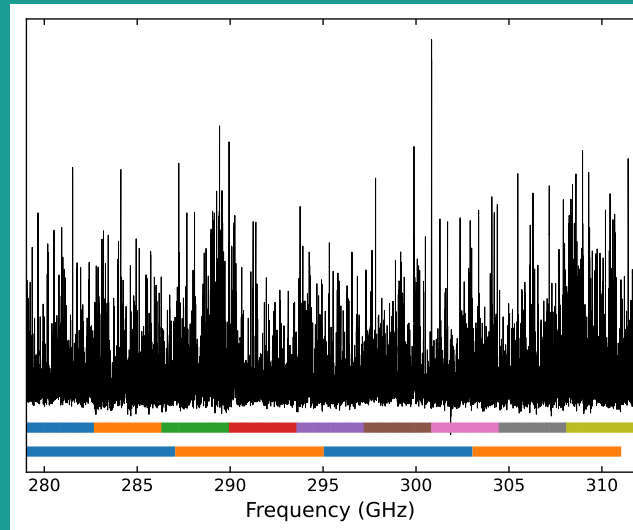
ALMA

High level of chemical complexity in young disks, but mostly frozen out



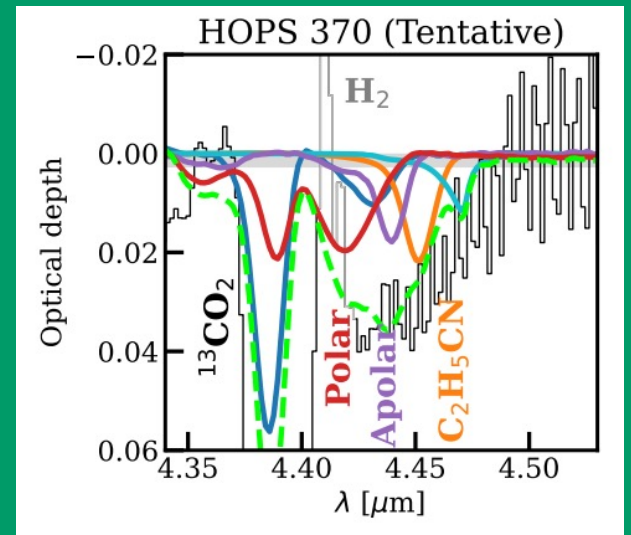
ALMA WSU

Much more efficient to study



JWST

Complementary info on chemical complexity in ice



YOUNG EMBEDDED DISKS:

The ALMA (WSU) and JWST era

**WE HAVE ONLY SCRATCHED
THE SURFACE!**